(A brief) Introduction to Neutrino Physics

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Physics Summer Lecture

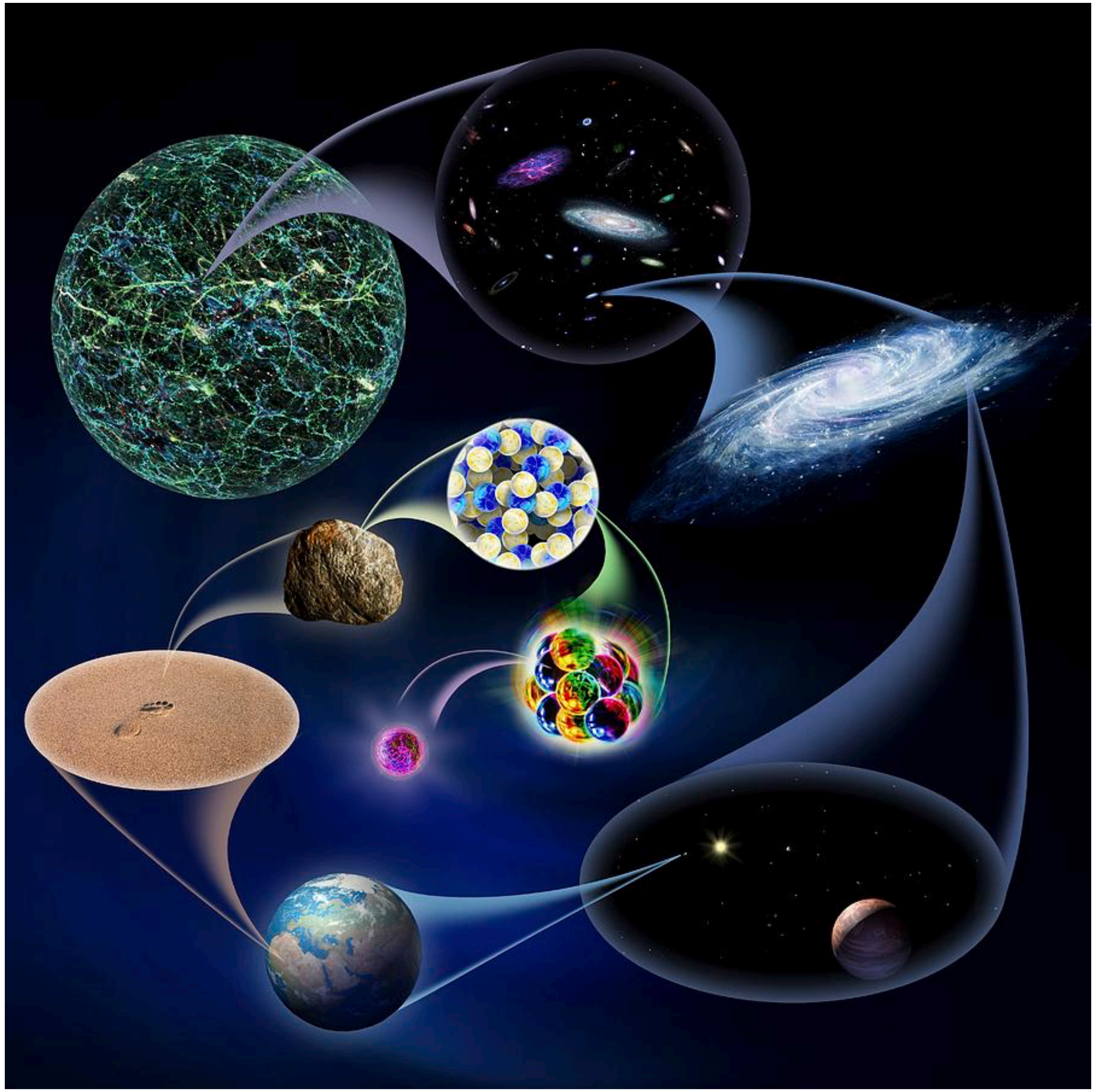
Brookhaven
National Laboratory

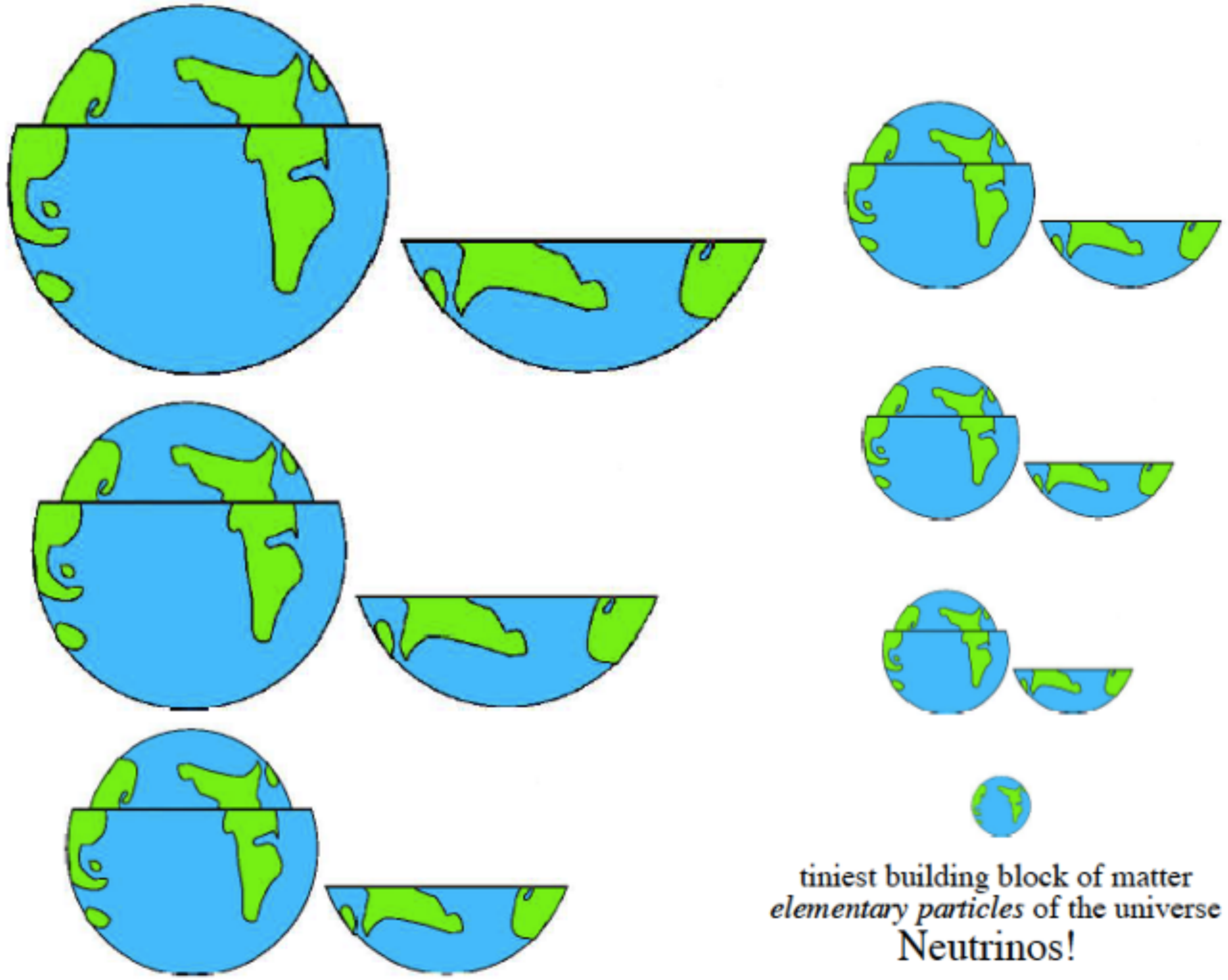
A brief bio...

- as a kid, I always was fascinated with the stars and galaxies (who isn't?), wanted to become an **astrophysicist** like Hawking
- in college, I started to become more interested in **theoretical particle physics**, with inspirations from Einstein, Feynman, Gell-Mann (who doesn't?)
- in graduate school, I started working on **experimental particle physics** (neutrino) as we started to have breakthroughs in experimental particle physics during the time (Higgs boson discovery, neutrino oscillation, gravitational wave...)
- as a postdoc, I also worked on dark matter search for few years, then came back to **experimental neutrino physics** as a faculty here at BNL
- now I work on experimental neutrino physics: how to detect, measure, and understand neutrinos and their nature

  - feel free to stop by 3-181 anytime to say hello or to ask/talk about anything
  - you can also email me with any questions: jjo@bnl.gov

Particle physicist trying to understand ordinary matter…



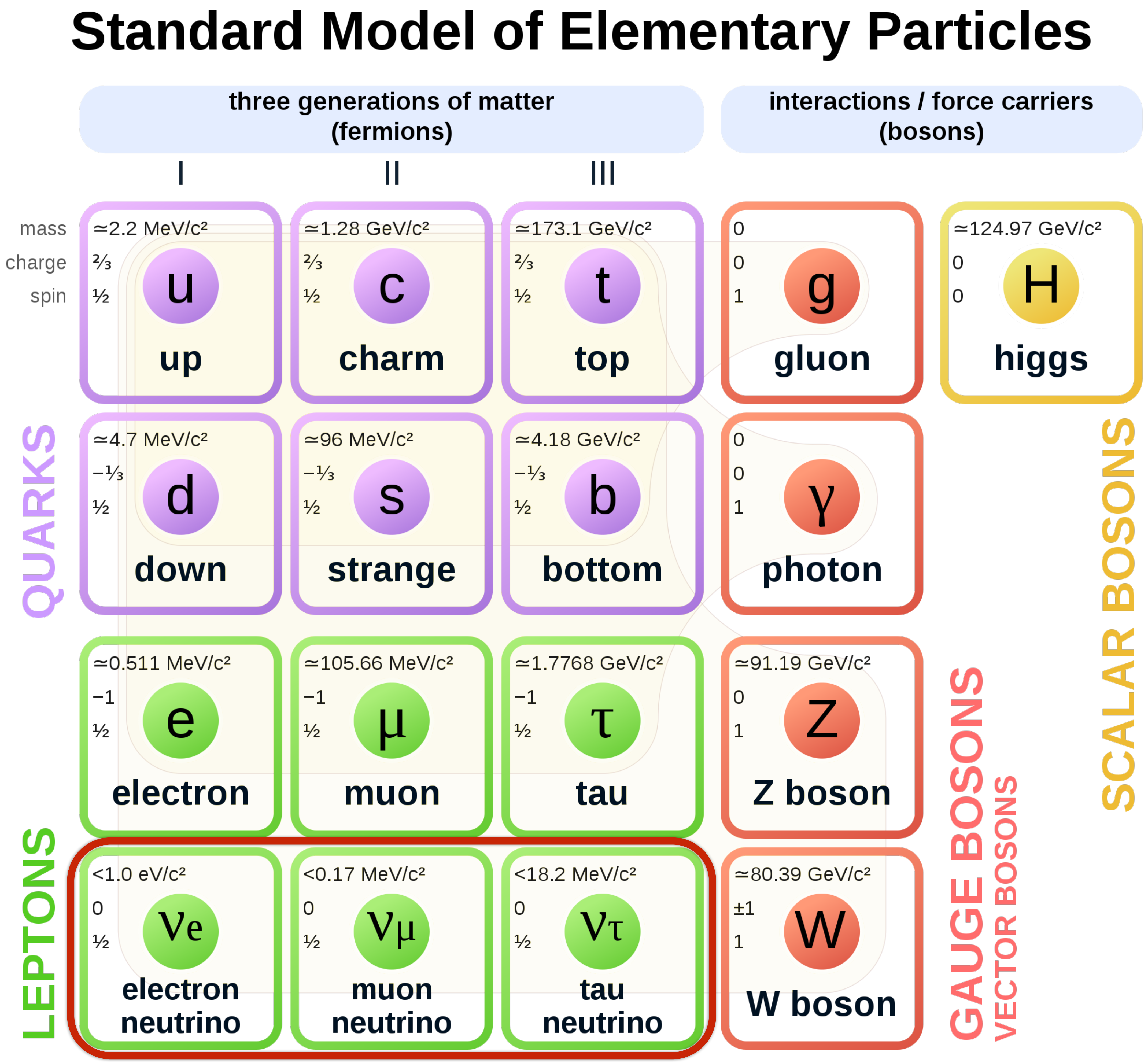


tiniest building block of matter:

elementary particles of the universe

Standard model of particle physics

- ordinary matter is well described by **12 building blocks of matter** and the force carriers through which they interact
- **neutrinos** make up three of the 12 building blocks, with special characteristics of:
  - neutral charge
  - tiny mass
  - weakly interacting only
- today, we will be talking about these neutrinos



Standard model of particle physics

Table 10.5: Principal Z pole observables and their SM predictions (*cf.* Table 10.4). The first M_Z is from LEP 1 [288] and the second from CDF [289]. The first $\bar{s}_\ell^2$ is the effective weak mixing angle extracted from the hadronic charge asymmetry at LEP 1 [288], the second is the combined value from the Tevatron [309], and the third is from the LHC [310–314]. The values of A_e are (i) from A_{LR} for hadronic final states [315]; (ii) from A_{LR} for leptonic final states and from polarized Bhabha scattering [316]; and (iii) from the angular distribution of the τ polarization at LEP 1 [288]. The A_τ values are from SLD [316], the total τ polarization from LEP [288], and from CMS [317], respectively. Note that the SM errors in Γ_Z, the R_ℓ, and $\sigma_{\rm had}$ are largely dominated by the uncertainty in α_s.

Quantity	Value	Standard Model	Pull
M_Z [GeV]	91.1876 ± 0.0021	91.1884 ± 0.0019	−0.4
	91.192 ± 0.007		0.6
Γ_Z [GeV]	2.4955 ± 0.0023	2.4940 ± 0.0009	0.7
$\sigma_{\rm had}$ [nb]	41.481 ± 0.033	41.481 ± 0.009	0.0
R_e	20.804 ± 0.050	20.736 ± 0.010	1.4
R_μ	20.784 ± 0.034	20.736 ± 0.010	1.4
R_τ	20.764 ± 0.045	20.781 ± 0.010	−0.4
R_b	0.21629 ± 0.00066	0.21583 ± 0.00002	0.7
R_c	0.1721 ± 0.0030	0.17221 ± 0.00003	0.0
$A_{FB}^{(0,e)}$	0.0145 ± 0.0025	0.01606 ± 0.00006	−0.6
$A_{FB}^{(0,\mu)}$	0.0169 ± 0.0013		0.6
$A_{FB}^{(0,\tau)}$	0.0188 ± 0.0017		1.6
$A_{FB}^{(0,b)}$	0.0996 ± 0.0016	0.1026 ± 0.0002	−1.8
$A_{FB}^{(0,c)}$	0.0707 ± 0.0035	0.0732 ± 0.0002	−0.7
$A_{FB}^{(0,s)}$	0.0976 ± 0.0114	0.1027 ± 0.0002	−0.4
$\bar{s}_\ell^2$	0.2324 ± 0.0012	0.23161 ± 0.00004	0.7
	0.23148 ± 0.00033		−0.4
	0.23145 ± 0.00028		−0.6
A_e	0.15138 ± 0.00216	0.1463 ± 0.0003	2.3
	0.1544 ± 0.0060		1.3
	0.1498 ± 0.0049		0.7
A_μ	0.142 ± 0.015		−0.3
A_τ	0.136 ± 0.015		−0.7
	0.1439 ± 0.0043		−0.6
	0.144 ± 0.015		−0.2
A_b	0.923 ± 0.020	0.9347	−0.6
A_c	0.670 ± 0.027	0.6674 ± 0.0001	0.1
A_s	0.895 ± 0.091	0.9356	−0.4

Table 10.6: Results derived from Table 10.5 and the corresponding covariance matrices [288,332], and the SM predictions for the partial and total Z decay widths [in MeV]. In the (second) third column lepton universality is (not) assumed.

Quantity	Value	Value (universal)	Standard Model
$\Gamma_{e^+e^-}$	83.87 ± 0.12	83.942 ± 0.085	83.955 ± 0.009
$\Gamma_{\mu^+\mu^-}$	83.95 ± 0.18	83.941 ± 0.085	83.955 ± 0.009
$\Gamma_{\tau^+\tau^-}$	84.03 ± 0.21	83.759 ± 0.085	83.772 ± 0.009
$\Gamma_{\rm inv}$	498.9 ± 2.5	500.5 ± 1.5	501.435 ± 0.045
$\Gamma_{u\bar{u}}$	—	—	299.87 ± 0.20
$\Gamma_{c\bar{c}}$	300.3 ± 5.3	300.0 ± 5.2	299.81 ± 0.20
$\Gamma_{d\bar{d}}, \Gamma_{s\bar{s}}$	—	—	382.75 ± 0.14
$\Gamma_{b\bar{b}}$	377.4 ± 1.3	377.0 ± 1.2	375.73 ∓ 0.18
$\Gamma_{\rm had}$	1744.8 ± 2.6	1743.2 ± 1.9	1740.88 ± 0.86
Γ_Z	2495.5 ± 2.3	2495.5 ± 2.3	2494.00 ± 0.87

Table 10.7: Principal electroweak SM fit result including mutual correlations.

M_Z [GeV]	91.1884 ± 0.0019	1.00	−0.08	0.00	0.00	0.02	0.02
$\hat{m}_t(\hat{m}_t)$ [GeV]	163.18 ± 0.54	−0.08	1.00	0.00	−0.12	−0.23	0.04
$\hat{m}_b(\hat{m}_b)$ [GeV]	4.180 ± 0.008	0.00	0.00	1.00	0.19	−0.02	0.01
$\hat{m}_c(\hat{m}_c)$ [GeV]	1.274 ± 0.009	0.00	−0.12	0.19	1.00	0.48	0.01
$\alpha_s(M_Z)$	0.1187 ± 0.0017	0.02	−0.23	−0.02	0.48	1.00	−0.04
$\Delta\alpha_{\rm had}^{(3)}(2\ {\rm GeV})$	0.00608 ± 0.00004	0.02	0.04	0.01	0.01	−0.04	1.00

S. Navas et al. (Particle Data Group), Phys. Rev. D 110, 030001 (2024)

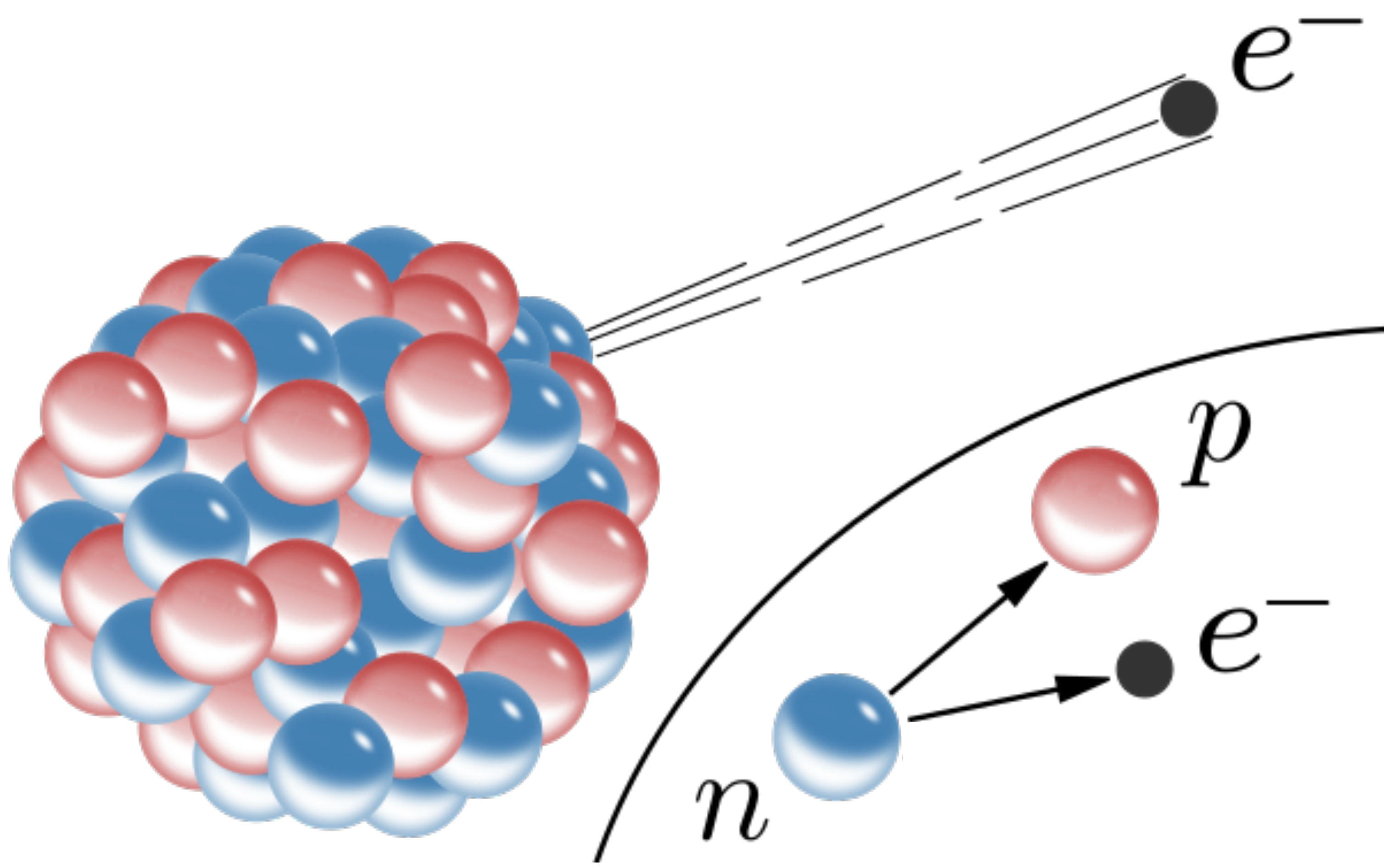
Contents

- **Beta decay**: how the neutrinos found
- **Missing neutrinos**: how the neutrinos change their flavors
- **Detecting invisible particles**: what and how we detect neutrinos

Beta decay: how the neutrinos were found

Beta decay: Theoretical prediction

- the beta decay is a radioactive decay in which **a proton in a nucleus is converted into a neutron** (or vice-versa)
- in the process, the nucleus **emits a beta particle** (electron or positron); hence *beta decay*



number of protons and neutrons

$$^{210}_{83}\text{Bi} \rightarrow {}^{210}_{84}\text{Po} + e^-$$

number of protons

$$n \rightarrow p^+ + e^-$$

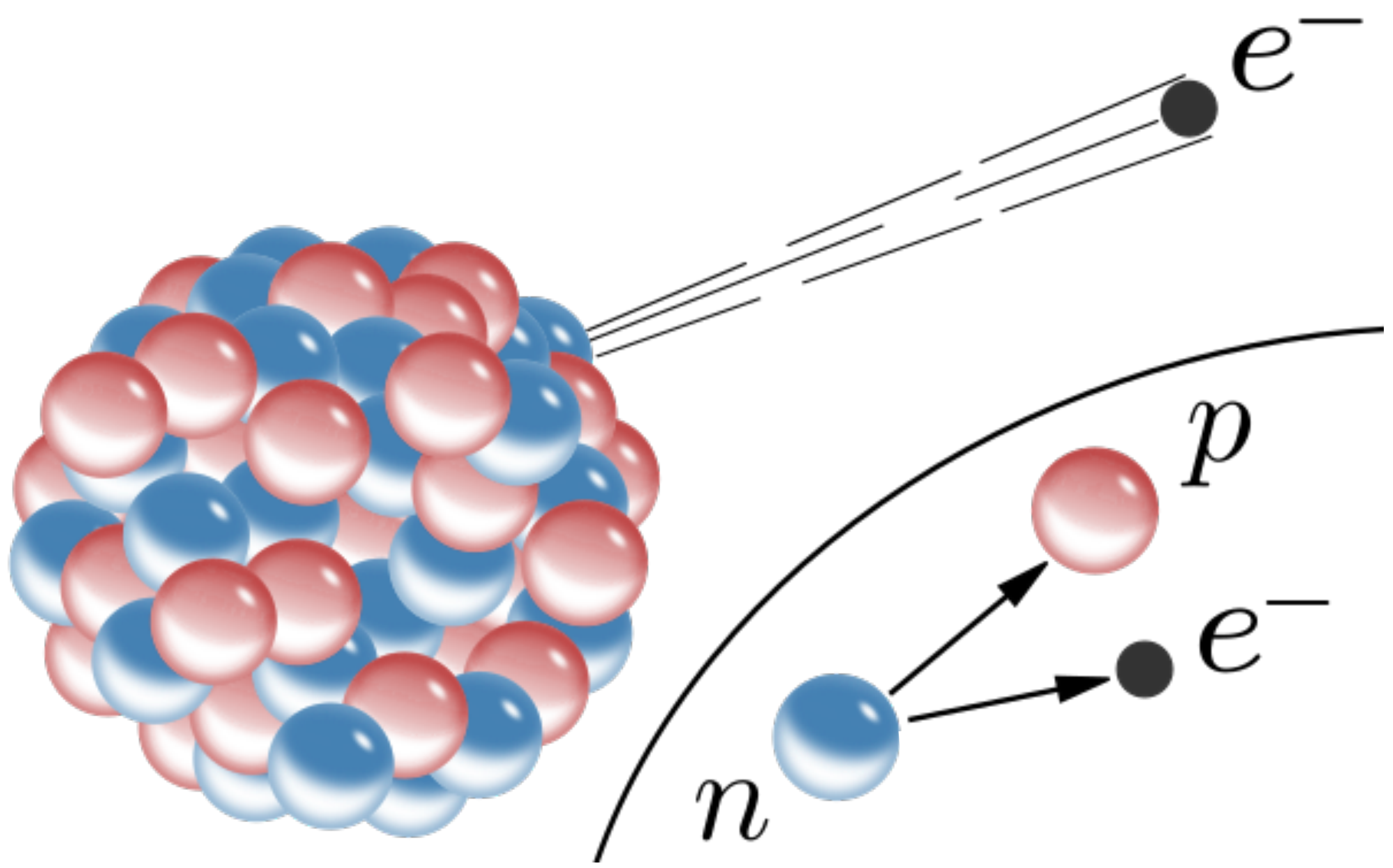
1899 – 1927
Rutherford, Meitner, Hahn, Chadwick, Ellis, Mott, *et. al*

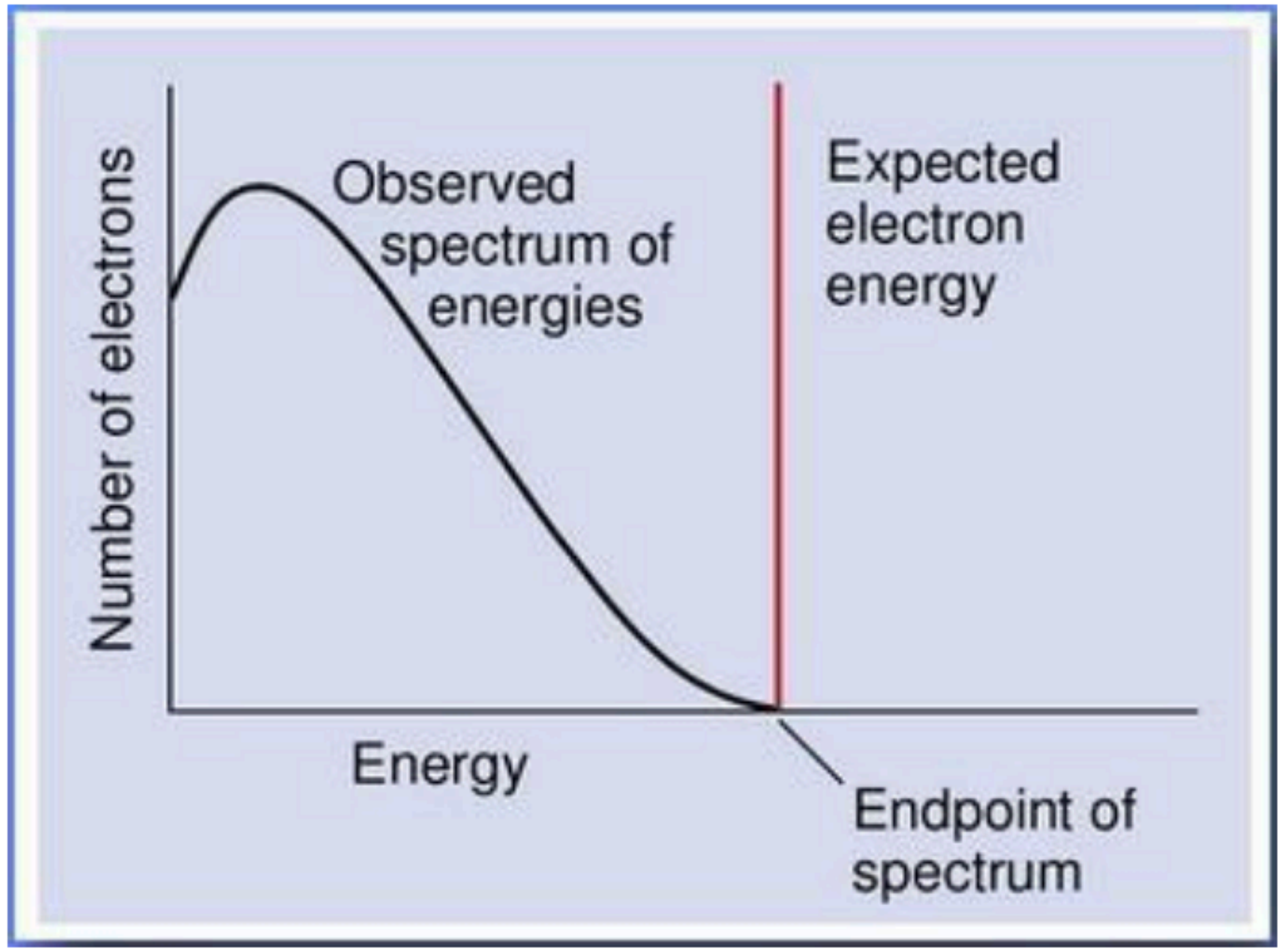
Beta decay: Theoretical prediction

- if the decay happens with the atom nucleus at rest, the energy of the electron is **expected to be always the same** considering energy conservation

- but instead **“spectrum” of energies was observed**, always less energy than expected

  - maybe there’s an invisible particle that takes away the energy?



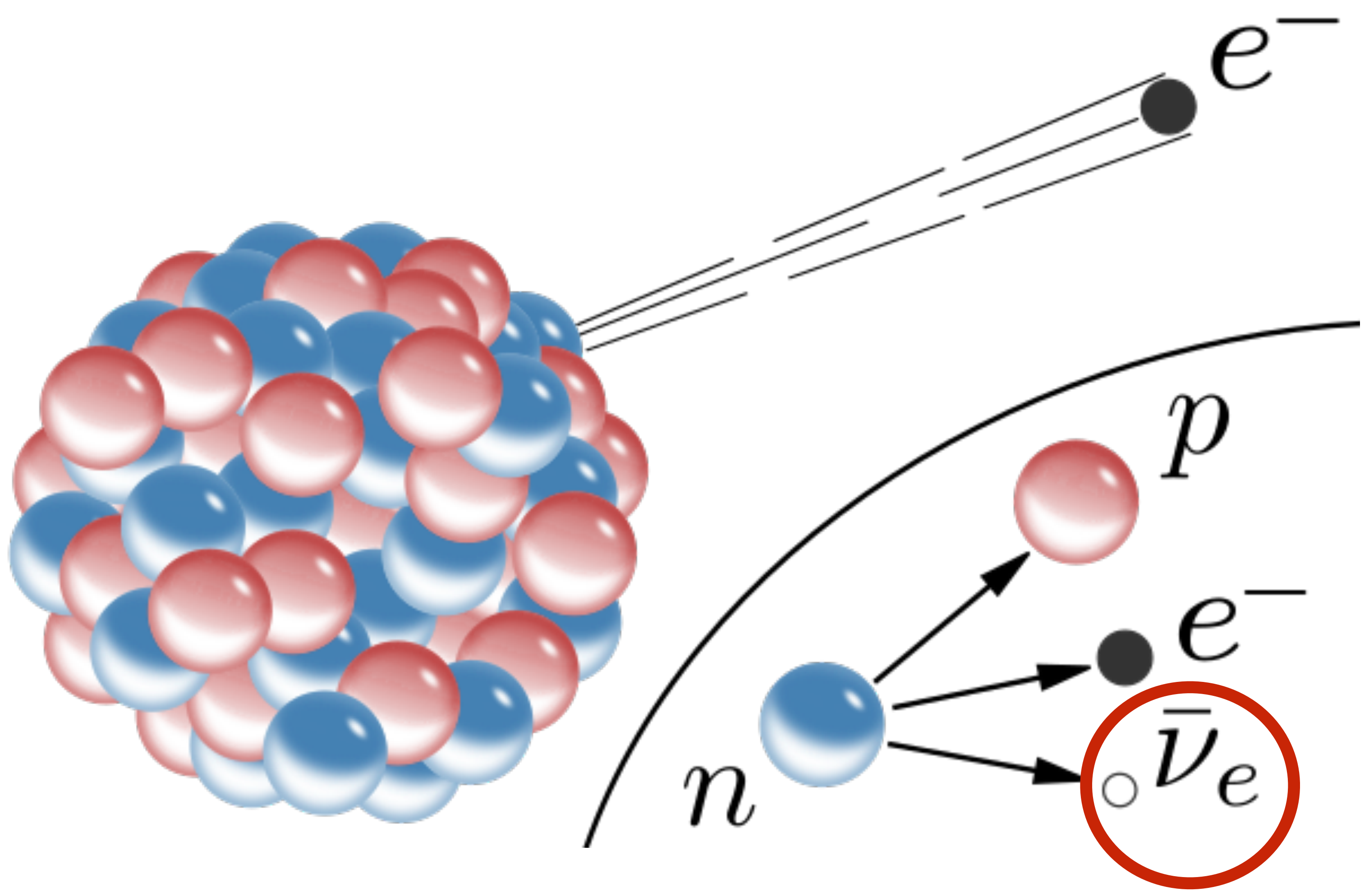


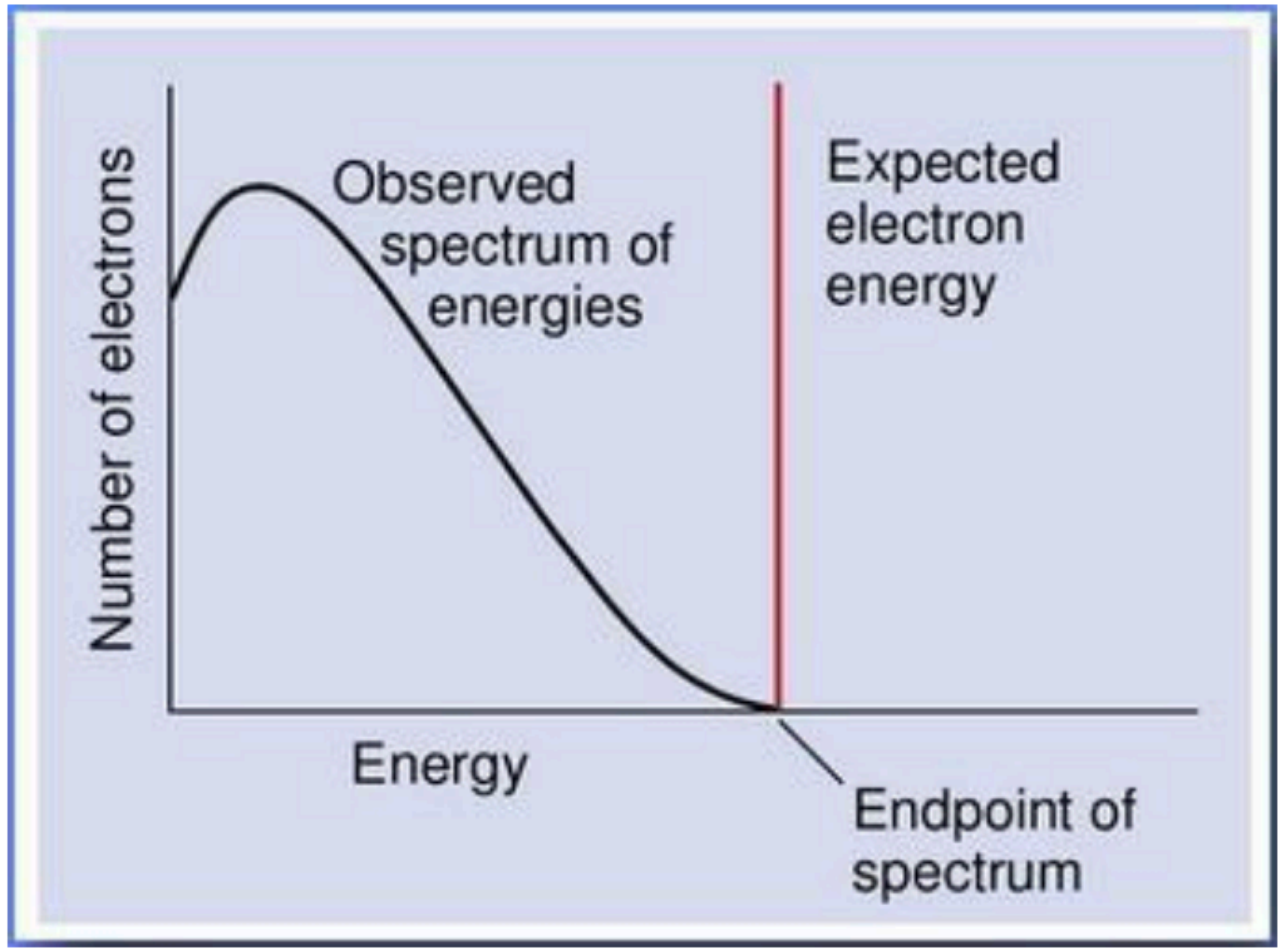
Beta decay: Theoretical prediction

- Wolfgang Pauli in 1930 postulated exactly that:
  an undetectable particle emitted during the decay, sharing the energy with the electron

- given the observation, this particle had to be electrically *neutral* and very *light*
  → ***neutrino***

$$^{210}_{83}\mathrm{Bi} \rightarrow {}^{210}_{84}\mathrm{Po} + e^{-} + \bar{\nu}_e$$





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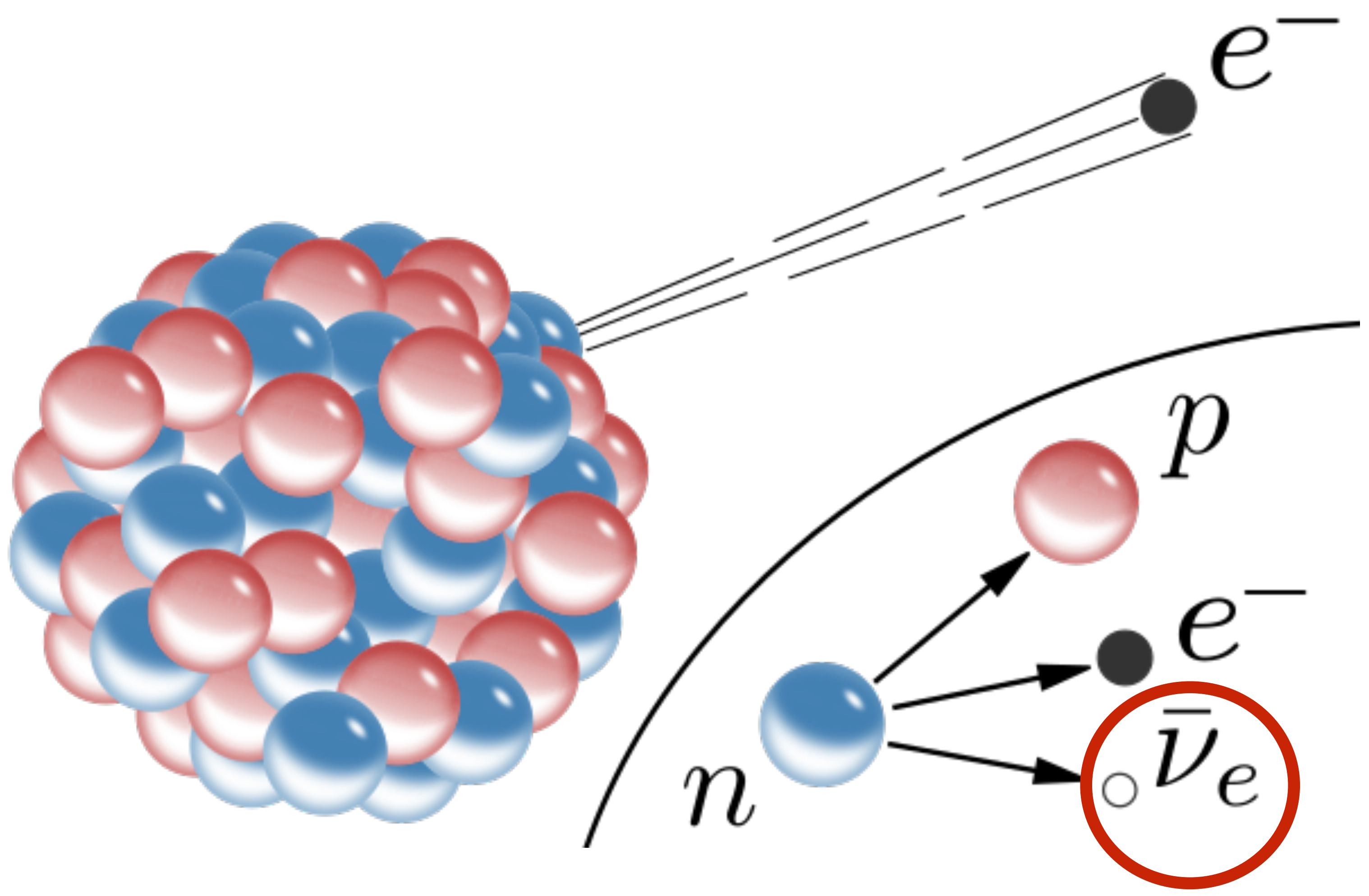

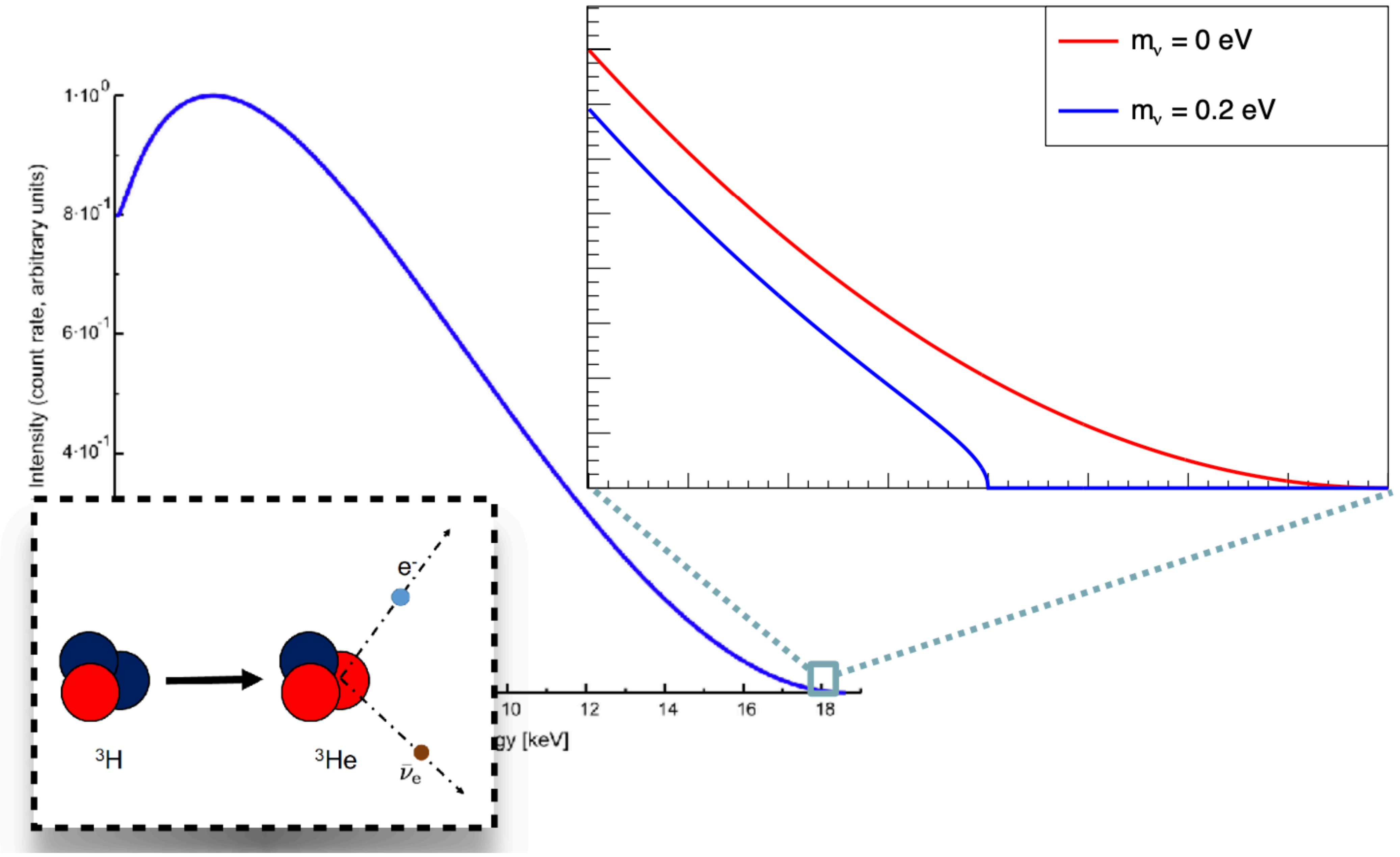

Beta decay: Theoretical prediction

Translation of the open letter sent by Wolfgang Pauli to Lise Meitner and Hans Geiger and a group of radioactive people at the Gauverein meeting in Tübingen.

Zürich, Dec. 4, 1930

Physics Institute of the ETH
Gloriastrasse

Zürich

Dear Radioactive Ladies and Gentlemen,

As the bearer of these lines, to whom I graciously ask you to listen, will explain to you in more detail, because of the "wrong" statistics of the N- and Li-6 nuclei and the continuous beta spectrum, I have hit upon a desperate remedy to save the "exchange theorem" (1) of statistics and the law of conservation of energy. Namely, the possibility that in the nuclei there could exist electrically neutral particles, which I will call neutrons, that have spin 1/2 and obey the exclusion principle and that further differ from light quanta in that they do not travel with the velocity of light. The mass of the neutrons should be of the same order of magnitude as the electron mass and in any event not larger than 0.01 proton mass. - The continuous beta spectrum would then make sense with the assumption that in beta decay, in addition to the electron, a neutron is emitted such that the sum of the energies of neutron and electron is constant.

Now it is also a question of which forces act upon neutrons. For me, the most likely model for the neutron seems to be, for wave-mechanical reasons (the bearer of these lines knows more), that the neutron at rest is a magnetic dipole with a certain moment μ. The experiments seem to require that the ionizing effect of such a neutron can not be bigger than the one of a gamma-ray, and then μ is probably not allowed to be larger than e • (10 -13cm).

But so far I do not dare to publish anything about this idea, and trustfully turn first to you, dear radioactive people, with the question of how likely it is to find experimental evidence for such a neutron if it would have the same or perhaps a 10 times larger ability to get through [material] than a gamma-ray.

I admit that my remedy may seem almost improbable because one probably would have seen those neutrons, if they exist, for a long time. But nothing ventured, nothing gained, and the seriousness of the situation, due to the continuous structure of the beta spectrum, is illuminated by a remark of my honored predecessor, Mr Debye, who told me recently in Bruxelles: "Oh, It's better not to think about this at all, like new taxes." Therefore one should seriously discuss every way of rescue. Thus, dear radioactive people, scrutinize and judge. - Unfortunately, I cannot personally appear in Tübingen since I amindispensable here in Zürich because of a ball on the night from December 6 to 7. With my best regards to you, and also to Mr. Back, your humble servant

W. Pauli

Wolfgang Ernst Pauli

Pauli in 1945

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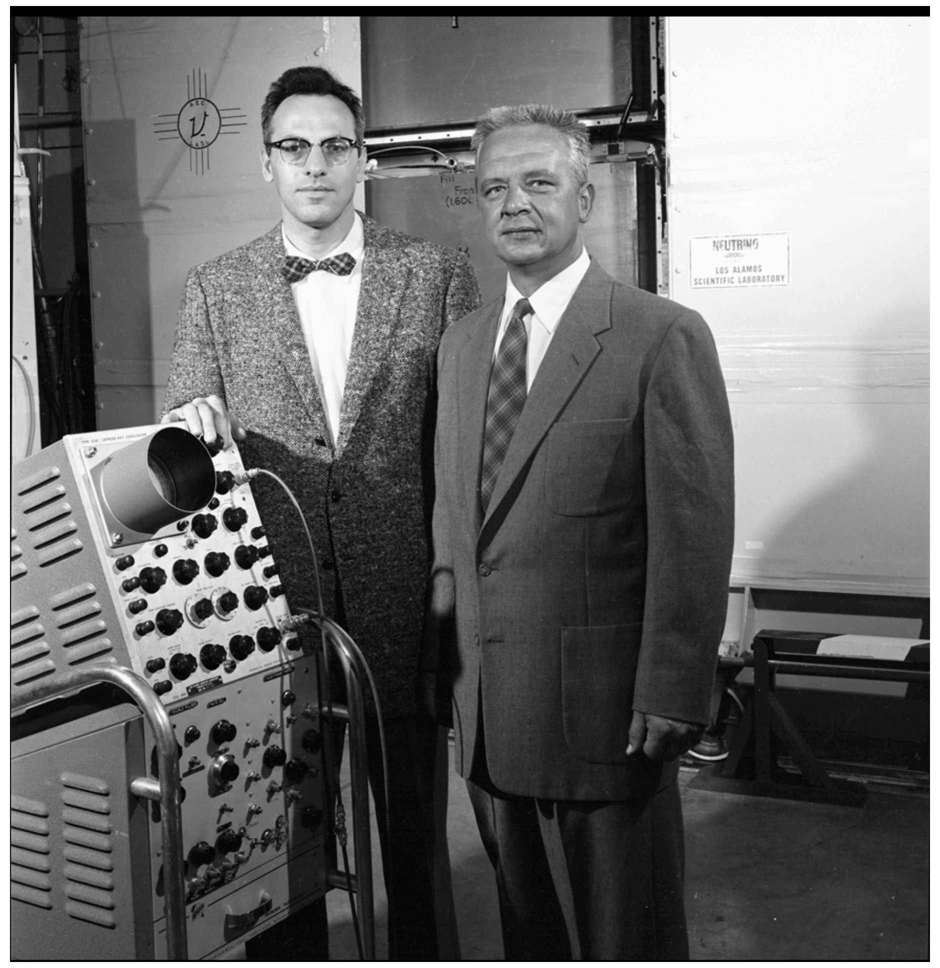
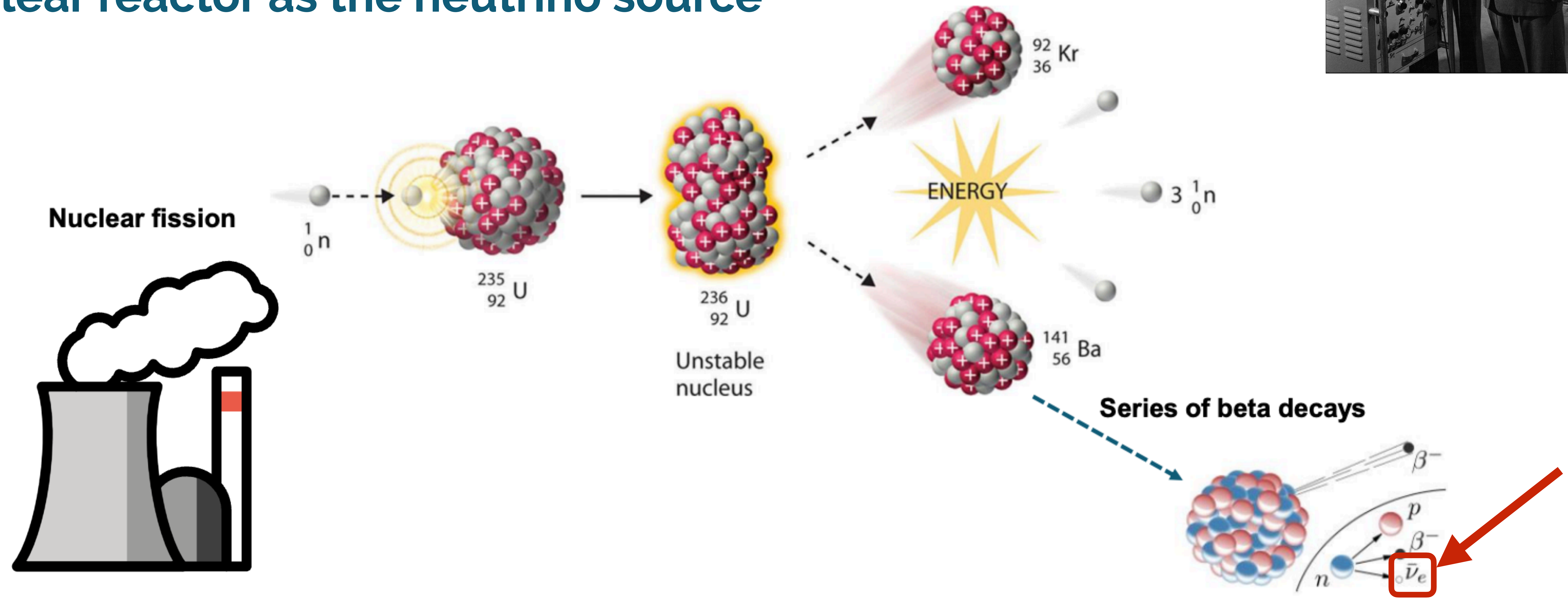
Detection of neutrinos

Nobel Prize 1995



- neutrinos were first detected experimentally in 1956 by Clyde Cowan and Frederick Reines

- the experiment took place in the Savannah River Plant using a **nuclear reactor as the neutrino source**



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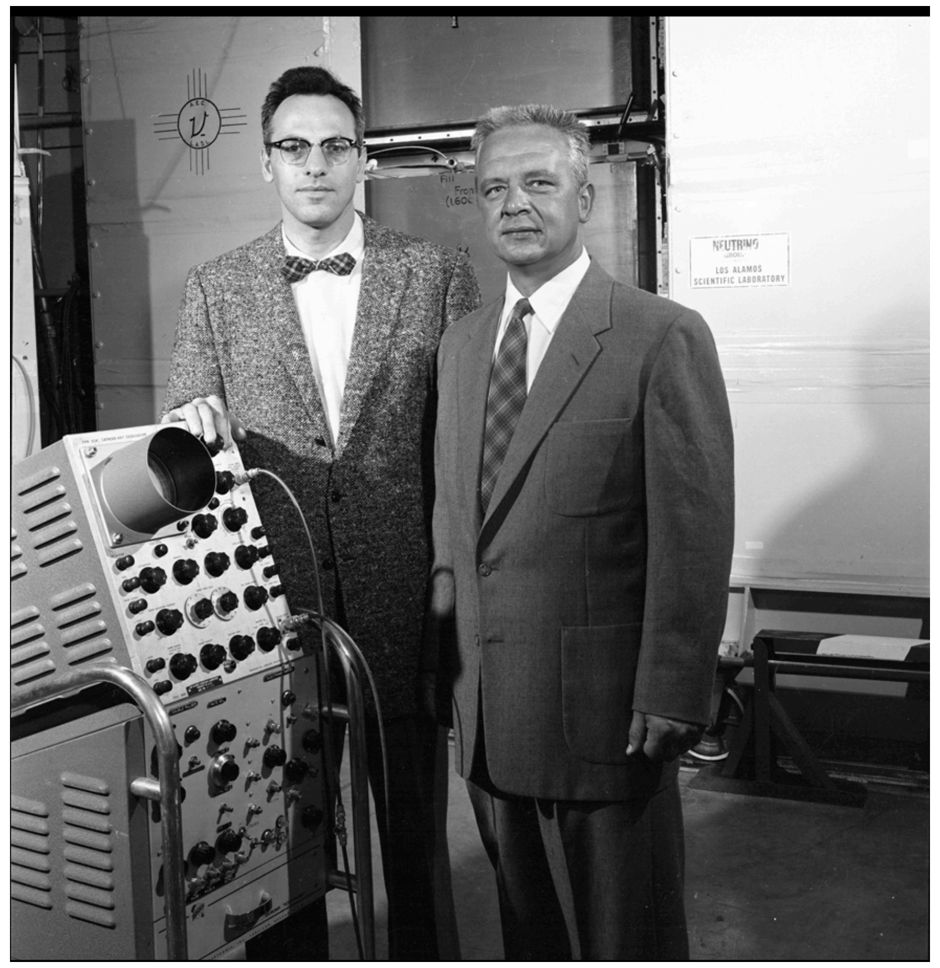

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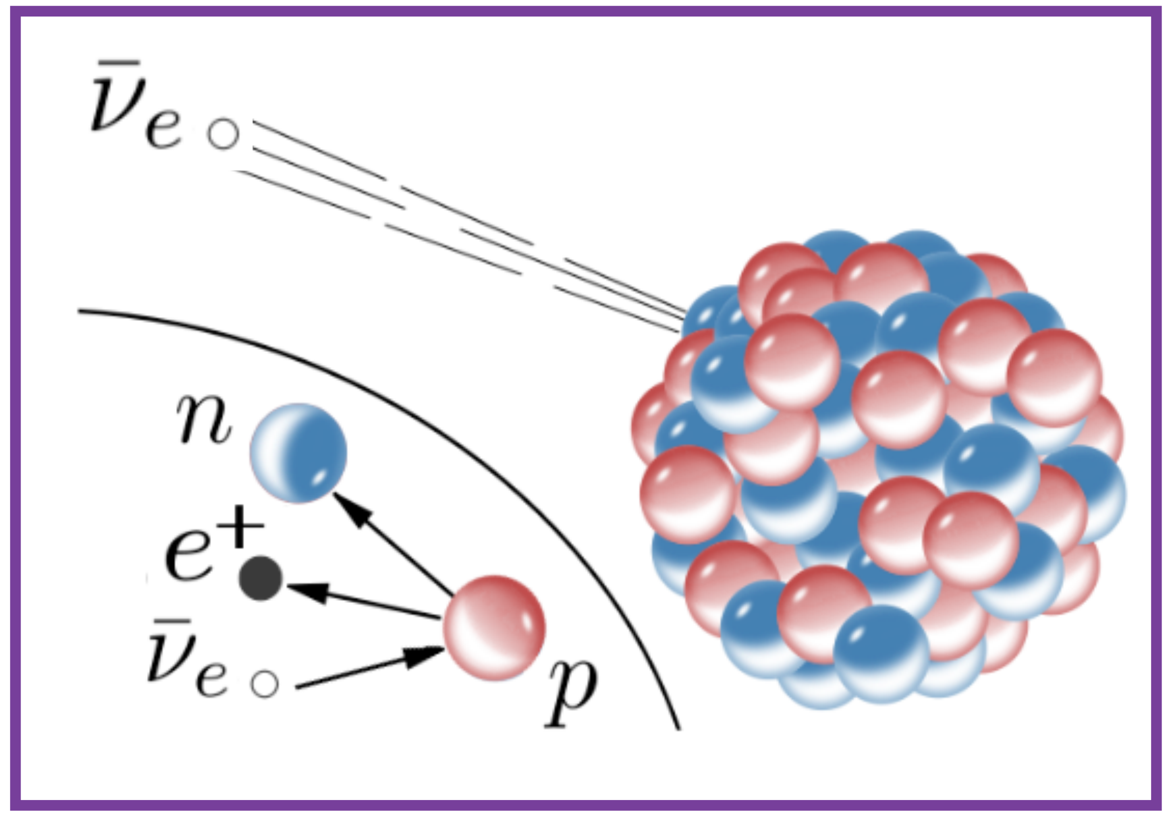
- the detection was made with the ***inverse beta decay*** process:
  neutrino interact with **proton**, turning it into **neutron** and emitting **positron**
  - using water (with large number of **protons**) with cadmium chloride
  - signal 1: **positron** interact immediately with electrons, create gamma rays
  - signal 2: **neutron** captured by cadmium, gives off a gamma ray

- *note: notice “beta” particles (electrons) are always associated here?*



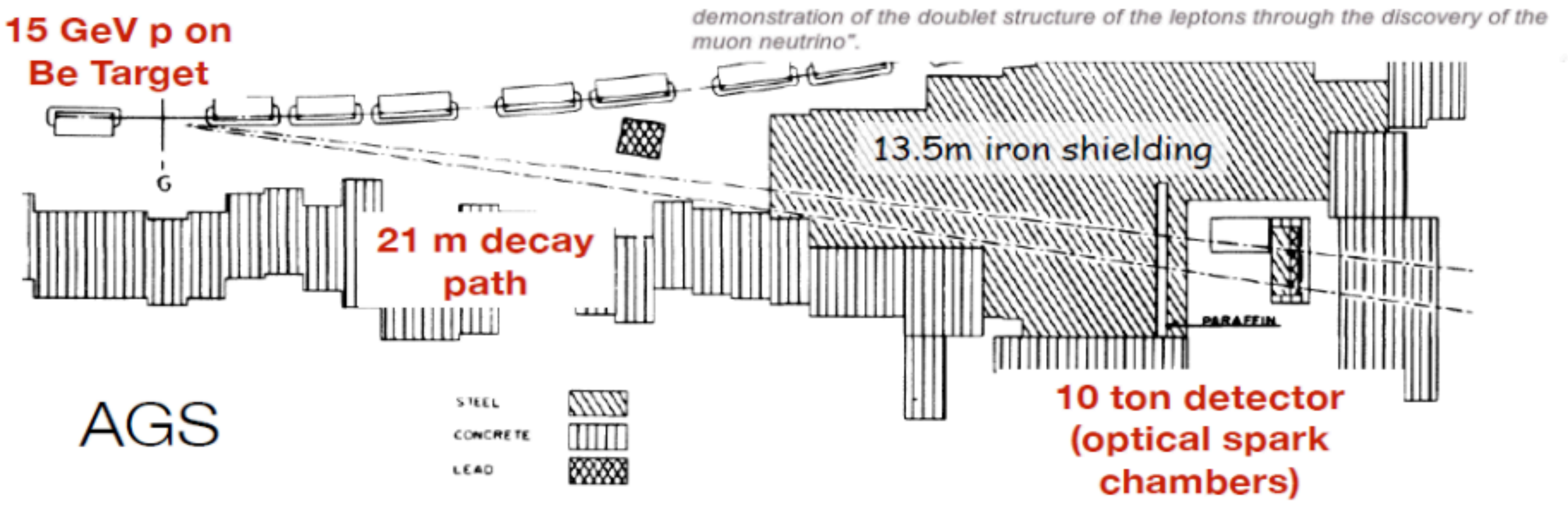
$$\bar{\nu}_e + p \rightarrow e^+ + n$$

Detection of neutrinos: more flavors

Nobel Prize 1988

- with following theories, different types of neutrinos were hypothesized
  - **muons** were discovered in cosmic rays in 1936; very similar to electrons, but heavier
  - muon neutrinos, **associated with muons instead of electrons**, then may exist
    - remember how beta decay always involves electrons?
  - at BNL in 1962, using proton beam on Be target, Lederman/Schwartz/Steinberger discovered neutrinos producing muons
- similarly, followed by tau lepton discovery, tau neutrino was discovered by DONUT experiment at Fermilab (2000)

World's first accelerator neutrino experiment



$$\nu_\mu + n \rightarrow \mu^- + p$$

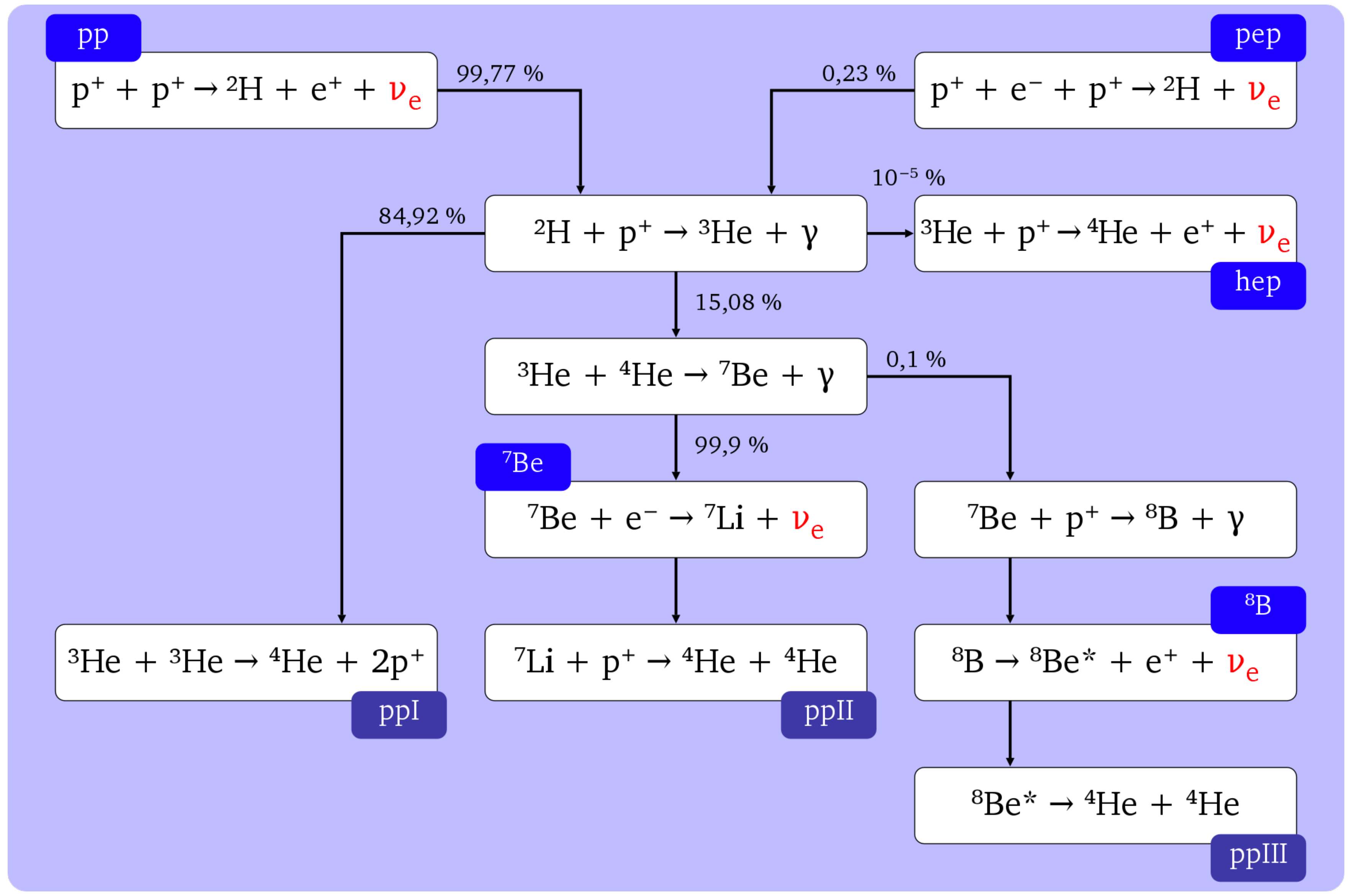
Missing neutrinos: how neutrinos change their flavors

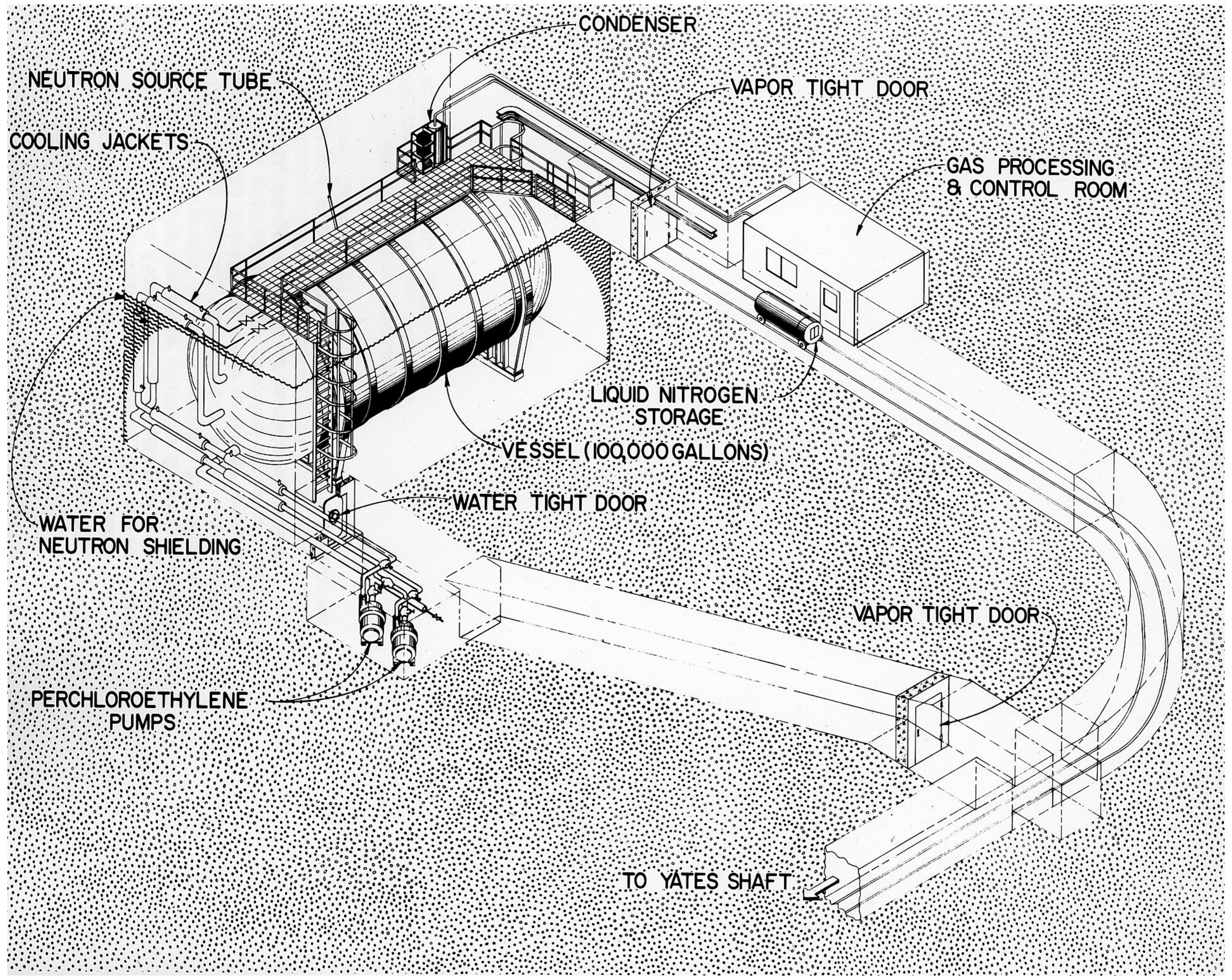
The solar neutrino problem

Nobel Prize 2002

- the Sun is a huge source of neutrinos from nuclear fusion
  - ~100 billion neutrinos from the Sun pass through your thumbnail every second!
- J. Bahcall calculated expected number of solar neutrinos expected to arrive at Earth
  - all the neutrinos generated in the Sun was to be electron neutrinos
- R. Davis used the Homestake experiment to detect these neutrinos in 1968
  - buried deep underground (1500m) to avoid cosmic ray background
  - one needs to detect chlorine-to-argon conversion, from an interaction from the electron neutrinos

$$\nu_e + {}^{37}_{17}Cl \rightarrow {}^{37}_{18}Ar + e^-$$

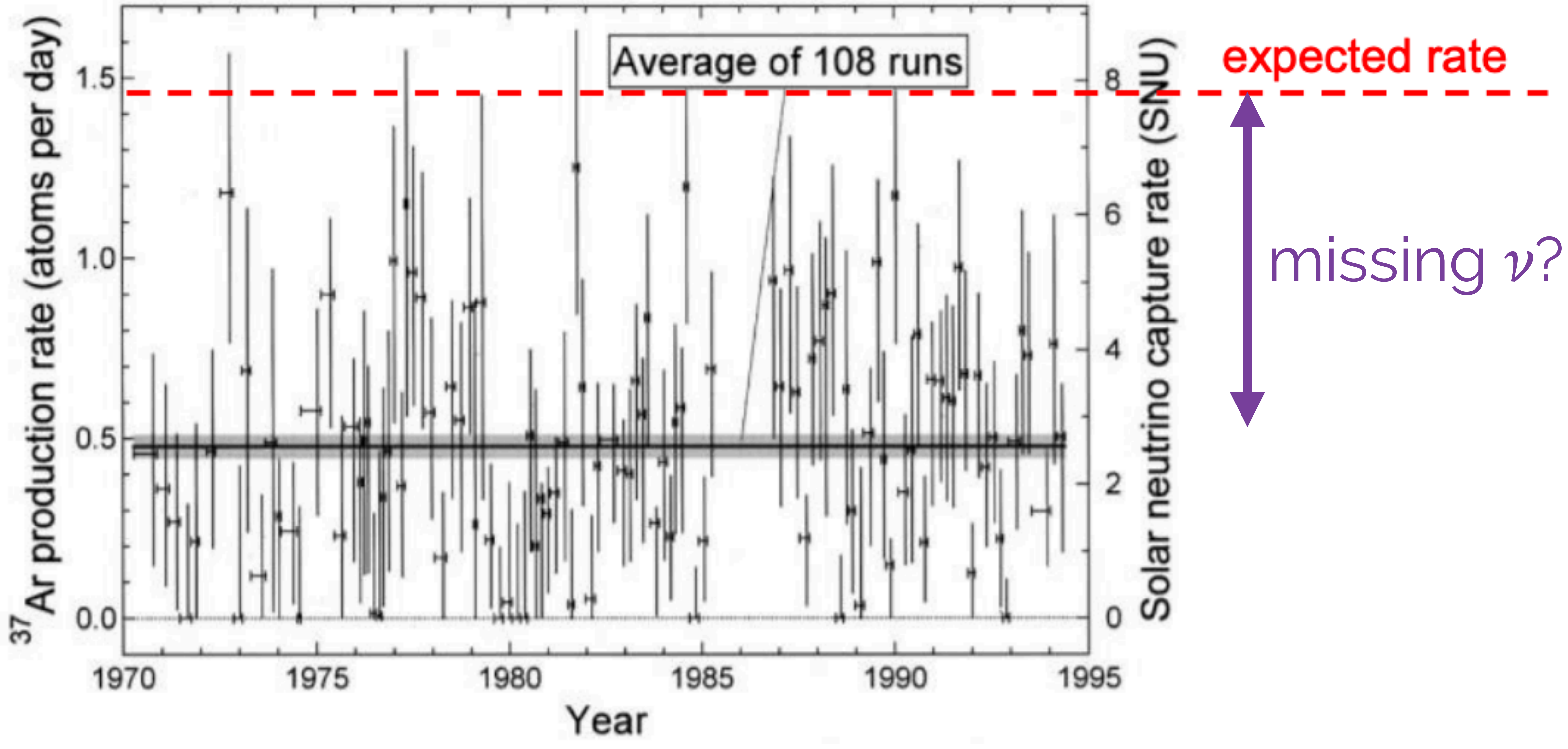




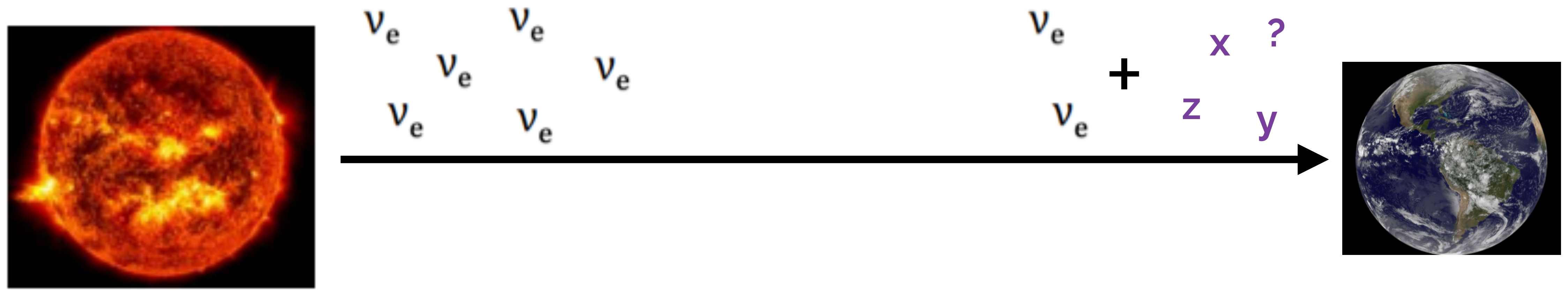
The solar neutrino problem

$$\nu_e + {}^{37}_{17}Cl \rightarrow {}^{37}_{18}Ar + e^-$$

- but detected number of (electron) neutrinos seem to be too small: **2/3 of them were missing!**
  - was the solar neutrino calculation off?
  - was it an error with the experiment?

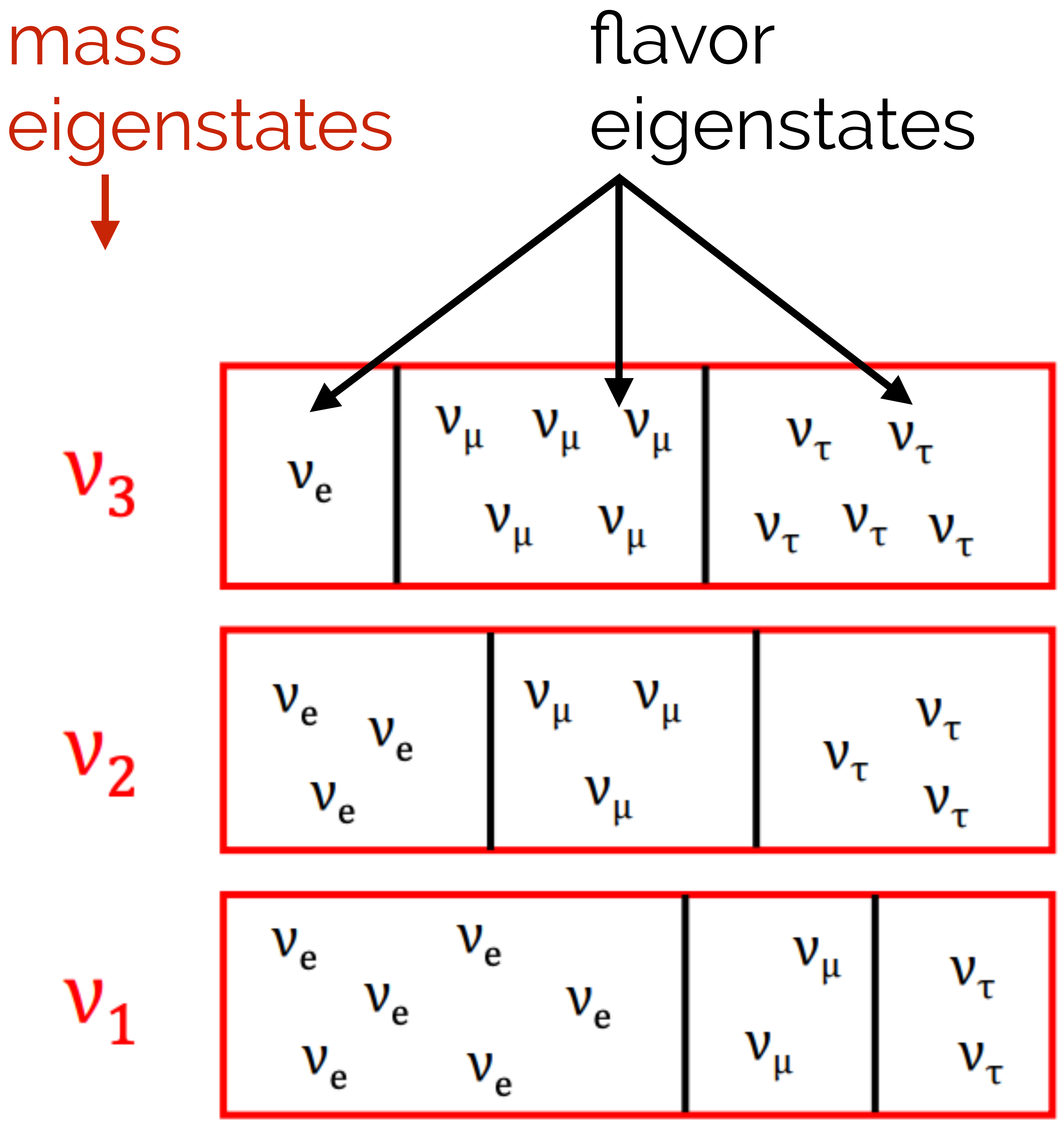


- B. Pontecorvo suggested in 1969, that **neutrinos could change in some way while traveling** from Sun to Earth



Neutrino oscillation: model

- **flavor mixing**: neutrinos are special, as they have two different eigenstates with weak/flavor & mass
  - **weak/flavor eigenstates**: states associated with the weak interaction, in which neutrinos are produced & detected; "interaction basis"
  - **mass eigenstates**: states of definite mass, which propagate through space-time; "propagation basis"
  - *note: in order to have this "oscillation", neutrinos need to have mass!*
- PMNS matrix can describe the mixing between these two eigenstates



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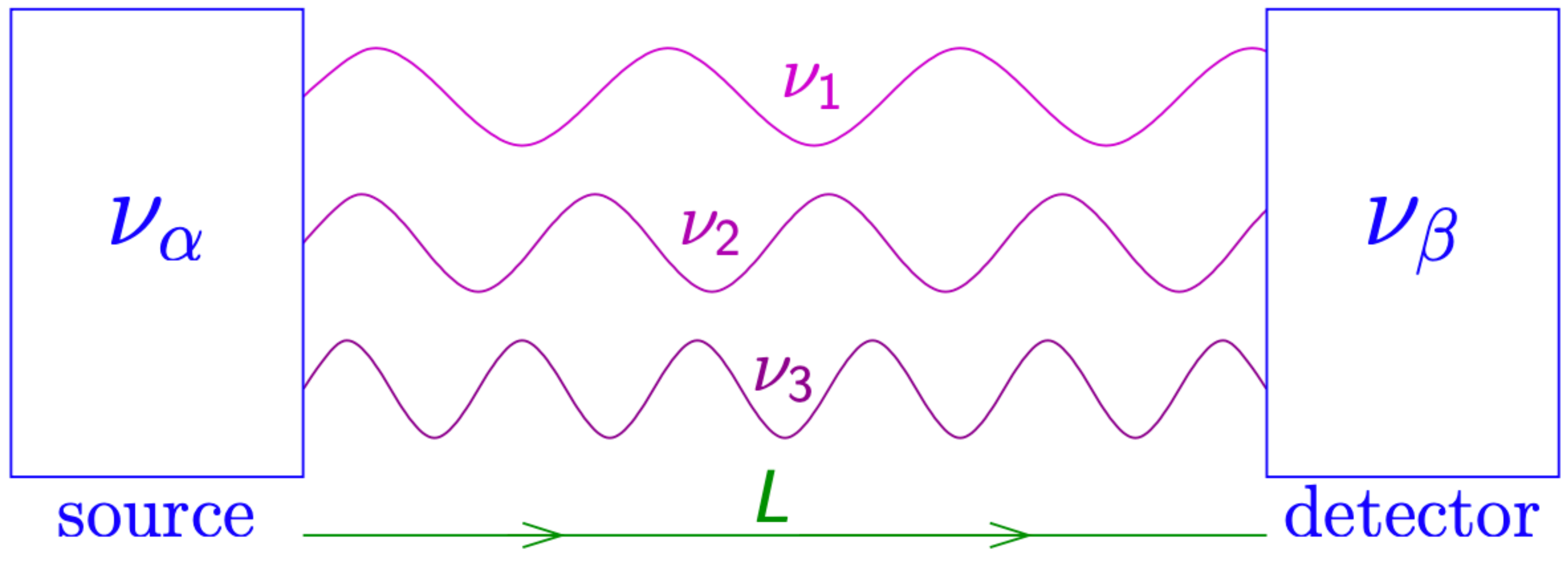

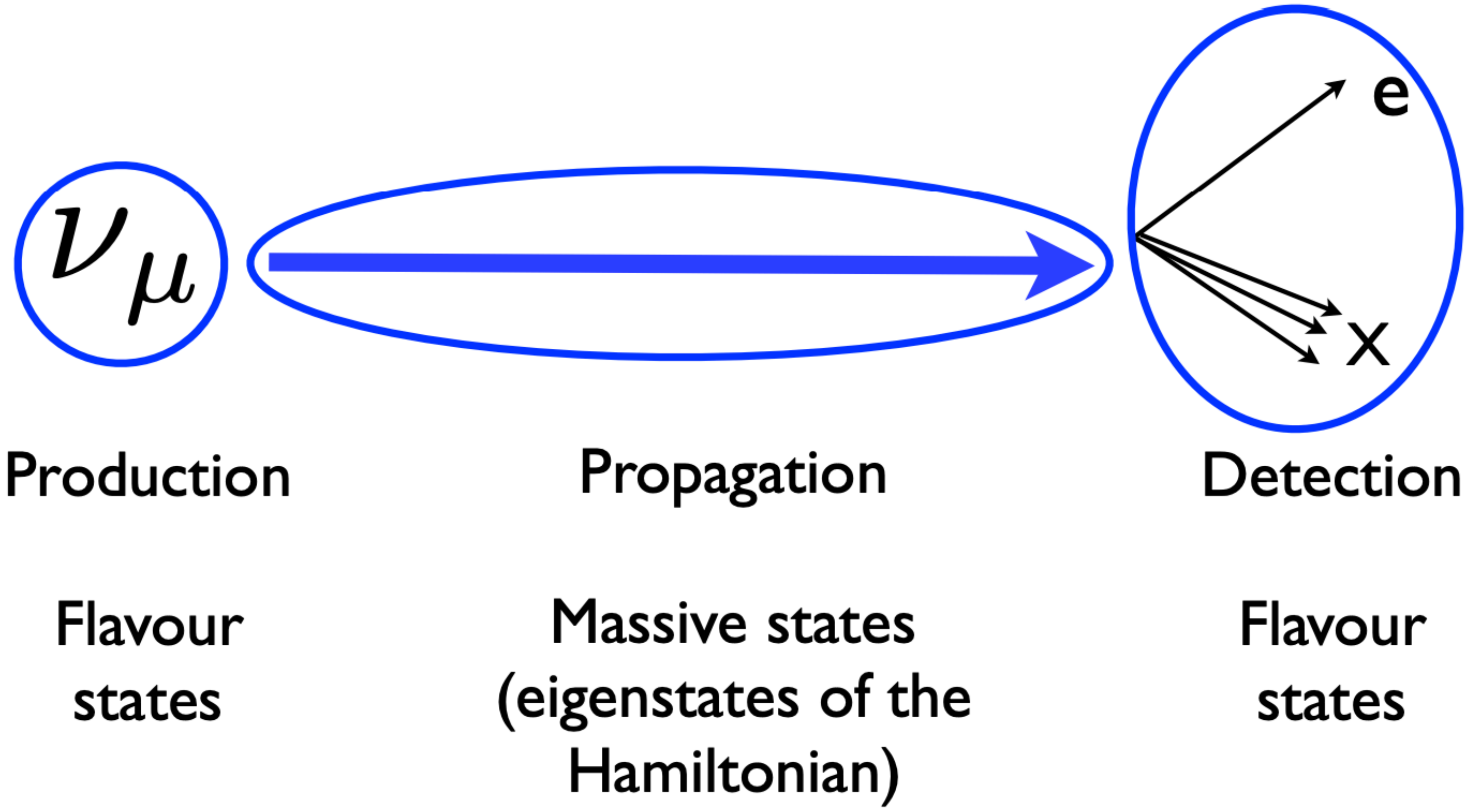

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flavor (α = e, μ, τ) ⇔ linear combinations ⇔ mass (i = 1, 2, 3)

$$|\nu_\alpha\rangle = \sum_i U^*_{\alpha i} |\nu_i\rangle \quad \Leftrightarrow \quad |\nu_i\rangle = \sum_\alpha U_{\alpha i} |\nu_\alpha\rangle$$

$$\begin{bmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{bmatrix} = \begin{bmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{bmatrix} \begin{bmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{bmatrix}.$$

Pontecorvo-Maki-Nakagawa-Sakata (PMNS) matrix

Neutrino oscillation: model

- PMNS matrix contains different parameters, such as mixing angles and δ_c
  - mixing angles represent the oscillation probabilities between two flavors of neutrinos

- δ_{cp} represents the difference in oscillation between neutrinos and antineutrinos
  - if δ_{cp} =0, neutrinos and antineutrinos will behave/oscillate in same fashion
  - if not, CP violation exist in the neutrino sector: which is a necessary condition for explaining matter-antimatter asymmetry in our Universe

$$\begin{bmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{bmatrix} = \underbrace{\begin{bmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{bmatrix}}_{\text{Pontecorvo-Maki-Nakagawa-Sakata (PMNS) matrix}} \begin{bmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{bmatrix}.$$

Mixing angles = (θ_{12}, θ_{23}, θ_{13}) , δ_{CP} is the CP-violation phase

$$U_{\mathrm{PMNS}} = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\substack{\text{I} \\ \text{atmospheric}}} \underbrace{\begin{pmatrix} c_{13} & 0 & e^{i\delta_{\mathrm{CP}}} s_{13} \\ 0 & 1 & 0 \\ -e^{i\delta_{\mathrm{CP}}} s_{13} & 0 & c_{13} \end{pmatrix}}_{\substack{\text{II} \\ \text{reactor}}} \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\substack{\text{III} \\ \text{solar}}}$$

where: $c_{\alpha\beta}$ = cos $\theta_{\alpha\beta}$; $s_{\alpha\beta}$ = sin $\theta_{\alpha\beta}$

Nonzero δ_{CP} ⟹ neutrinos and antineutrinos oscillate different

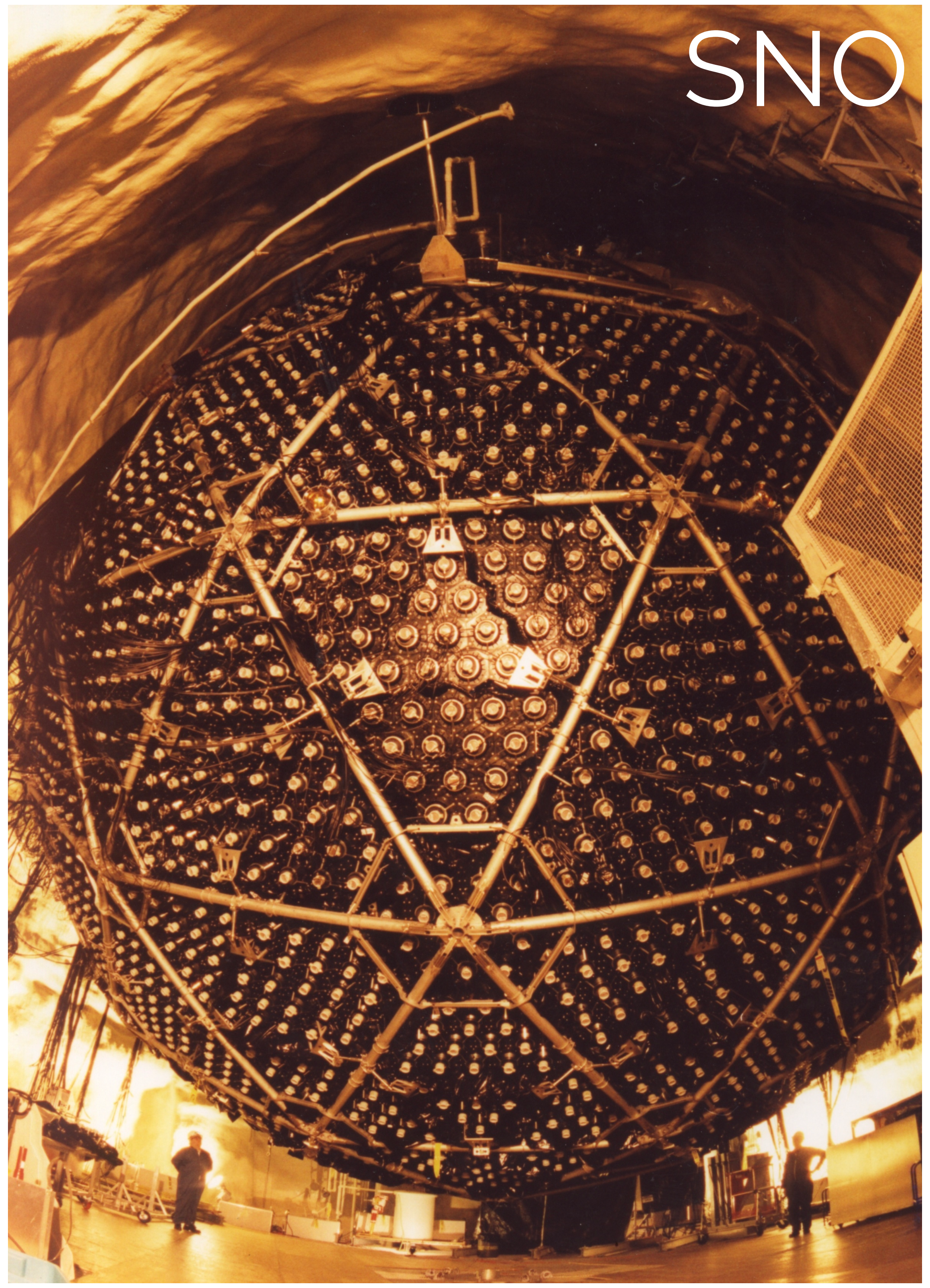
Neutrino oscillation: confirmation

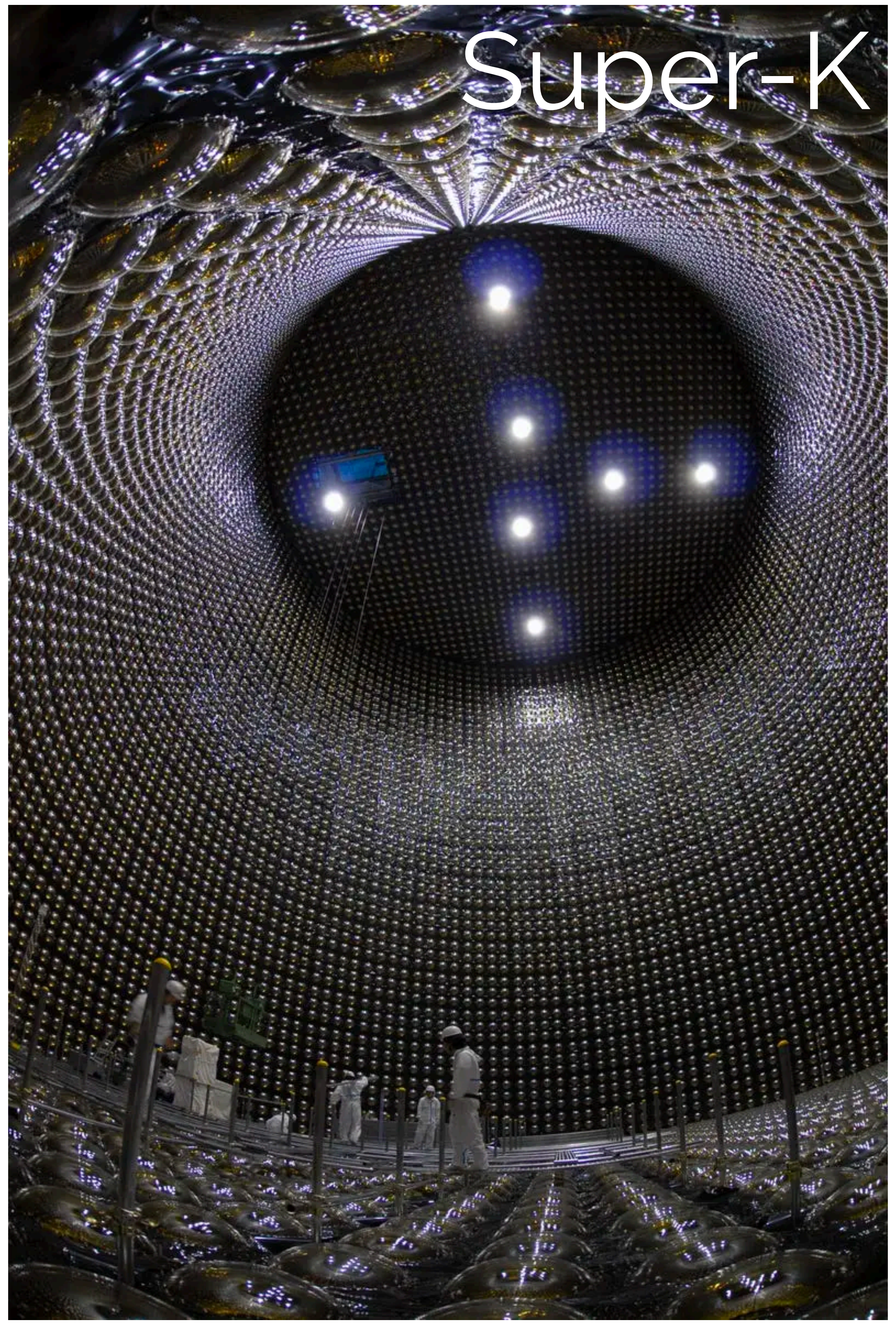
- neutrino changing its flavor during travel (neutrino oscillation) can explain why we detect only 1/3 of expected electron neutrinos; but can this neutrino oscillation be experimentally confirmed?

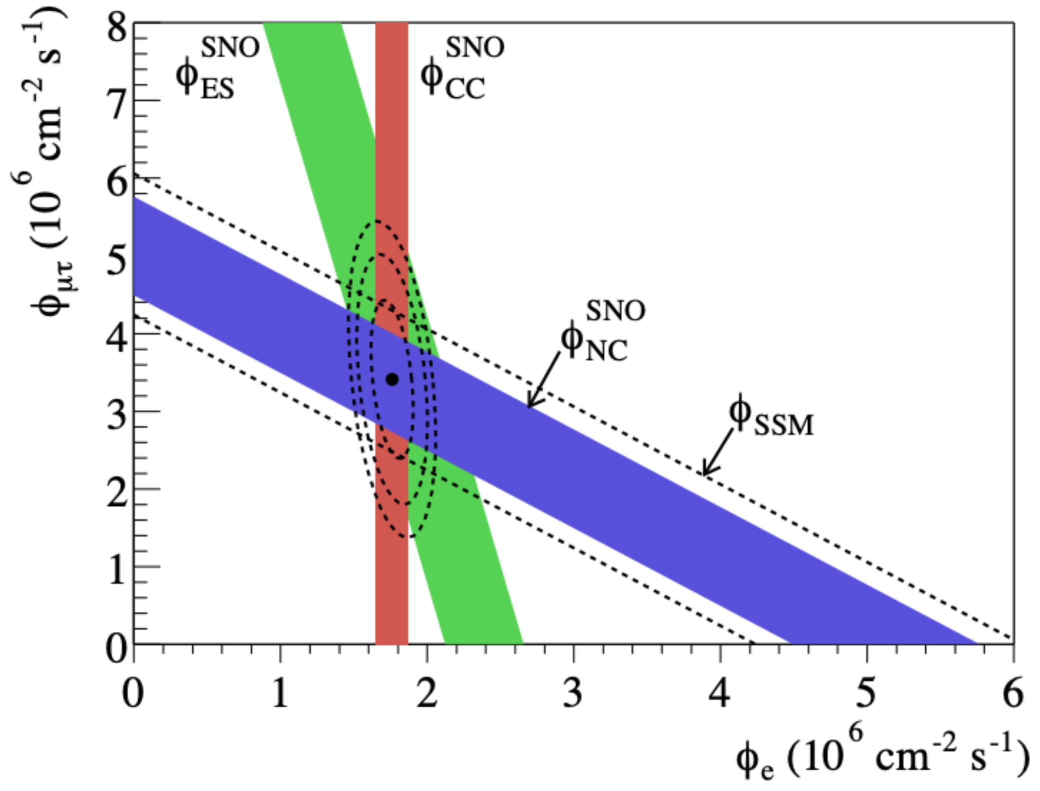
- two separate experiments, Super-Kamiokande and SNO, confirmed this:

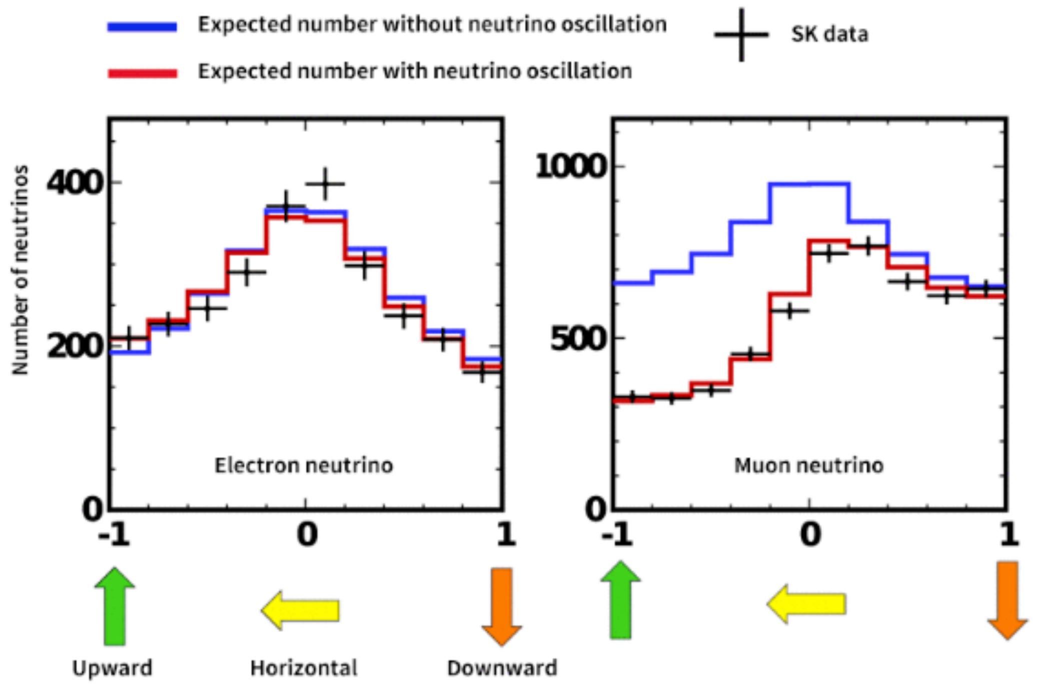
  - **SNO** detected different interactions from solar neutrinos, sensitive to different neutrino flavors: confirming **electron neutrino indeed changed into muon/tau neutrinos**

  - **SK** observed a deficit of muon neutrinos coming from the opposite side of the Earth (longer travel distance), compared to those coming from right above (shorter travel distance): **confirming muon neutrinos oscillating into tau neutrinos**





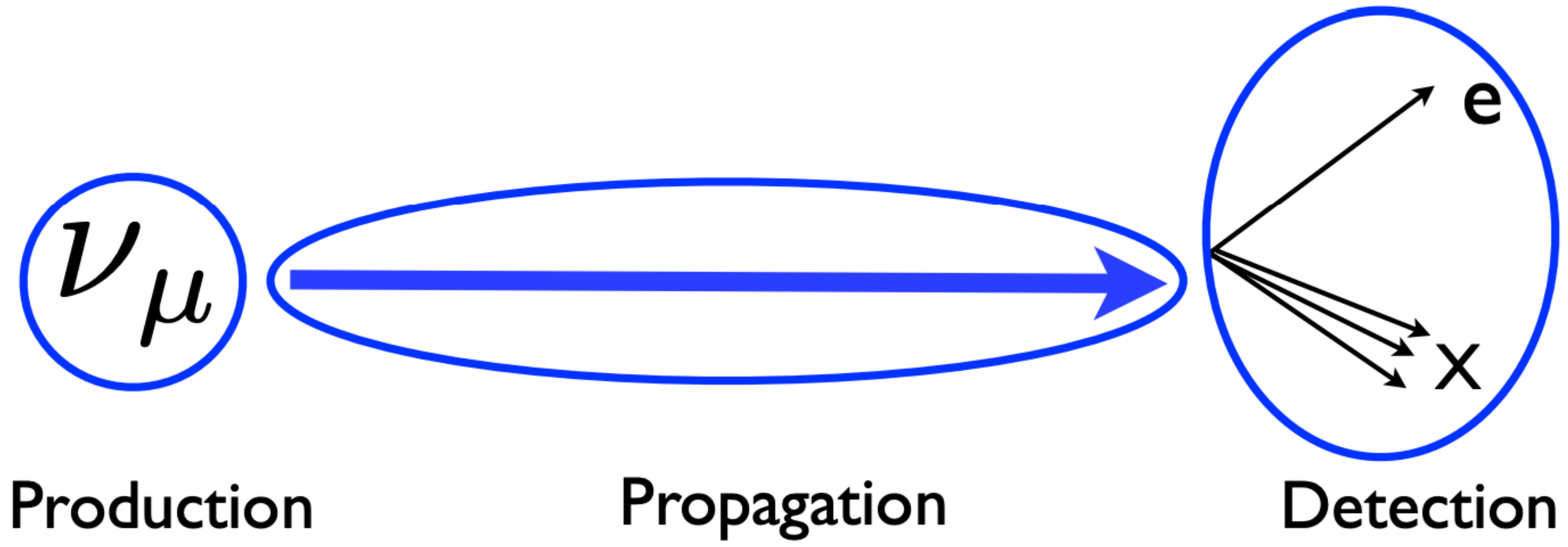


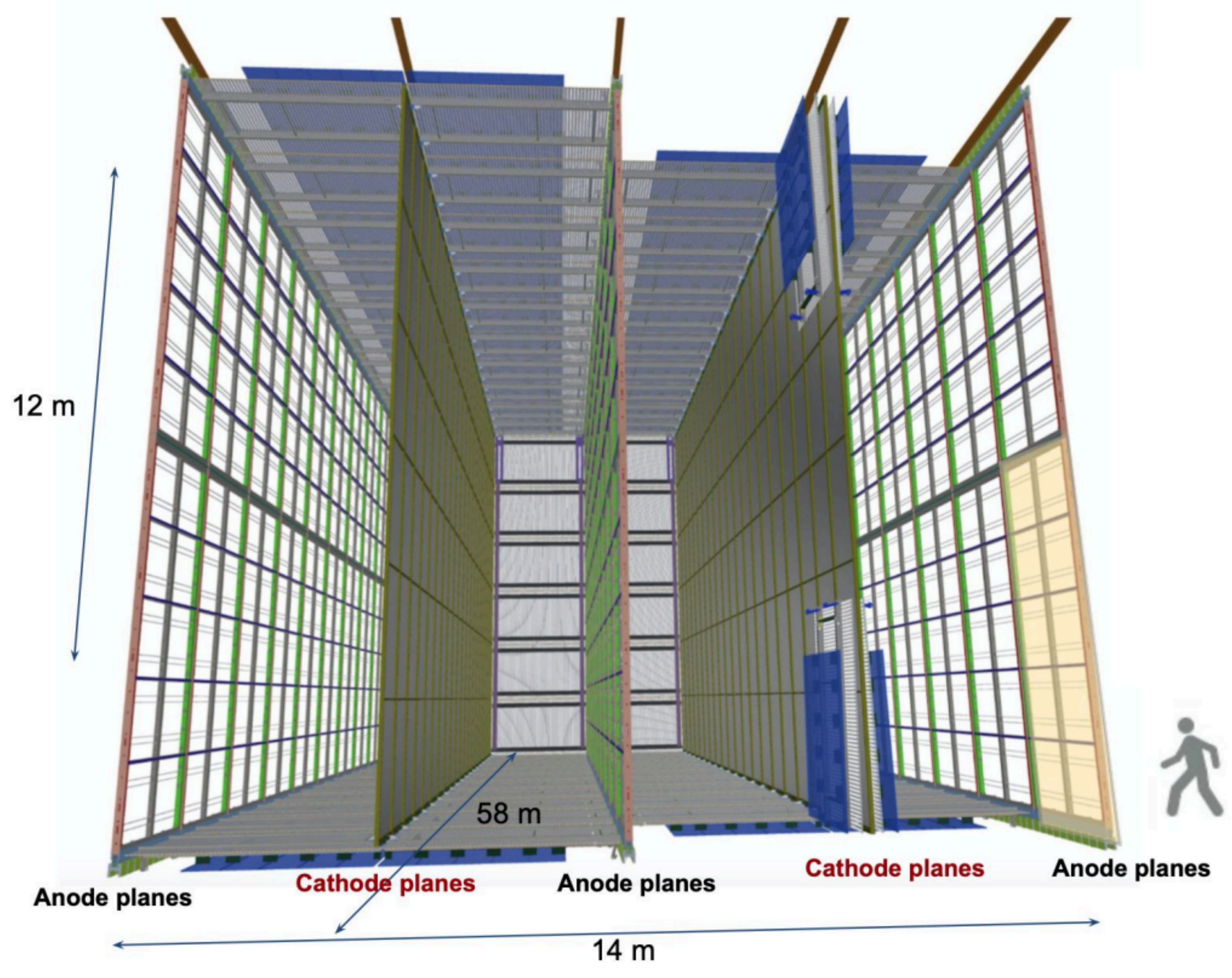


Detecting invisible particles: what, how, & why we detect neutrinos

What do we detect

- neutrinos are really hard to detect: they carry no charge, interact very rarely, and only through the Weak force
- but **we can detect what comes out** of the interaction of the neutrinos
  - electrons, muons, hadrons, photons, …
- what do we need?
  - need a **source** of neutrinos: either natural or artificial
  - need a **big detector**, to increase probability of the neutrino interaction
  - need to detect the **outgoing particles** precisely
  - need good **theoretical predictions/models** of neutrino production, propagation, and interaction





How do we detect

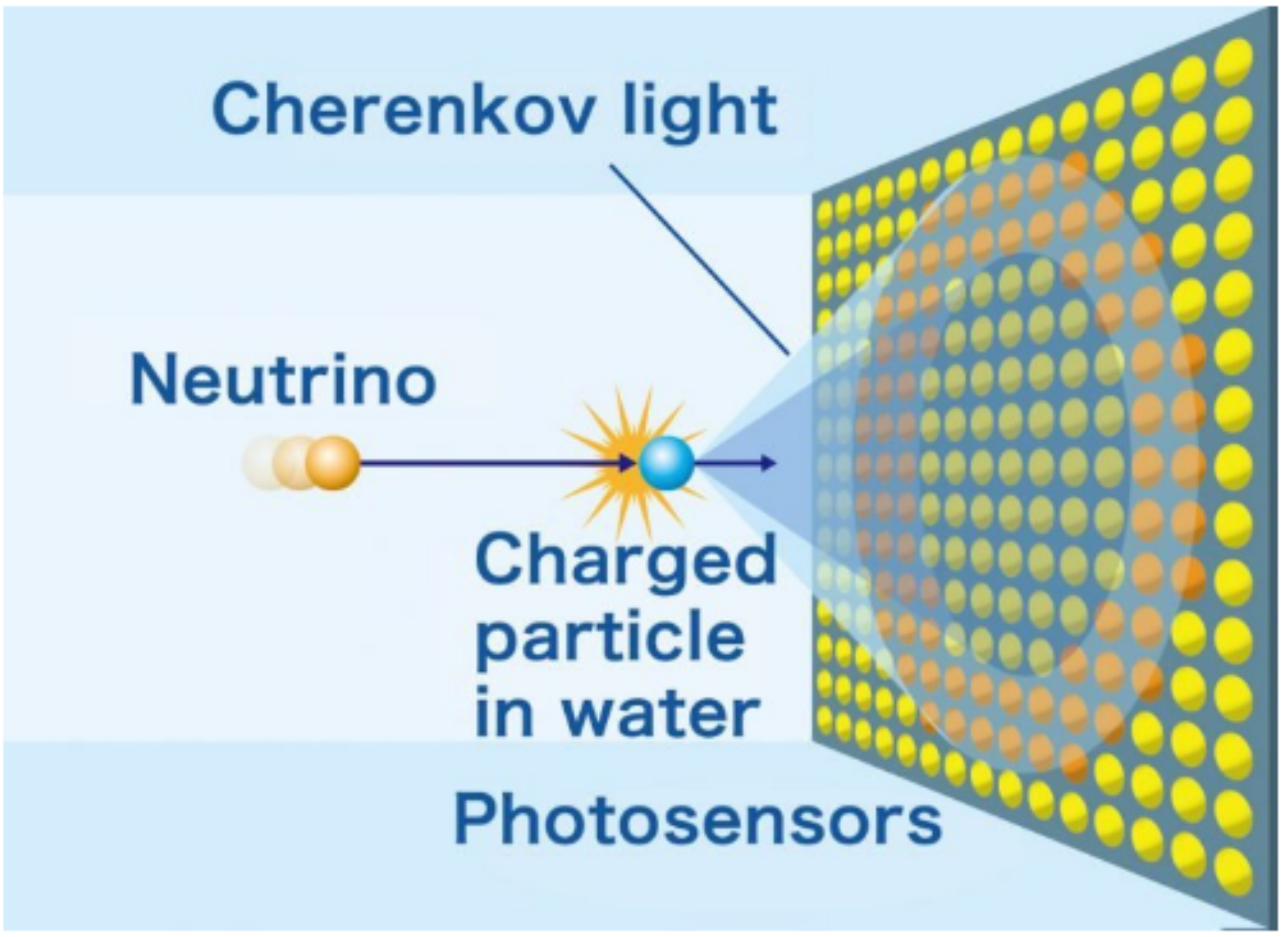
- at the end of the day, we want to (precisely) detect outgoing particles: their type, direction, energy, …

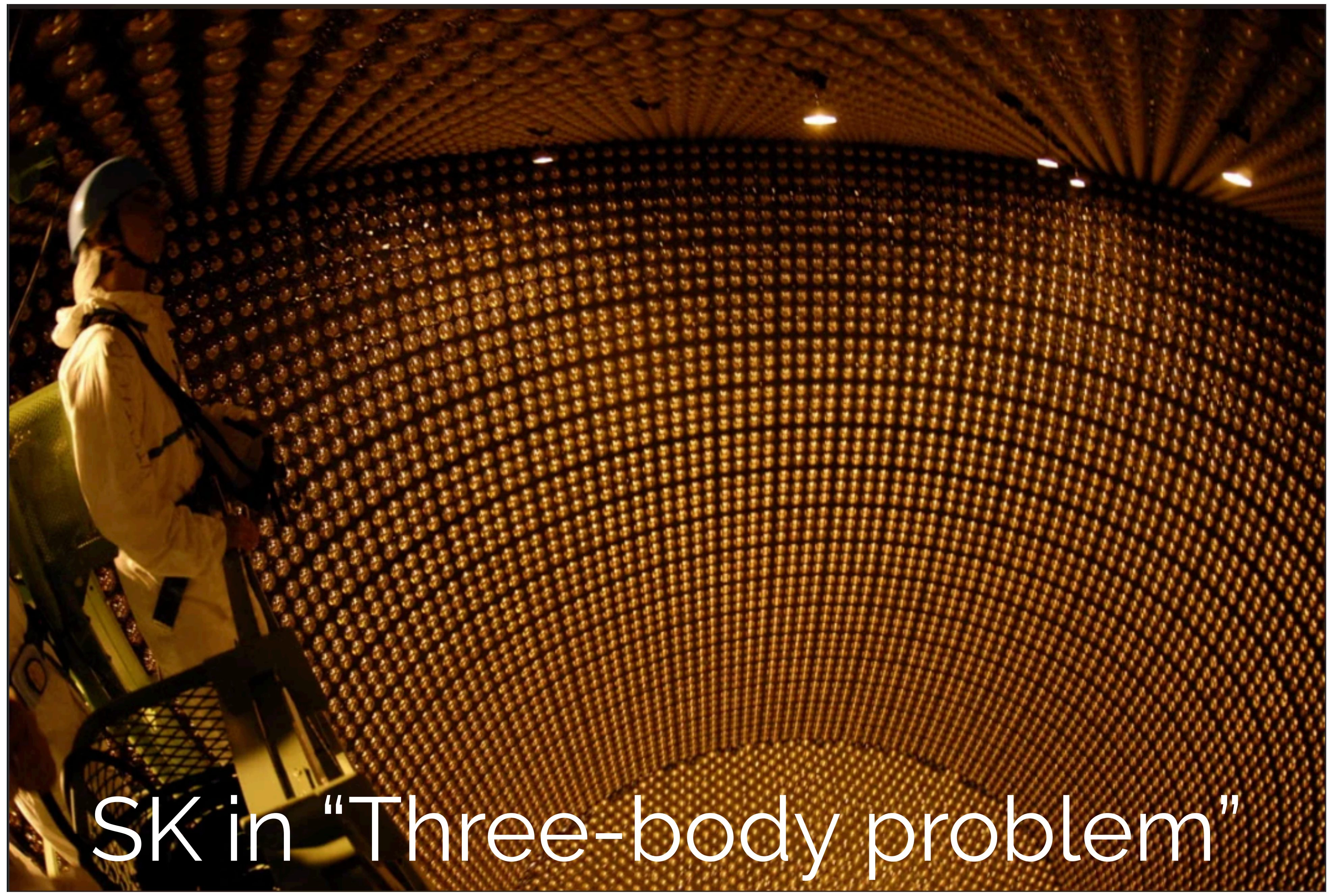
- as technology develops, we began to gather more and more information of these particles

- an example: **water Cherenkov detector** of Super-Kamiokande, where charged particle generating a "light shock wave" as it travels through water

- here we take a look at state-of-the-art neutrino detection technology: **Liquid Argon Time Projection Chamber (LArTPC)**

physicsopenlab.org





How do we detect: Liquid Argon Time Projection Chamber

- Liquid argon (LAr) as total absorption calorimeter
  - dense, abundant, cheap
  - ionization and scintillation signals
- Time Projection Chamber (TPC) as 4π charged particle detector
  - 3D reconstruction with a fully active volume
- LAr+TPC: fine-grained 3D tracking with local dE/dx information and fully active target medium

NUCLEAR INSTRUMENTS AND METHODS 120 (1974) 221–236; © NORTH-HOLLAND PUBLISHING CO.

LIQUID-ARGON IONIZATION CHAMBERS AS TOTAL-ABSORPTION DETECTORS*

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Received 14 May 1974

1974

The Time-Projection Chamber
- A new 4π detector for charged particles

David R. Nygren

Lawrence Berkeley Laboratory
Berkeley, California 97420

1976

THE LIQUID-ARGON TIME PROJECTION CHAMBER:

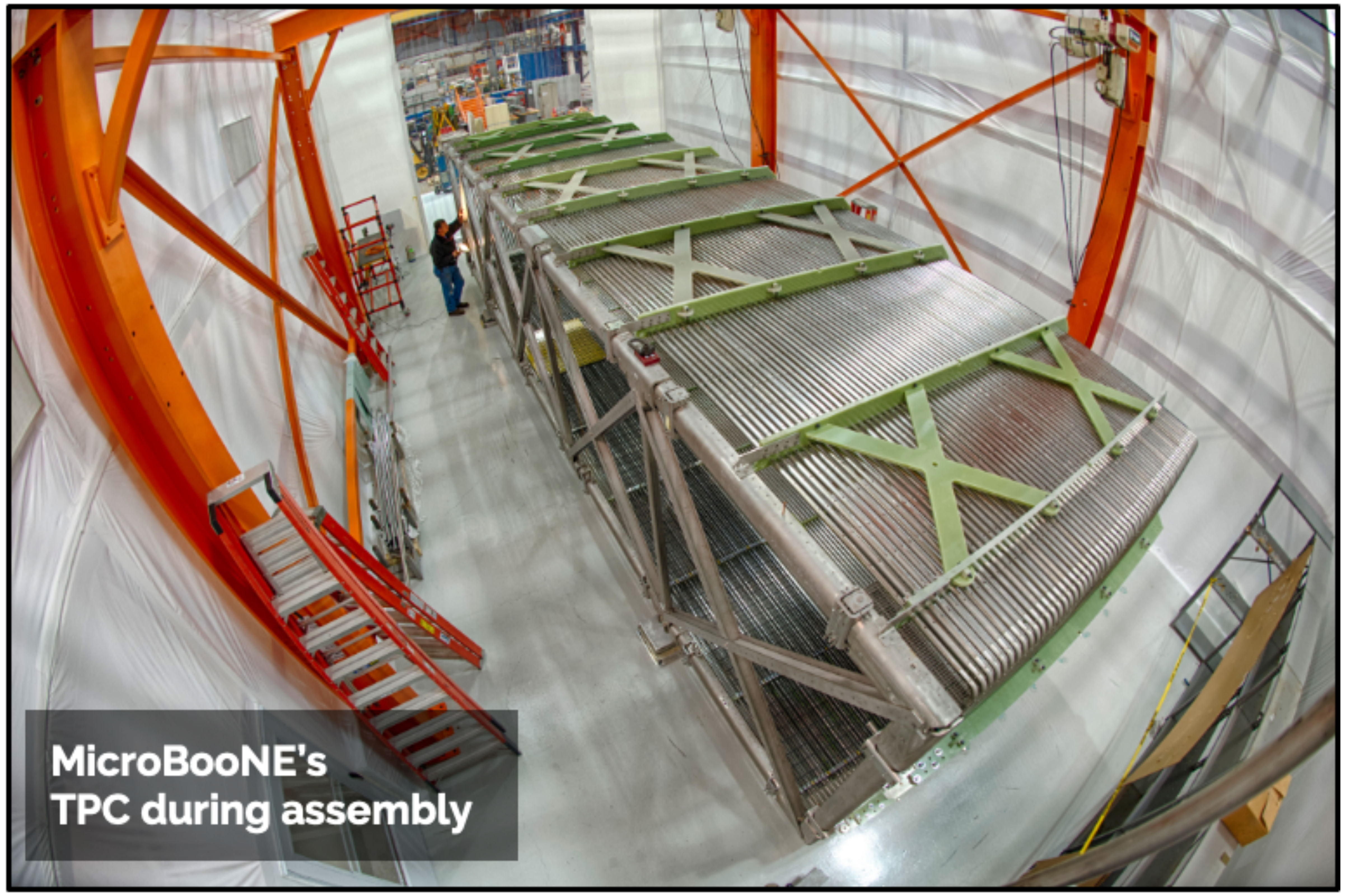
A NEW CONCEPT FOR NEUTRINO DETECTORS

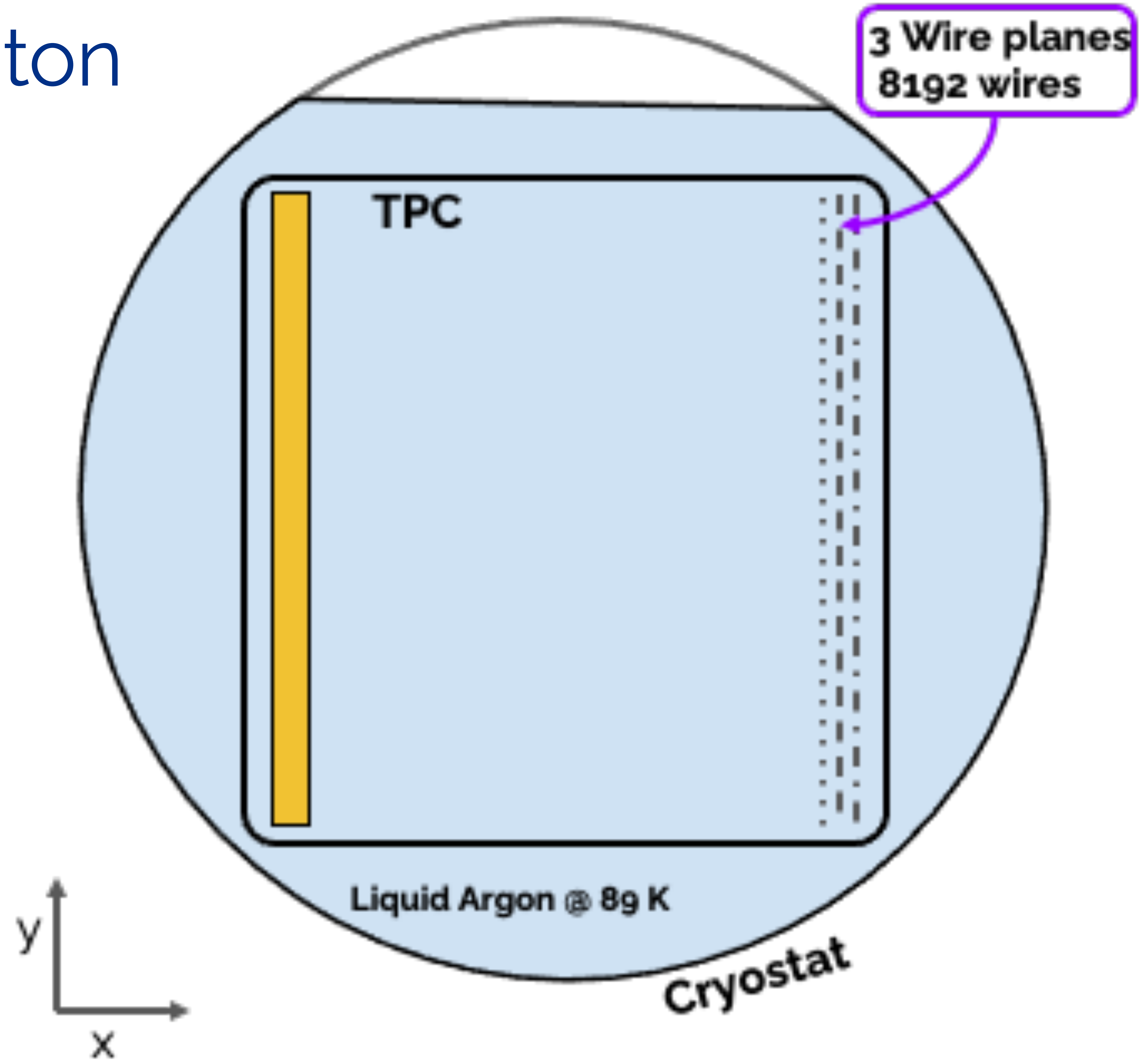
C. Rubbia

1977

LArTPC principle: the MicroBooNE detector

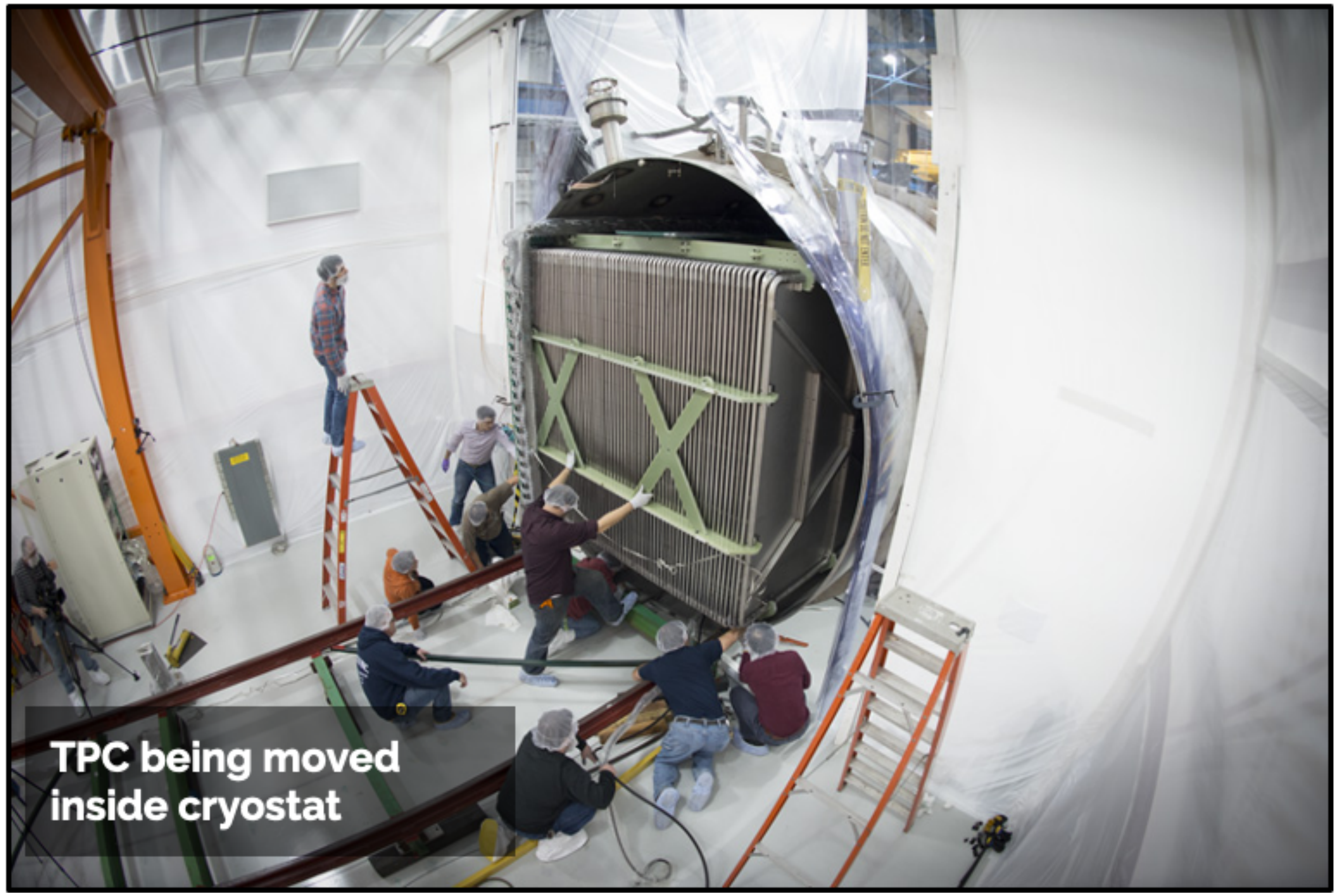
at MicroBooNE's core is an 85 ton LArTPC

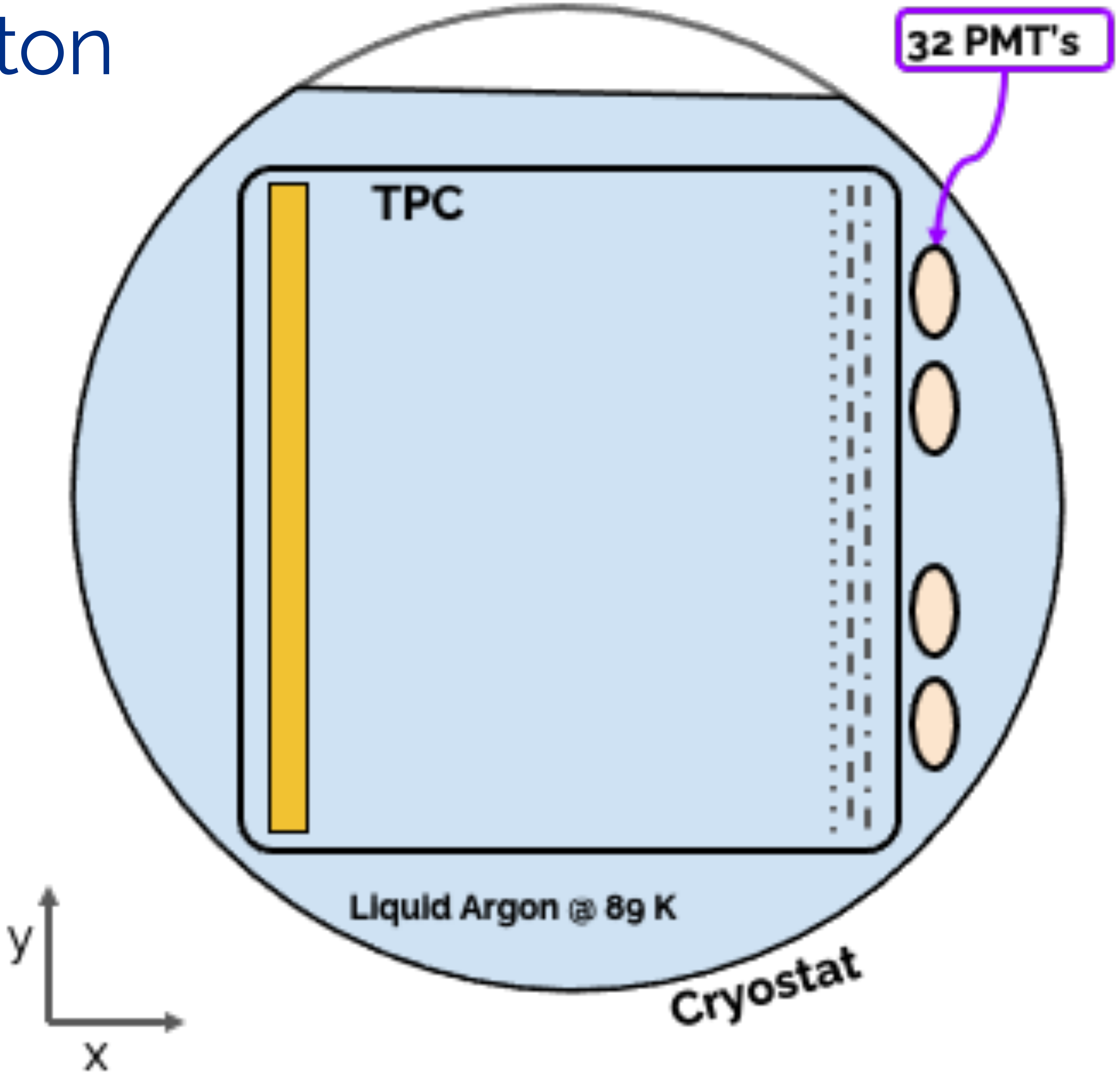




LArTPC principle: the MicroBooNE detector

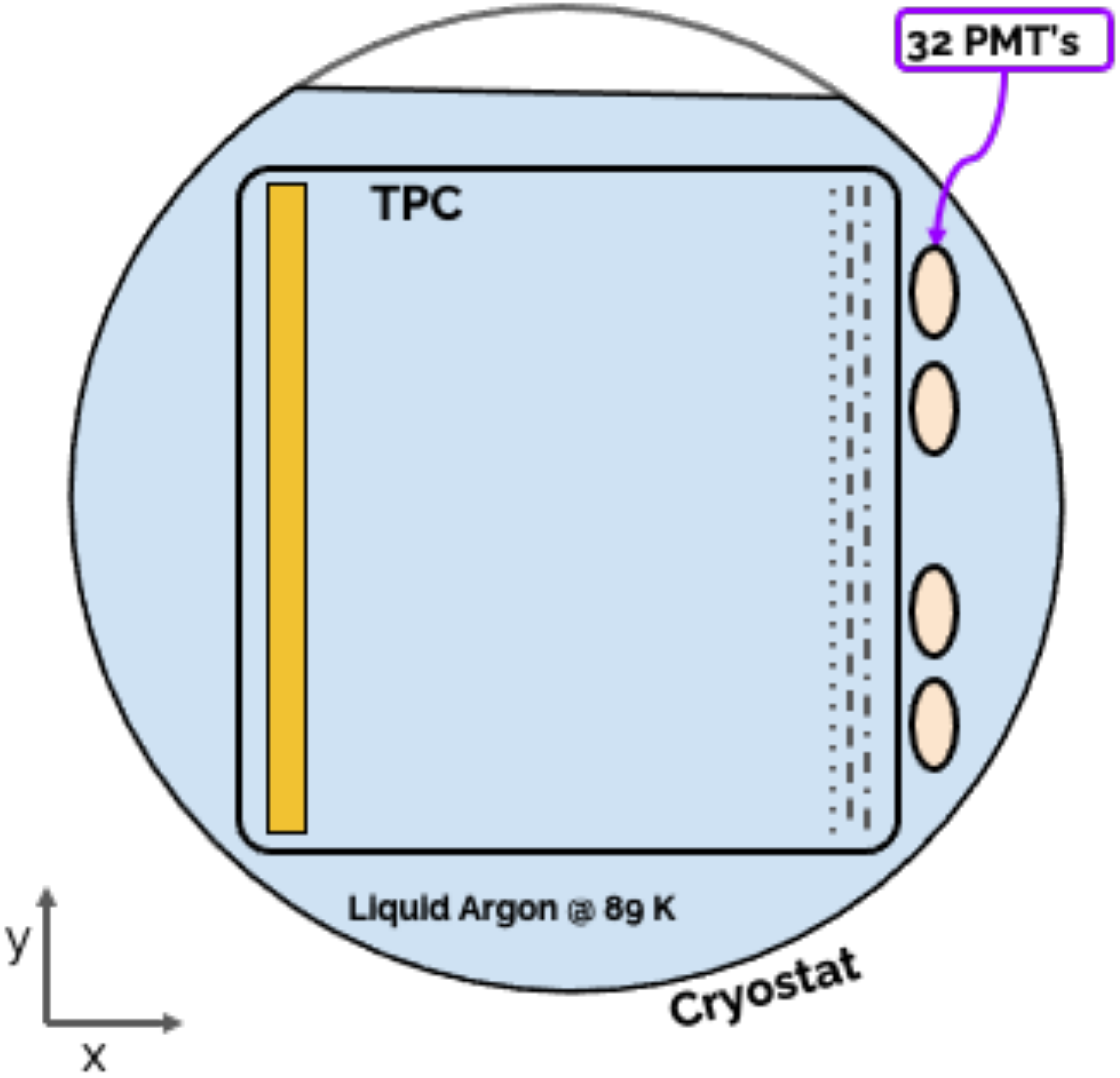
at MicroBooNE's core is an 85 ton LArTPC

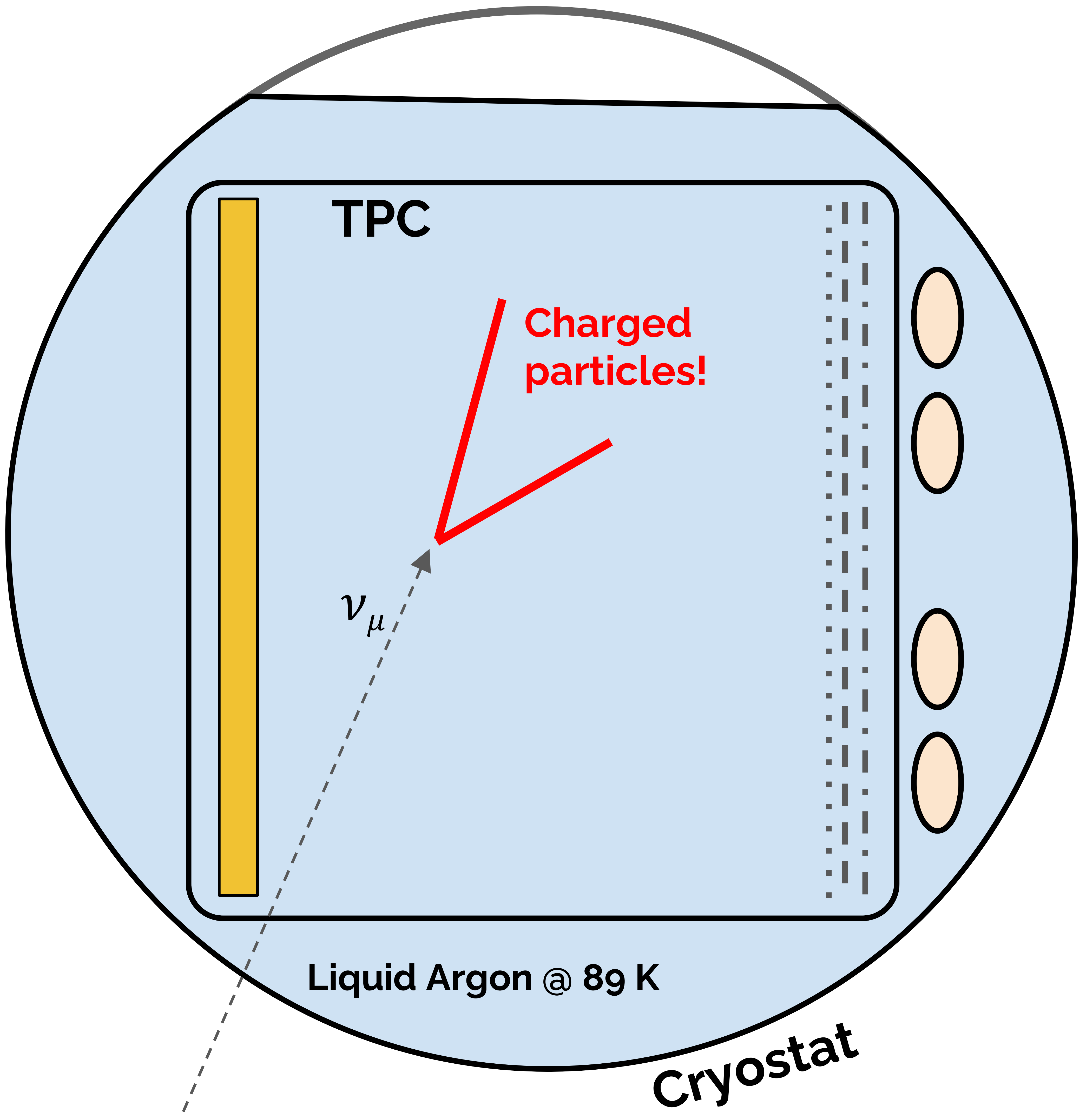




LArTPC principle: the MicroBooNE detector

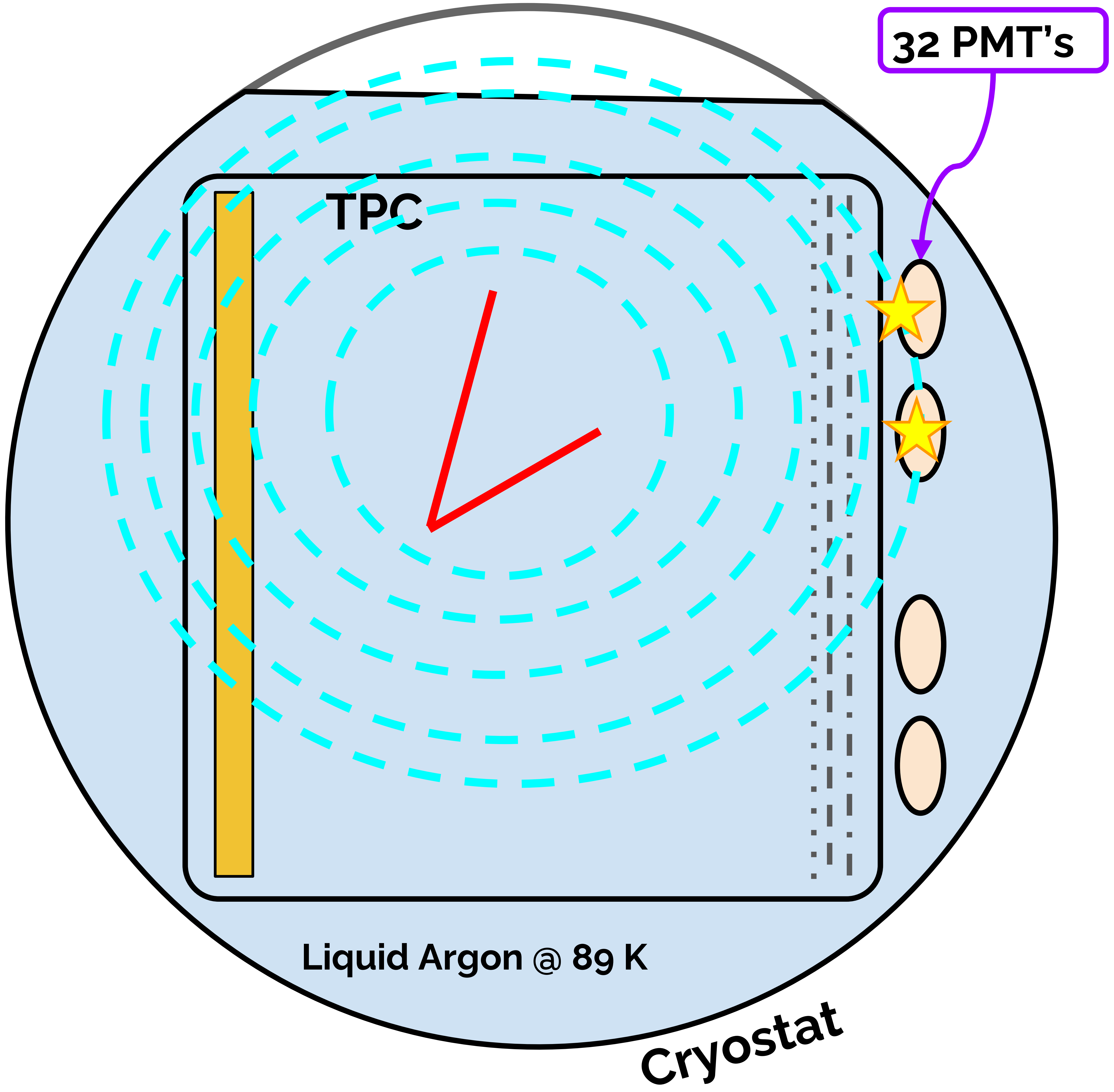
in addition there is a **light detection system** consisting of 32 8-inch PMTs



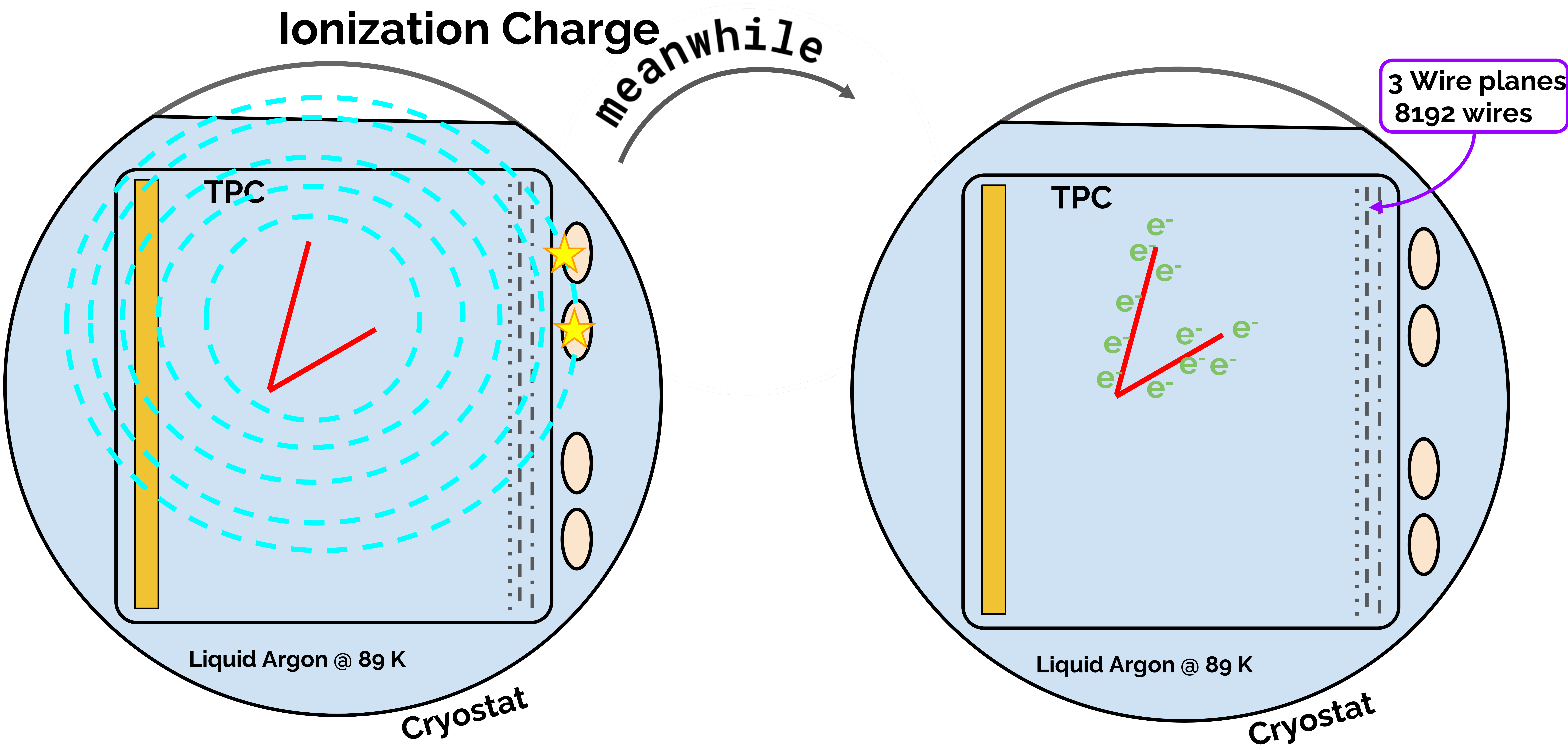


TPC
Charged particles!
ν_μ
Liquid Argon @ 89 K
Cryostat

Scintillation light



Scintillation light
Ionization Charge



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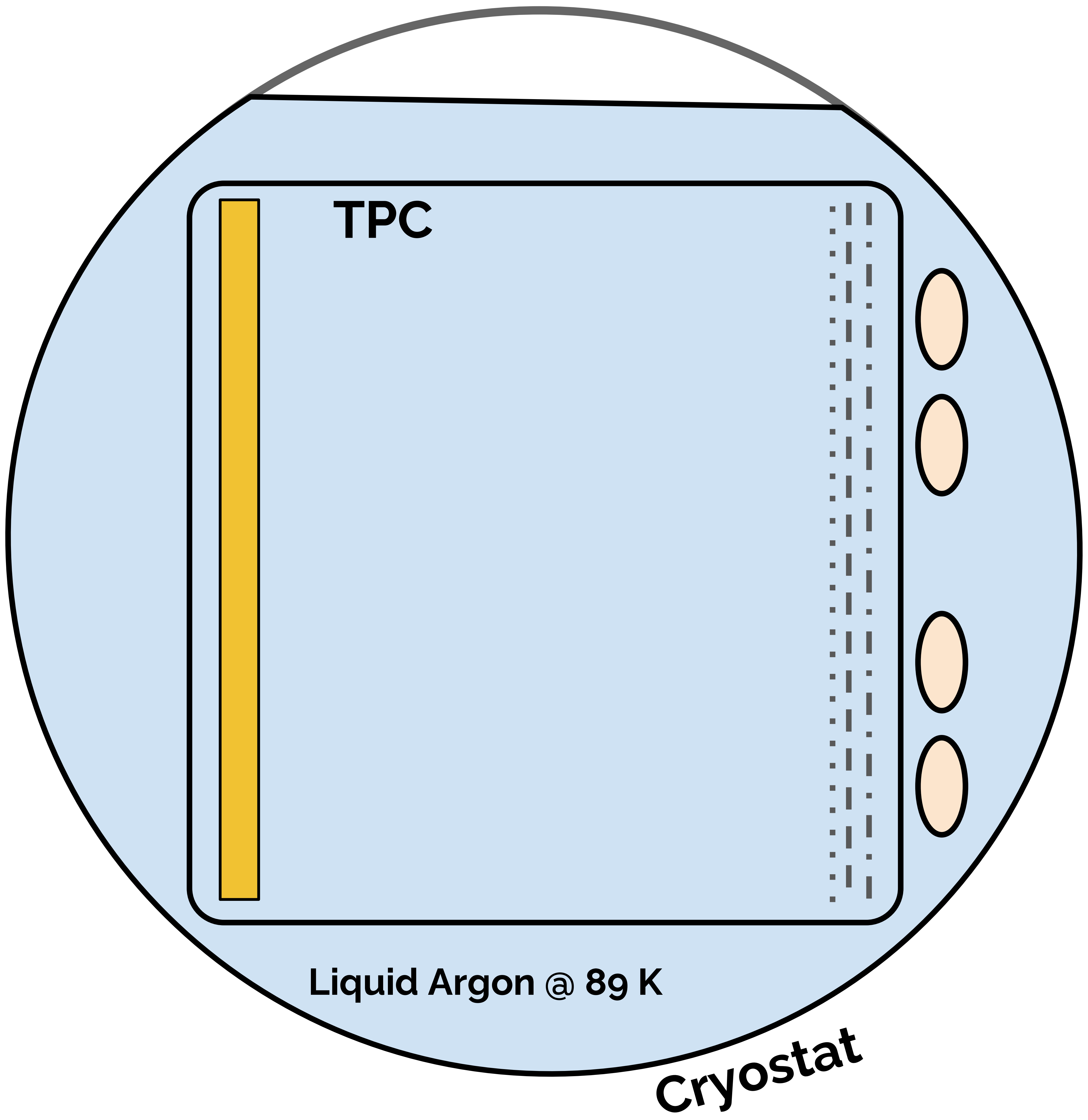

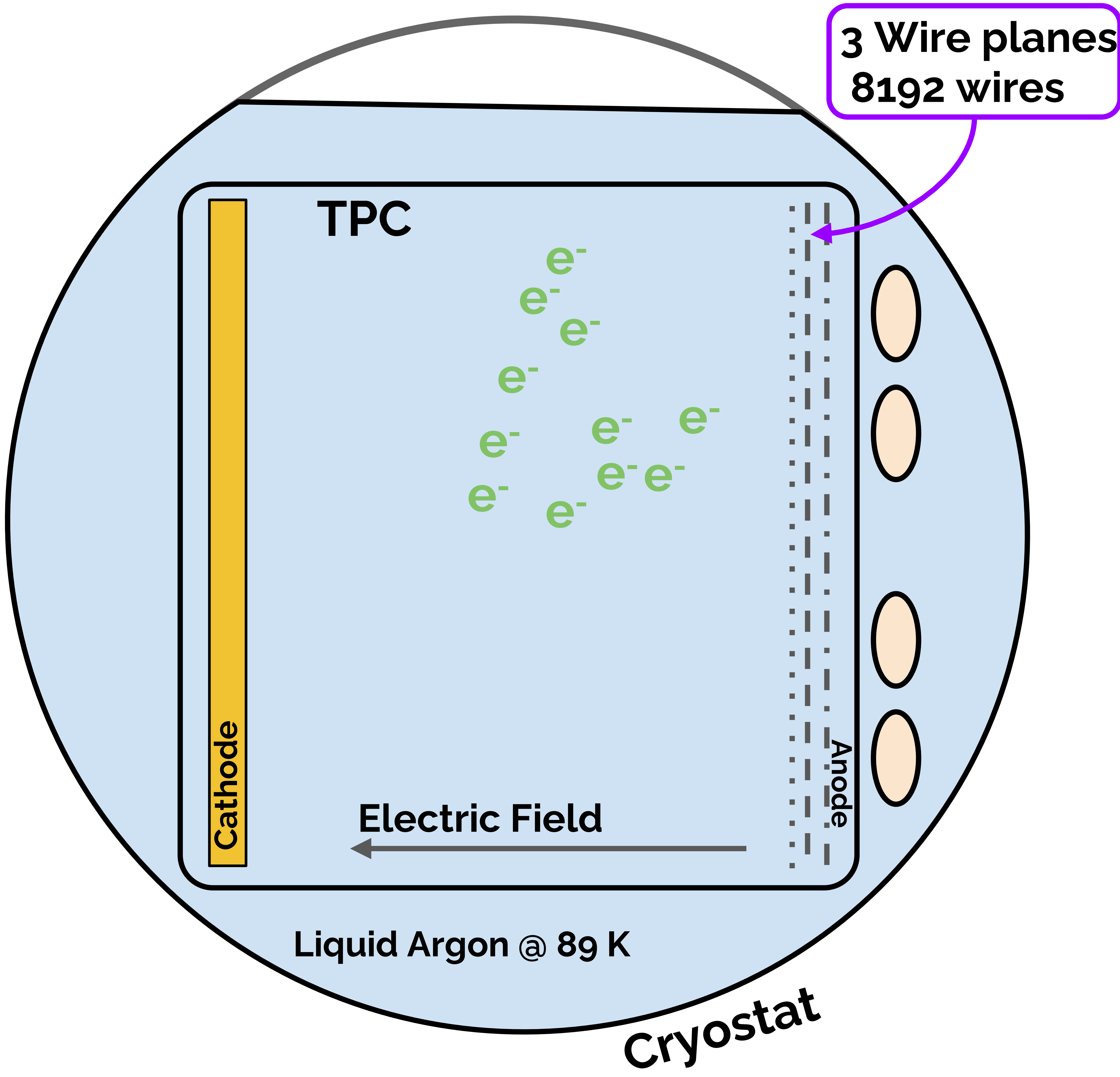

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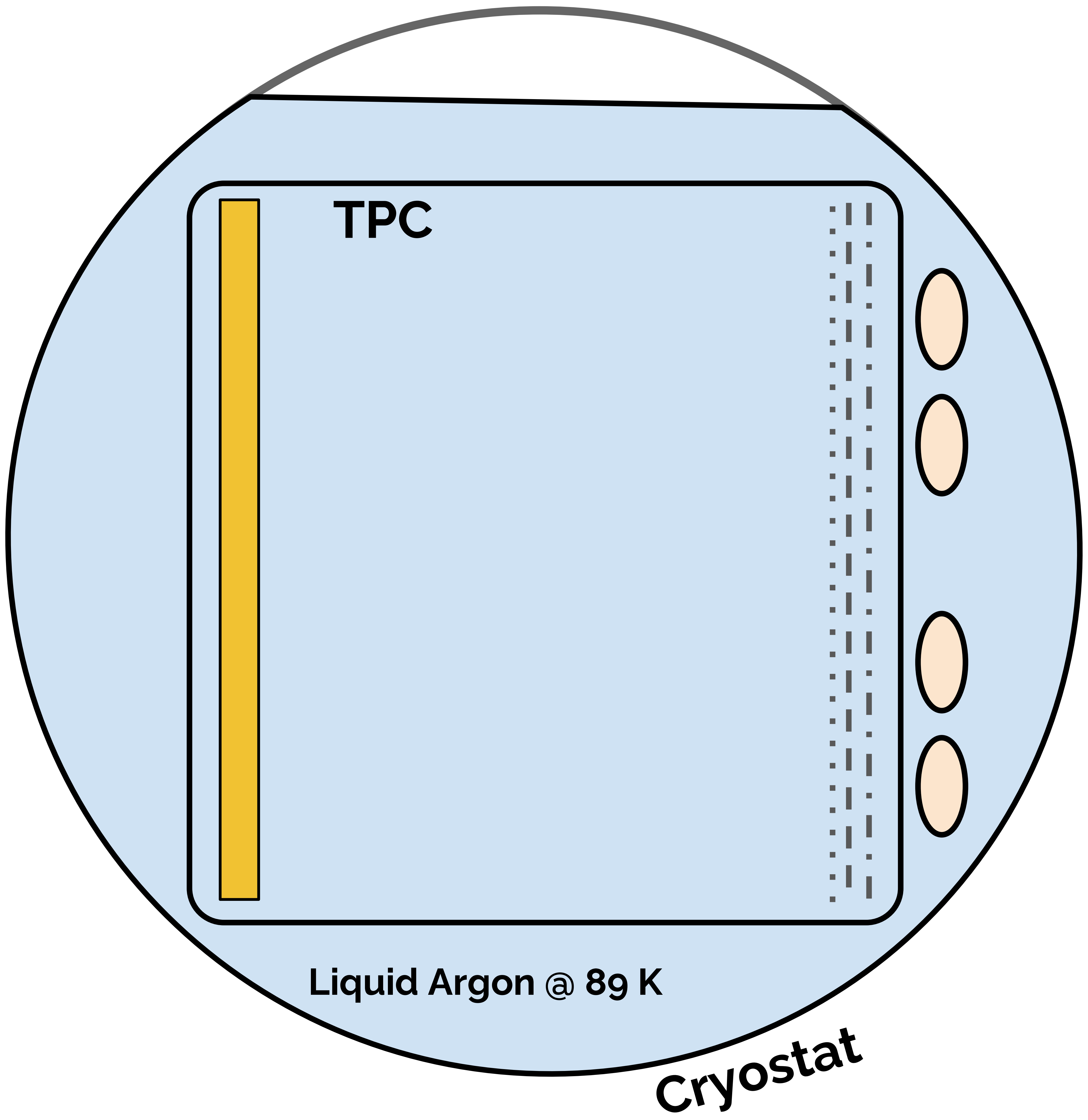

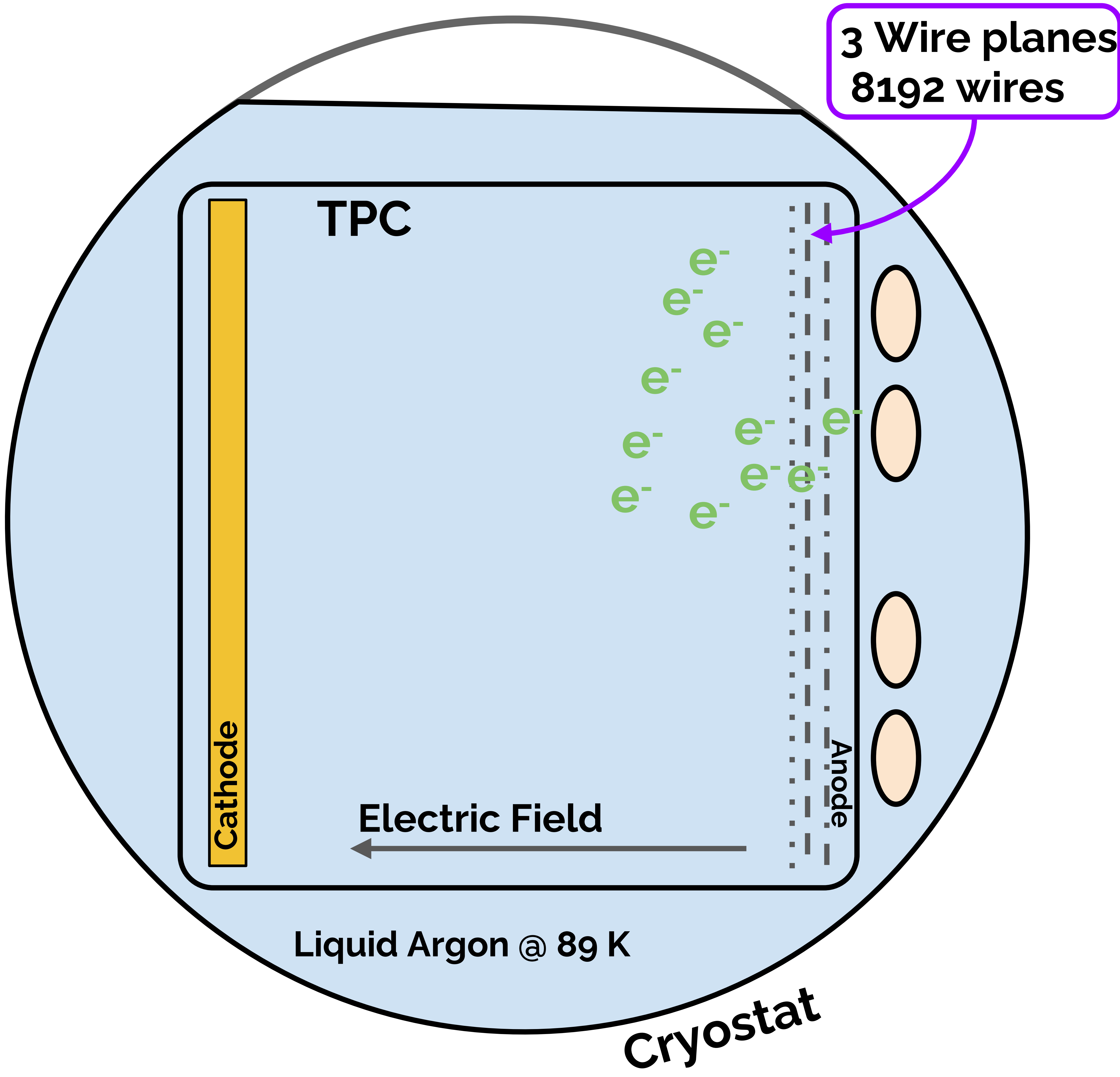

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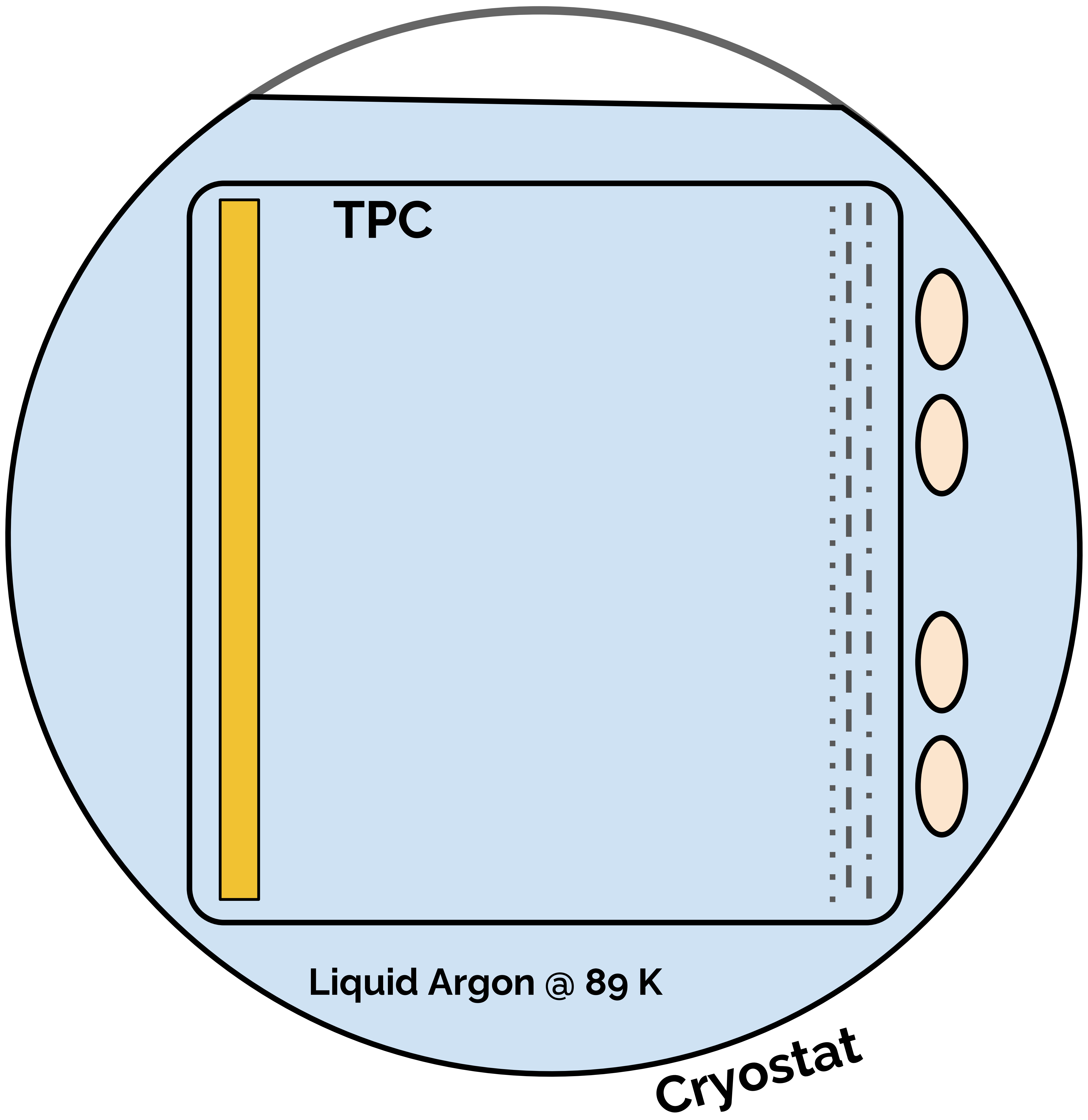

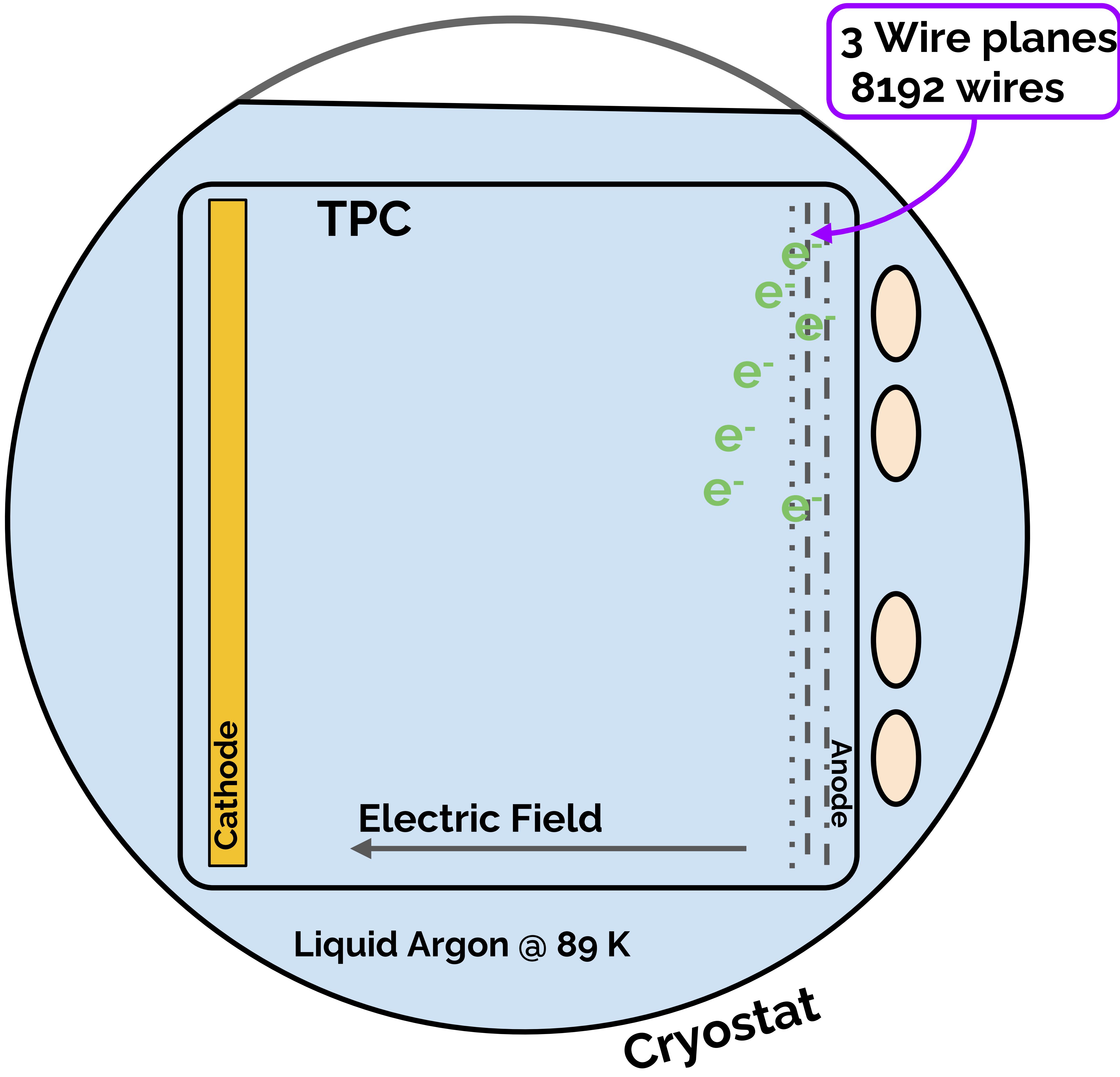

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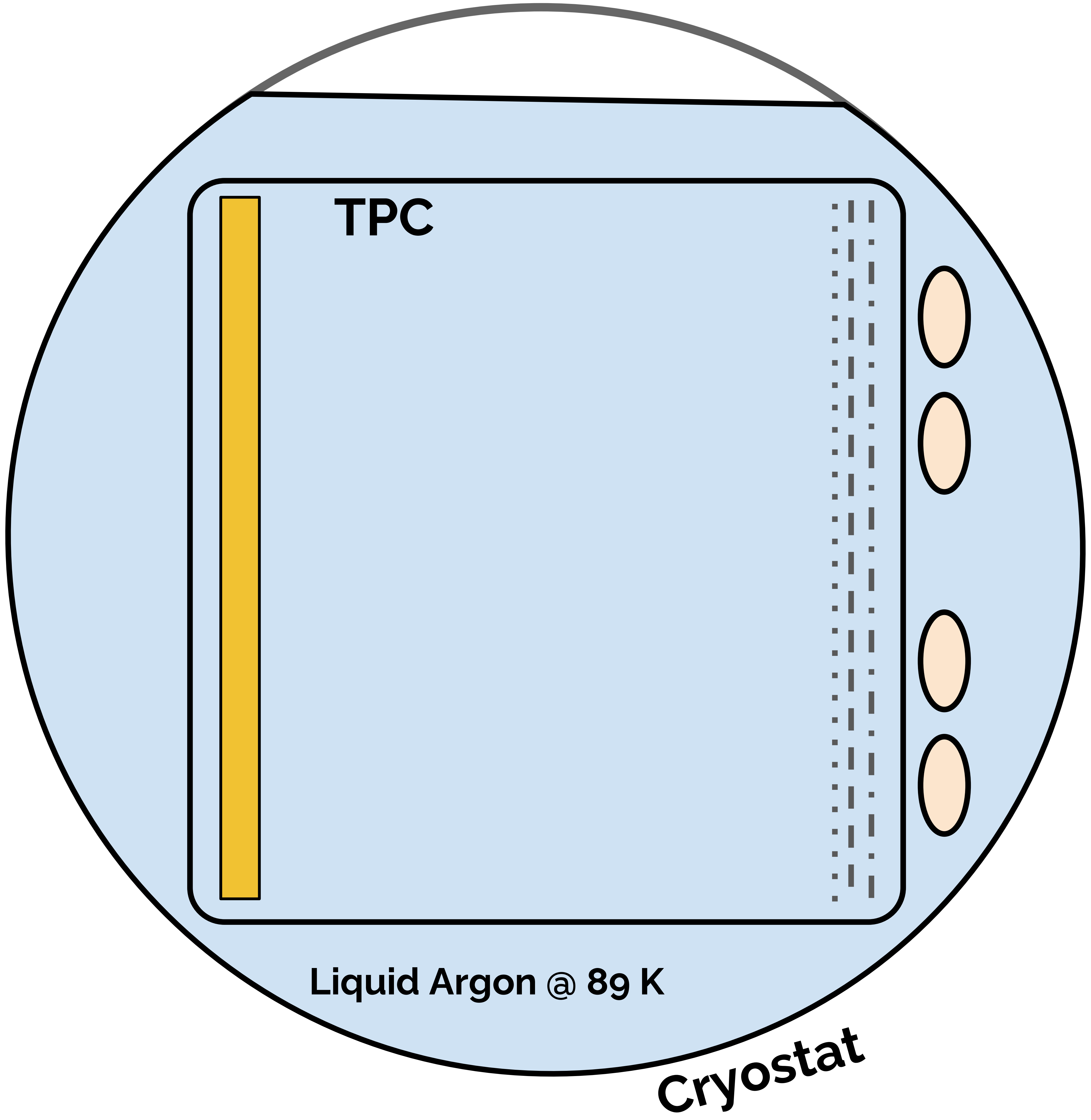

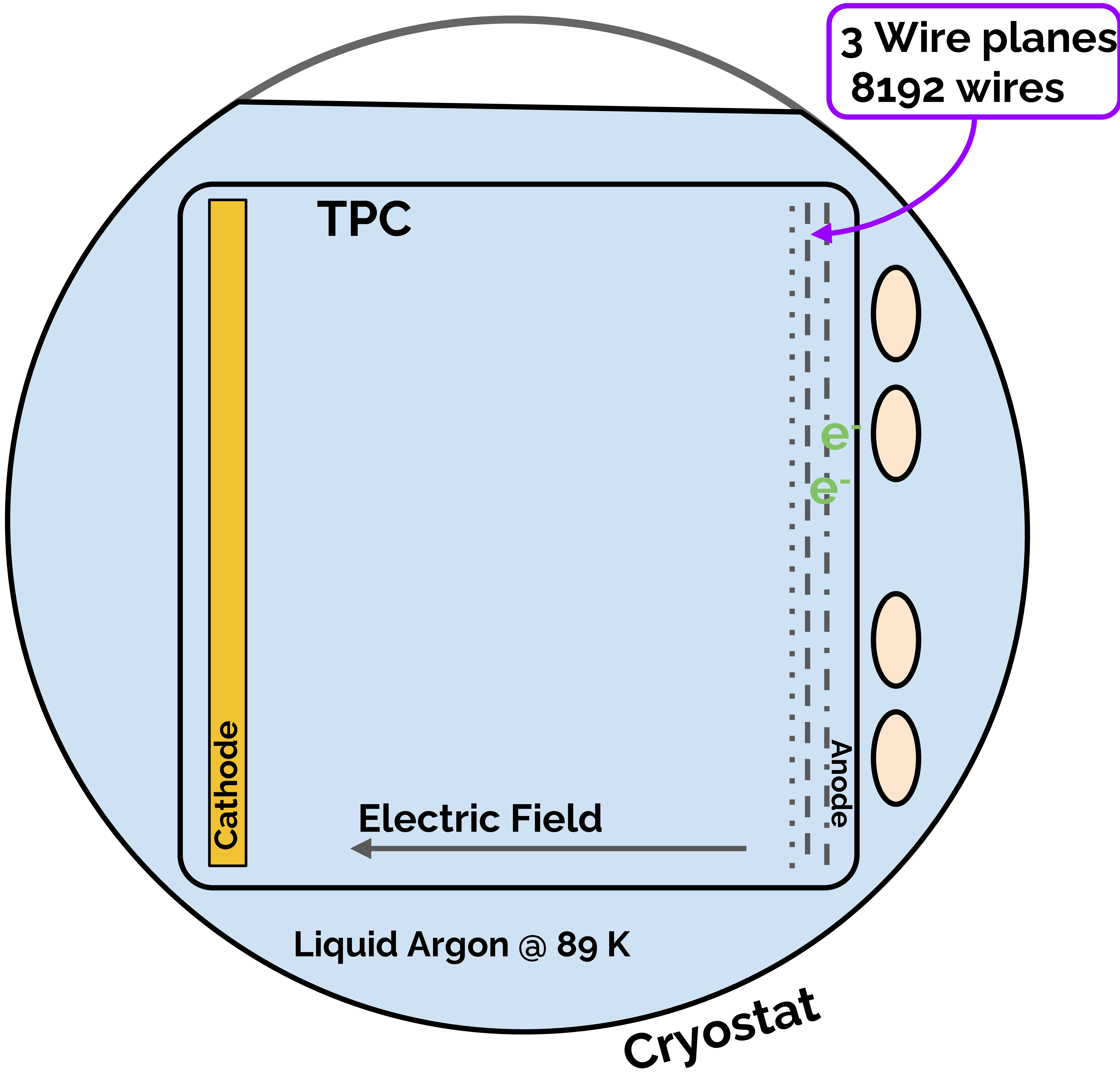

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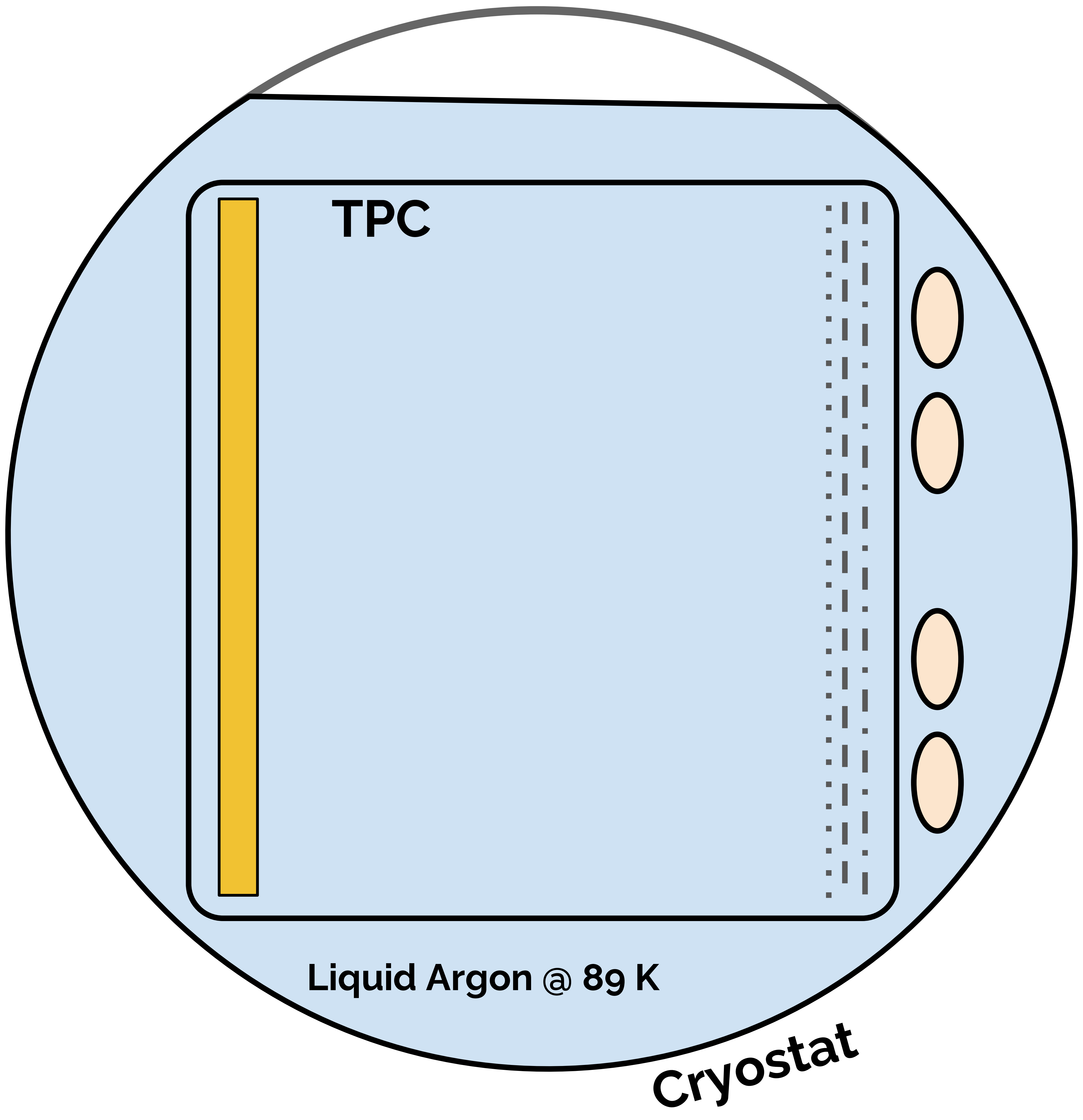

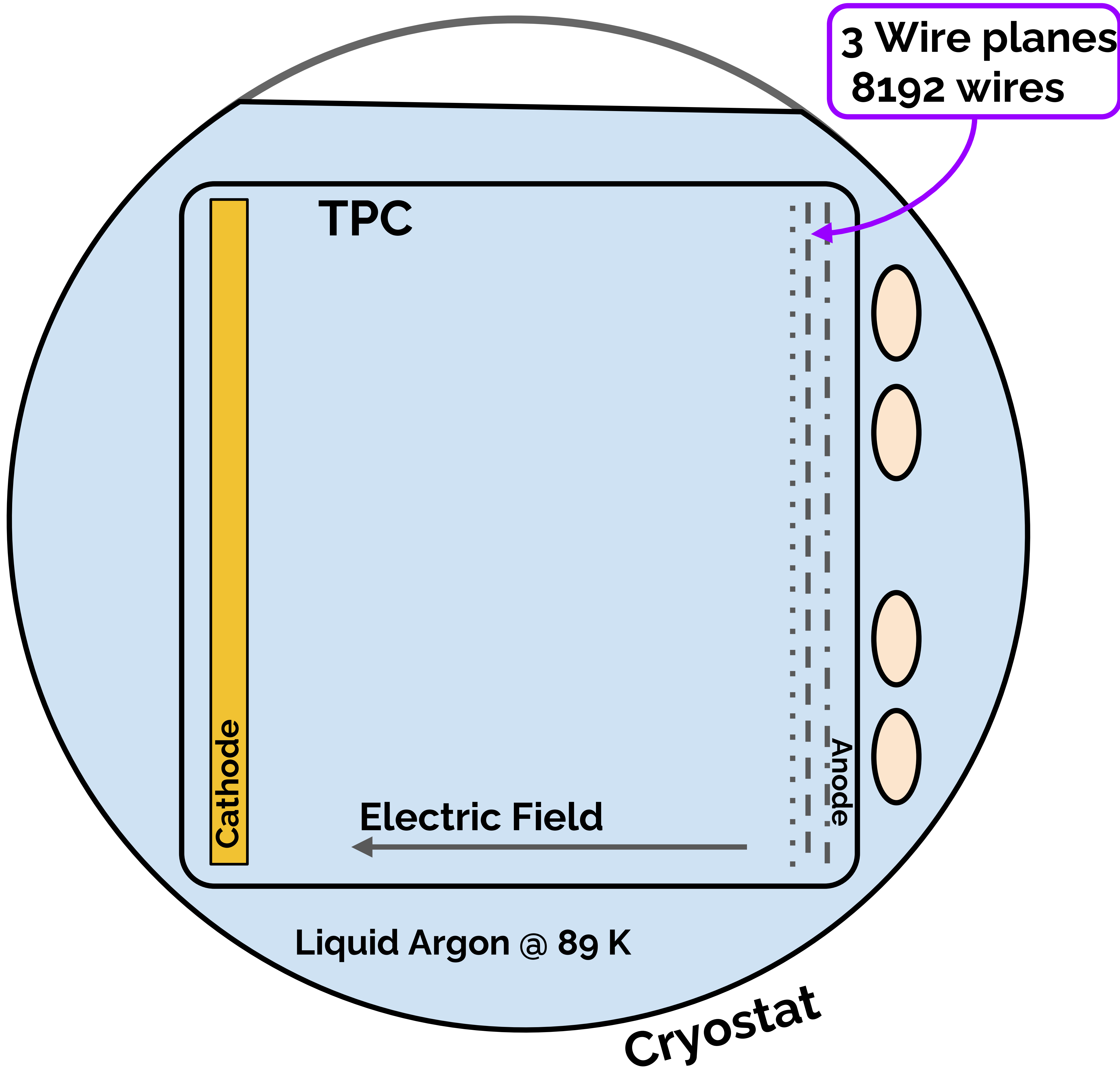

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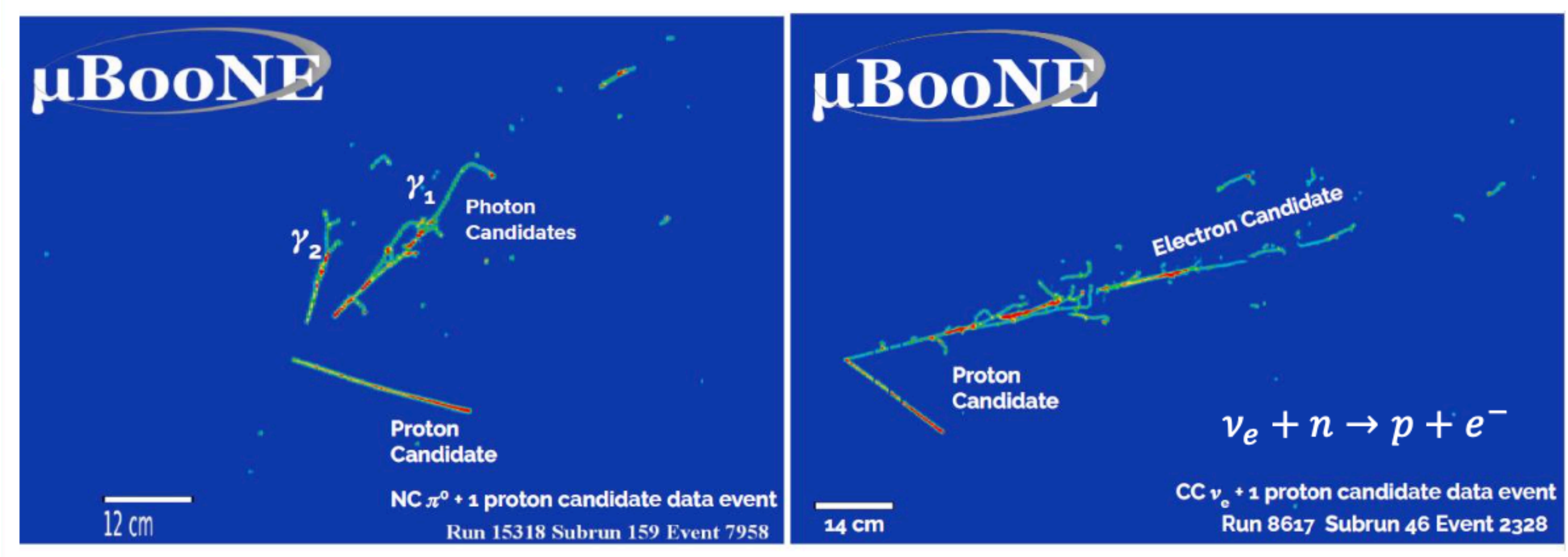
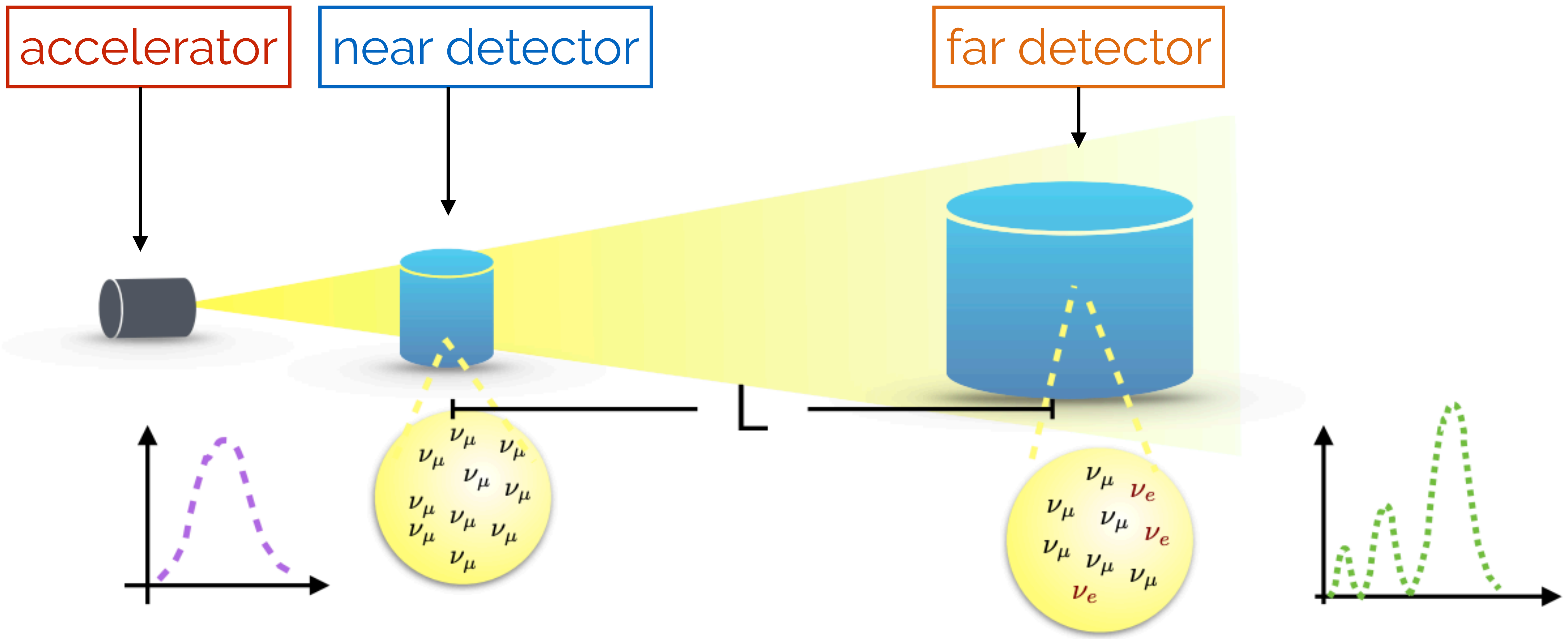


- capable of separating electrons from photons, with gap and calorimetry information

Long-baseline neutrino experiment principles

- artificial neutrino beam generated at an **accelerator**
- measure rate of neutrino events ir the **near detector**
  - use the measurement to predict the neutrino flux at far detector
- measure rate of (un/oscillated) neutrino events in the **far detector**



$$N_\mu(E_\nu) = \sigma(E_\nu)\Phi_\nu(E_\nu)\epsilon(E_\nu)$$

Interaction cross section

Neutrino flux

Detector effects

$$N_\mu(E_\nu) = P(\nu_\mu \to \nu_\mu)\sigma(E_\nu)\Phi_\nu(E_\nu)\epsilon(E_\nu)$$

$$N_e(E_\nu) = P(\nu_\mu \to \nu_e)\sigma(E_\nu)\Phi_\nu(E_\nu)\epsilon(E_\nu)$$

Oscillation probability

$$P(\nu_\mu \to \nu_x) \sim \sin^2 2\theta \, \sin^2\left(\frac{\Delta m^2 L}{4E_\nu^{true}}\right)$$

DUNE: next-generation long-baseline neutrino experiment

Far Site

- *1300km from the proton source*
- *very large LAr TPCs (each 17 ktons)*
- *underground in South Dakota*

Near Site

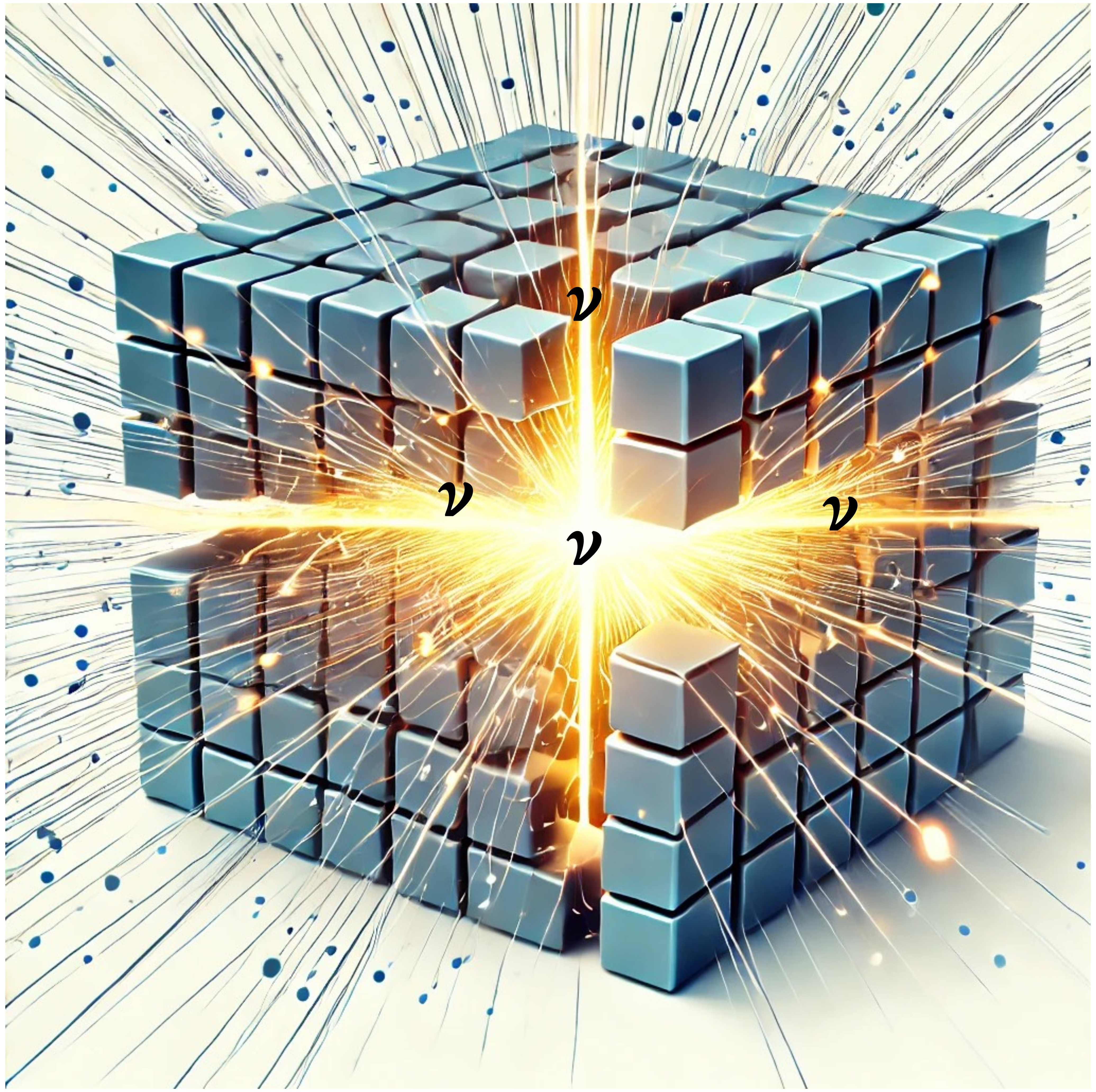
- *550m from proton source*
- *on-site at Fermilab*
- *both stationary & moveable detectors*

Why do we want to detect neutrinos?

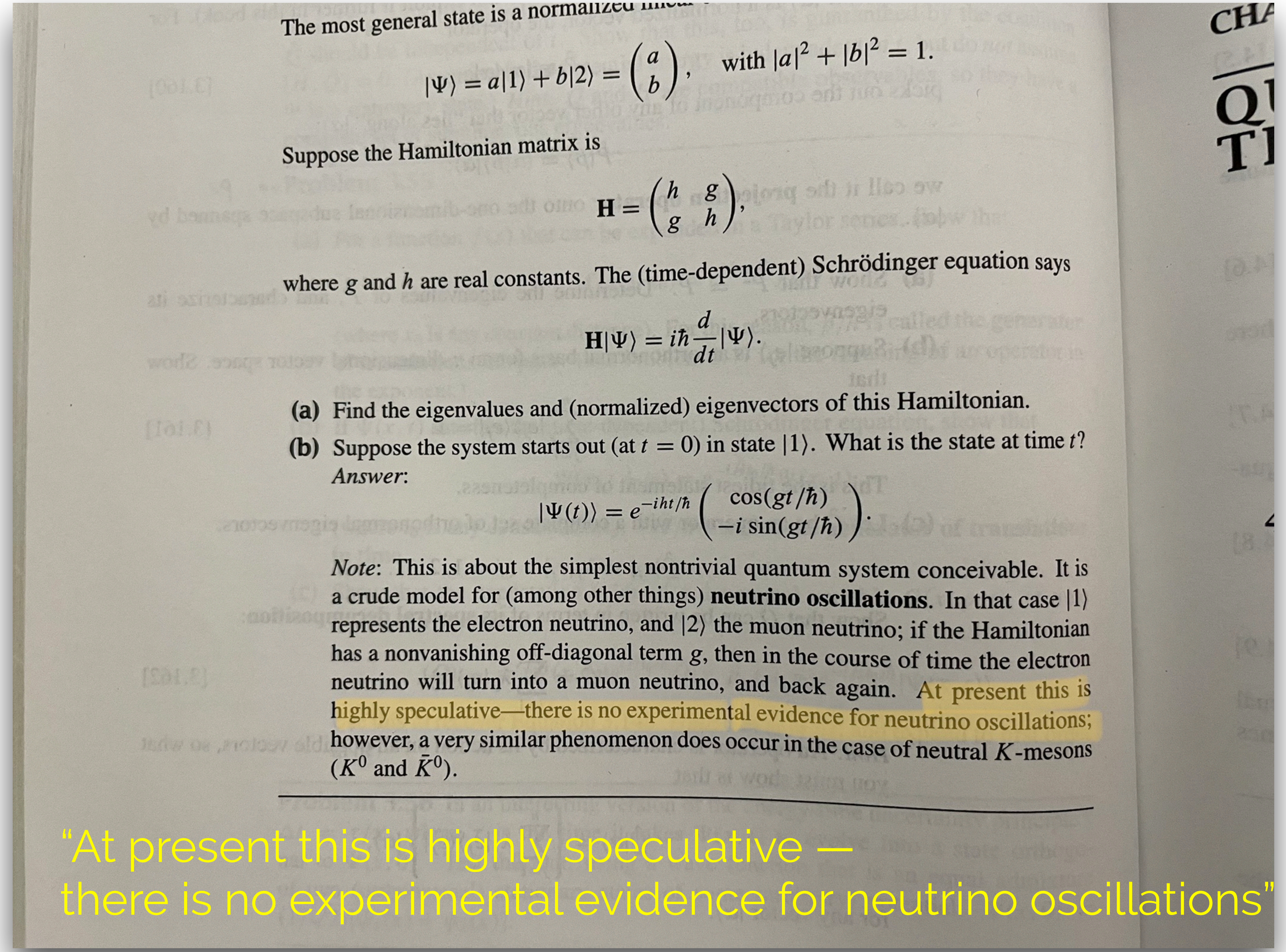
- remember neutrinos need to have ***mass*** to oscillate?

  - the Standard Model, which predicts hundreds of properties of all the particles we can measure precisely, assumes the neutrinos to be massless

  - discovery of neutrino oscillation, starts to make a crack in this incredibly successful model of particle physics



Why do we want to detect neutrinos?


The most general state is a normalized linear ...

$$|\Psi\rangle = a|1\rangle + b|2\rangle = \begin{pmatrix} a \\ b \end{pmatrix}, \quad \text{with } |a|^2 + |b|^2 = 1.$$

Suppose the Hamiltonian matrix is

$$\mathbf{H} = \begin{pmatrix} h & g \\ g & h \end{pmatrix},$$

where g and h are real constants. The (time-dependent) Schrödinger equation says

$$\mathbf{H}|\Psi\rangle = i\hbar \frac{d}{dt}|\Psi\rangle.$$

(a) Find the eigenvalues and (normalized) eigenvectors of this Hamiltonian.

(b) Suppose the system starts out (at $t = 0$) in state $|1\rangle$. What is the state at time t?

Answer:

$$|\Psi(t)\rangle = e^{-iht/\hbar} \begin{pmatrix} \cos(gt/\hbar) \\ -i\sin(gt/\hbar) \end{pmatrix}.$$

Note: This is about the simplest nontrivial quantum system conceivable. It is a crude model for (among other things) **neutrino oscillations**. In that case $|1\rangle$ represents the electron neutrino, and $|2\rangle$ the muon neutrino; if the Hamiltonian has a nonvanishing off-diagonal term g, then in the course of time the electron neutrino will turn into a muon neutrino, and back again. At present this is highly speculative—there is no experimental evidence for neutrino oscillations; however, a very similar phenomenon does occur in the case of neutral K-mesons (K^0 and $\bar{K}^0$).

"At present this is highly speculative —
there is no experimental evidence for neutrino oscillations"

D. J. Griffith, *Introduction to Quantum Mechanics* (p.120, **1995**)

Why do we want to detect neutrinos?

standard model

could **CP violation** in neutrino interactions explain the matter/antimatter asymmetry?

what is the **ordering of the neutrino mass**?

beyond the standard model

are there **new interactions** we could discover via neutrino?

are there **additional neutrinos** beyond known three types?

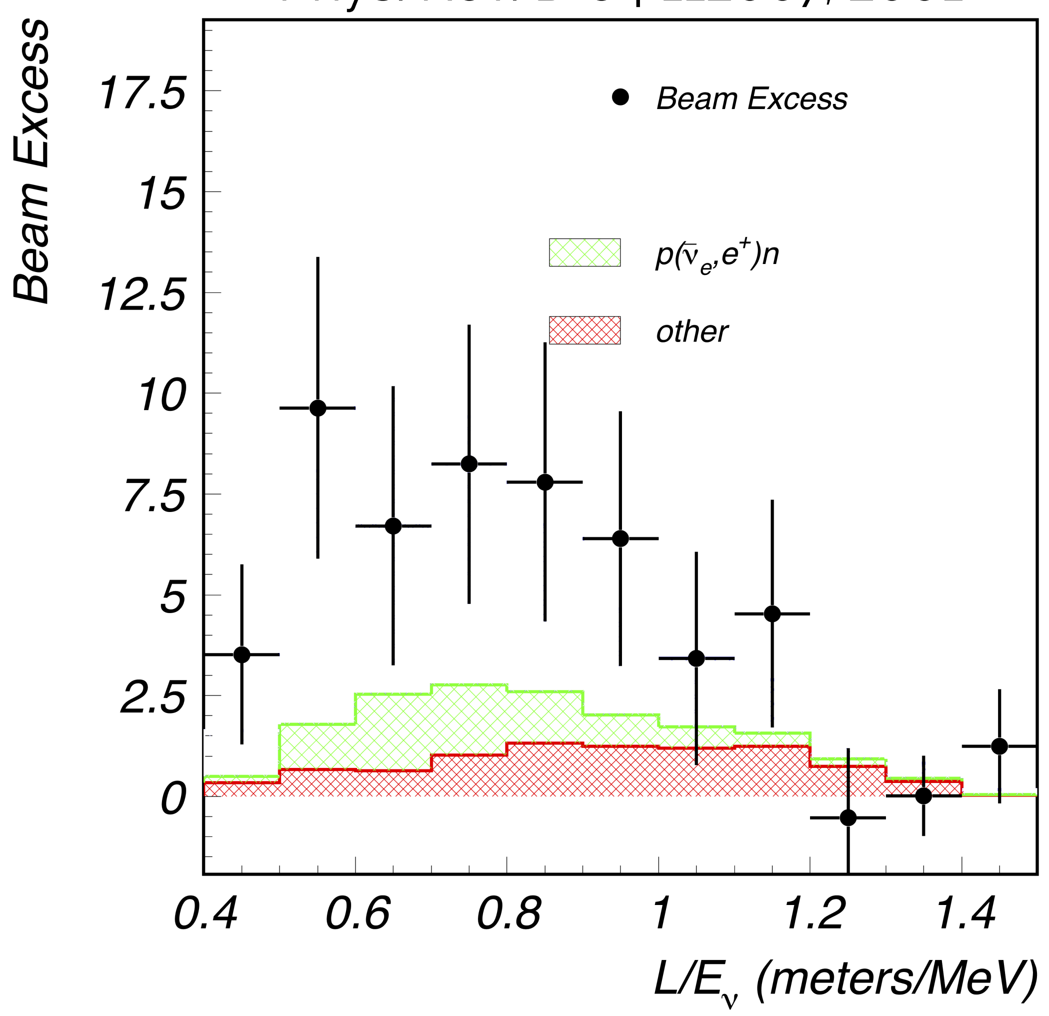
what is neutrino mass?

is the neutrino its own anti particle?

Why do we want to detect neutrinos?

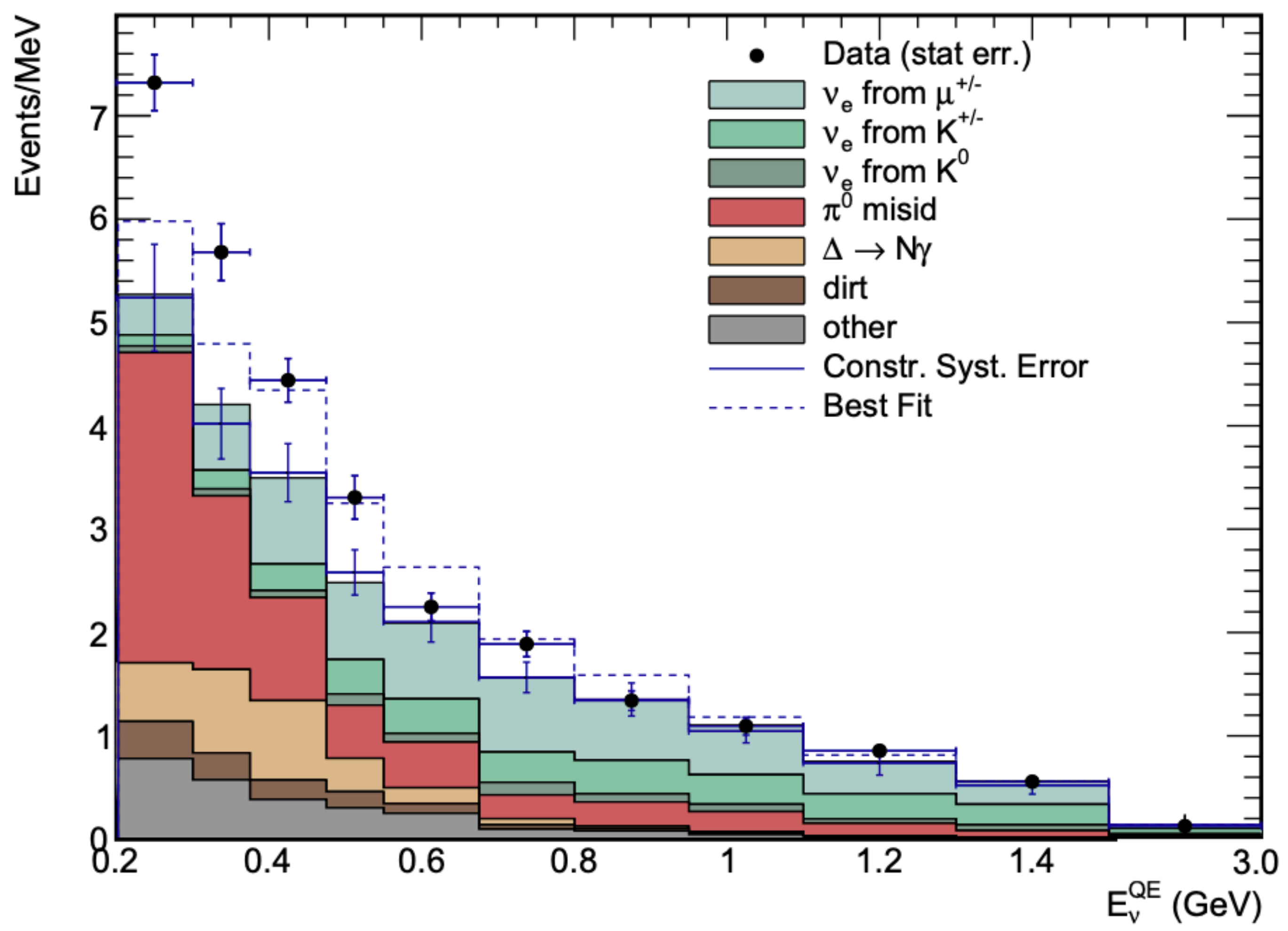
- we are already starting to see few "anomalies" in the neutrino physics
  - seeing results where the measurements and our best prediction of neutrinos start to disagree
  - remember the Solar neutrino problem?
  - will these lead to a discovery of new physics?

Phys. Rev. D 64 112007, 2001

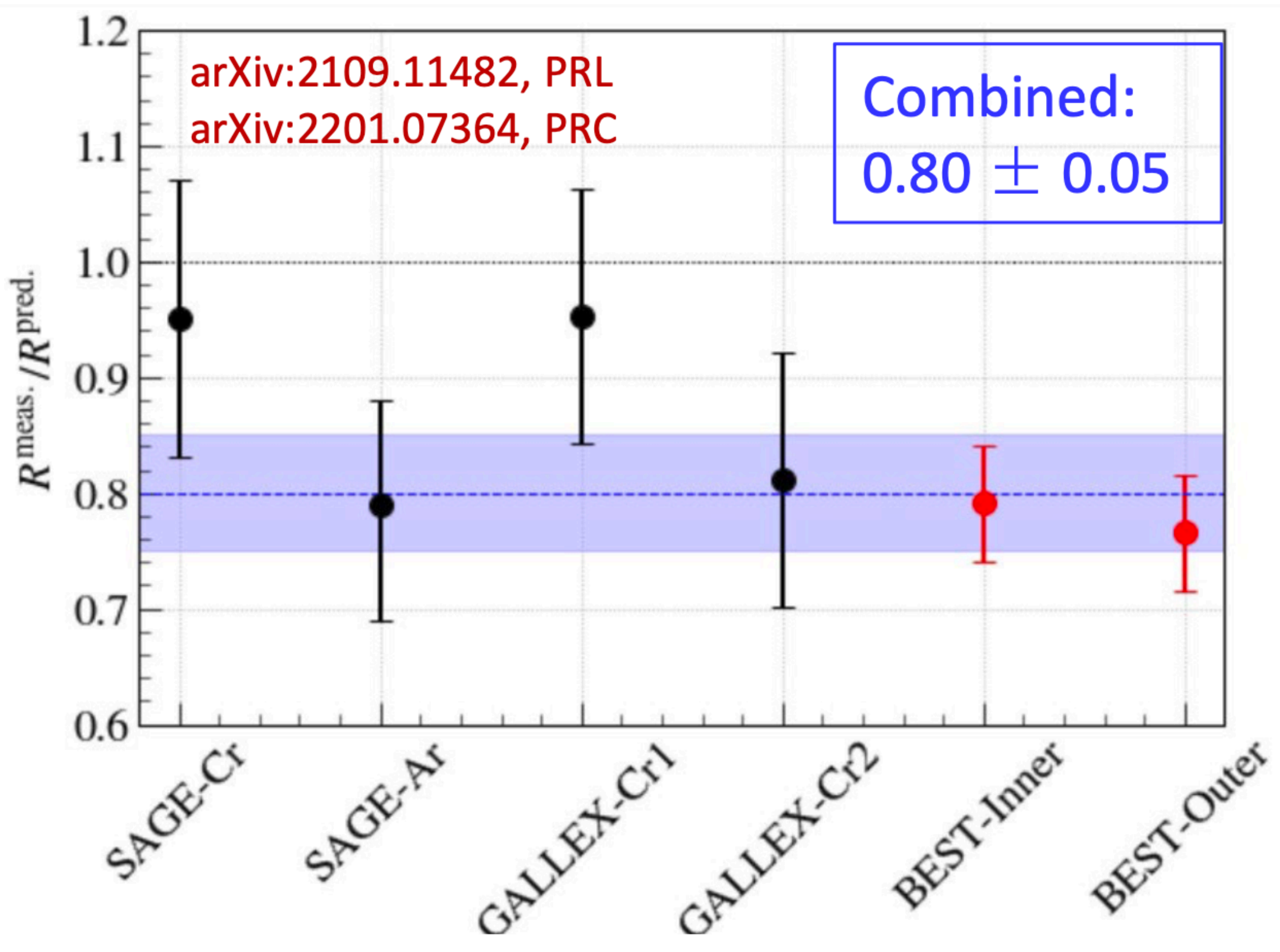


LSND anomaly

Phys. Rev. D 103, 052002 (2021)



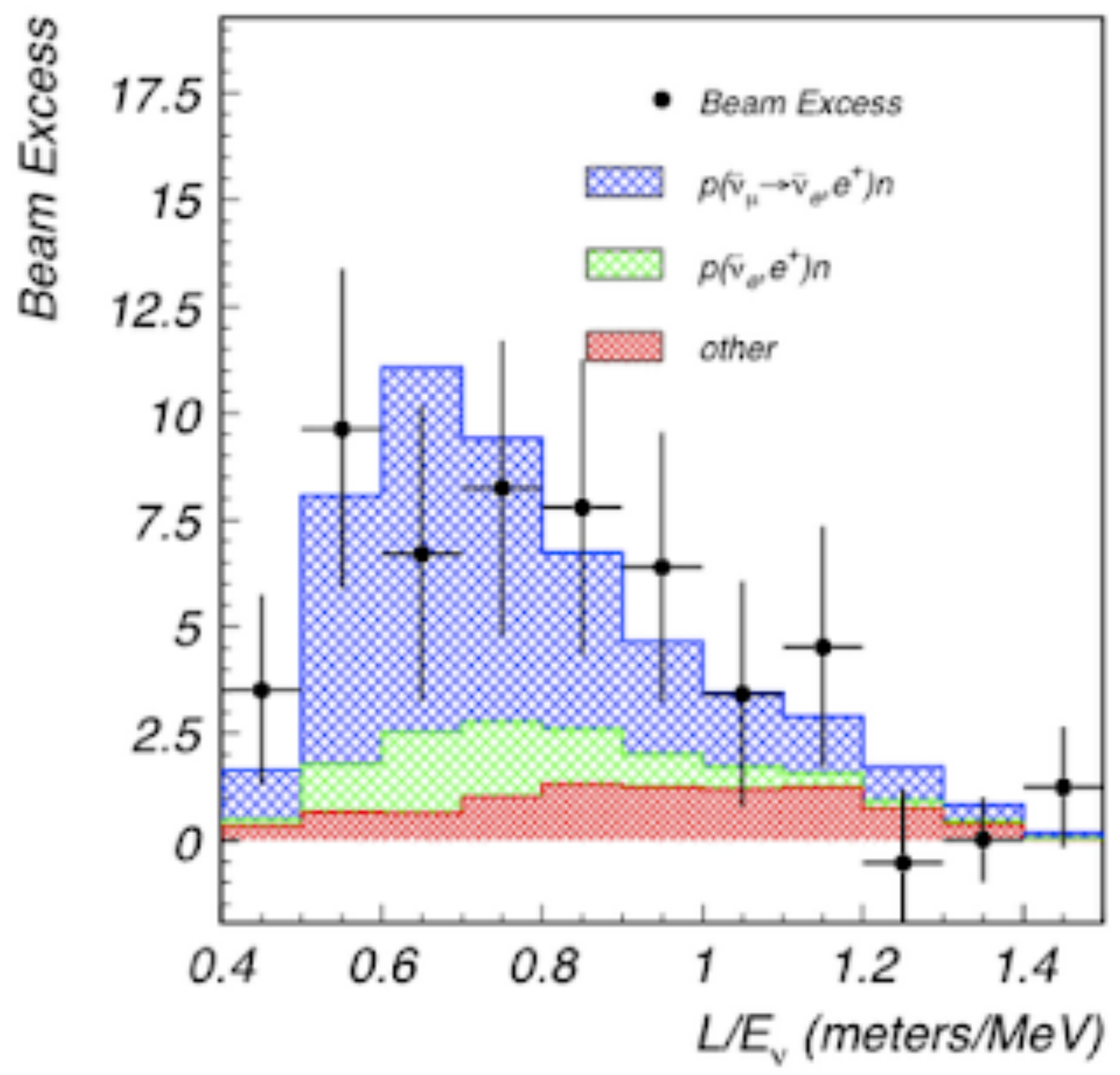
MiniBooNE anomaly

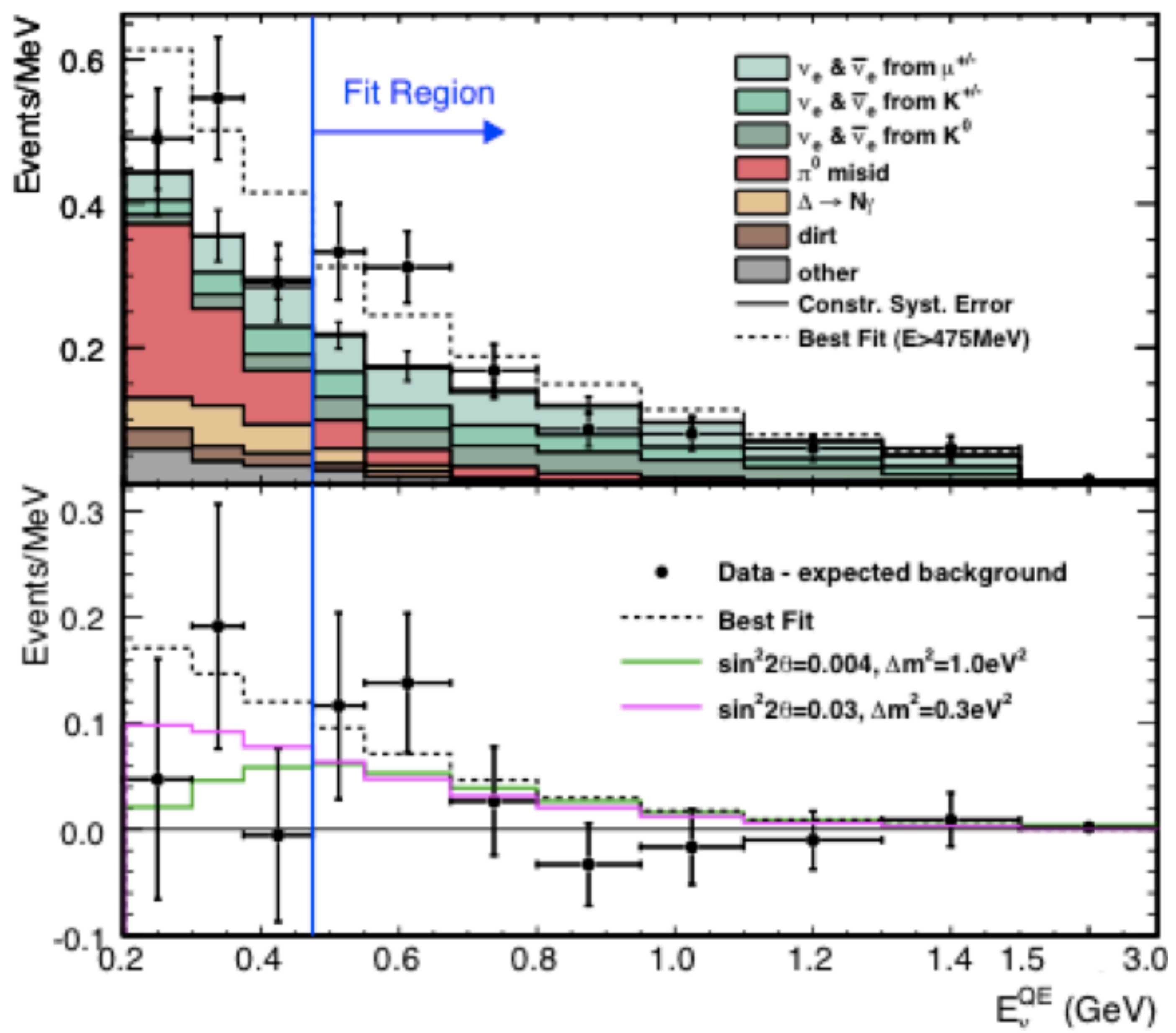


Gallium anomaly

Why do we want to detect neutrinos? Example of “extra” neutrino

- maybe adding an extra, “**sterile**” neutrino help resolving these anomalies
- potentially detectable through impact on neutrino oscillations
- *Q: can this new type of neutrino be solution to these anomalies?*
- *A: unfortunately, it’s not so simple… there are severe tension between different measurements & channels*





Why do we want to detect neutrinos? Example of "extra" neutrino

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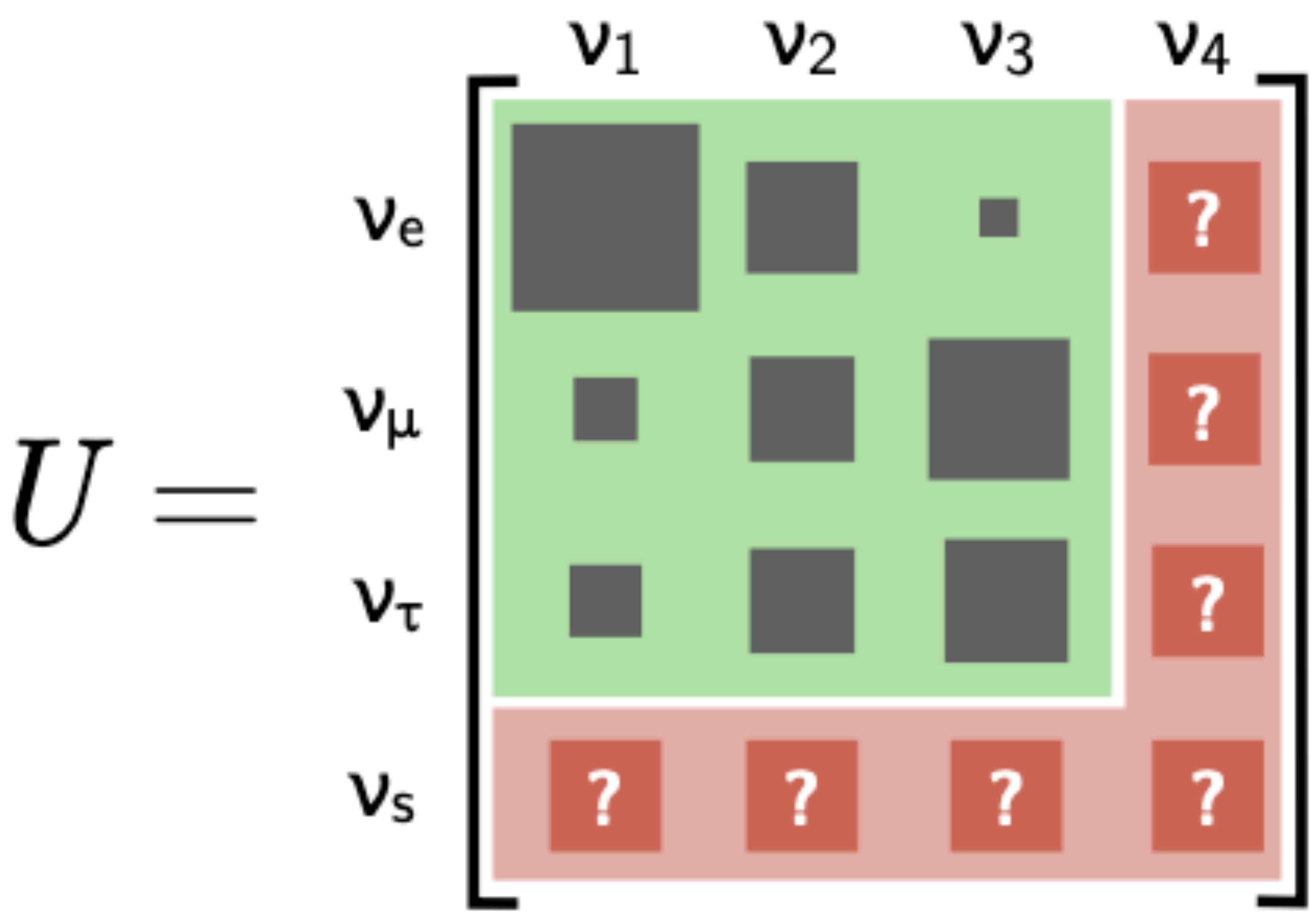


Flavor transitions via this new mixing:

$$P_{\alpha\beta} = 4|U_{\alpha 4}|^2 |U_{\beta 4}|^2 \sin^2\left(1.27\frac{\Delta m_{41}^2 L}{E}\right)$$

Summary

- neutrino physics is relatively young, but started to have a big impact in our understanding of the Universe
- at the heart of neutrino physics, there is massive neutrino & neutrino oscillation
- detecting neutrino is challenging, yet we can do it… and pretty well!
- precise measurements of neutrino's behavior will open a new era of particle physics

Backup

Properties of neutrinos

- extremely weakly interacting with matter
- can pass through most materials without being detected
- travel close to the speed of light
- important in various astrophysical processes and fundamental to our understanding of universe

The solar neutrino problem: neutrino oscillation

- **flavor mixing**: mismatch between weak/flavor eigenstates and mass eigenstates of fermions, due to coexistence of 2 types of interactions
  - **weak eigenstates**: members of weak isospin doublets transforming into each other through the interaction with the W boson
  - **mass eigenstates**: states of definite masses that are created by the interaction with the Higgs boson (Yukawa coupling)

flavor (α = e, μ, τ) ⇔ linear combinations ⇔ mass (i = 1, 2, 3)

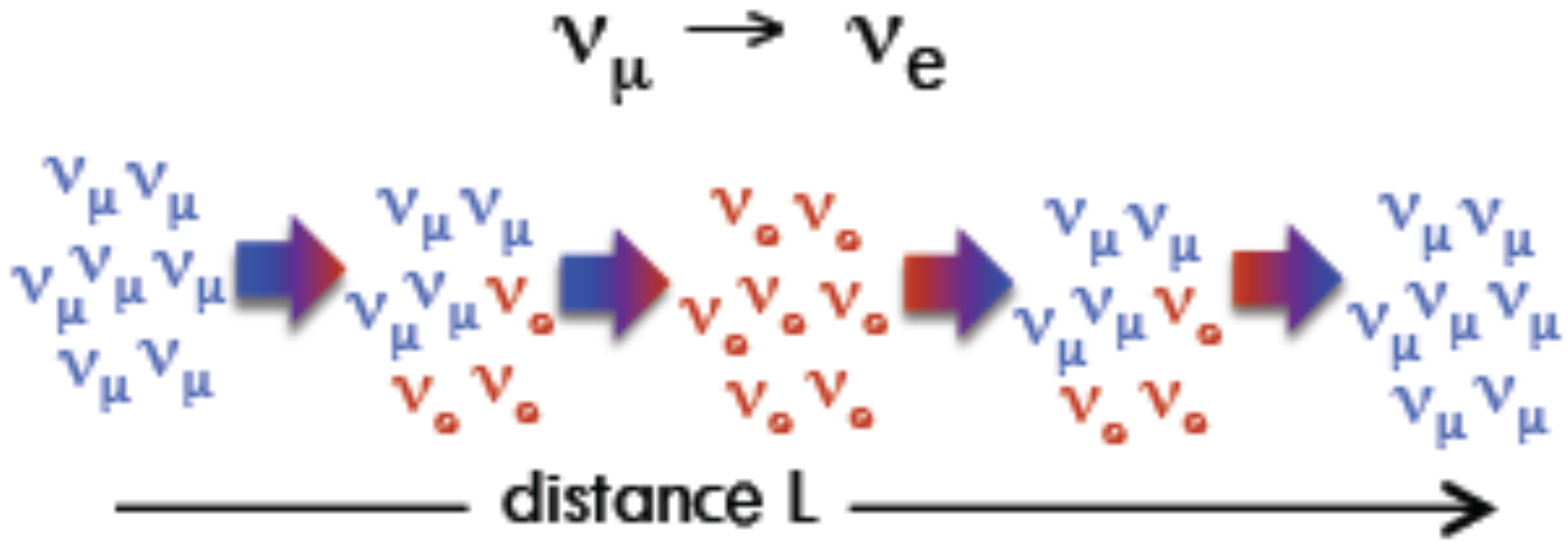
$$|\nu_\alpha\rangle = \sum_i U^*_{\alpha i} |\nu_i\rangle \quad \Longleftrightarrow \quad |\nu_i\rangle = \sum_\alpha U_{\alpha i} |\nu_\alpha\rangle$$

$$\begin{bmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{bmatrix} = \begin{bmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{bmatrix} \begin{bmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{bmatrix}.$$

Pontecorvo-Maki-Nakagawa-Sakata (PMNS) matrix

Crack in the Standard Model: Massive neutrinos

- the Nobel Prize in Physics 2015 was awarded jointly to Takaaki Kajita and Arthur B. McDonald "for the discovery of neutrino oscillations, which showed that **neutrinos have mass**"



- neutrino morph into another kind & back again: quantum mechanical effect
  - **if neutrinos oscillate, they must have mass**
  - depend on neutrino flavor and neutrino energy

Why add extra neutrino?

- since the detection of neutrino and oscillation, many experiments start to collect & analyze neutrino data

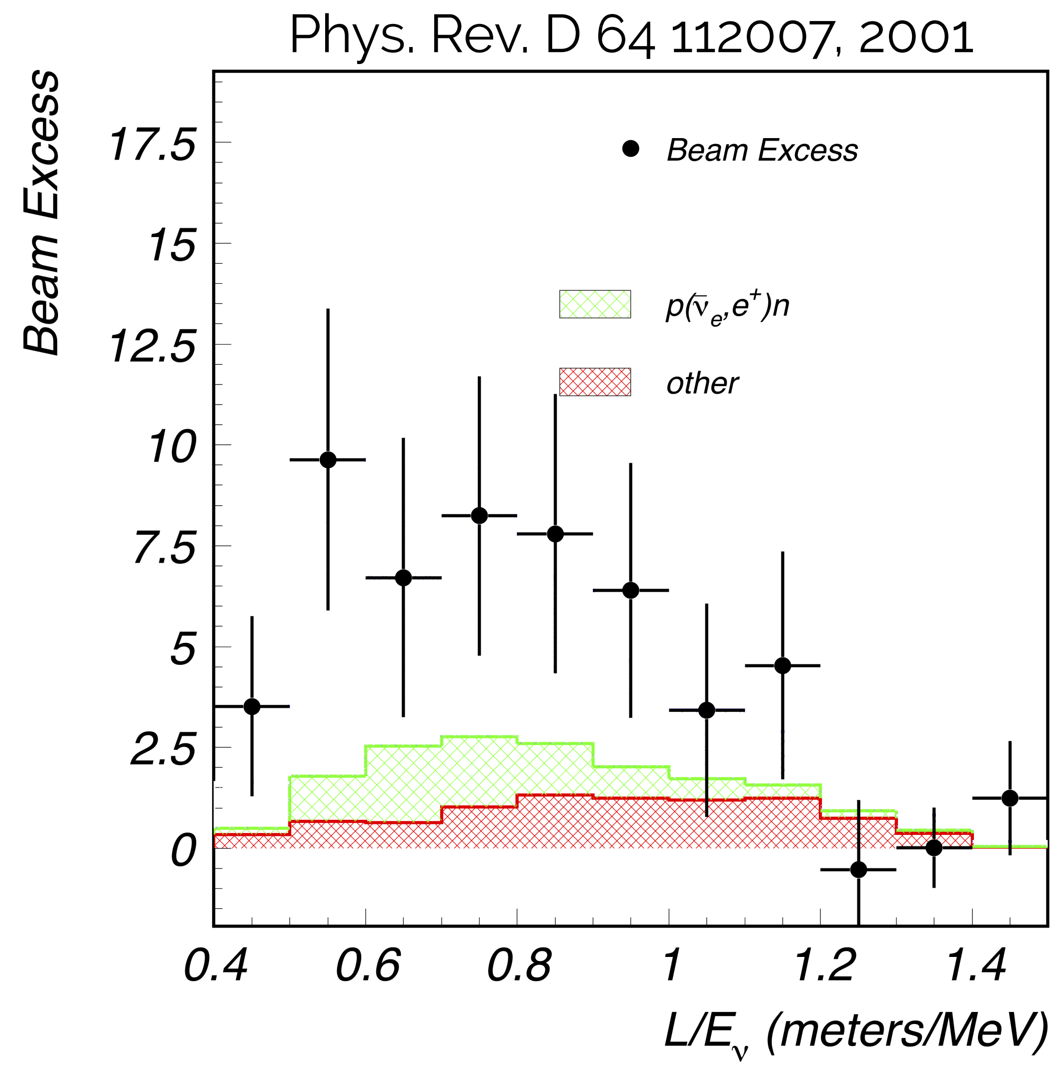
- several experiments have found series of anomalous results

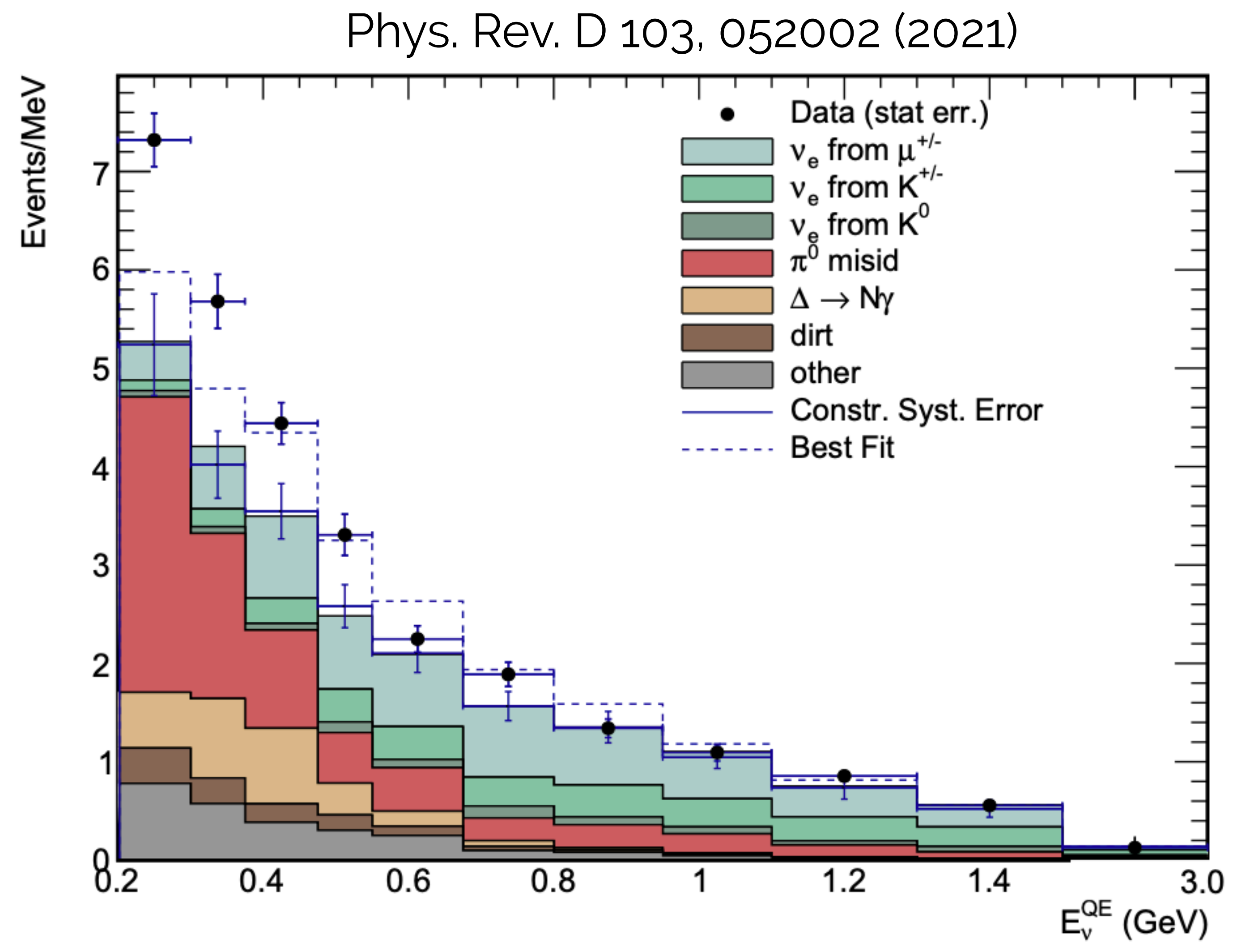
  - anomalous in a way that "observation" (detected/measured data) does not agree with "prediction" (simulation/model generated with the current best of our knowledge)

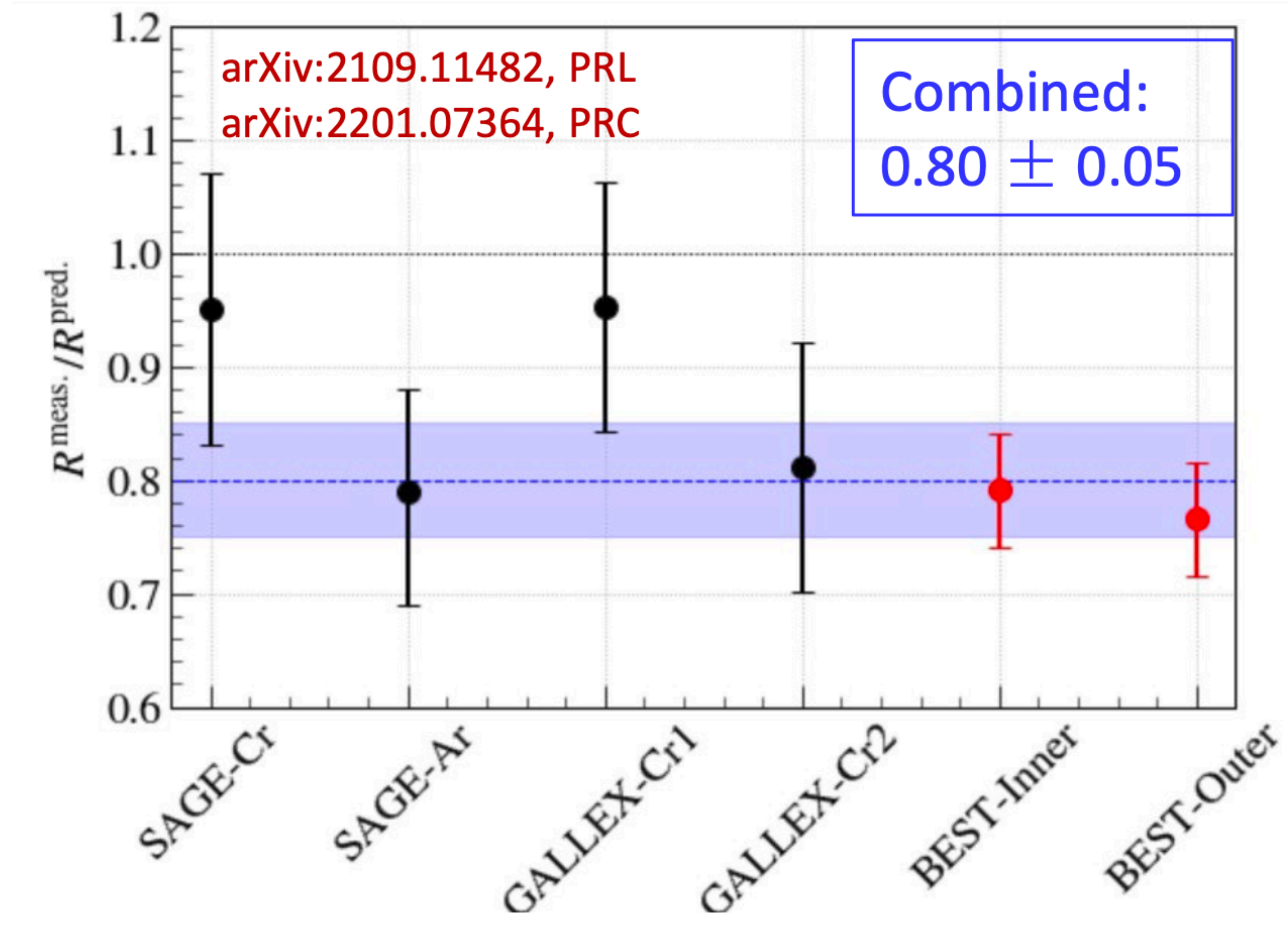
  - LSND: measured more ν_e than predicted

  - MiniBooNE: measured more ν_e than predicted

  - GALLEX/SAGE/BEST: measured less ν_e than predicted







Why add extra neutrino?

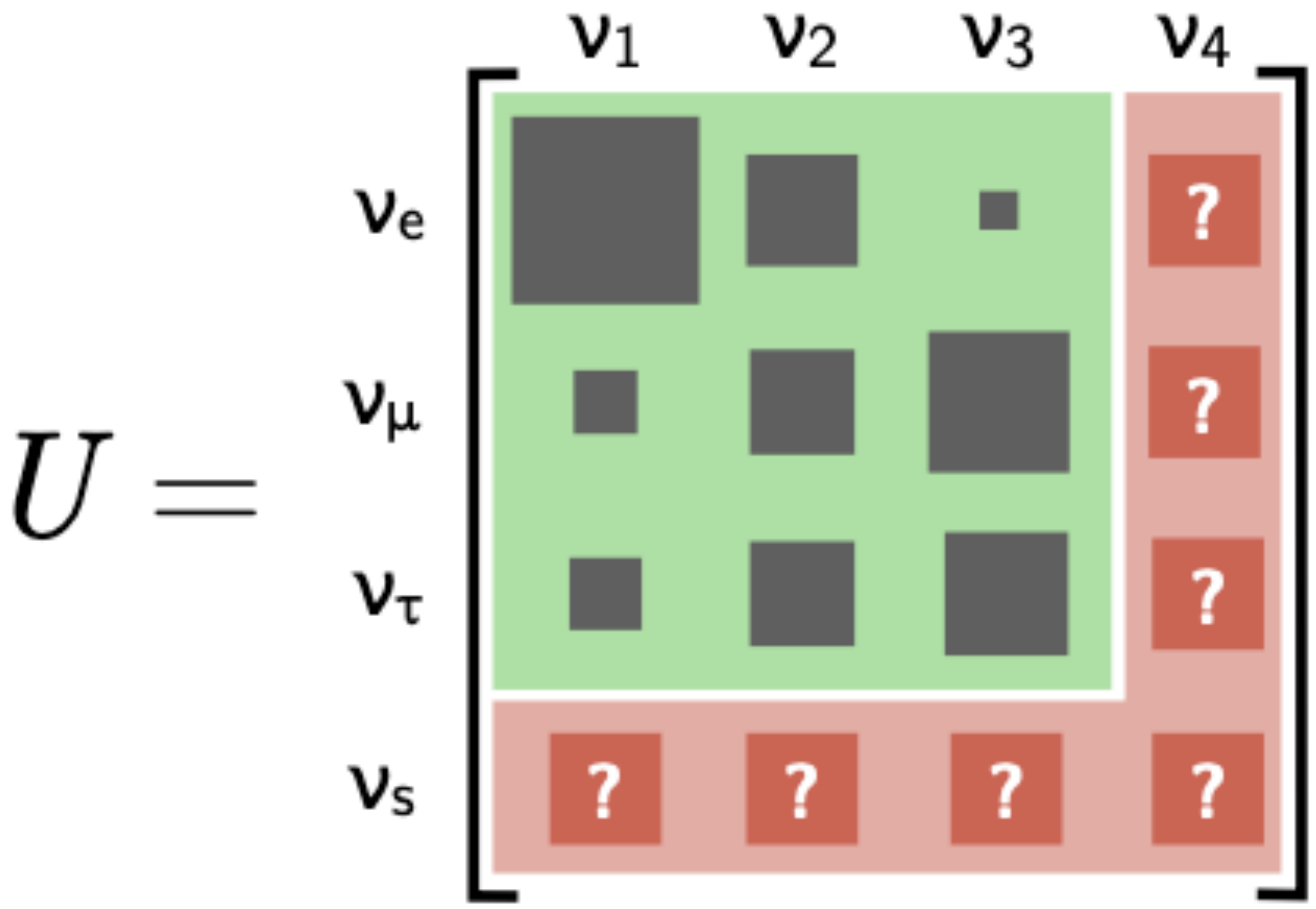
Experiment	Type	Channel	Significance
LSND	DAR	$\bar{\nu}_\mu \to \bar{\nu}_e$ CC	3.8σ
MiniBooNE	SBL accelerator	$\nu_\mu \to \nu_e$ CC	3.4σ
MiniBooNE	SBL accelerator	$\bar{\nu}_\mu \to \bar{\nu}_e$ CC	2.8σ
GALLEX/SAGE	Source - e capture	ν_e disappearance	2.8σ
Reactors	Beta-decay	$\bar{\nu}_e$ disappearance	3.0σ

K. N. Abazajian et al. "Light Sterile Neutrinos: A Whitepaper"
arxiv:1204.5379

- taken individually, each anomaly is not significant enough to be convincing: but they all are pointing toward the similar thing
- most commonly interpreted as hint for one or more new "sterile" neutrino (oscillates but does not interact weakly)

Why add extra neutrino?

- the number of *weakly interacting* "**active**" neutrino flavors is fixed to three, by the *Z* width measurements (LEP)
- but additional, *non-interacting* "**sterile**" neutrino states could still exist
- potentially detectable through impact on neutrino oscillations
- *Q: can this new type of neutrino be solution to these anomalies?*
- *A: unfortunately, it's not so simple… there are severe tension between different measurements & channels*



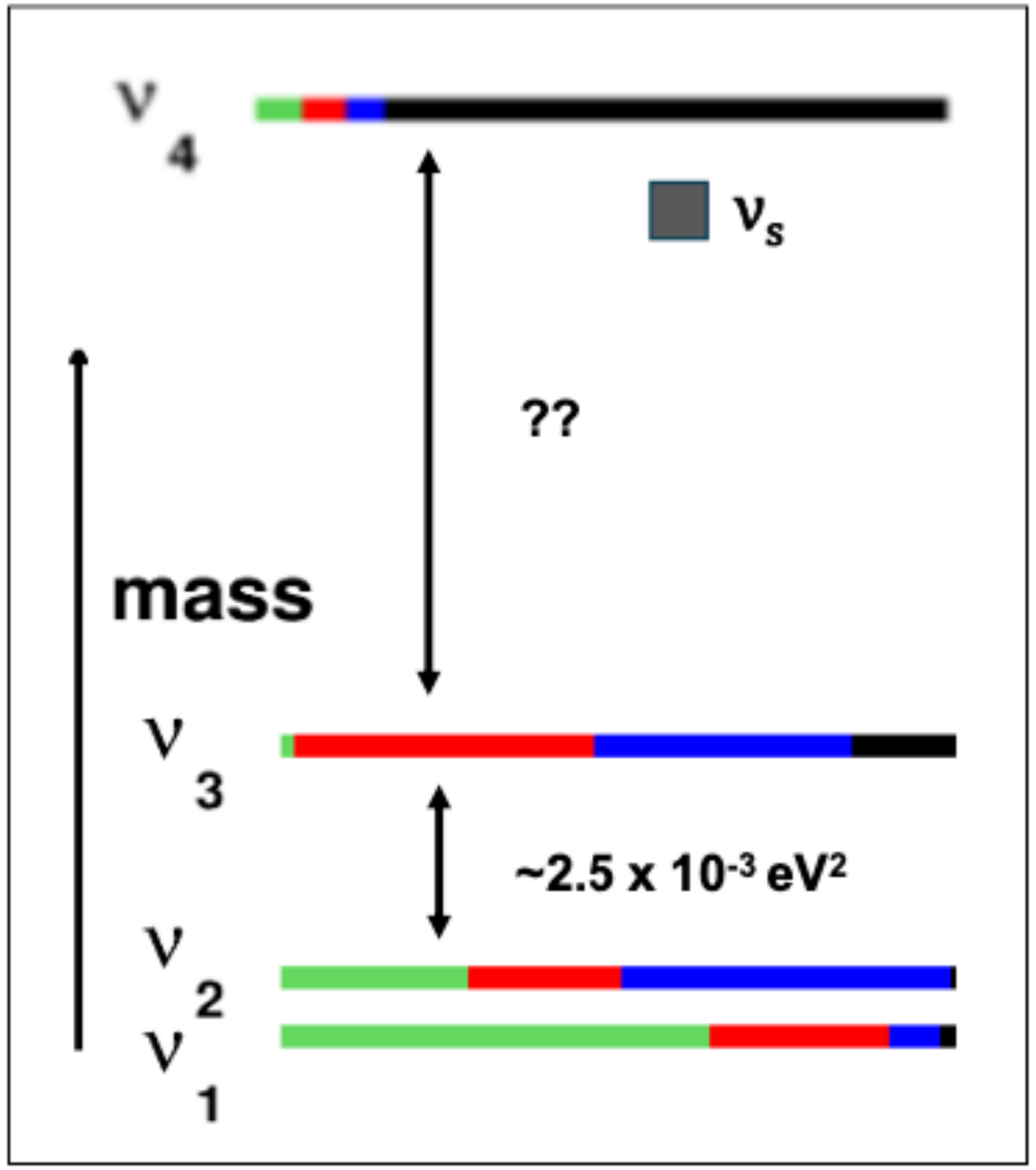
Flavor transitions via this new mixing:

$$P_{\alpha\beta} = 4|U_{\alpha 4}|^2|U_{\beta 4}|^2 \sin^2\left(1.27\frac{\Delta m^2_{41} L}{E}\right)$$

How can we detect sterile neutrino?

- sterile neutrino does not interact weakly, only experience gravity: no way to *directly* detect it
- but it still oscillates like other neutrino species, hence affecting neutrino oscillation pattern
  - oscillation probability of how one neutrino state morphs into the other state will be different if extra neutrino exists (i.e. PMNS matrix changes)
  - ν_e disappearance channel: $\nu_e \rightarrow \nu_e$
    - how many ν_e has been oscillated into other (including ν_s) neutrino types?
  - ν_e appearance channel: $\nu_u \rightarrow \nu_e$
    - how many ν_e has been oscillated from ν_u?

3+1 mixing model



The MiniBooNE Anomaly: Low Energy Excess (LEE)

- MiniBooNE observed low-energy excess (LEE) of electron-neutrino-like events
  - LEE: more events measured/detected than predicted, in the low energy region
- eV-scale sterile neutrino could explain this excess
  - the excess is due to sterile neutrino oscillated into electron neutrino
  - prediction is lower than observed because the prediction is made based on 3-neutrino paradigm

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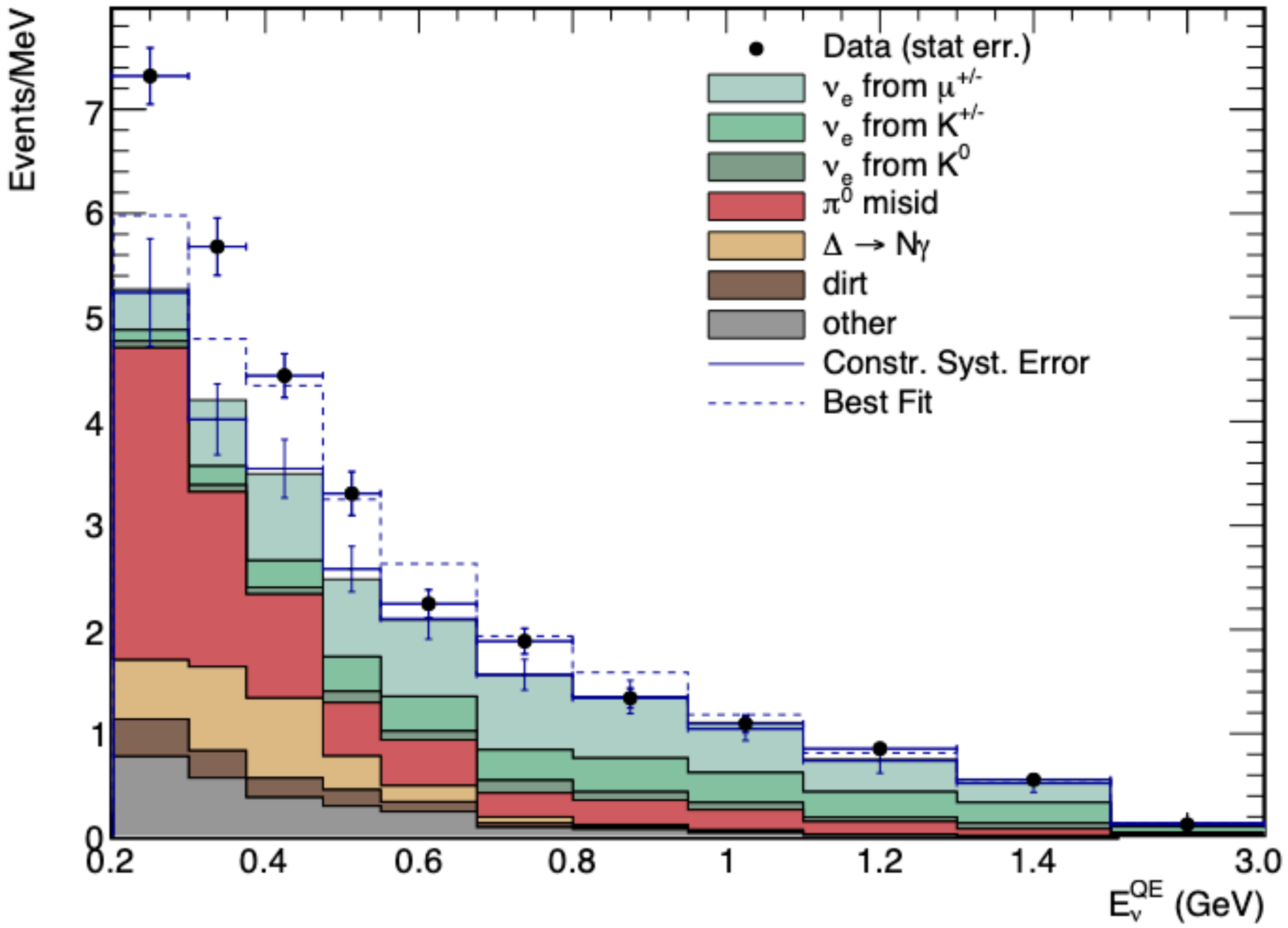
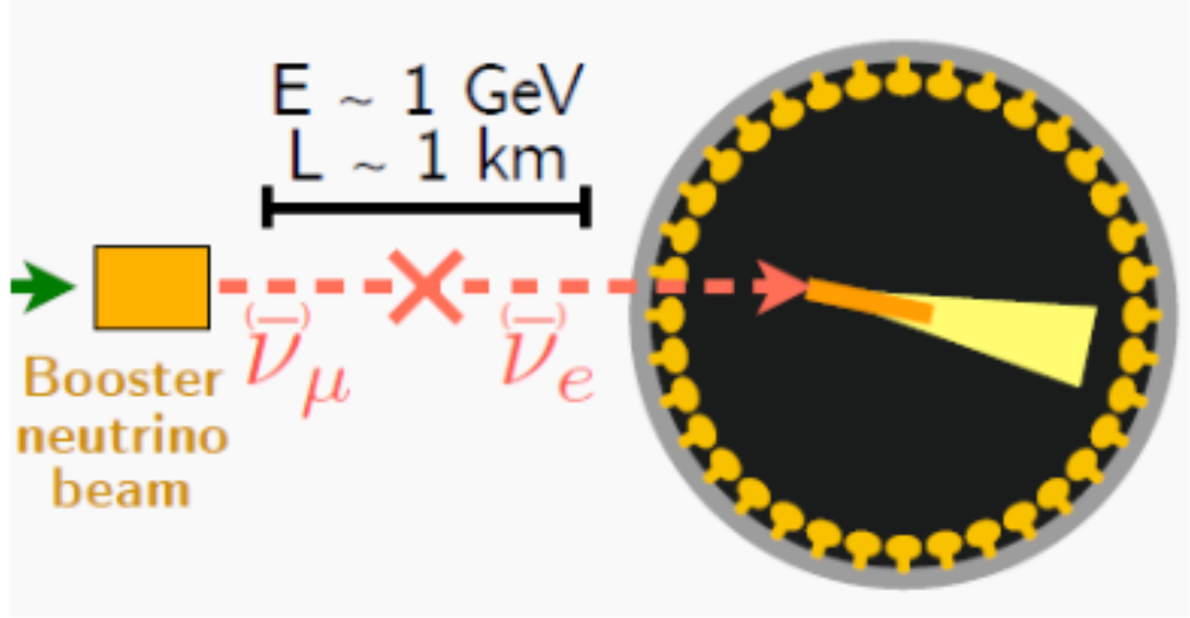
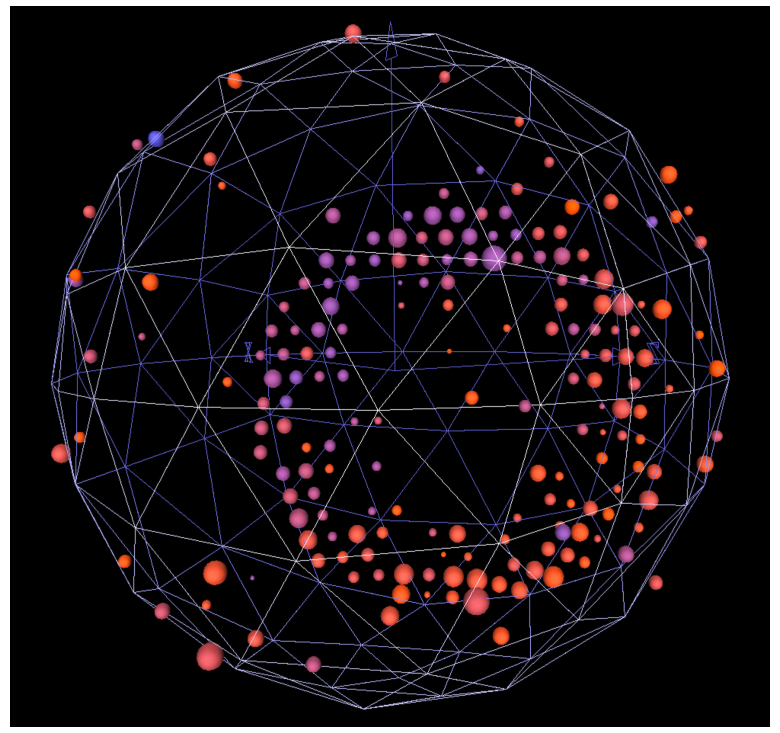
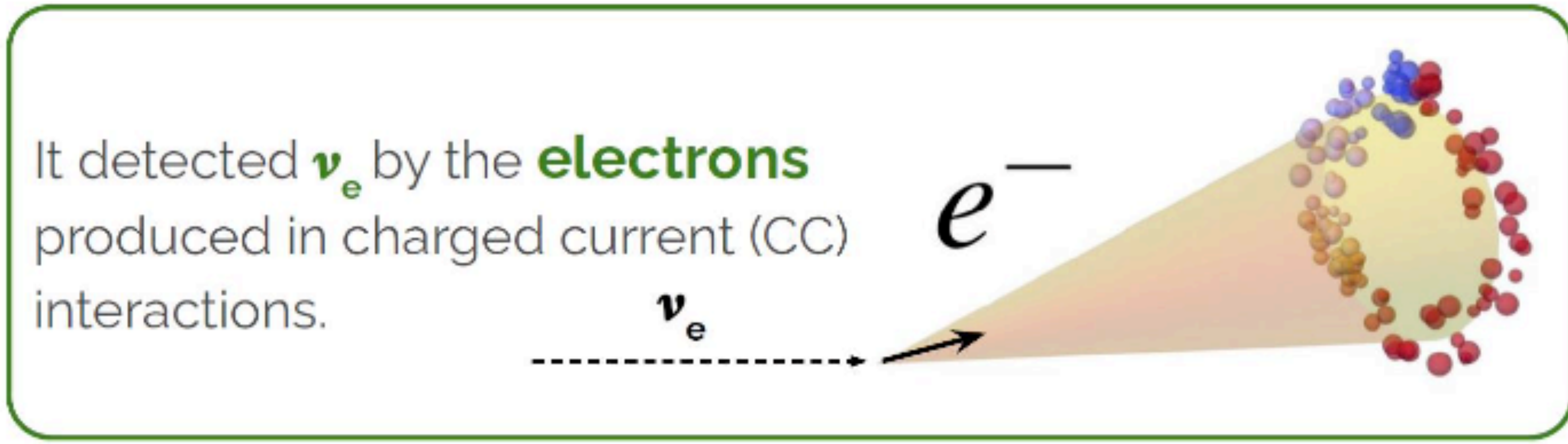


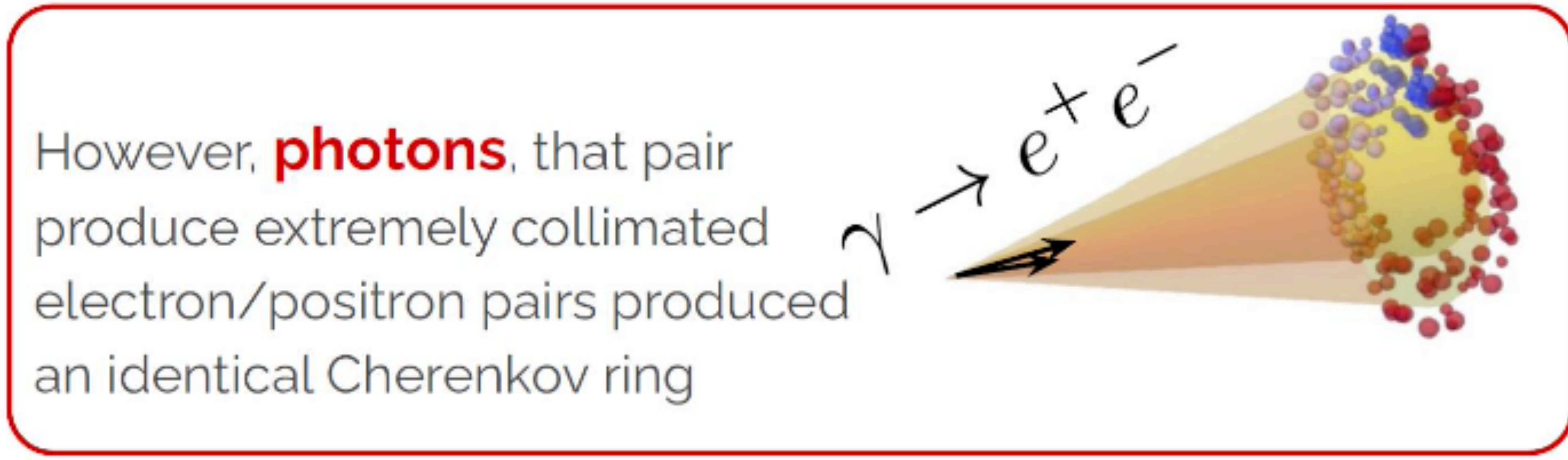
The MiniBooNE Anomaly: Low Energy Excess (LEE)

- MiniBooNE is a Cherenkov detector
  - mostly detecting outgoing leptons (electrons, muons, etc)
  - cannot distinguish between electrons and photons
- this limitation makes it hard to interpret the origin of LEE
  - if electrons, this can be explained by sterile neutrino oscillated into electron neutrinos
  - if photons, this can be explained by underestimated prediction of single-photon-producing SM process









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LIQUID-ARGON IONIZATION CHAMBERS AS TOTAL-ABSORPTION DETECTORS*

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1974

The Time-Projection Chamber
- A new 4π detector for charged particles

David R. Nygren

Lawrence Berkeley Laboratory
Berkeley, California 97420

1976

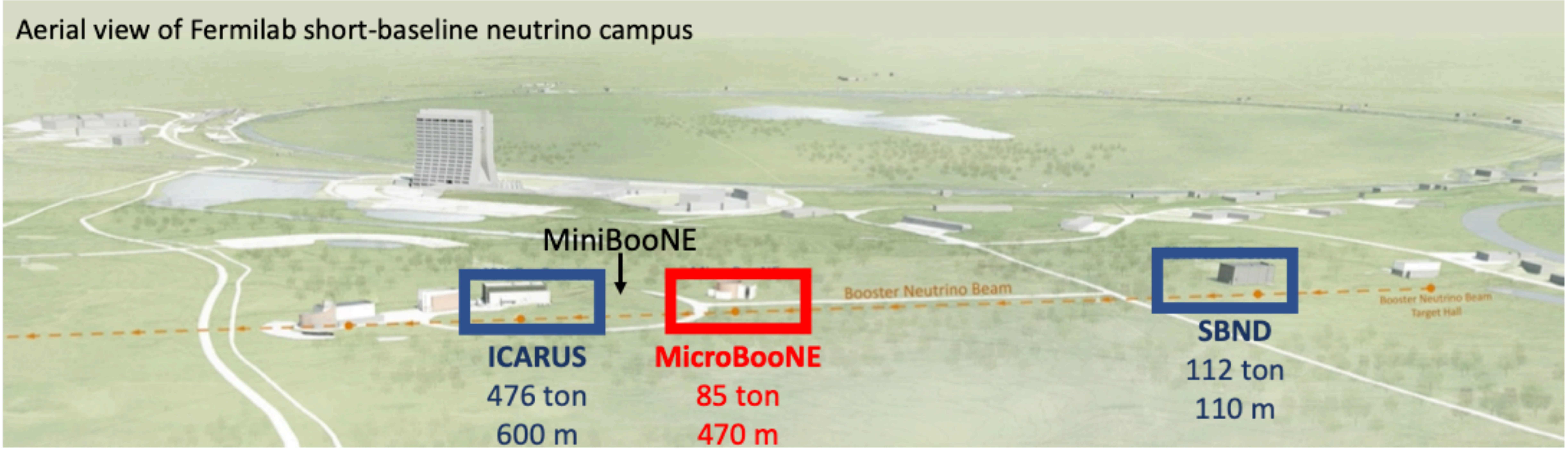
THE LIQUID-ARGON TIME PROJECTION CHAMBER:

A NEW CONCEPT FOR NEUTRINO DETECTORS

C. Rubbia

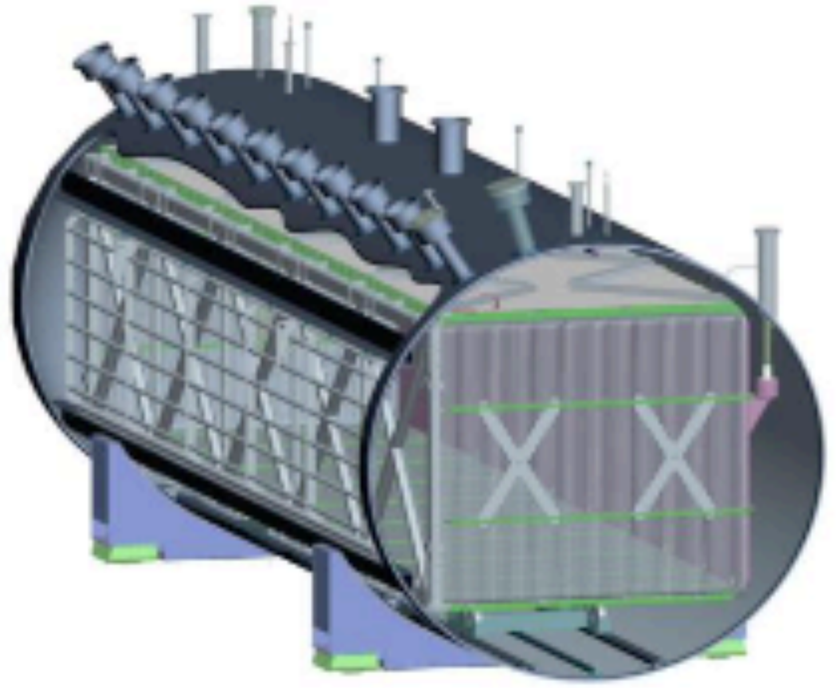
1977

Fermilab SBN program

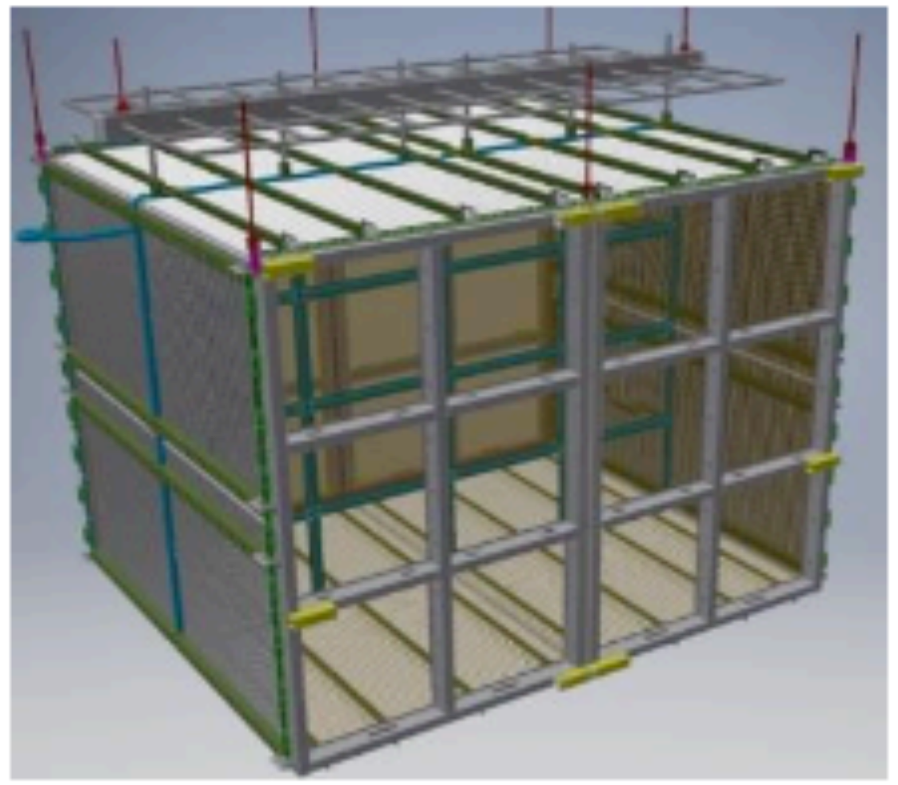


- three LArTPC detectors, with same neutrino beamline and different baseline
- reduce statistical uncertainties with large mass far detector (ICARUS)
- reduce systematic uncertainties with same LArTPC detector technology

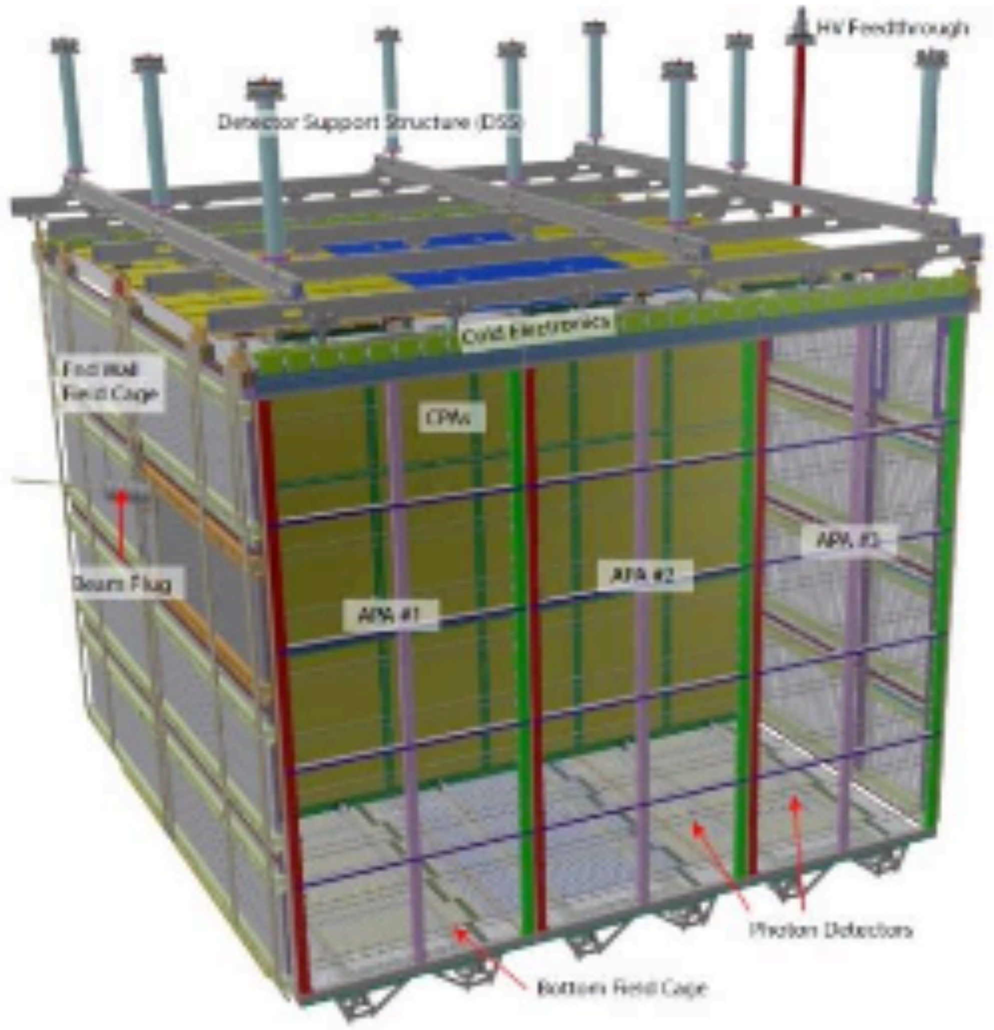
Fermilab SBN program



MicroBooNE, 87 ton
2.3m x 2.5m x 10.4m



SBND, 112 ton
4m x 4m x 5m



ICARUS, 476 ton
1.5m x 2.2m x 18m x 4

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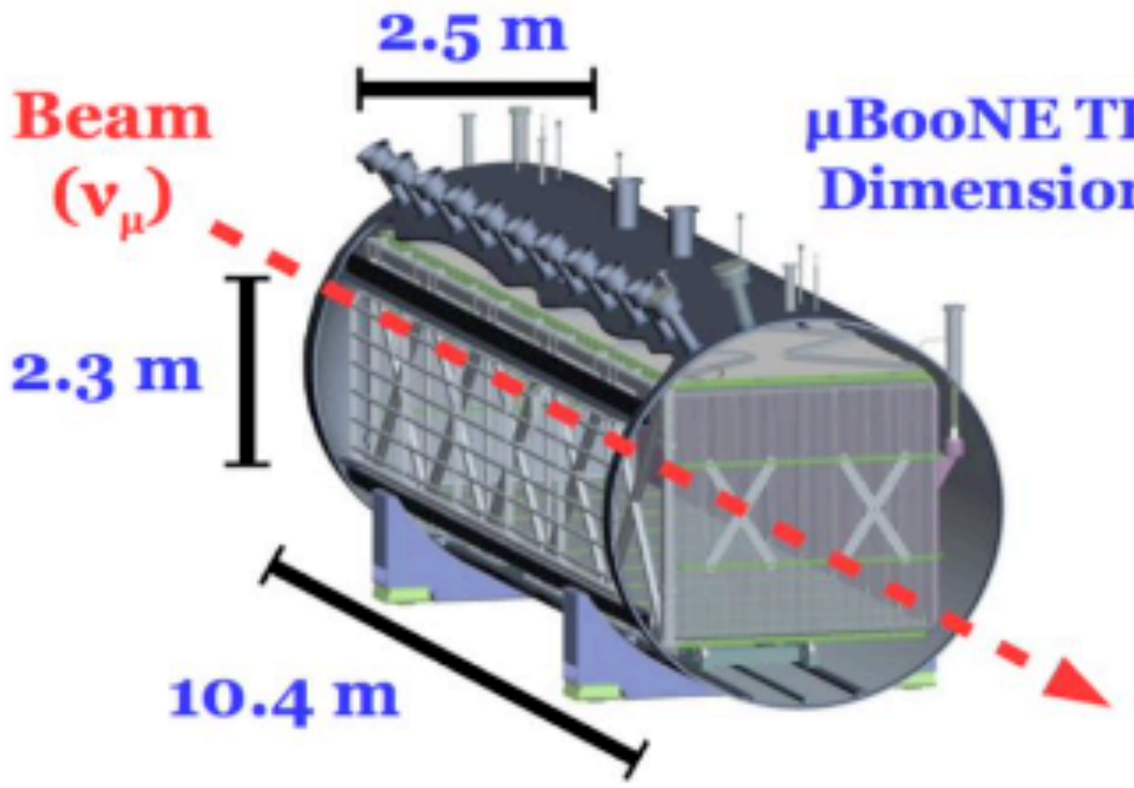
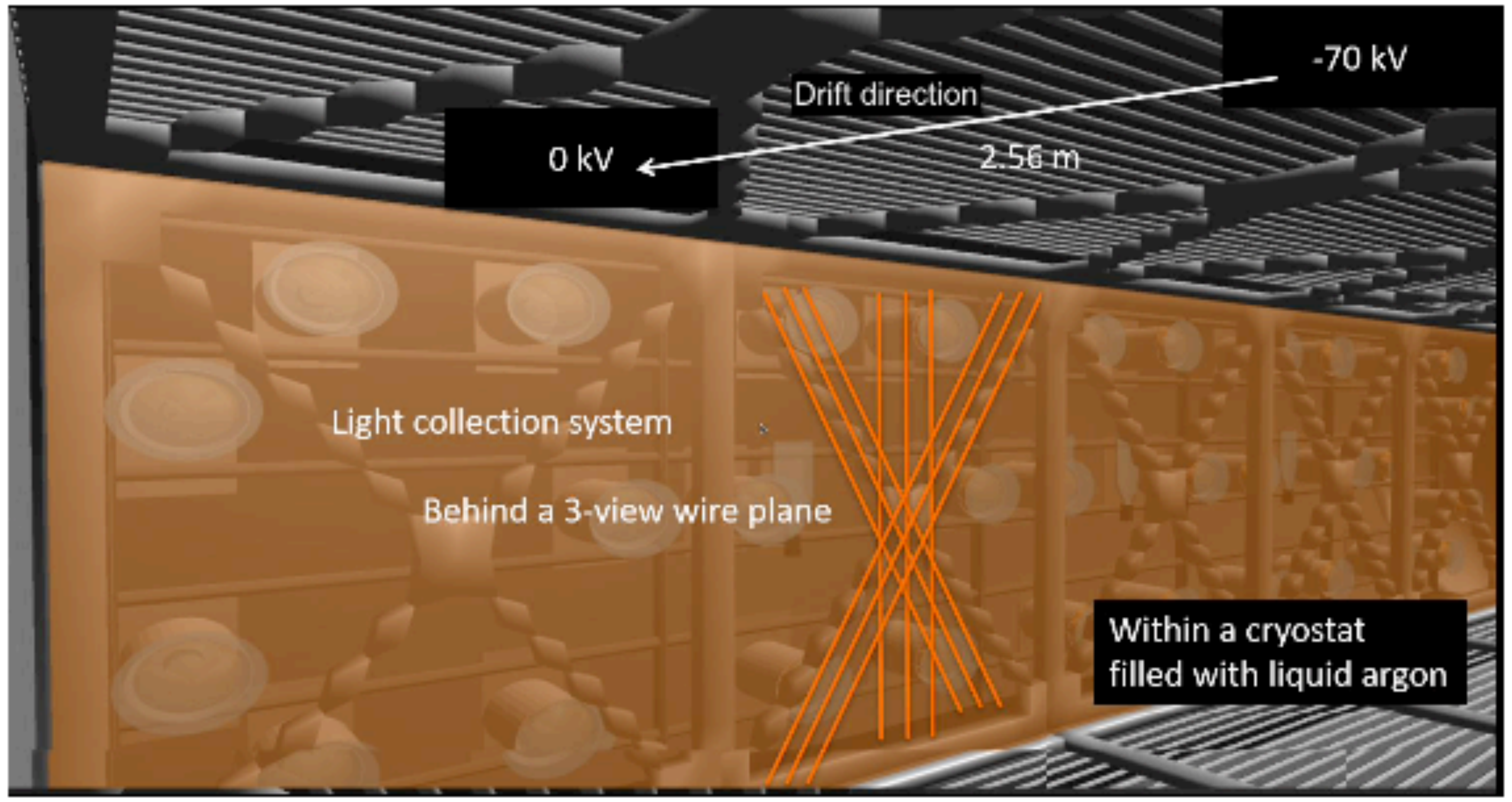
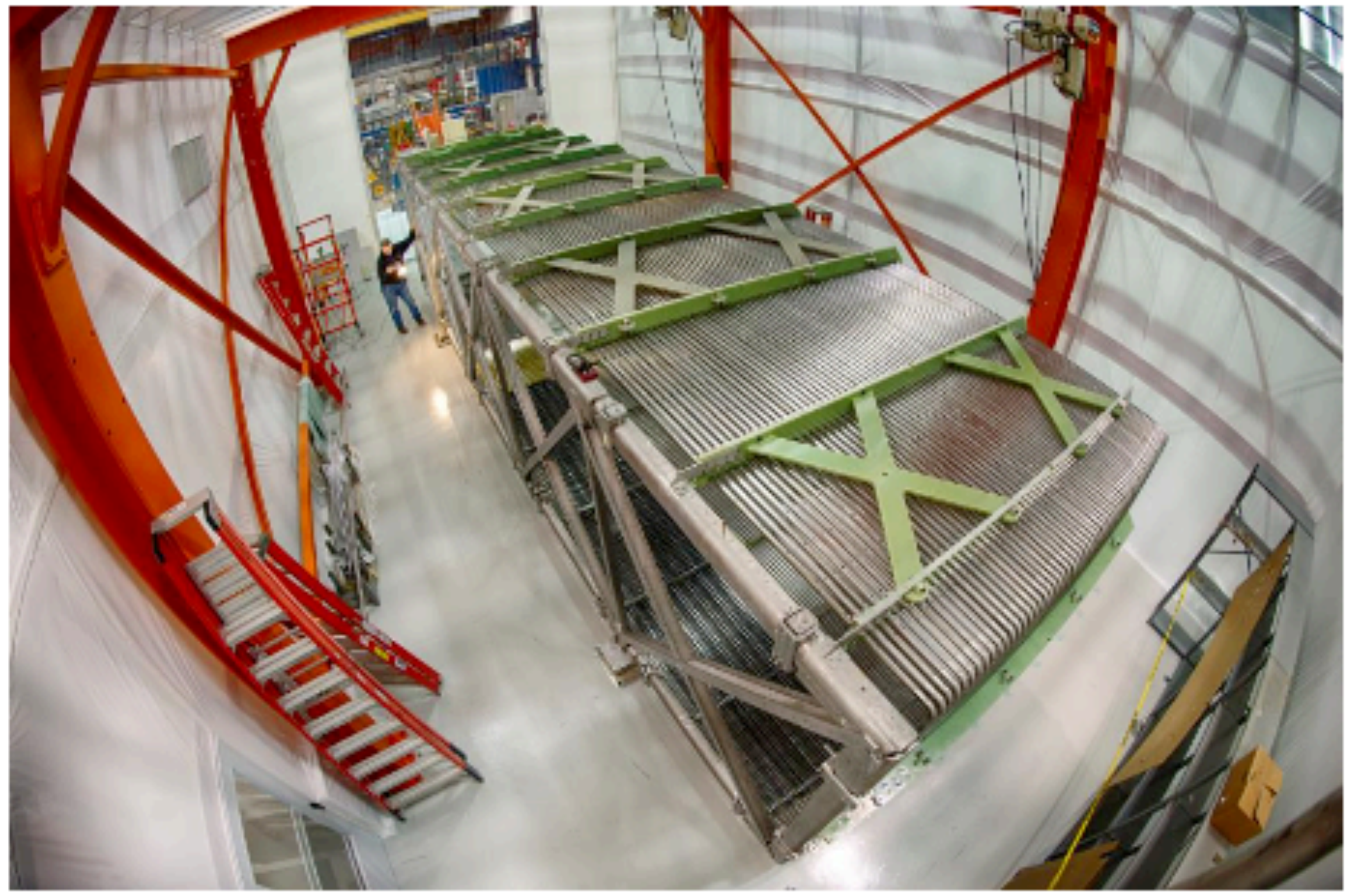
MicroBooNE



2.5 m
Beam
(ν_μ)
μBooNE TPC
Dimensions
2.3 m
10.4 m

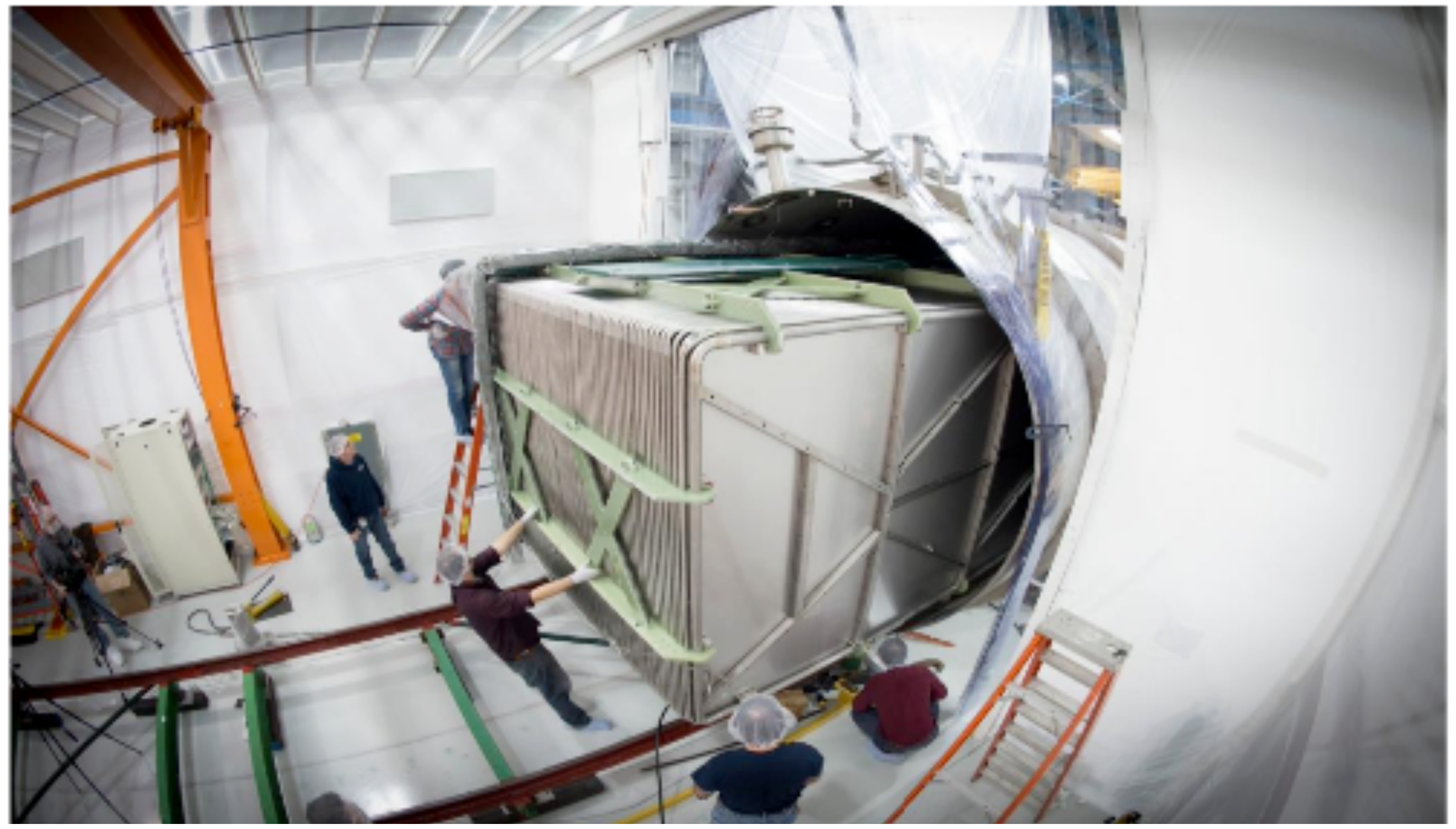


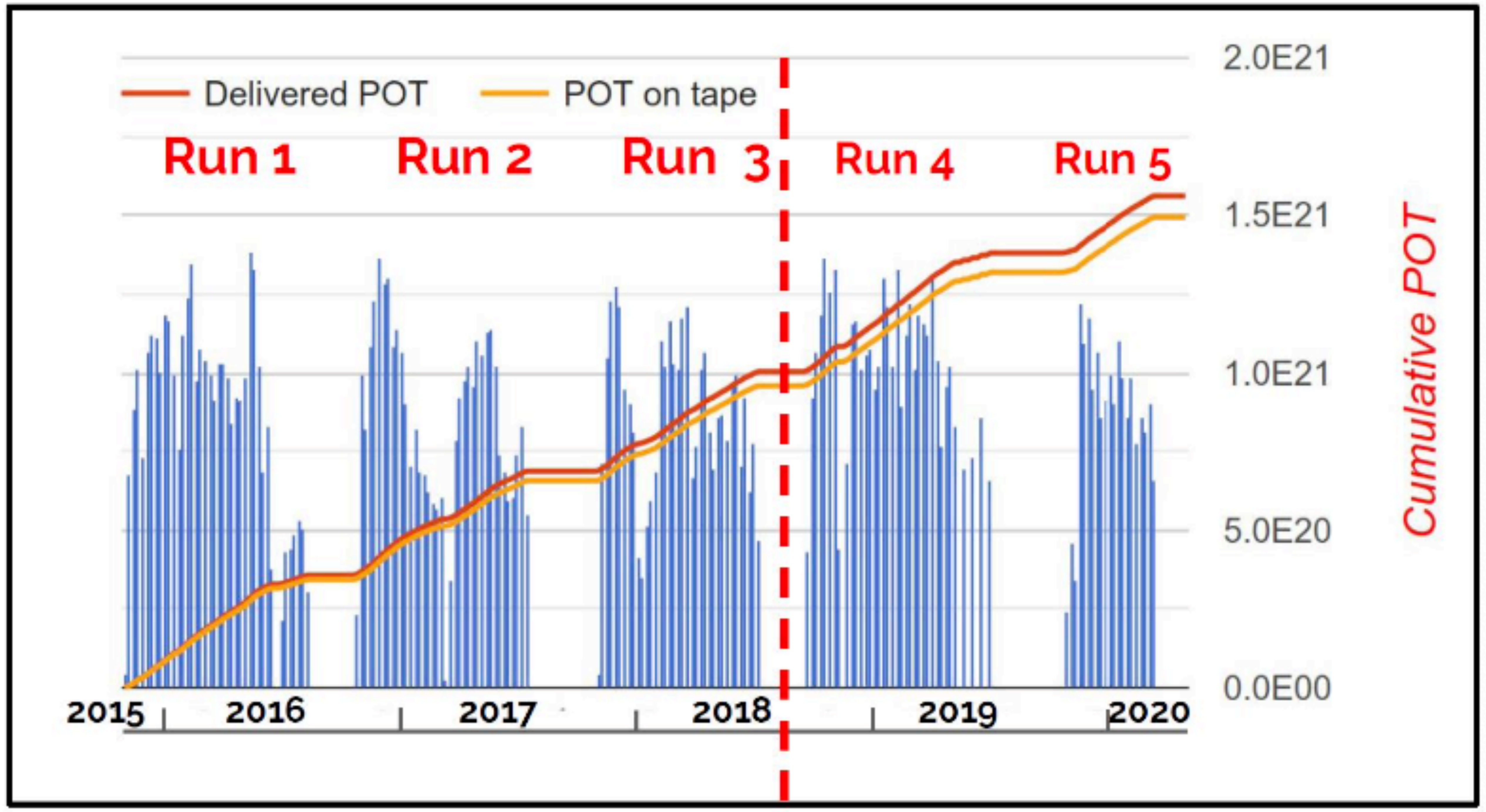
-70 kV
Drift direction
0 kV
2.56 m
Light collection system
Behind a 3-view wire plane
Within a cryostat
filled with liquid argon



MicroBooNE

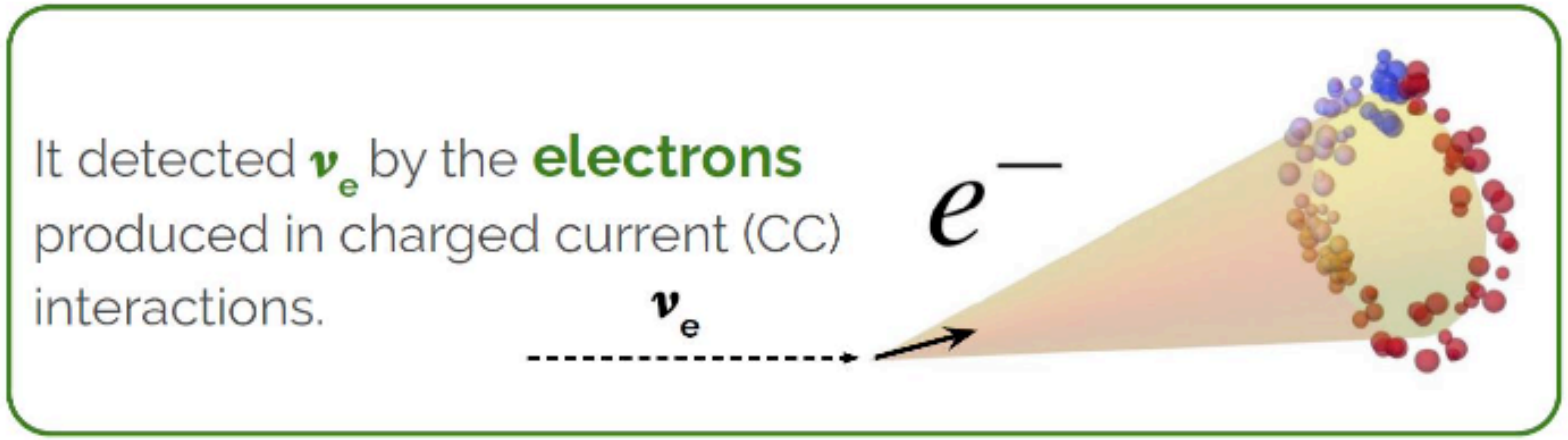
- started taking data since 2015
- finished operation in 2021
- accumulated the world's largest sample of neutrino interaction on argon
- one of the first LArTPC detectors with many new features
  - cold, low noise electronics *(see Shanshan's talk at 3pm)*
  - excellent LAr purity
  - pioneered LArTPC detector physics
  - stable & long-term running

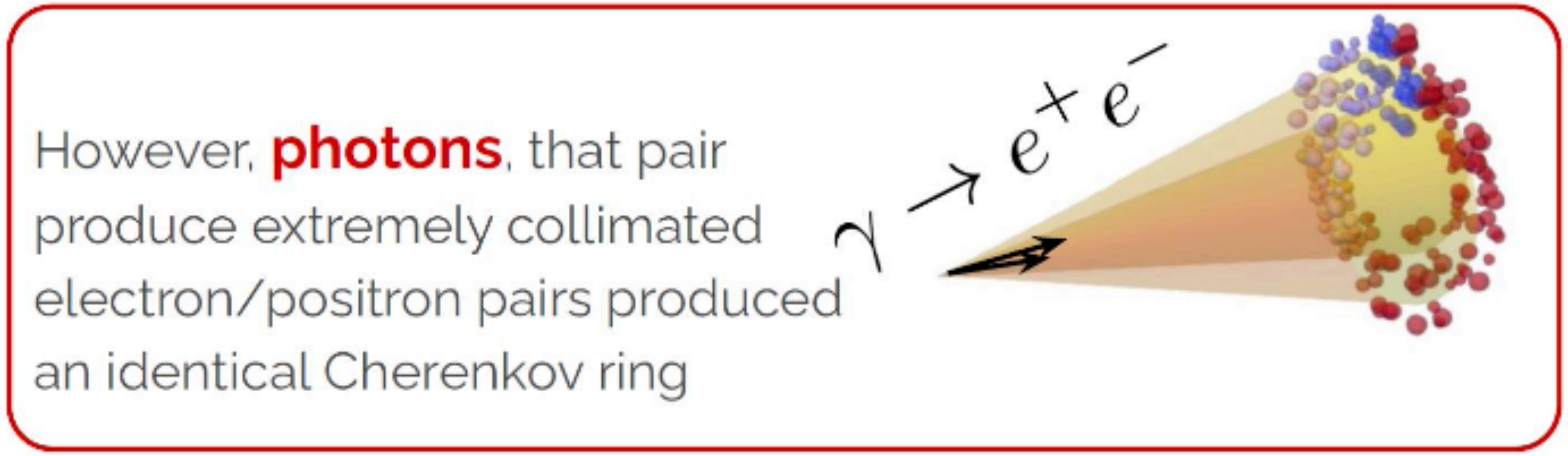




The MiniBooNE Anomaly: recap

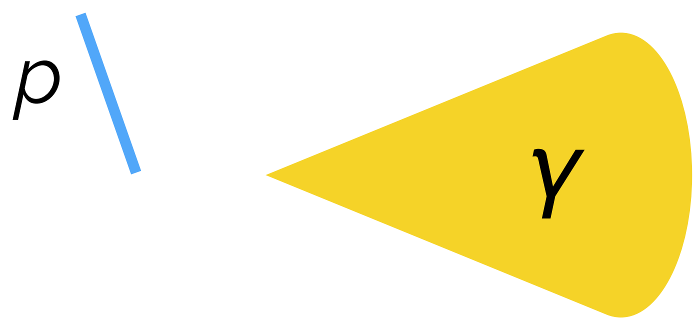
- this limitation makes it hard to interpret the LEE
  - if electrons, this can be explained by sterile neutrino oscillated into electron neutrinos
  - if photons, this can be explained by underestimated prediction of single-photon-producing SM process

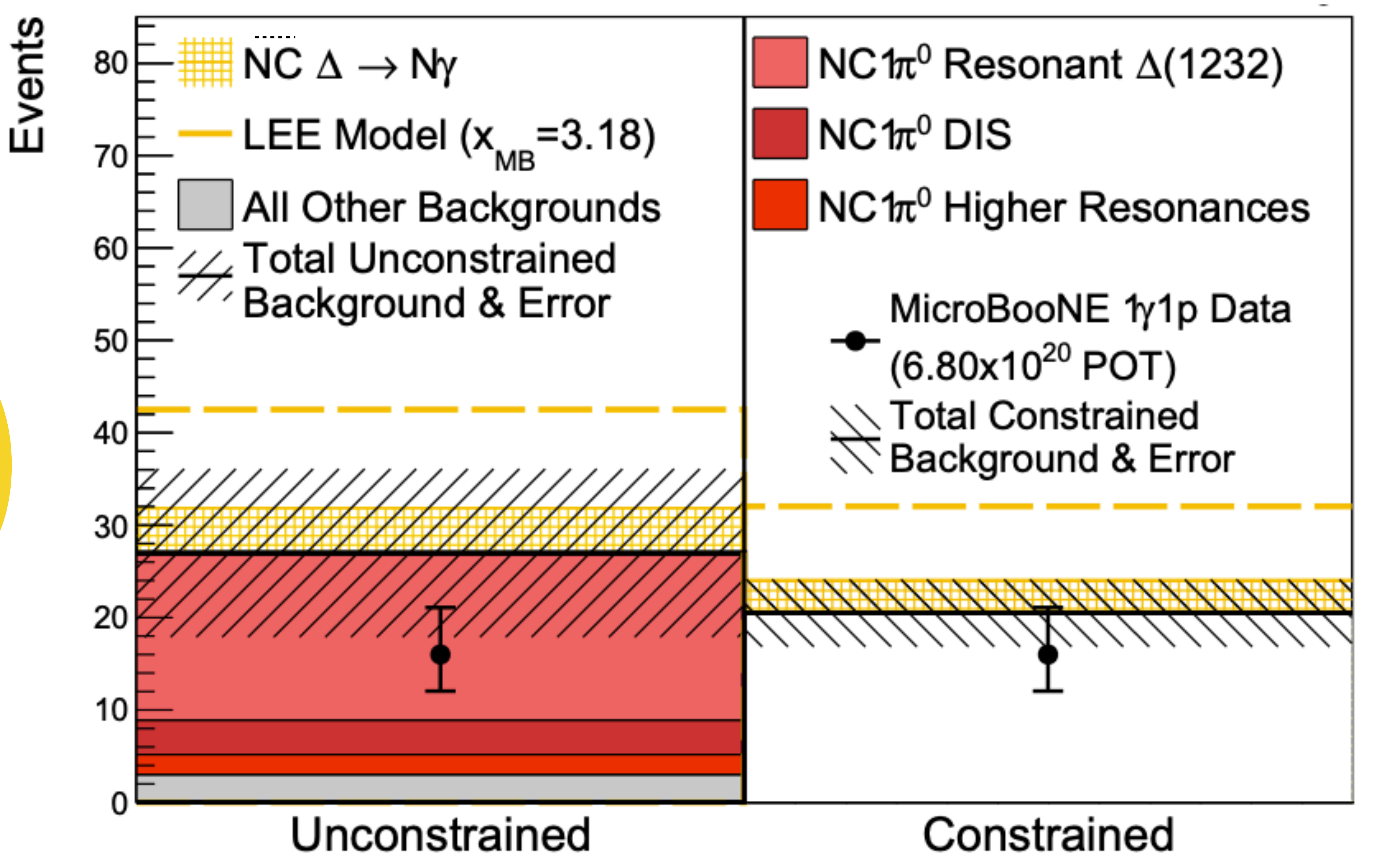


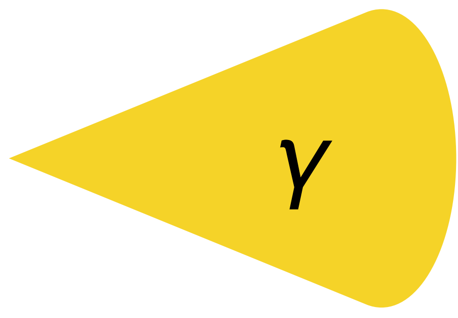


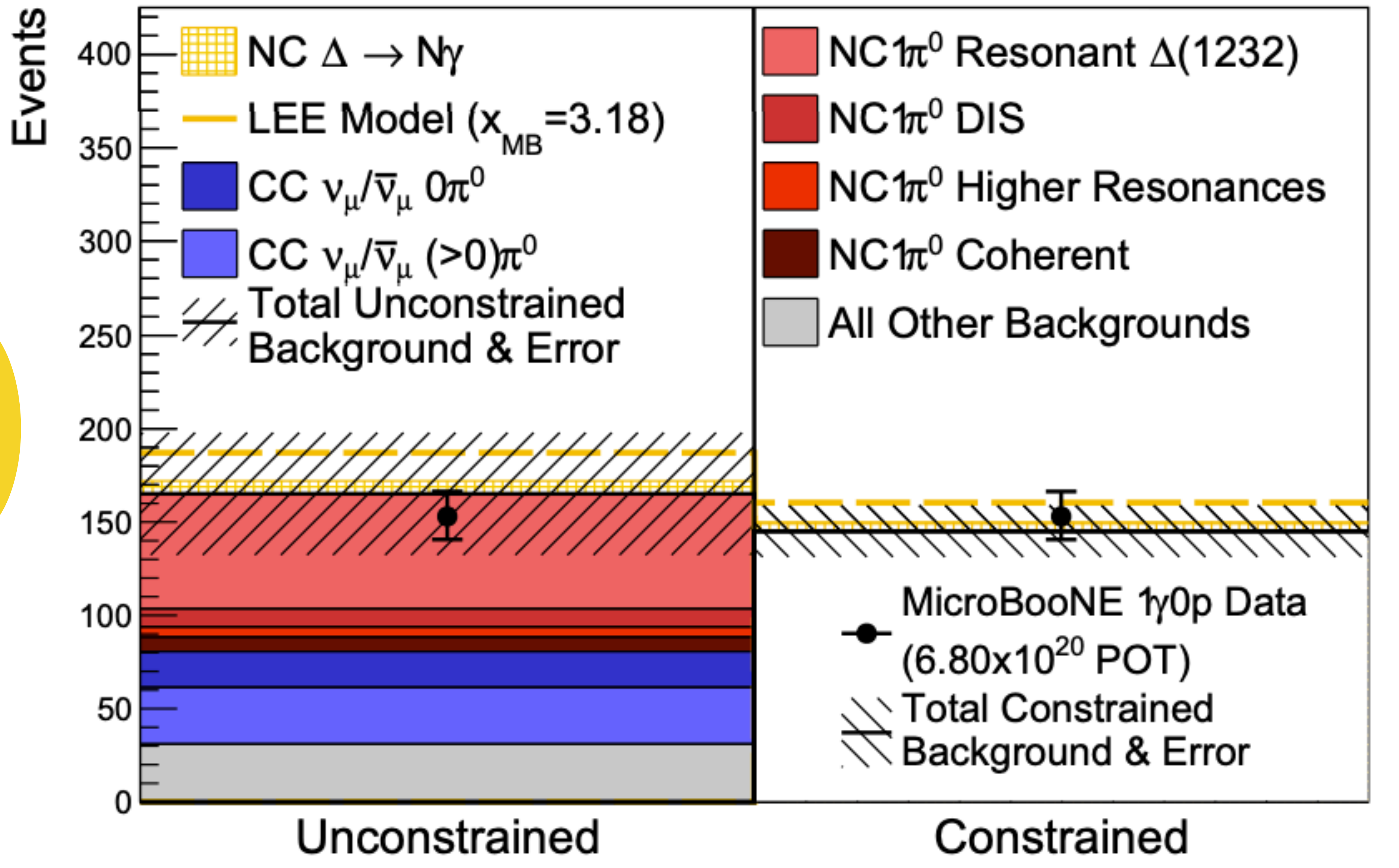
MicroBooNE LEE result

- first MicroBooNE result probed both electron-like and photon-like signals, with LArTPC's ability of e/γ separation
- photon analysis targets NC Δ—>Nγ channel
  - test if this channel is underestimated in the standard model
  - result shows no evidence for enhanced rate of single photons from NCΔ decay





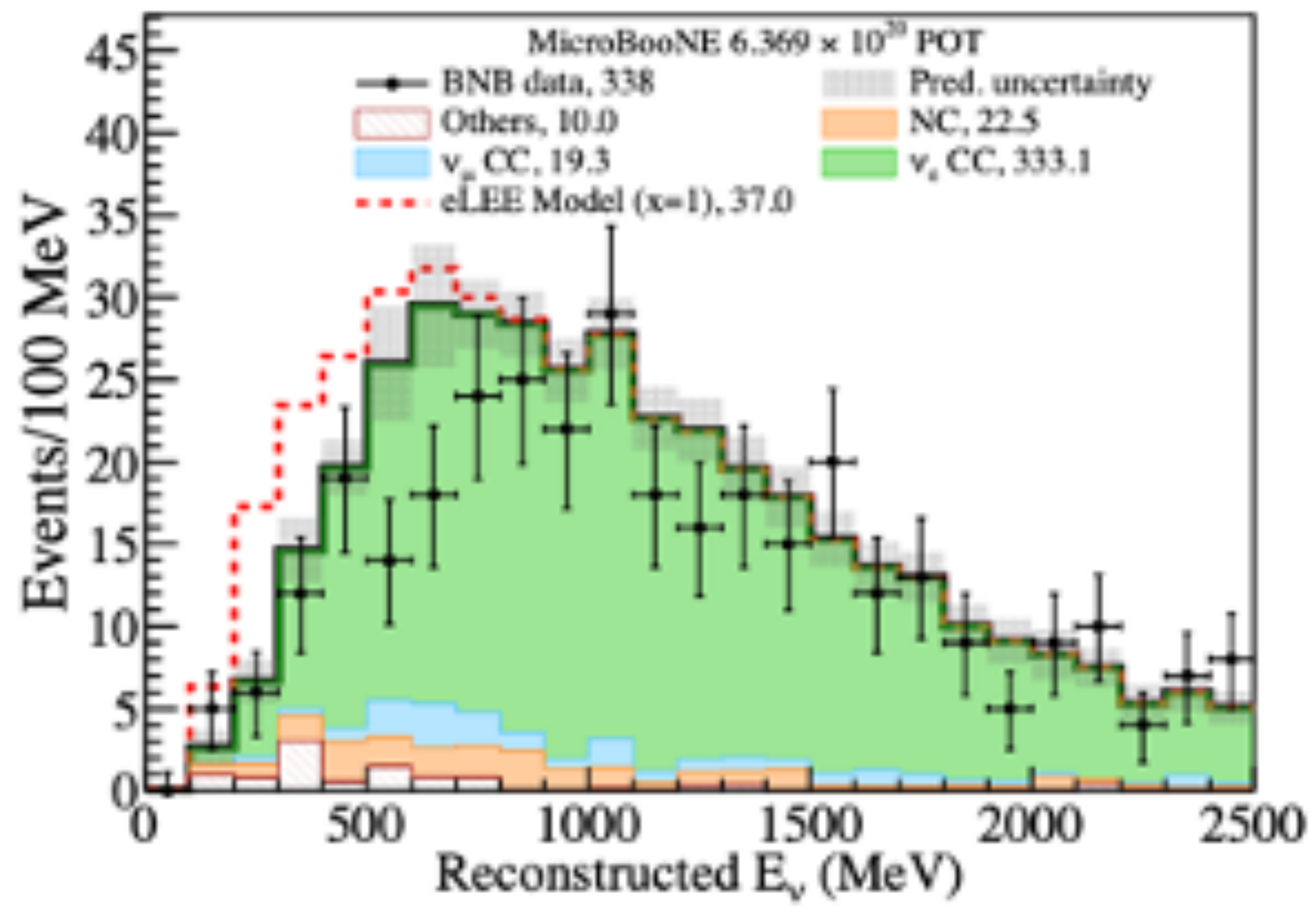




MicroBooNE LEE result

- first MicroBooNE result probed both electron-like and photon-like signals, with LArTPC's ability of e/γ separation
- electron analysis selects electron neutrino events
- test if the MiniBooNE low energy excess can be seen
  - probes 4 different topologies
  - result shows the observation is in agreement with prediction, no sign of MiniBooNE LEE

Phys. Rev. Lett. 128, 241801 (2022)



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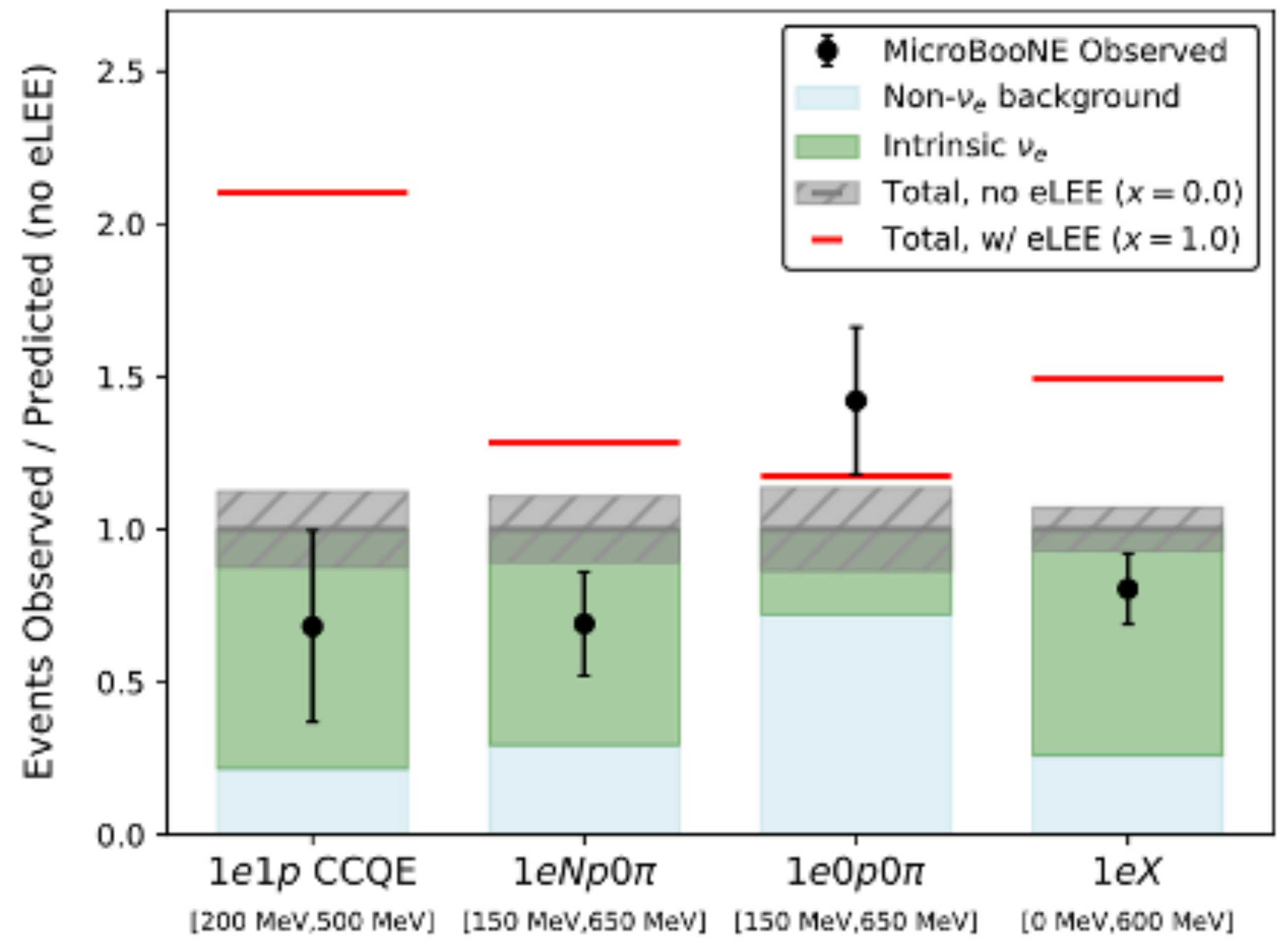
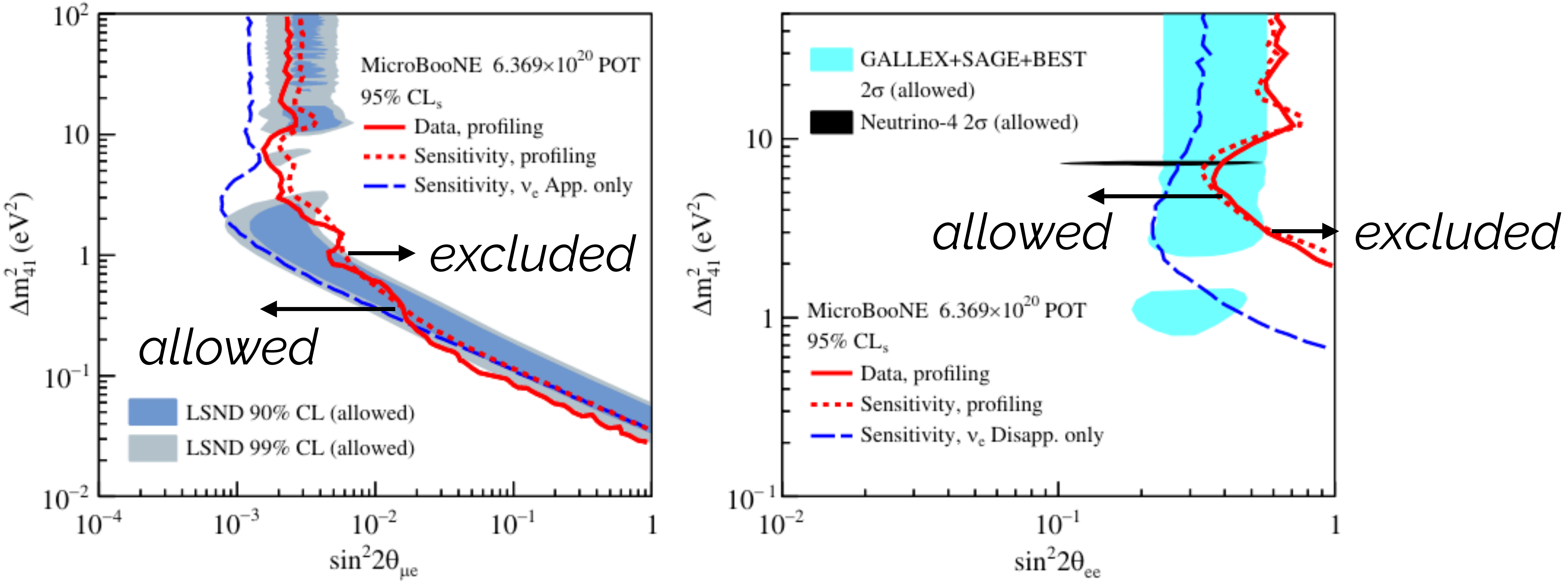


MicroBooNE sterile neutrino search

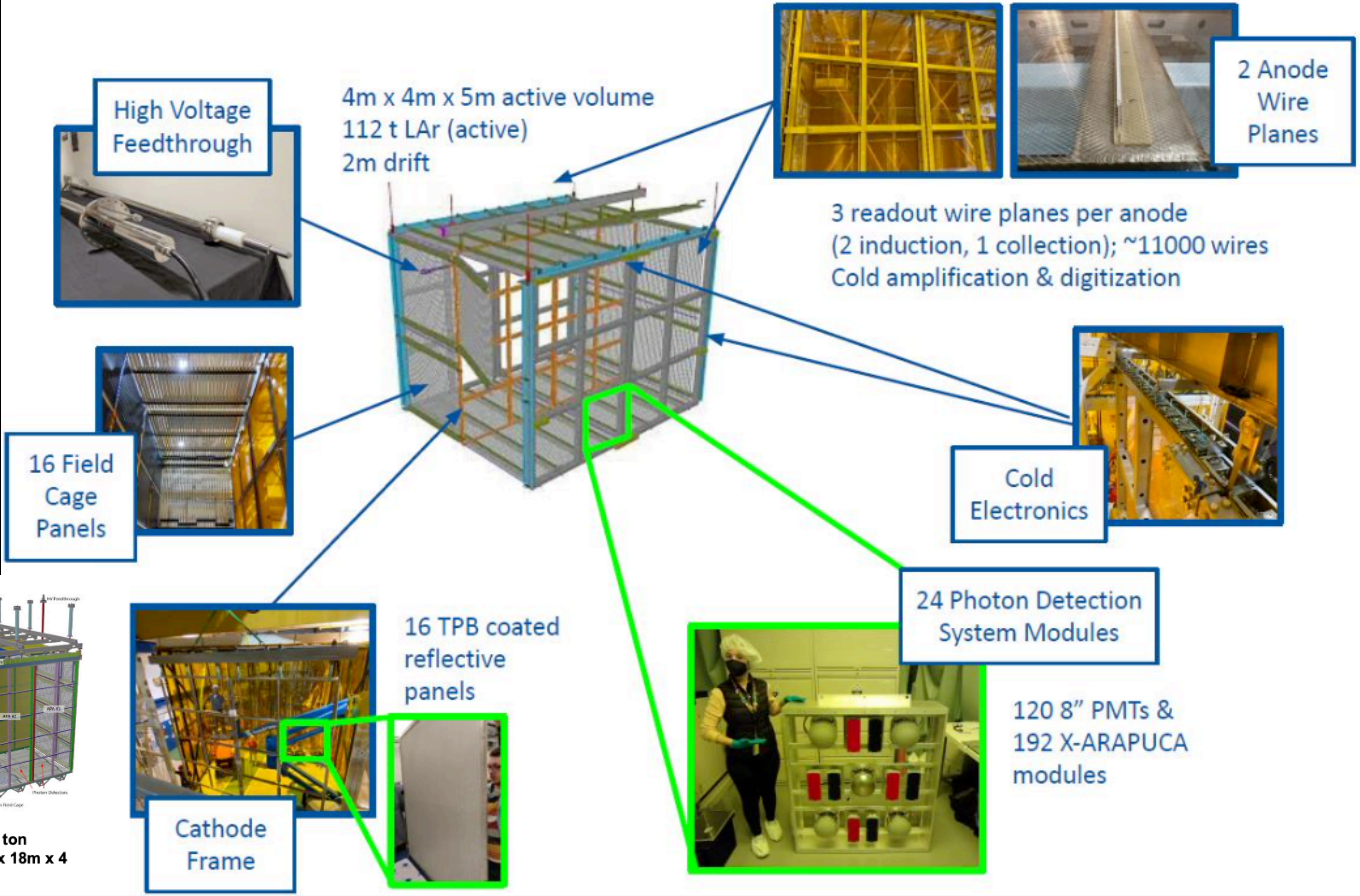
Phy. Rev. Lett. 130 011801 (2023)

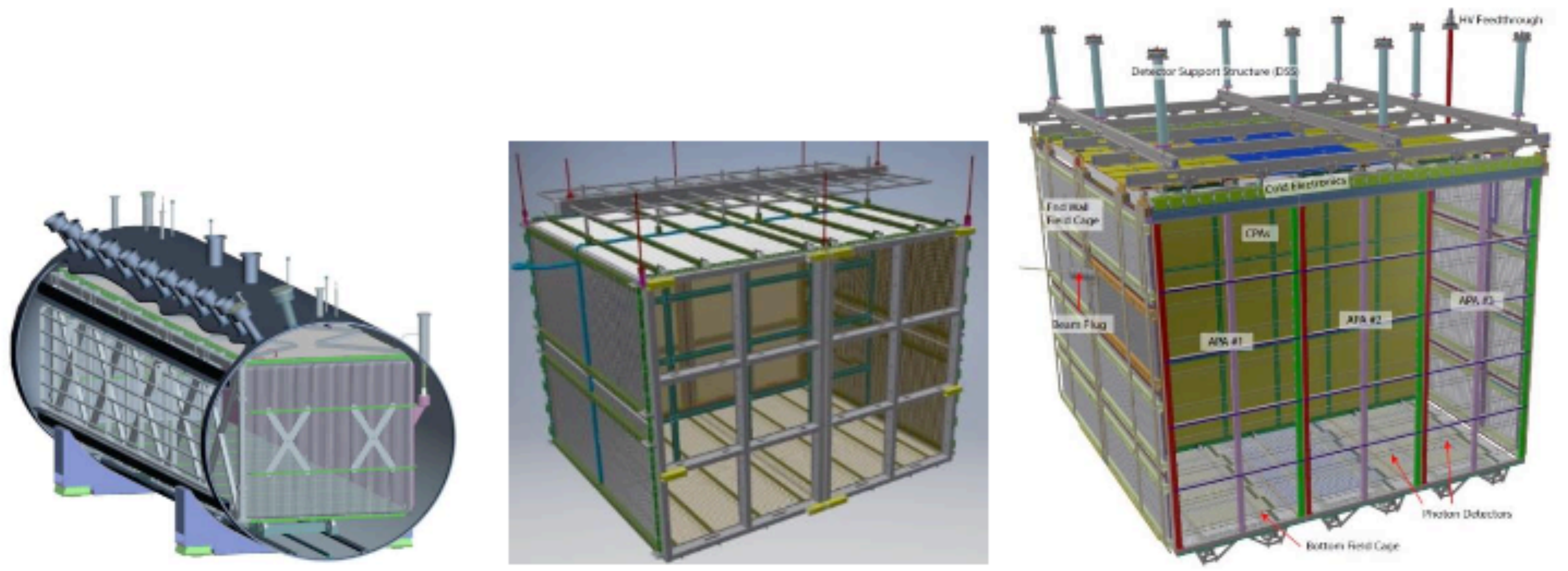


- LEE results are re-interpreted under a sterile neutrino oscillation hypothesis
- MicroBooNE could reject some portion of LSND and GALLEX/SAGE/BEST allowed region
- updated result is aiming to exclude most of the allowed region

Short Baseline Neutrino Detector: SBND

- same LArTPC technology with some upgrade
- closer to the target: much more neutrinos detected
- will start taking data in late 2023





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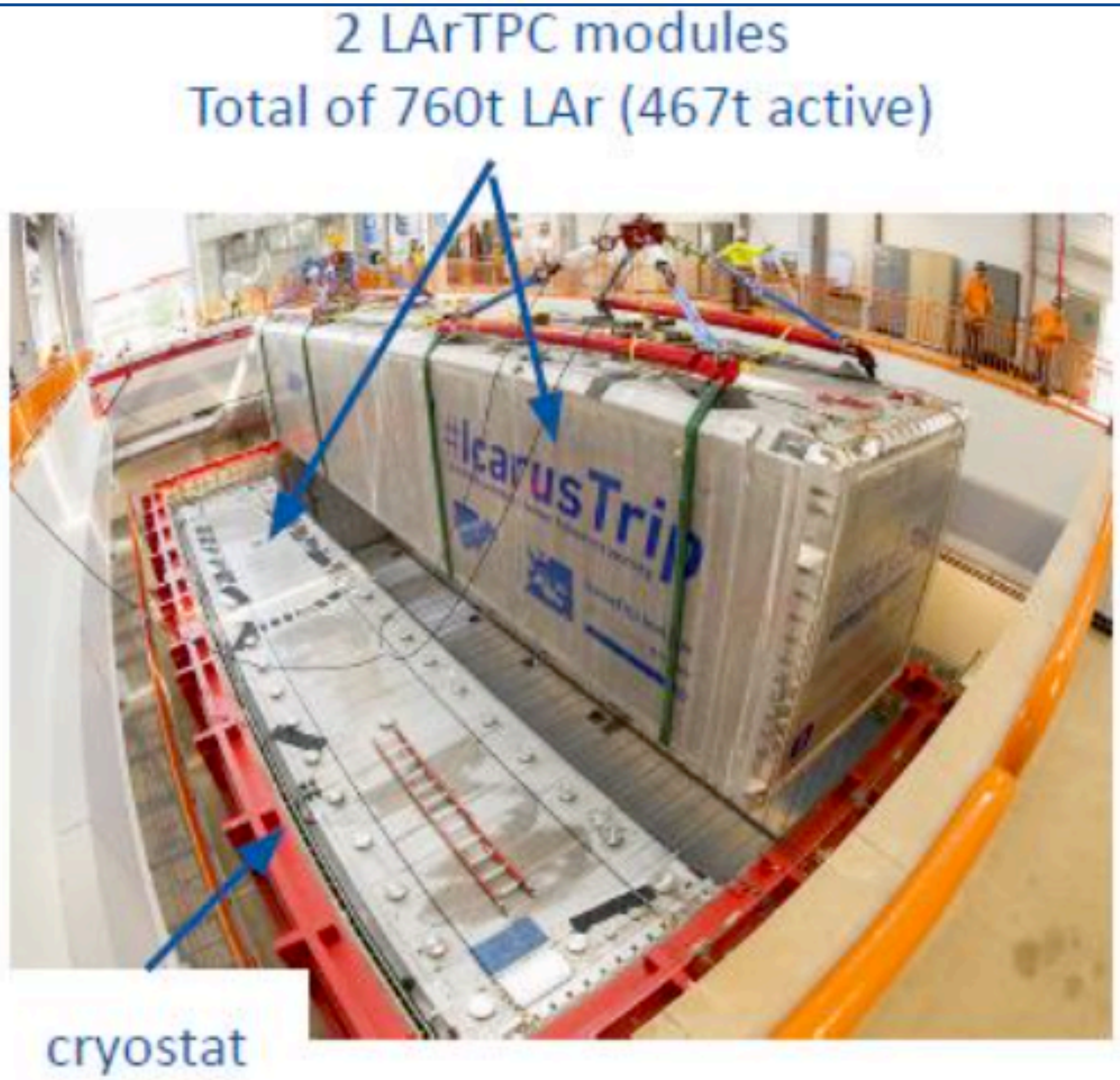
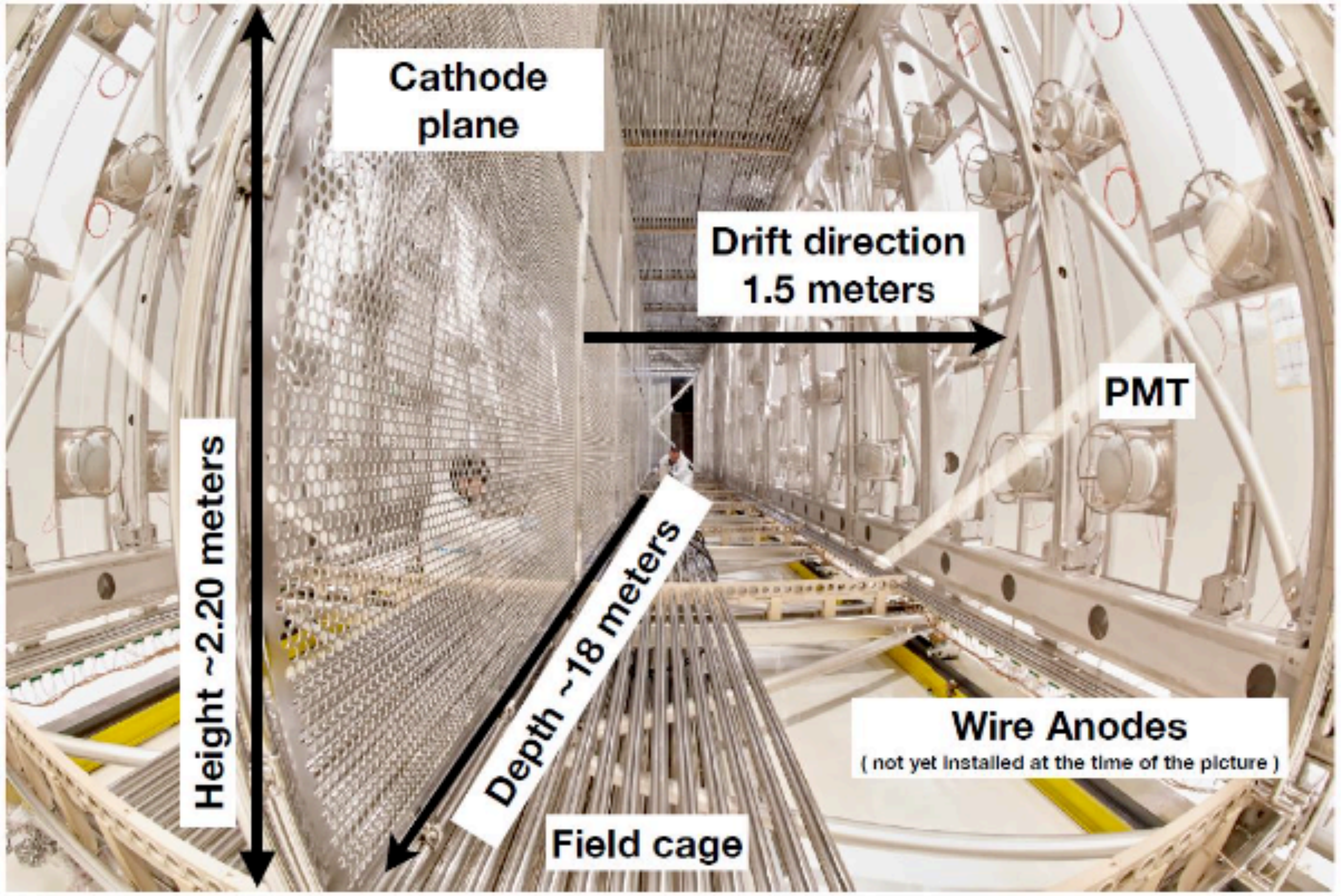
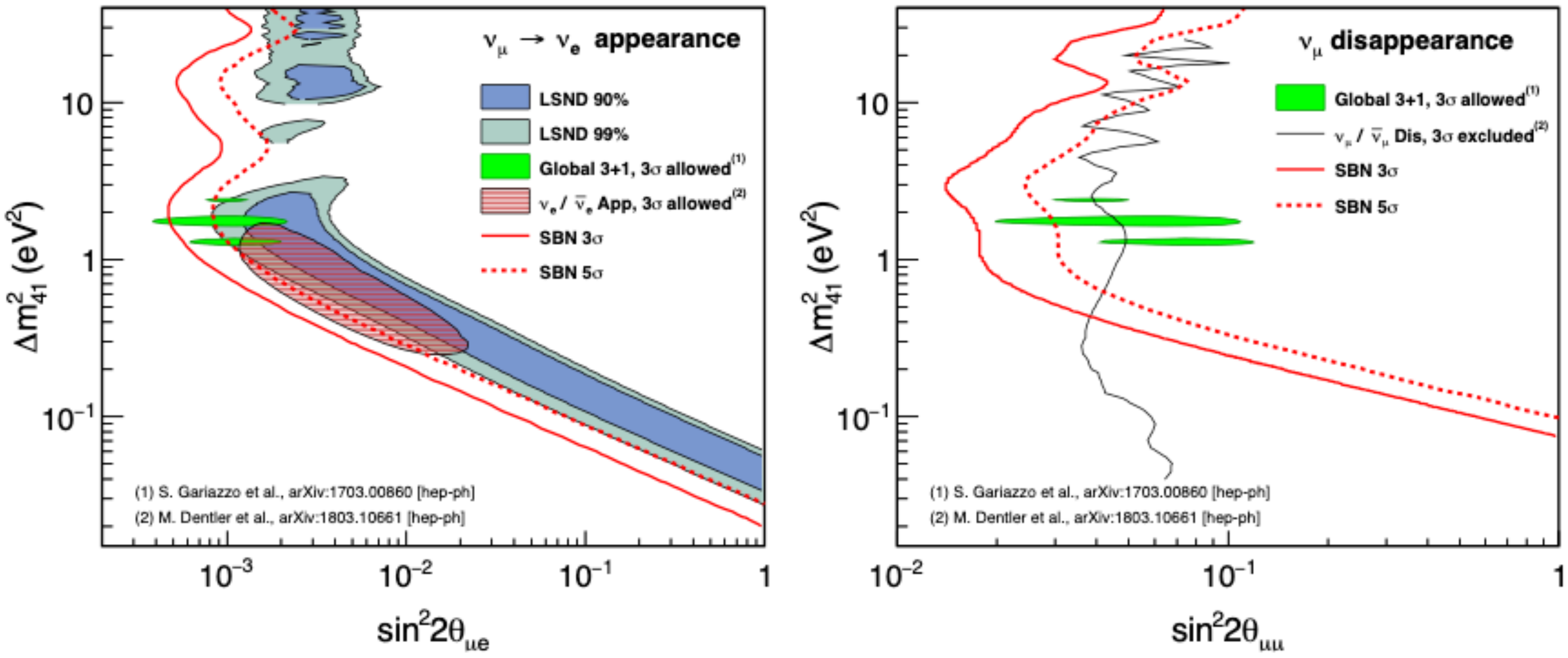
ICARUS





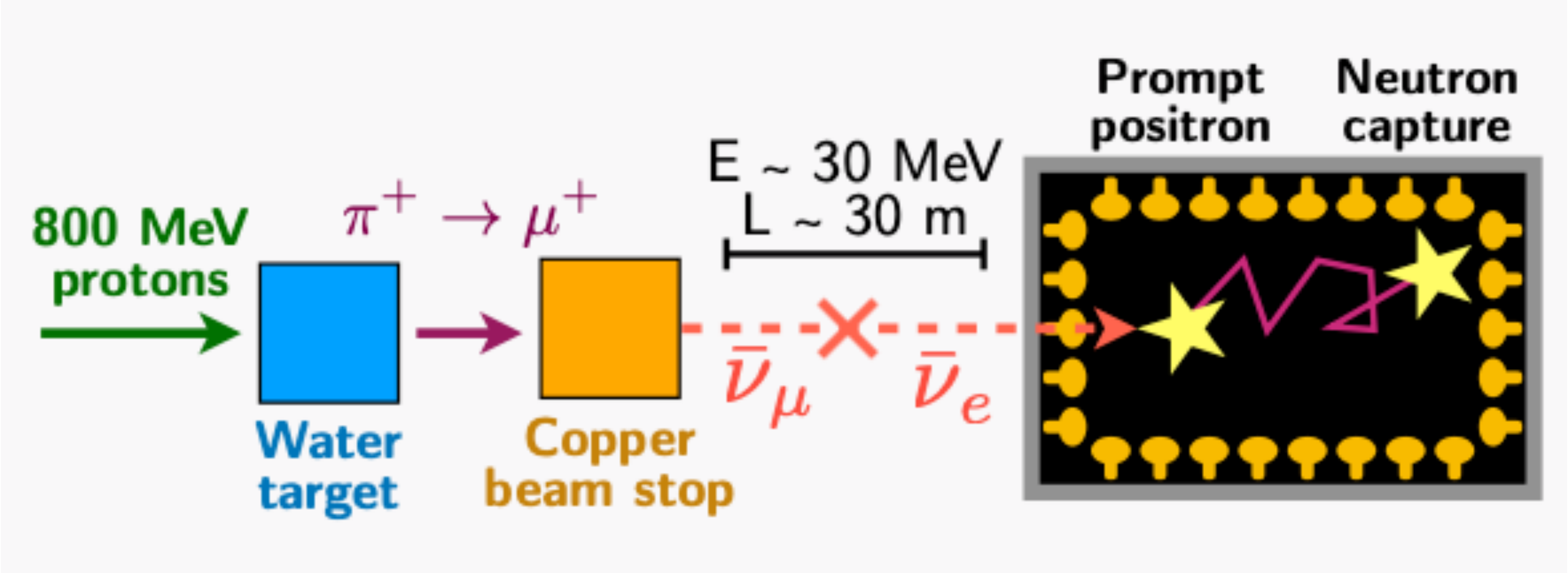
- shipped from Europe (LSNG), refurbished & upgraded
- farther away from the target, but much larger volume
- started taking neutrino data since 2021

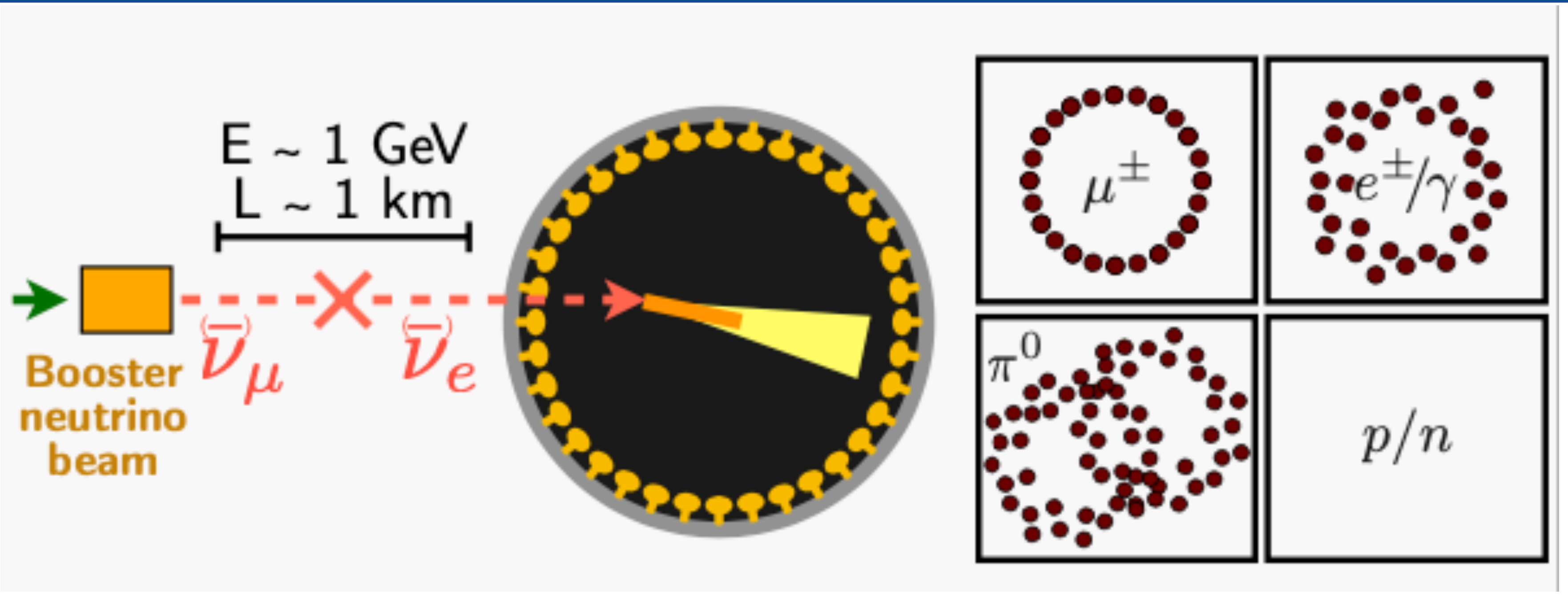
SBN program

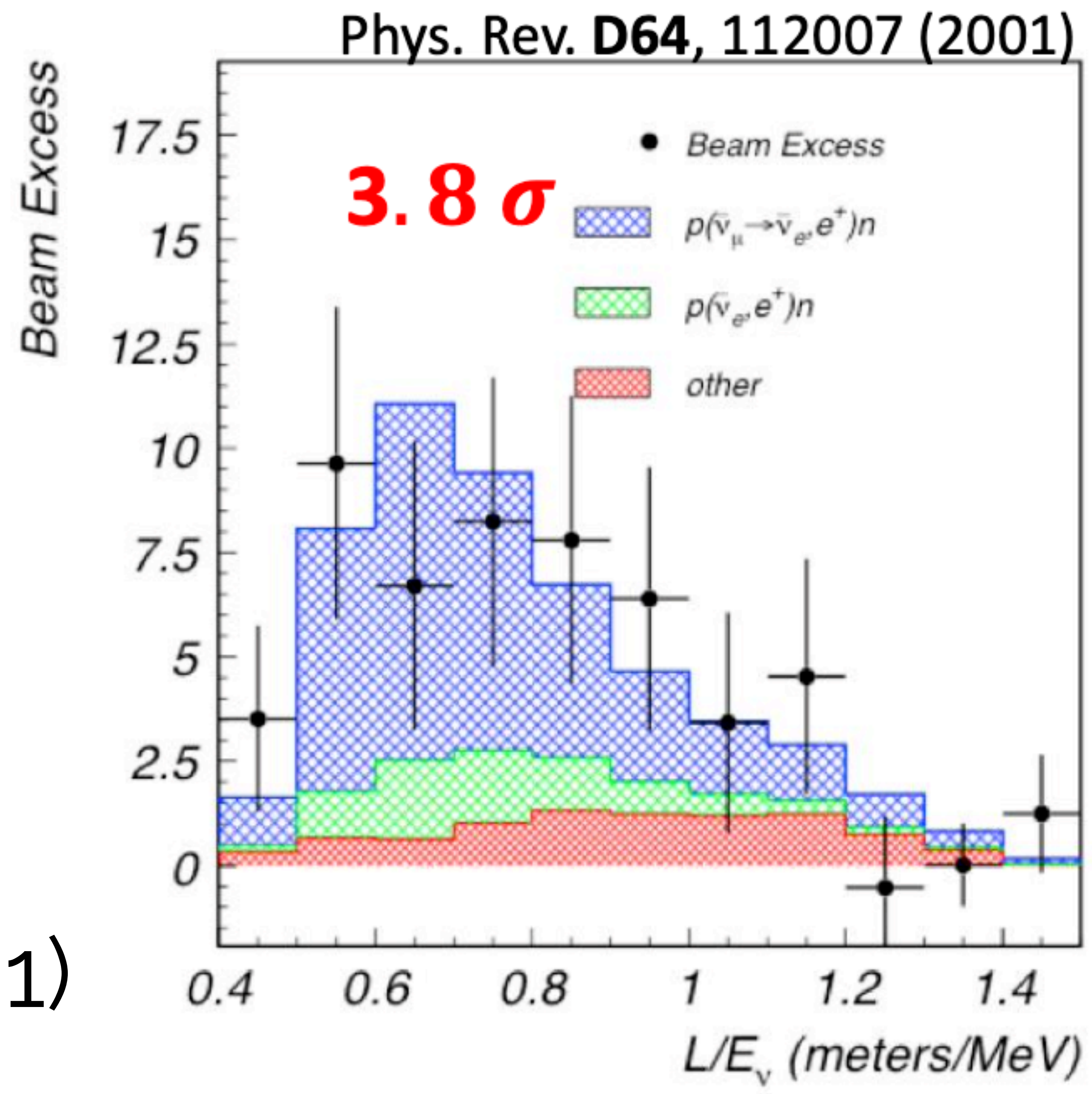


- main goal is to definitively test sterile neutrino hypothesis
  - confirm or dispute anomalies that can be explained by sterile neutrino hypothesis
- also will measure & study how neutrino interacts with argon: important input to future DUNE experiment

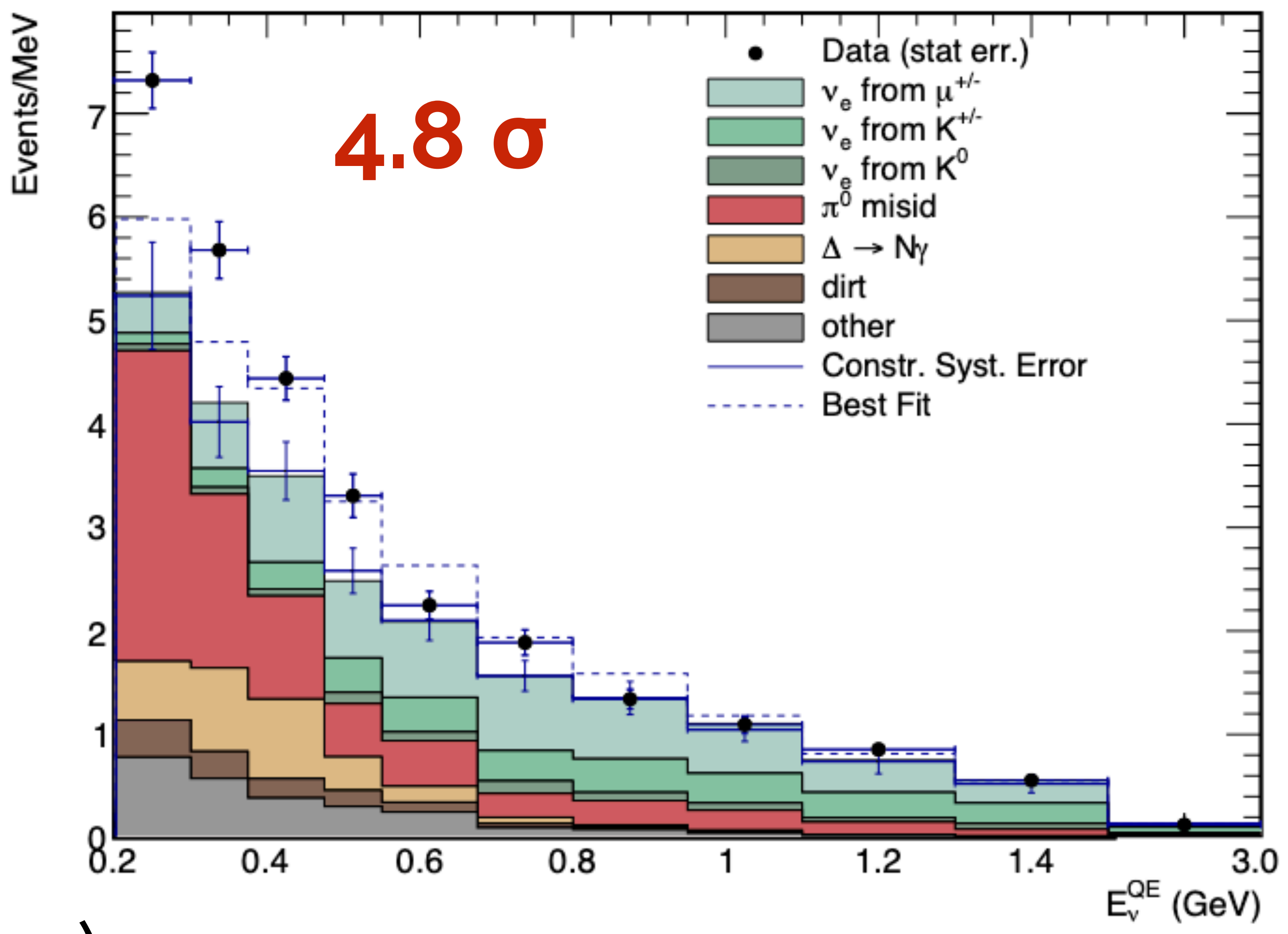
LSND & MiniBooNE anomaly







arxiv:2006.16883



- **LSND** (1990-2001)
- $\bar{\nu}_\mu$ -> $\bar{\nu}_e$ excess over background suggests evidence for oscillation at Δm^2 ~ 1eV^2

- **MiniBooNE** (1998-2020)
- measured ν_μ-> ν_e and $\bar{\nu}_\mu$-> ν_e appearance
- the excess of events at low energy

LArTPC: Liquid Argon Time Projection Chamber

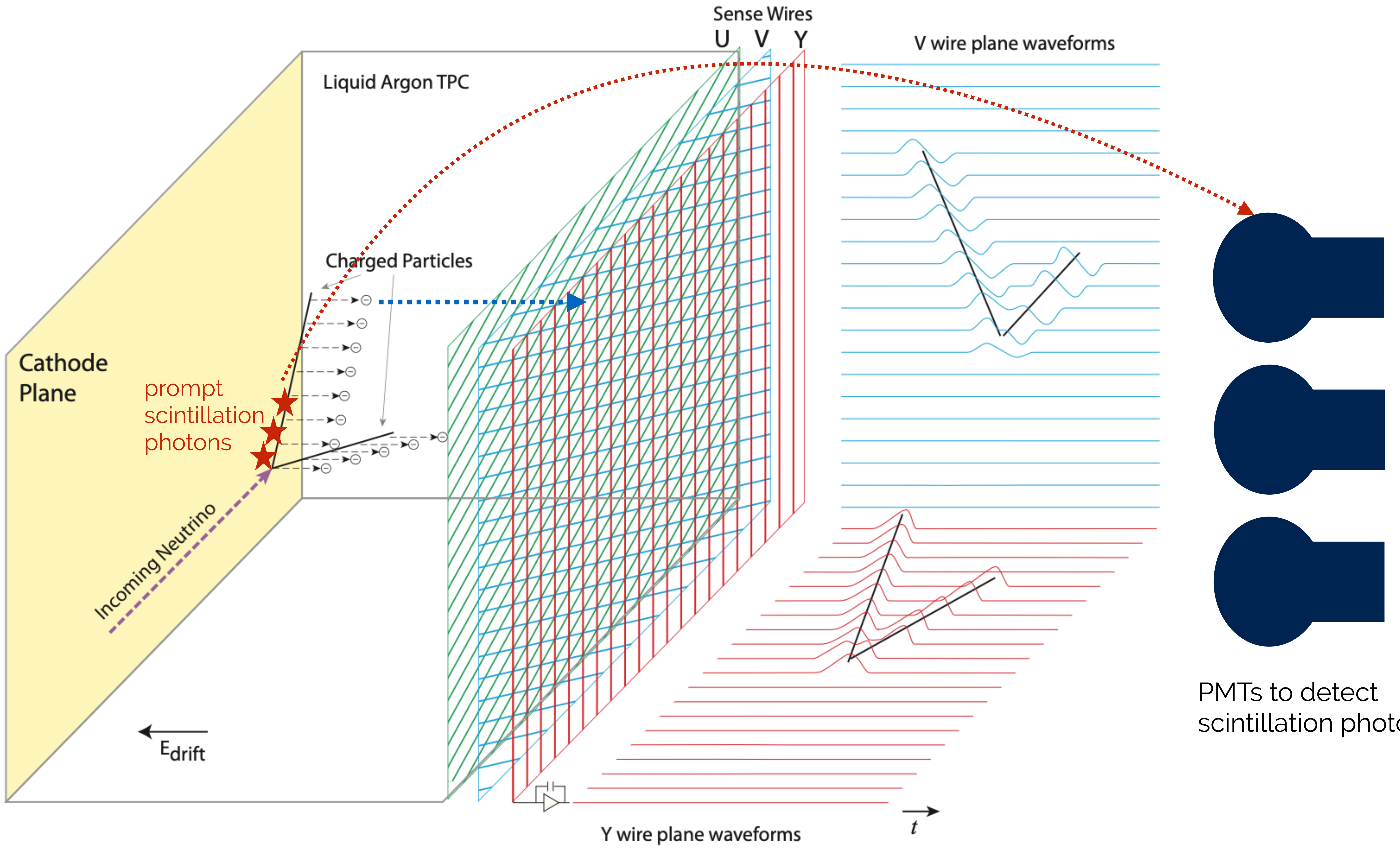
charged particle enters detector

⬇

scintillation light emitted by excited Ar, detected by PMTs

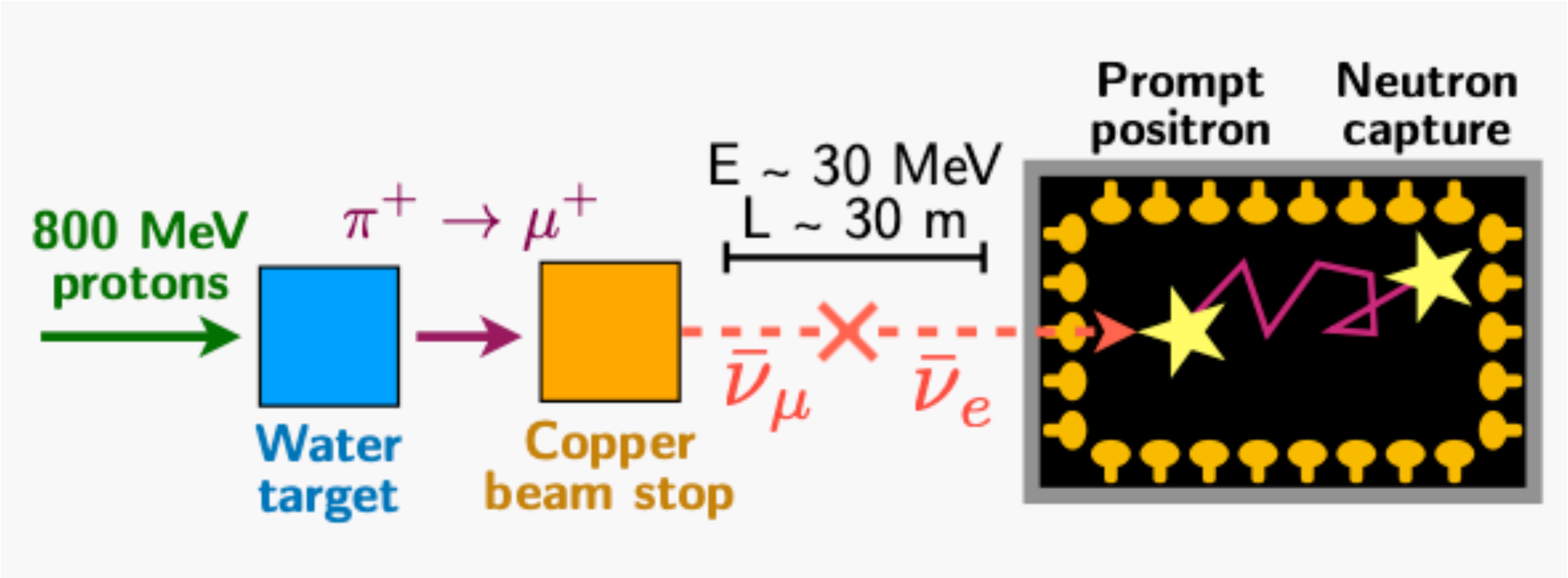
⬇

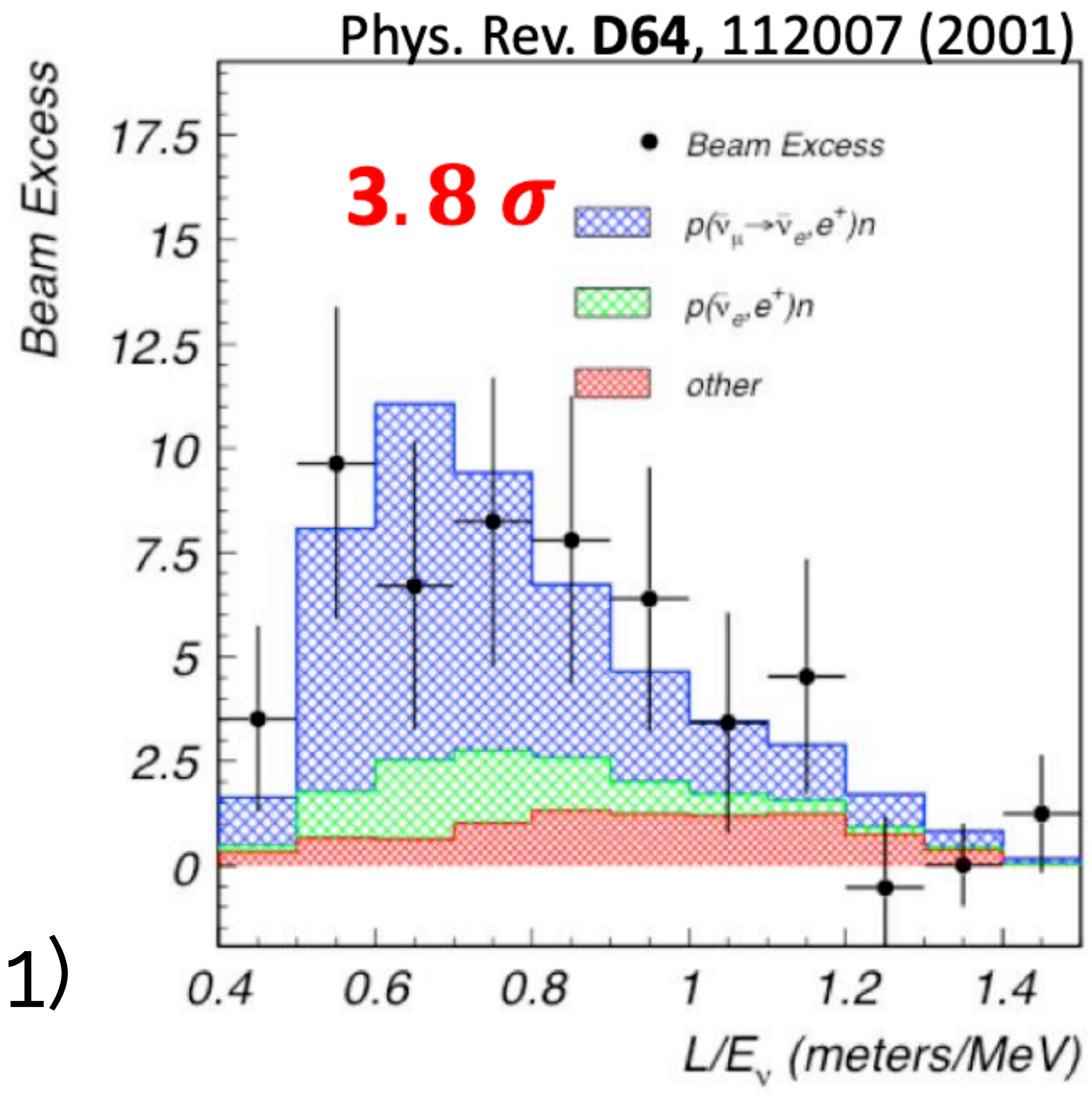
ionization electrons drift to anode plane, detected by sense wires

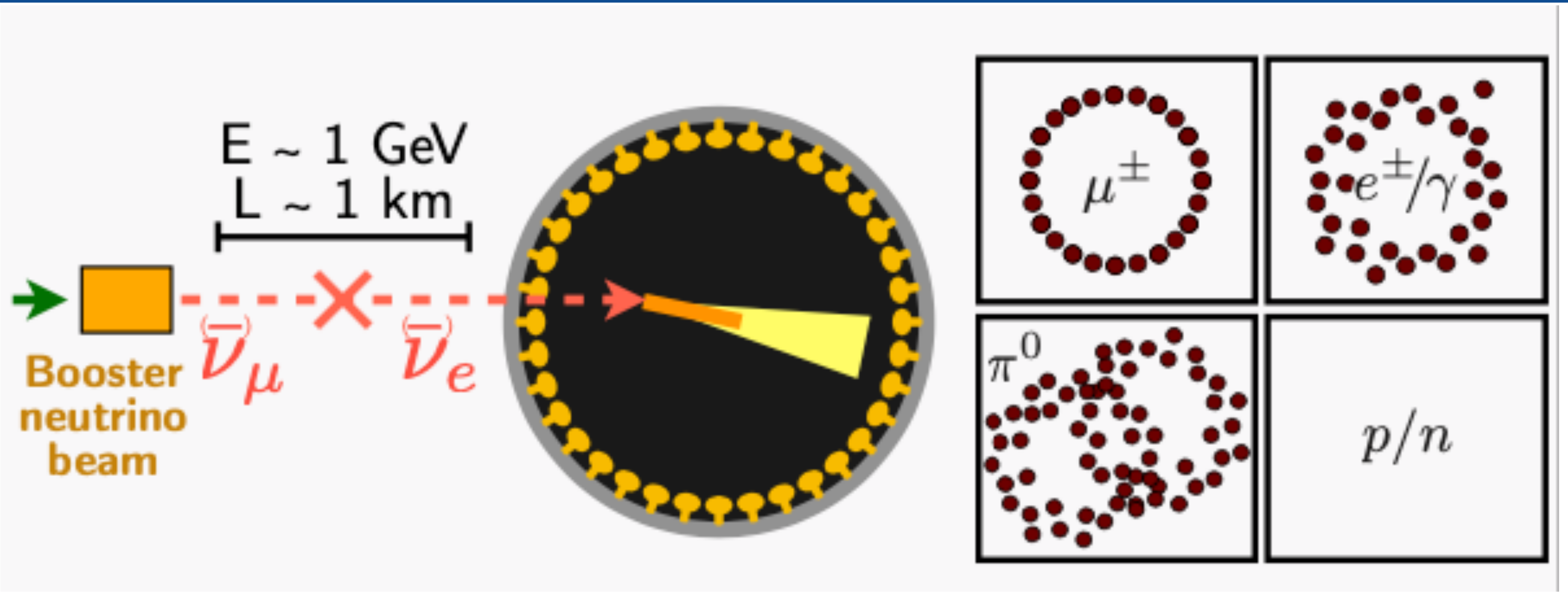


PMTs to detect scintillation photons

LSND & MiniBooNE anomaly







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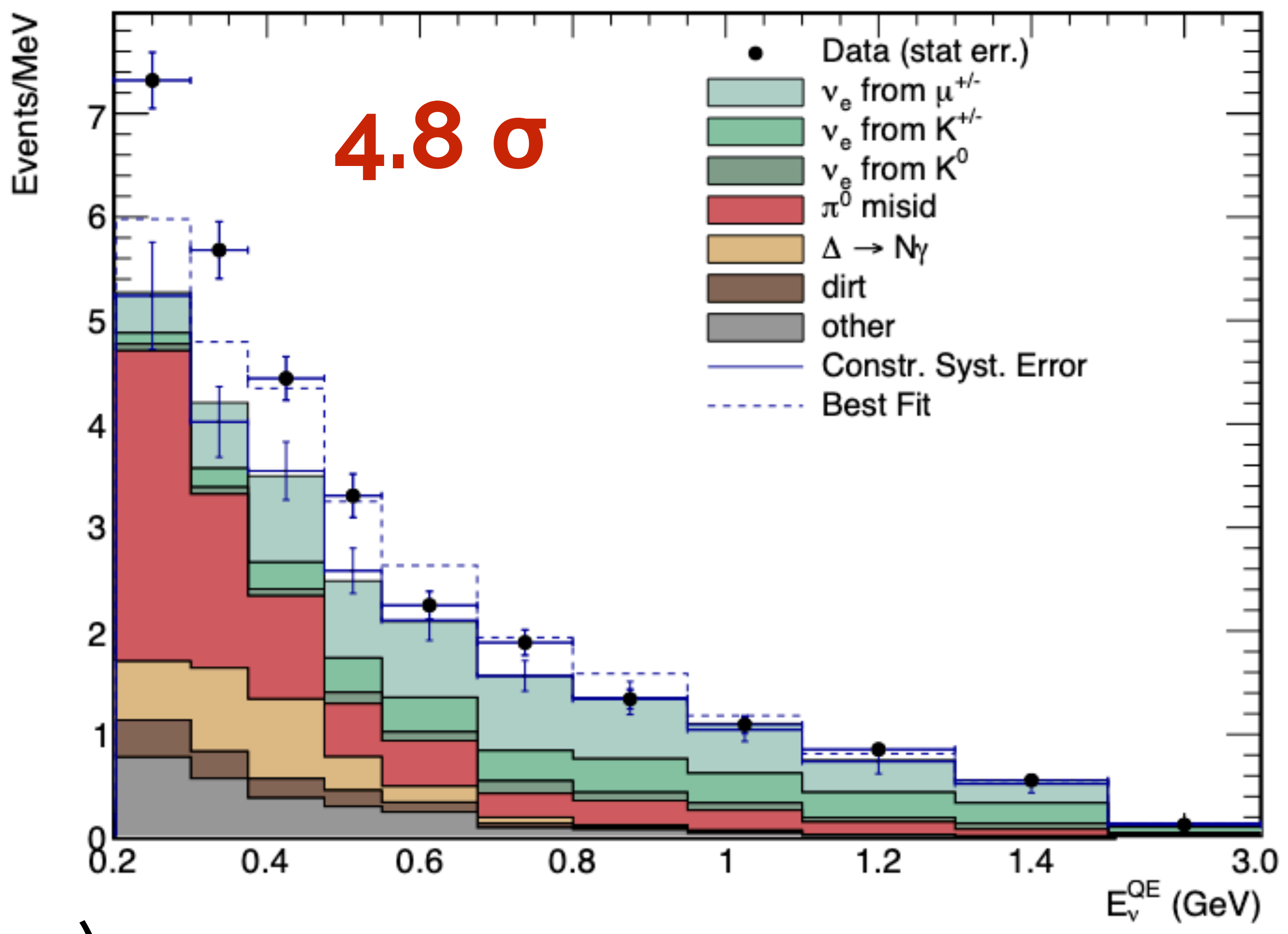

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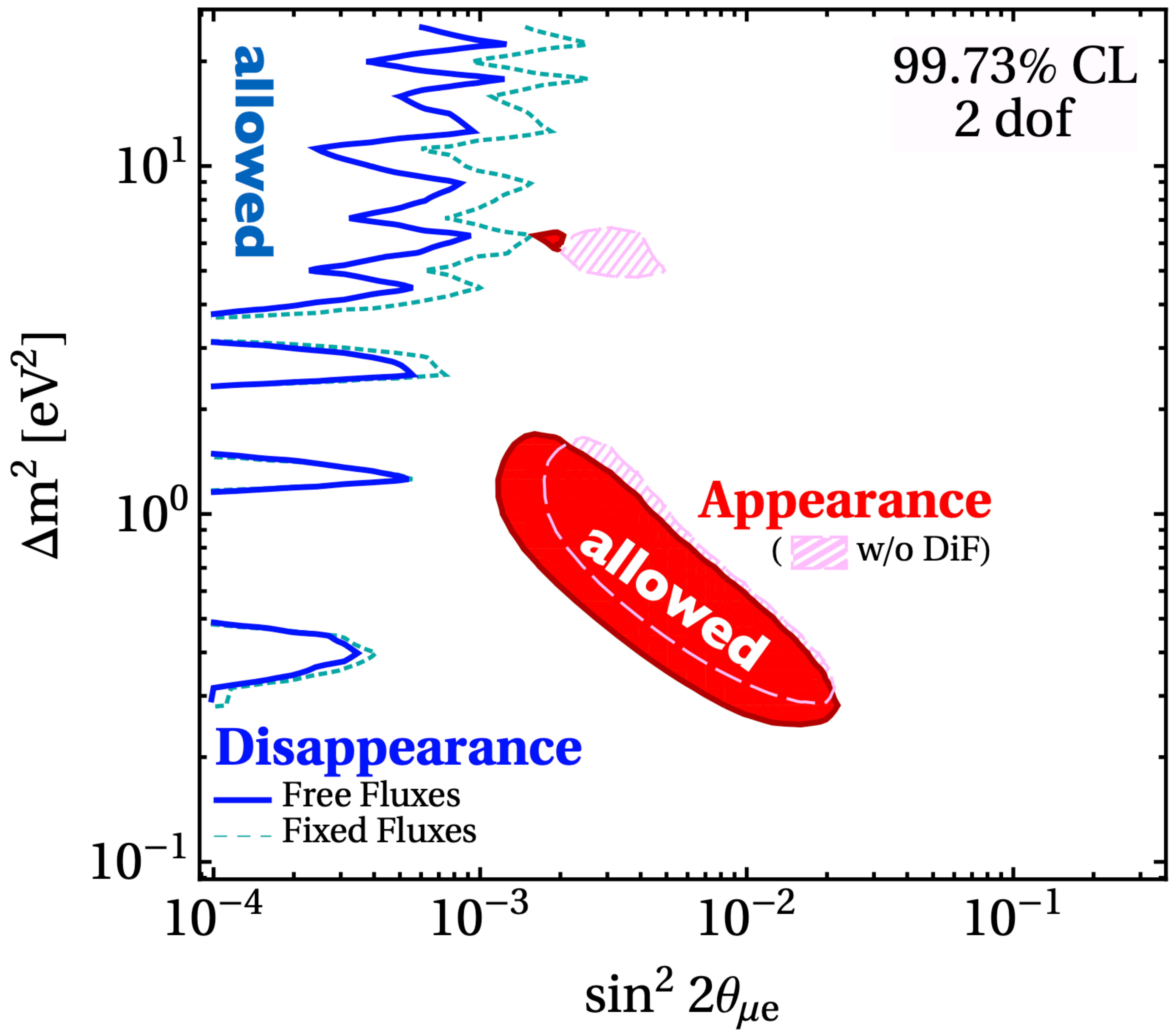
tension in global picture



- *unfortunately, it's more complicated than that…*
- significant tension between ν_e appearance and ν_e and ν_μ disappearance
- lots of different independent observations currently unexplained
- *we need to understand the anomalies better!*

From Pedro Machado's Neutrino 2020 talk: Sterile Neutrino Global Picture

short-baseline anomalies

PMNS mixing matrix

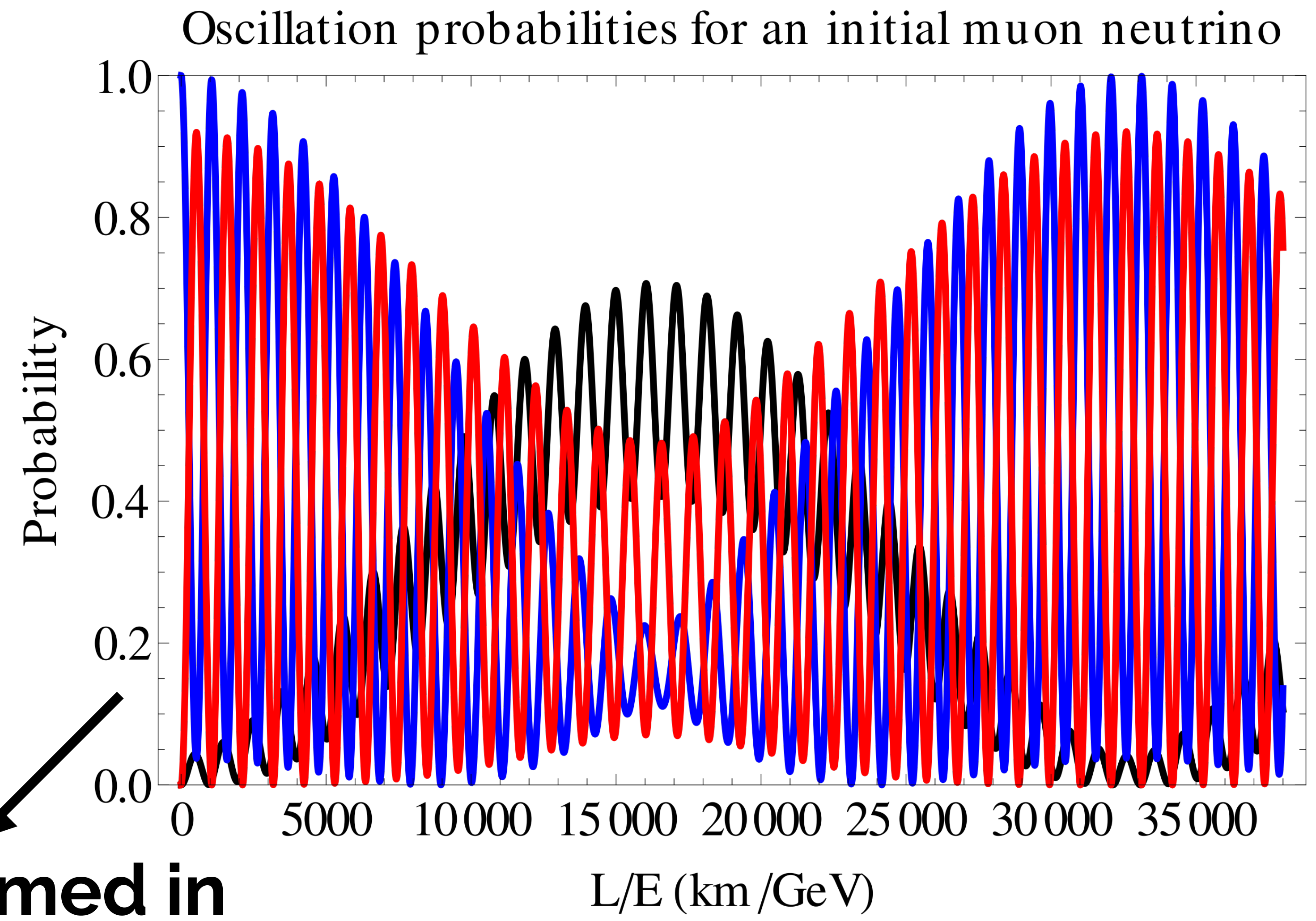
$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & +c_{23} & +s_{23} \\ 0 & -s_{23} & +c_{23} \end{pmatrix} \begin{pmatrix} +c_{13} & 0 & +s_{13}e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta} & 0 & +c_{13} \end{pmatrix} \begin{pmatrix} +c_{12} & +s_{12} & 0 \\ -s_{12} & +c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix}$$

flavor eigenstates | Atmospheric | Reactor/Accelerator | Solar | mass eigenstates

$$(c_{ij} = \cos\theta_{ij}, s_{ij} = \sin\theta_{ij})$$

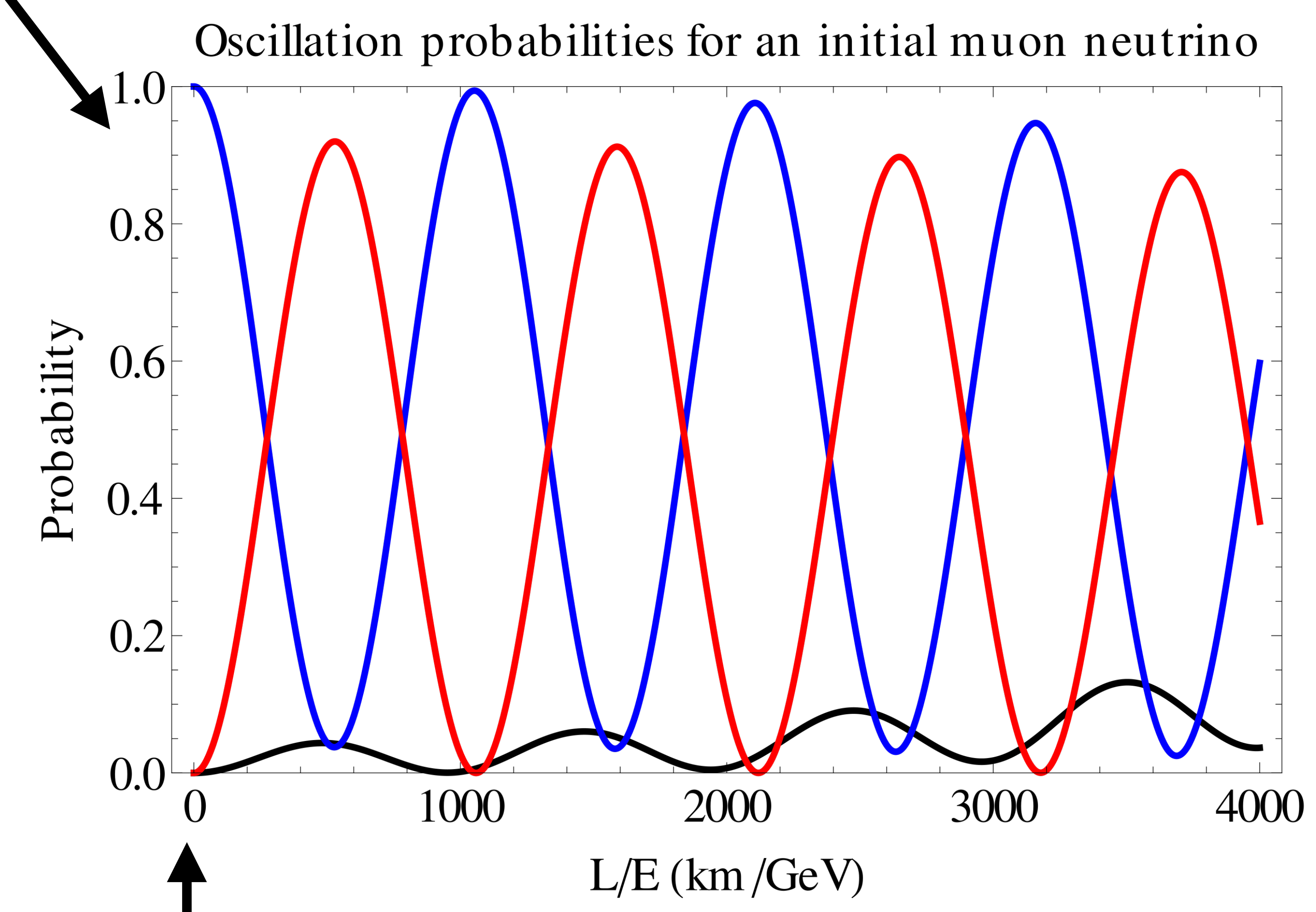
- three flavor neutrino states is well established by neutrino oscillation physics in solar, atmospheric, reactor, and accelerator domains
- puzzling collection of short-baseline anomalies: reactor anomaly, gallium anomaly, LSND & MiniBooNE anomaly
  - possible portal for new physics: the holy grail of the particle physics community
  - correctly estimating backgrounds/oscillation is important for the future neutrino program such as DUNE
  - need to resolve the anomalies -> MicroBooNE & SBN program

Neutrino Oscillations



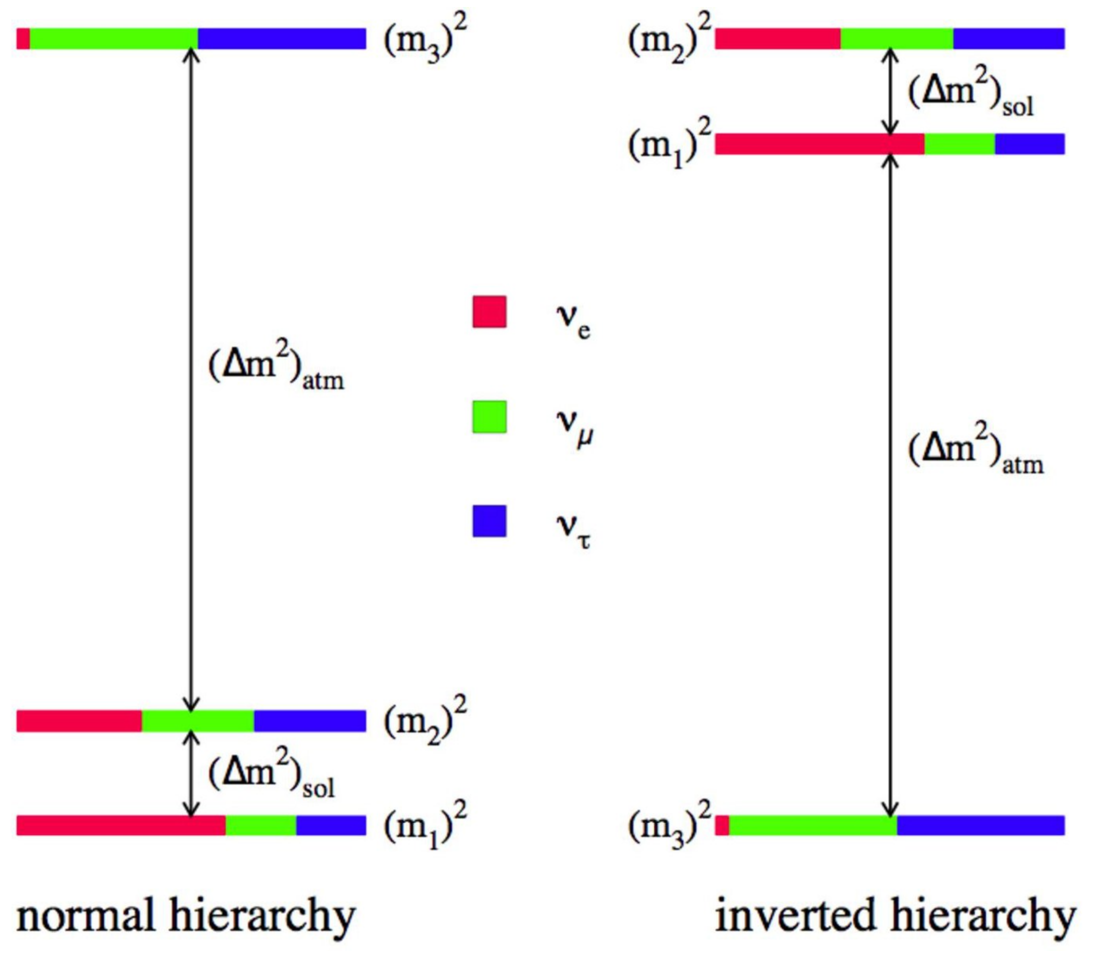
zoomed in

ν_μ
ν_e
ν_τ



MicroBooNE: ~0.5 km / ~1GeV, negligible neutrino oscillation expected

from https://en.wikipedia.org/wiki/Neutrino_oscillation



- Neutrino flavor eigenstates are not the same as the mass eigenstates
- Neutrinos generally are produced in a flavor eigenstate, which is a superposition of three mass eigenstates
- These mass eigenstates have different energies, and therefore change phase over time at different rates according to Schrodinger's equation
- This leads to neutrino oscillations when viewed in the flavor basis
- The existence of sterile neutrinos (additional mass eigenstates) would change the details of this picture