The Little Neutral One

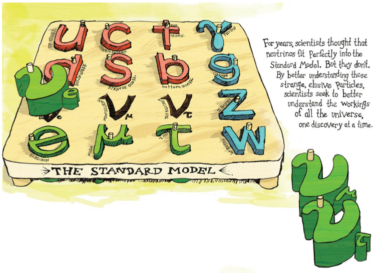
History and Overview

Physics Department Summer Lecture Series 2019

Mary Bishai
Brookhaven National Laboratory

Dec 5th, 2018

About Neutrinos



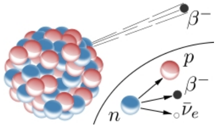
From Symmetry Magazine, Feb 2013

Cosmic Gall

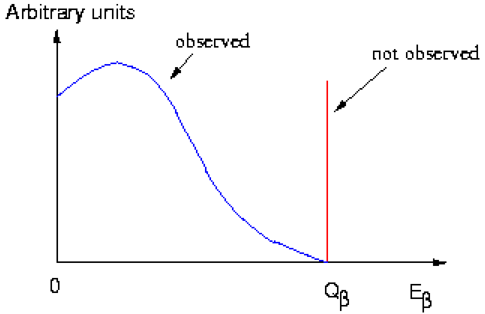
by John Updike

Neutrinos, they are very small.
They have no charge and have no mass
And do not interact at all.
The earth is just a silly ball
To them, through which they simply pass,
Like dustmaids down a drafty hall
Or photons through a sheet of glass.
They snub the most exquisite gas,
Ignore the most substantial wall,
Cold-shoulder steel and sounding brass,
Insult the stallion in his stall,
And, scorning barriers of class,
Infiltrate you and me! Like tall
And painless guillotines, they fall
Down through our heads into the grass.
At night, they enter at Nepal
And pierce the lover and his lass
From underneath the bed—you call
It wonderful; I call it crass.

NEUTRINO CONCEPTION



Before 1930's: beta decay spectrum continuous - is this energy non-conservation?



Neutrino Conception

Dec 1930: Wolfgang Pauli's letter to physicists at a workshop in Tubingen:

Wolfgang Pauli

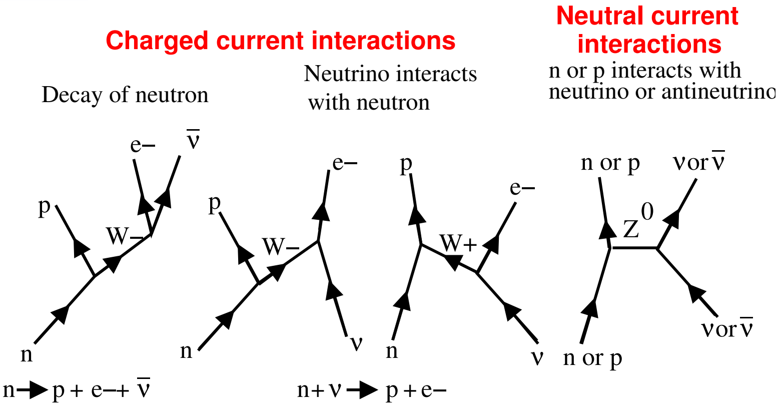
Dear Radioactive Ladies and Gentlemen,

........., I have hit upon a desparate remedy to save the "exchange theorem" of statistics and the law of conservation of energy. Namely, the possibility that there could exist in the nuclei electrically neutral particles, that I wish to call neutrons.... The mass of the neutrons should be of the same order of magnitude as the electron mass and in any event not larger than 0.01 proton masses. The continuous beta spectrum would then become understandable by the assumption that in beta decay a neutron is emitted in addition to the electron such that the sum of the energies of the neutron and the electron is constant.........
Unfortunately, I cannot appear in Tubingen personally since I am indispensable here in Zurich because of a ball on the night of 6/7 December. With my best regards to you, and also to Mr Back.
Your humble servant

. W. Pauli

The Theory of Weak Interactions

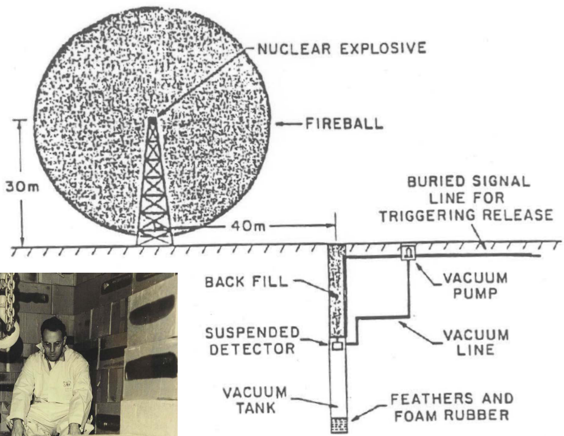
$\geq$ **1933:** **Fermi builds his theory of weak interactions and beta decay**



NEUTRINO DISCOVERY: NUCLEAR REACTORS

Finding Neutrinos.... 1^{st} attempt

1950's: Fredrick Reines, protege of Richard Feynman proposes to find neutrinos

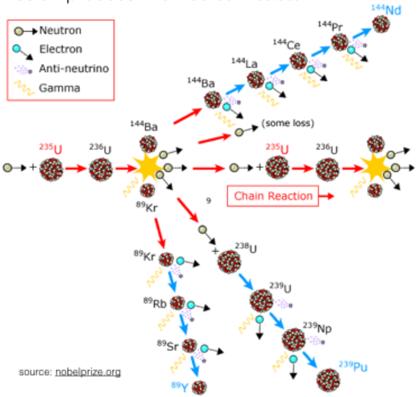


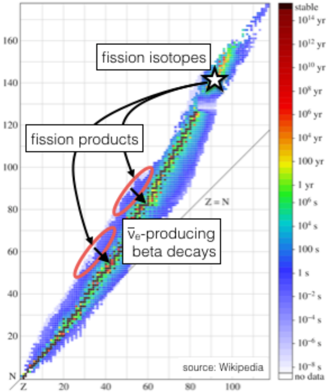
NOT APPROVED !

Finding Neutrinos.... 2^{nd} attempt

1950's: Fred Reines at Los Alamos and Clyde Cowan propose to use the Hanford nuclear reactor (1953) and the new Savannah River nuclear reactor (1955) to find neutrinos.

fission process in a nuclear reactor





BROOKHAVEN NATIONAL LABORATORY

Finding Neutrinos.... 2^{nd} attempt

1950's: Fred Reines at Los Alamos and Clyde Cowan propose to use the Hanford nuclear reactor (1953) and the new Savannah River nuclear reactor (1955) to find neutrinos.

THE UNIVERSITY OF CHICAGO
CHICAGO 37 · ILLINOIS
INSTITUTE FOR NUCLEAR STUDIES

October 8, 1952

Dr. Fred Reines
Los Alamos Scientific Laboratory
P.O. Box 1663
Los Alamos, New Mexico

Dear Fred:

Thank you for your letter of October 4th by Clyde Cowan and yourself. I was very much interested in your new plan for the detection of the neutrino. Certainly your new method should be much simpler to carry out and have the great advantage that the measurement can be repeated any number of times. I shall be very interested in seeing how your 10 cubic foot scintillation counter is going to work, but I do not know of any reason why it should not.

Good luck.

Sincerely yours,

Enrico

Enrico Fermi

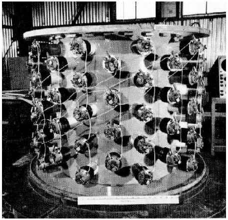
EF:vr

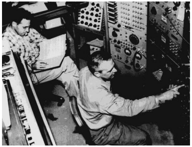
Finding Neutrinos.... 2^{nd} attempt

1950's: Fred Reines at Los Alamos and Clyde Cowan propose to use the Hanford nuclear reactor (1953) and the new Savannah River nuclear reactor (1955) to find neutrinos.
A detector filled with water with $CdCl_2$ in solution was located 11 meters from the reactor center and 12 meters underground.

The detection sequence was as follows:

1. $\bar{\nu}_e + p \rightarrow n + e^+$
2. $e^+ + e^- \rightarrow \gamma\gamma$
3. $n +^{108} Cd \rightarrow^{109} Cd* \rightarrow^{109} Cd + \gamma$ $(\tau = 5\mu s)$.





Neutrinos first detected using a nuclear reactor!

Reines shared 1995 Nobel for work on neutrino physics.

ν: A Truly Elusive Particle!

Reines and Cowan were the first to estimate the interaction strength of neutrinos. The cross-section is $\sigma \sim 10^{-43}\mathrm{cm}^2$ per nucleon (N = n or p).

$$\nu \text{ mean free path} = \frac{1}{\sigma \times \text{number of nucleons per cm}^3}$$

ν Exercise: What is the mean free path of a neutrino in lead?

(use Table of atomic and nuclear properties)

ν: A Truly Elusive Particle!

Reines and Cowan were the first to estimate the interaction strength of neutrinos. The cross-section is $\sigma \sim 10^{-43}\mathrm{cm}^2$ per nucleon (N = n or p).

$$\nu \text{ mean free path} = \frac{1}{\sigma \times \text{number of nucleons per cm}^3}$$

ν Exercise: What is the mean free path of a neutrino in lead?

(use Table of atomic and nuclear properties)

$$\begin{aligned} &= \frac{1}{10^{-43}\mathrm{cm}^2 \times 11.4\mathrm{g/cm}^3 \times 6.02 \times 10^{23}\mathrm{nucleons/g}} \\ &\approx 1.5 \times 10^{16} m \end{aligned}$$

How many light years is that? How does it compare to the distance from the sun to the moon?

ν: A Truly Elusive Particle!

Reines and Cowan were the first to estimate the interaction strength of neutrinos. The cross-section is $\sigma \sim 10^{-43}\text{cm}^2$ per nucleon (N = n or p).

$$\nu \text{ mean free path} = \frac{1}{\sigma \times \text{number of nucleons per cm}^3}$$

ν Exercise: What is the mean free path of a neutrino in lead?

(use Table of atomic and nuclear properties)

$$= \frac{1}{10^{-43}\text{cm}^2 \times 11.4\text{g/cm}^3 \times 6.02 \times 10^{23}\text{nucleons/g}}$$

$$\approx 1.5 \times 10^{16} m$$

How many light years is that? How does it compare to the distance from the sun to the moon?

$$= \text{1.6 LIGHT YEARS OF LEAD}$$

$$= 100,000 \text{ distance earth to sun}$$

A proton has a mean free path of 10cm in lead

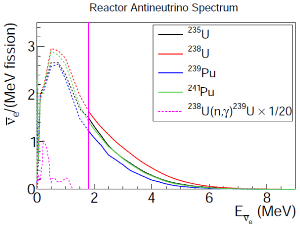
Reactor power and neutrinos

ν Exercise:

The following table shows the breakdown of energy released per fission from ^{235}U:

Fission fragment	Energy (MeV)
Fission products	175
$\langle 2.44 \rangle$ neutrons	5
γ from fission	7
γs and βs from *beta* decay	13
$\langle 6 \rangle$ neutrinos	10
Total	210

5% of a reactor's power is in neutrinos !



How many neutrinos are emitted per second from a 1 Gigawatt (thermal) reactor?

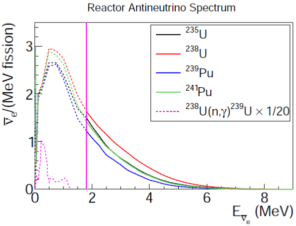
Reactor power and neutrinos

ν Exercise:

The following table shows the breakdown of energy released per fission from ^{235}U:

Fission fragment	Energy (MeV)
Fission products	175
$\langle 2.44 \rangle$ neutrons	5
γ from fission	7
γs and βs from *beta* decay	13
$\langle 6 \rangle$ neutrinos	10
Total	210

5% of a reactor's power is in neutrinos !



How many neutrinos are emitted per second from a 1 Gigawatt (thermal) reactor?

$$
\begin{aligned}
1 \times 10^9 \text{ Joules/sec} &= 6.242 \times 10^{18} \text{ GeV/sec} \\
&= 3 \times 10^{19} \text{ fissions/sec} \\
&\sim 2 \times 10^{20}\ \nu\text{/sec} \\
&= 1.6 \times 10^{13}\text{/m}^2\text{/sec at 1 km}
\end{aligned}
$$

Reactor Power and Neutrinos

ν Exercise:

Using the rate of neutrinos emitted from a reactor ($= 2 \times 10^{20}$/sec/GW) and the average cross-section of the inverse beta decay process ($\bar{\nu}_e + p \to e^+ + n$) is $\sigma = 10^{-43}\mathrm{cm}^2$/proton, what is the rate of neutrino interactions per day in a detector containing 100 tons of scintillator (CH_2) located 1km from a 1GW reactor? Note that the IBD process only happens on free protons (H)

Reactor Power and Neutrinos

ν Exercise:

Using the rate of neutrinos emitted from a reactor ($= 2 \times 10^{20}$/sec/GW) and the average cross-section of the inverse beta decay process ($\bar{\nu}_e + p \rightarrow e^+ + n$) is $\sigma = 10^{-43}\mathrm{cm}^2$/proton, what is the rate of neutrino interactions per day in a detector containing 100 tons of scintillator (CH_2) located 1km from a 1GW reactor? Note that the IBD process only happens on free protons (H)

 # interactions/day = flux (ν/cm^2/day) $\times$ σ (cm^2/p) $\times$ protons/Nucleons $\times$ Nucleons/gram $\times$ 10^8 g/100tons

Reactor Power and Neutrinos

ν Exercise:

Using the rate of neutrinos emitted from a reactor ($= 2 \times 10^{20}$/sec/GW) and the average cross-section of the inverse beta decay process ($\bar{\nu}_e + p \rightarrow e^+ + n$) is $\sigma = 10^{-43}\mathrm{cm}^2$/proton, what is the rate of neutrino interactions per day in a detector containing 100 tons of scintillator (CH_2) located 1km from a 1GW reactor? Note that the IBD process only happens on free protons (H)

 # interactions/day = flux (ν/cm^2/day) $\times$ σ (cm^2/p) $\times$ protons/Nucleons $\times$ Nucleons/gram $\times$ 10^8 g/100tons

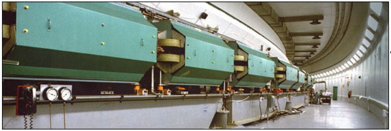
 # interactions/day = 118

Precision ν expt: need 1 GW nuclear reactor ($1B) + 100's tons

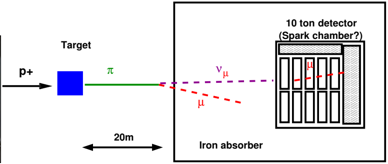
DISCOVERY OF NEUTRINO FLAVOR

Producing Neutrinos from an Accelerator

1962: Leon Lederman, Melvin Schwartz and Jack Steinberger use a proton beam from BNL's Alternating Gradient Synchrotron (AGS) to produce a beam of neutrinos using the decay $\pi \rightarrow \mu\nu_x$

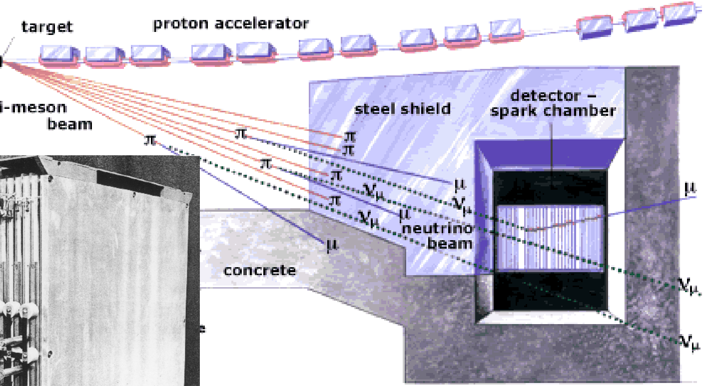


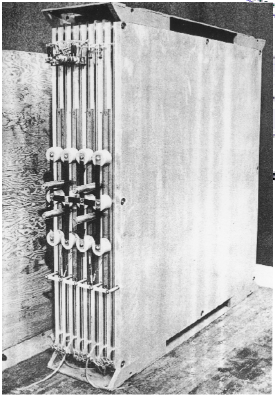
The AGS



Making ν's

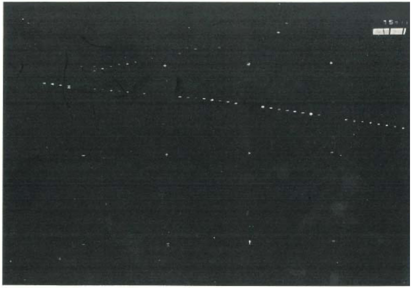
The Two-Neutrino Experiment





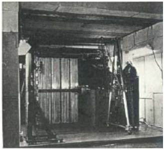
BROOKHAVEN NATIONAL LABORATORY

The Two-Neutrino Experiment



Neutrino Event

Columbia (Nevis)



BNL

Classification of "Events"

Single Tracks

p_μ < 300 MeV/c[a]	49
p_μ > 300	34
> 400	19
> 500	8
> 600	3
> 700	2
Total "single Muon Events"	34

Vertex Events

Visible Energy Released < 1 BeV	15
Visible Energy Released > 1 BeV	7
Total vertex events	22

"Shower" Events

Energy of "electron" = 200 ± 100 MeV	3
220	1
240	1
280	1
Total "shower events"[b]	6

[a] These are not included in the "event" count.

[b] The two shower events which are so located that their potential energy release in the chamber corresponds to muons of less than 300 MeV/c are not included here.

The first event!

The Two-Neutrino Experiment

Result: **40 neutrino interactions recorded in the detector, 6 of the resultant particles where identified as background and 34 identified as $\mu \Rightarrow \nu_x = \nu_\mu$**

The first successful accelerator neutrino experiment was at Brookhaven Lab.

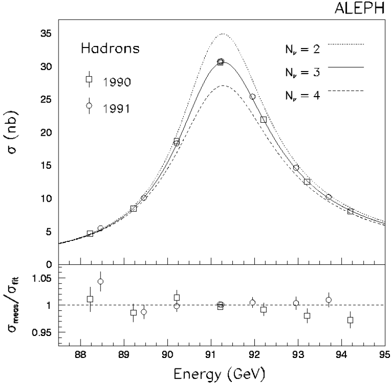
1988 NOBEL PRIZE

Number of Neutrino Flavors: Particle Colliders

1980's - 90's: The number of neutrino types is precisely determined from studies of Z^0 boson properties produced in e^+e^- colliders.

The LEP e^+e^- collider at CERN, Switzerland

The 27km LEP ring was reused to build the Large Hadron Collider



$$N_\nu = 2.984 \pm 0.008$$

Neutrinos from Accelerators

To produce neutrinos from accelerators

$p^{+} + A \rightarrow \pi^{\pm} + X, \quad \pi^{\pm} \rightarrow \mu^{\pm} + \nu_{\mu}/\bar{\nu}_{\mu}$

where A = Carbon (Graphite), Berillyium, Tungsten, X is other particles

ν Exercise: The Main Injector accelerator at Fermilab produces 4.86×10^{13} 120 GeV protons in a 10 microsecond pulse every 1.33 seconds to the NuMI beamline. What is the average power of the proton beam delivered in megawatts?

Neutrinos from Accelerators

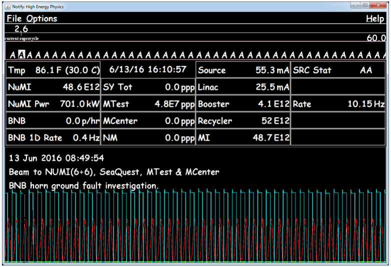
To produce neutrinos from accelerators

$p^+ + A \rightarrow \pi^\pm + X, \quad \pi^\pm \rightarrow \mu^\pm + \nu_\mu/\bar{\nu}_\mu$

where A = Carbon (Graphite), Berillyium, Tungsten, X is other particles

ν Exercise: The Main Injector accelerator at Fermilab produces 4.86×10^{13} 120 GeV protons in a 10 microsecond pulse every 1.33 seconds to the NuMI beamline. What is the average power of the proton beam delivered in megawatts?

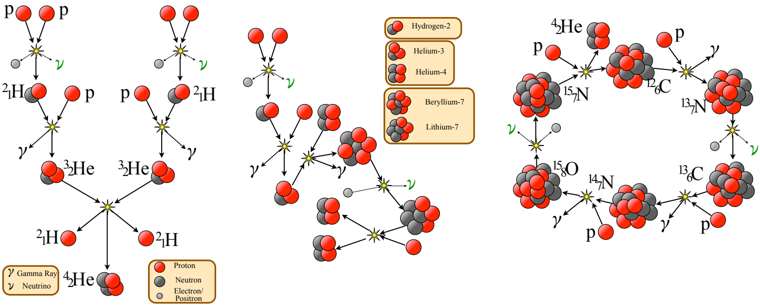
Power = 120 GeV $\times$ 4.86 10^{13} protons $\times$ 1.6 10^{-10} Joules/GeV $\times$ 1/1.33s = 702 kW



NEUTRINO MIXING: SOLAR

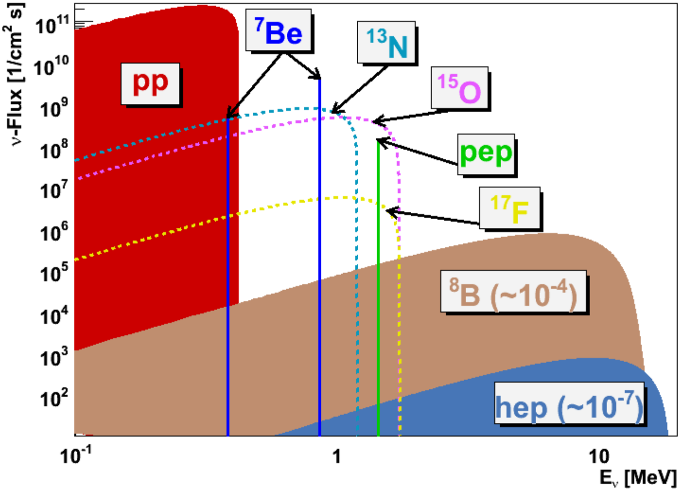
Solar Neutrinos

Fusion of nuclei in the Sun produces solar energy and neutrinos



Solar Neutrinos

Fusion of nuclei in the Sun produces solar energy and neutrinos

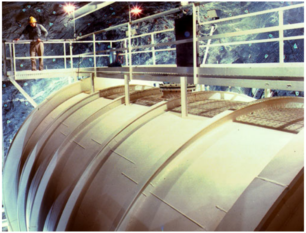


The Homestake Experiment

1967: Ray Davis from BNL installs a large detector, containing 615 tons of tetrachloroethylene (cleaning fluid), 1.6km underground in Homestake mine, SD.

1. $\nu_e^{sun} + ^{37}CL \rightarrow e^- + ^{37}Ar$, $\tau(^{37}Ar) = 35$ days.
2. Number of Ar atoms $\approx$ number of ν_e^{sun} interactions.

Ray Davis



Results: 1969 - 1993 Measured 2.5 ± 0.2 SNU (1 SNU = 1 neutrino interaction per second for 10^{36} target atoms) while theory predicts 8 SNU. This is a ν_e^{sun} deficit of 69%.

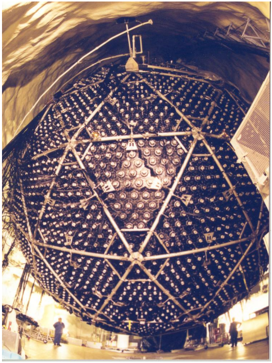
Where did the suns ν_e's go?

RAY DAVIS SHARES 2002 NOBEL PRIZE

SNO Experiment: Solar ν Measurments

$1 \leftrightarrow 2$ mix ing

2001-02: Sudbury Neutrino Observatory. Water Čerenkov detector with 1 kT heavy water (0.5 B$ worth on loan from Atomic Energy of Canada Ltd.) located 2Km below ground in INCO's Creighton nickel mine near Sudbury, Ontario. Can detect the following ν^{sun} interactions:



$$1)\ \nu_e + d \rightarrow e^- + p + p \ \text{(CC)}.$$
$$2)\ \nu_x + d \rightarrow p + n + \nu_x \ \text{(NC)}.$$
$$3)\ \nu_x + e^- \rightarrow e^- + \nu_x \ \text{(ES)}.$$

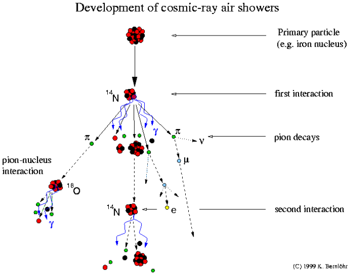
SNO measured:

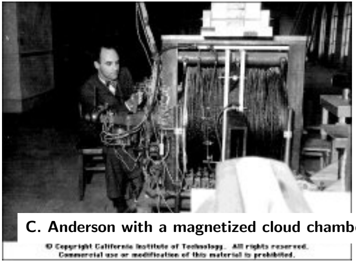
$$\phi^{CC}_{SNO}(\nu_e) = 1.75 \pm 0.07(\text{stat})^{+0.12}_{-0.11}(\text{sys.}) \pm 0.05(\text{theor}) \times 10^6 cm^{-2}s^{-1}$$
$$\phi^{ES}_{SNO}(\nu_x) = 2.39 \pm 0.34(\text{stat})^{+0.16}_{-0.14}(\text{sys.}) \pm \times 10^6 cm^{-2}s^{-1}$$
$$\phi^{NC}_{SNO}(\nu_x) = 5.09 \pm 0.44(\text{stat})^{+0.46}_{-0.43}(\text{sys.}) \pm \times 10^6 cm^{-2}s^{-1}$$

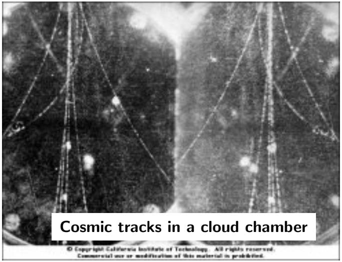
All the solar ν's are there but ν_e appears as ν_x!

Discovery of the Muon (μ)

1936: Carl Andersen, Seth Neddermeyer observed an unknown charged particle in cosmic rays with mass between that of the electron and the proton - called it the μ meson (now muons).




C. Anderson with a magnetized cloud chamber


Cosmic tracks in a cloud chamber

The Lepton Family and Flavors

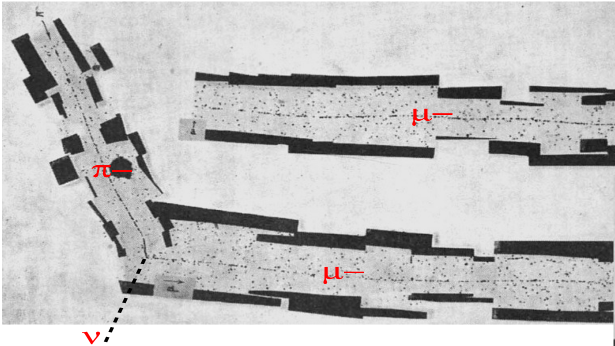
The muon and the electron are *different "flavors" of the same family of elementary particles called leptons*.

Generation	I	II	III
Lepton	e^-	μ	τ
Mass (GeV)	0.000511	0.1057	1.78
Lifetime (sec)	stable	2.2×10^{-6}	2.9×10^{-13}

Neutrinos are neutral leptons. **Do ν's have flavor too?**

Discovery of the Pion: 1947

Cecil Powell, Cesar Lattes and Giusseppe Occhialini collect emulsion photos of cosmic rays on top of mountains and aboard high altitude RAF flights. A charged particle is found decaying to a muon:



$\text{mass}_{\pi^-} = 0.1396\ \text{GeV/c}^2$, $\tau = 26$ ns.

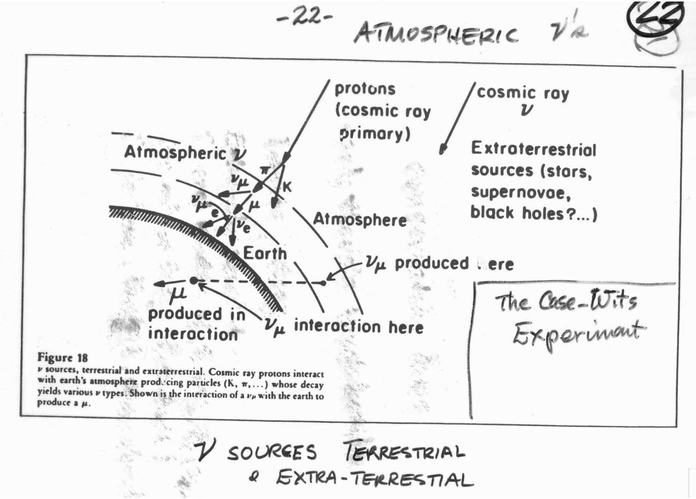
Pions are composed of $q\bar{q}'$ pairs. Weak decays produce neutrinos like in beta decay.

1950 Nobel prize for Powell

Cesar Lattes (born 1924, Curitiba, Brazil)

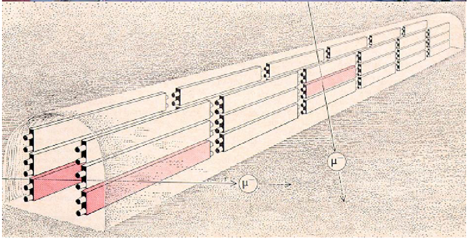
Proposal to find Atmospheric Neutrinos

Slide to find atmospheri neutrinos by Fred Reines (Case Western Institute):



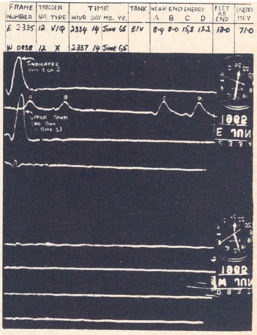
The CWI-SAND Experiment

1964: The Case Western Institute-South Africa Neutrino Detector (CWI-SAND) and a search for atmospheric ν_μ at the East Rand gold mine in South Africa at 3585m depth

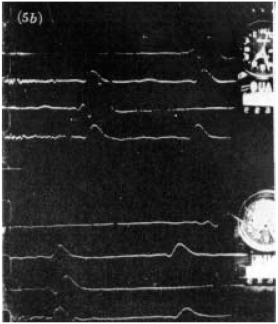


The CWI-SAND Experiment

1964: The Case Western Institute-South Africa Neutrino Detector (CWI-SAND) and a search for atmospheric ν_μ at the East Rand gold mine in South Africa at 3585m depth



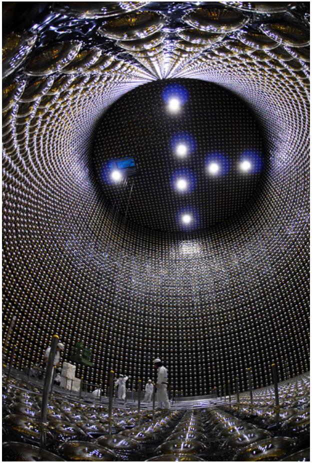
Downward-going Muon (background)



Horizontal Muon (neutrino signal)

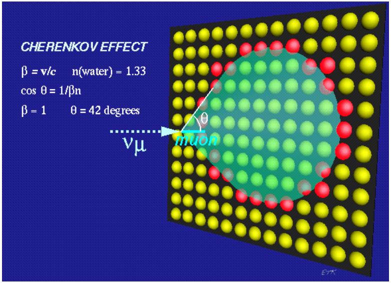
Detection of the first neutrino in nature!

The Super-Kamiokande Experiment. Kamioka Mine, Japan

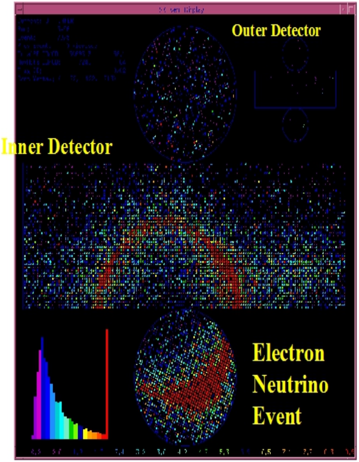


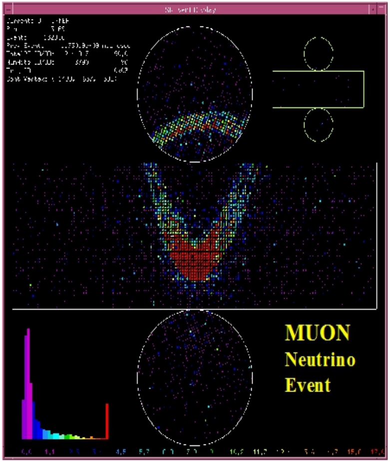
50kT double layered tank of ultra pure water surrounded by 11,146 20" diameter photomultiplier tubes.

Atmospheric neutrinos are identified by using CC interaction $\nu_{\mu,e} \rightarrow e^{\pm}, \mu^{\pm}X$. The lepton produces Cherenkov light as it goes through the detector:

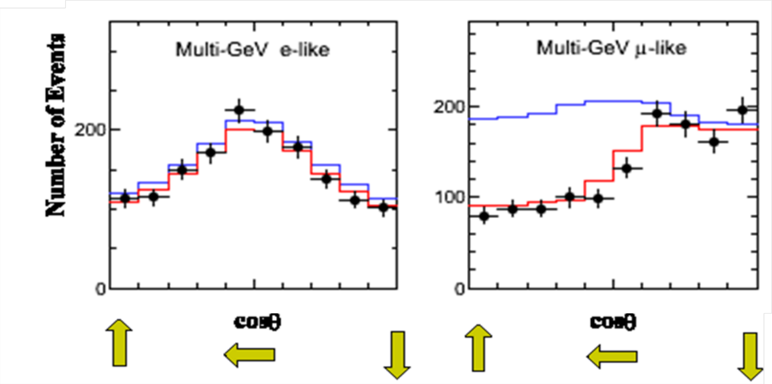


The Super-Kamiokande Experiment. Kamioka Mine, Japan





More Disappearing Neutrinos!!



All the ν_e are there! But what happened to the ν_μ ??

Quantum Mechanics

1924: Louis-Victor-Pierre-Raymond, 7th duc de Broglie proposes in his doctoral thesis that all matter has wave-like and particle-like properties.

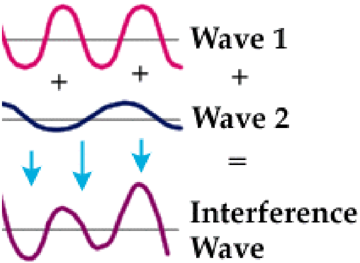
For highly relativistic particles : energy ≈ momentum

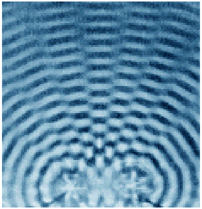
De Broglie

$$\textbf{Wavelength (nm)} \approx \frac{\mathbf{1.24 \times 10^{-6}\ GeV.nm}}{\textbf{Energy (GeV)}}$$

Neutrino Mixing

1957,1967: B. Pontecorvo proposes that neutrinos of a particular flavor are a mix of quantum states with different masses that propagate with different phases:





The interference of water waves coming from two sources.

The inteference pattern depends on the difference in masses

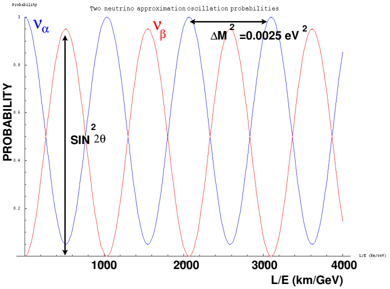
Neutrino Mixing ⇒ Oscillations

$$\begin{pmatrix} \nu_a \\ \nu_b \end{pmatrix} = \begin{pmatrix} \cos(\theta) & \sin(\theta) \\ -\sin(\theta) & \cos(\theta) \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix}$$

$$\begin{aligned} \nu_a(t) &= \cos(\theta)\nu_1(t) + \sin(\theta)\nu_2(t) \\ P(\nu_a \to \nu_b) &= |<\nu_b|\nu_a(t)>|^2 \\ &= \sin^2(\theta)\cos^2(\theta)|e^{-iE_2t} - e^{-iE_1t}|^2 \end{aligned}$$

$$P(\nu_a \to \nu_b) = \sin^2 2\theta \sin^2 \frac{1.27\Delta m^2_{21} L}{E}$$

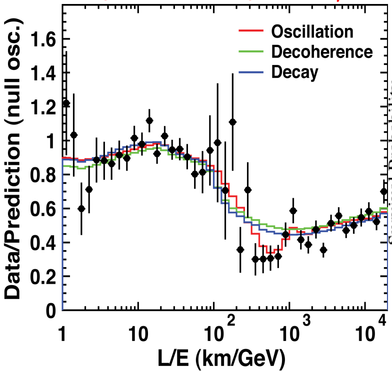
where $\Delta m^2_{21} = (m_2^2 - m_1^2)$ in eV^2, L (km) and E (GeV).



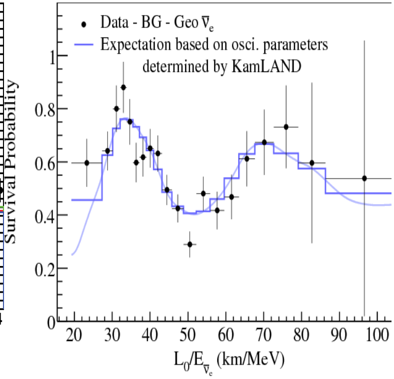
Observation of oscillations implies non-zero mass eigenstates

Two Different Mass Scales!

Super-K, atmospheric ν_μ



KamLAND, reactor $\bar{\nu}_e$



Global fit 2013:

$$\Delta m^2_{\mathrm{atm}} = 2.43^{+0.06}_{-0.10} \times 10^{-3}\ \mathrm{eV}^2$$

$$\sin^2\theta_{\mathrm{atm}} = 0.386^{+0.24}_{-0.21}$$

Atmospheric L/E $\sim$ 500 km/GeV

Global fit 2013:

$$\Delta m^2_{\mathrm{solar}} = 7.54^{+0.26}_{-0.22} \times 10^{-5}\ \mathrm{eV}^2$$

$$\sin^2\theta_{\mathrm{solar}} = 0.307^{+0.18}_{-0.16}$$

Solar L/E $\sim$ 15,000 km/GeV

2015 Nobel Prize

Takaaki Kajita
University of Tokyo, Japan
(SuperKamiokande)

Arthur B. MacDonald
Queens University, Canada
(SNO)

The Nobel Prize in Physics 2015 was awarded jointly to Takaaki Kajita and Arthur B. McDonald *"for the discovery of neutrino oscillations, which shows that neutrinos have mass"*

The Implications of 3-Neutrino Mixing

We know now of 3 flavours of neutrinos: The 3 flavour PMNS mixing matrix was developed in 1962 by Maki-Nakagawa-Sakata based on Pontecorvo's earlier work:

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} = \underbrace{\begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix}}_{U_{PMNS}} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix}$$

Commonly paramterized as $U_{\text{PMNS}} =$

$$\underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\nu_\mu \text{ disappearance}} \underbrace{\begin{pmatrix} c_{13} & 0 & e^{i\delta_{\text{CP}}} s_{13} \\ 0 & 1 & 0 \\ -e^{i\delta_{\text{CP}}} s_{13} & 0 & c_{13} \end{pmatrix}}_{\nu_\mu \to \nu_e,\ \text{reactor } \bar{\nu_e} \text{ disappear}} \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{solar } \nu_e,\ \bar{\nu_e} \text{ disappear}}$$

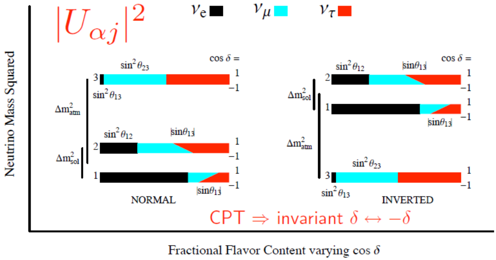
where $c_{ij} = \cos\theta_{ij}$ and $s_{ij} = \sin\theta_{ij}$.

$\sin^2\theta_{13}$: Amount of ν_e in ν_3 **$\tan^2\theta_{23}$: Ratio of $\frac{\nu_\mu}{\nu_\tau}$ in ν_3**

$\tan^2\theta_{12}$: $\frac{\text{Amount of } \nu_e \text{ in } \nu_2}{\text{Amount of } \nu_e \text{ in } \nu_1}$

There are 3 quantum states mixing $\Rightarrow$ there is an overall phase: δ_{CP}.

Neutrino Mixing: 3 flavors, 3 amplitudes, 2 mass scales, 1 phase

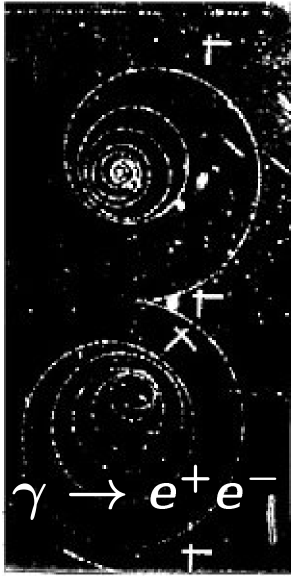


- What is the neutrino mass hierarchy? ($\delta m^2_{31} \equiv m^2_3 - m^2_1 > 0$)
- Is ν_3 mostly ν_μ or ν_τ ? $(\theta_{23} < \pi/4 \ \text{ or } > \pi/4)$
- Is CP Violated in Neutrino Oscillations? ($\delta \neq 0, \pi$)

Charge-Parity Symmetry

Charge-parity symmetry: laws of physics are the same if a particle is interchanged with its anti-particle and left and right are swapped.

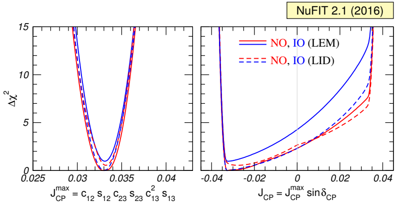
A violation of CP $\Rightarrow$ matter/anti-matter asymmetry.



CP Violation in PMNS (leptons) and CKM (quarks)

In 3-flavor mixing the degree of CP violation is determined by the Jarlskog invariant:

$$J_{CP}^{\mathrm{PMNS}} \equiv \frac{1}{8} \sin 2\theta_{12} \sin 2\theta_{13} \sin 2\theta_{23} \cos \theta_{13} \sin \delta_{\mathrm{CP}}.$$



(JHEP 11 (2014) 052, arXiv:1409.5439)

Given the current best-fit values of the ν mixing angles :

$$J_{CP}^{\mathrm{PMNS}} \approx 3 \times 10^{-2} \sin \delta_{\mathrm{CP}}.$$

For CKM (mixing among the 3 quark generations):

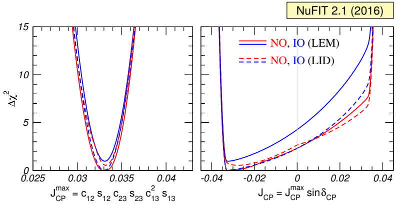
$$J_{CP}^{\mathrm{CKM}} \approx 3 \times 10^{-5},$$

despite the large value of $\delta_{CP}^{\mathrm{CKM}} \approx 70^\circ$.

CP Violation in PMNS (leptons) and CKM (quarks)

In 3-flavor mixing the degree of CP violation is determined by the Jarlskog invariant:

$$J_{CP}^{\mathrm{PMNS}} \equiv \frac{1}{8} \sin 2\theta_{12} \sin 2\theta_{13} \sin 2\theta_{23} \cos \theta_{13} \sin \delta_{\mathrm{CP}}.$$



(JHEP 11 (2014) 052, arXiv:1409.5439)

Given the current best-fit values of the ν mixing angles :

$$J_{CP}^{\mathrm{PMNS}} \approx 3 \times 10^{-2} \sin \delta_{\mathrm{CP}}.$$

For CKM (mixing among the 3 quark generations):

$$J_{CP}^{\mathrm{CKM}} \approx 3 \times 10^{-5},$$

despite the large value of $\delta_{CP}^{\mathrm{CKM}} \approx 70°$.

$\nu_\mu \to \nu_e$ Oscillations in the 3-flavor ν SM

In the ν 3-flavor model matter/anti-matter asymmetries in neutrinos are best probed using $\nu_\mu/\bar{\nu}_\mu \to \nu_e/\bar{\nu}_e$ oscillations (or vice versa). With terms up to second order in $\alpha \equiv \Delta m^2_{21}/\Delta m^2_{31} = 0.03$ and $\sin^2\theta_{13} = 0.02$, (M. Freund. Phys. Rev. D 64, 053003):

$$P(\nu_\mu \to \nu_e) \cong P(\nu_e \to \nu_\mu) \cong \underbrace{P_0}_{\theta_{13}} + \underbrace{P_{\sin\delta}}_{\text{CP violating}} + \underbrace{P_{\cos\delta}}_{\text{CP conserving}} + \underbrace{P_3}_{\text{solar oscillation}}$$

where **for oscillations in vacuum:**

$$P_0 = \sin^2\theta_{23}\sin^2 2\theta_{13}\sin^2(\Delta),$$

$$P_{\sin\delta} = \alpha\, 8J_{cp}\sin^3(\Delta),$$

$$P_{\cos\delta} = \alpha\, 8J_{cp}\cot\delta_{\rm CP}\cos\Delta\sin^2(\Delta),$$

$$P_3 = \alpha^2\cos^2\theta_{23}\sin^2 2\theta_{12}\sin^2(\Delta),$$

where $\Delta = 1.27\Delta m^2_{31}(eV^2)L(km)/E(GeV)$

$$\text{For } \bar{\nu}_\mu \to \bar{\nu}_e,\ \underbrace{P_{\sin\delta} \to -P_{\sin\delta}}_{\text{CP asymmetry}},$$

$\nu_\mu \to \nu_e$ Oscillations in the 3-flavor ν SM

In the ν 3-flavor model matter/anti-matter asymmetries in neutrinos are best probed using $\nu_\mu/\bar{\nu}_\mu \to \nu_e/\bar{\nu}_e$ oscillations (or vice versa). With terms up to second order in $\alpha \equiv \Delta m^2_{21}/\Delta m^2_{31} = 0.03$ and $\sin^2\theta_{13} = 0.02$, (M. Freund. Phys. Rev. D 64, 053003):

$$P(\nu_\mu \to \nu_e) \cong P(\nu_e \to \nu_\mu) \cong \underbrace{P_0}_{\theta_{13}} + \underbrace{P_{\sin\delta}}_{\text{CP violating}} + \underbrace{P_{\cos\delta}}_{\text{CP conserving}} + \underbrace{P_3}_{\text{solar oscillation}}$$

where **for oscillations in matter with constant density:**

$$\begin{aligned}
P_0 &= \sin^2\theta_{23}\frac{\sin^2 2\theta_{13}}{(A-1)^2}\sin^2[(A-1)\Delta], \\
P_{\sin\delta} &= \alpha\frac{8J_{cp}}{A(1-A)}\sin\Delta\sin(A\Delta)\sin[(1-A)\Delta], \\
P_{\cos\delta} &= \alpha\frac{8J_{cp}\cot\delta_{CP}}{A(1-A)}\cos\Delta\sin(A\Delta)\sin[(1-A)\Delta], \\
P_3 &= \alpha^2\cos^2\theta_{23}\frac{\sin^2 2\theta_{12}}{A^2}\sin^2(A\Delta),
\end{aligned}$$

where $\Delta = 1.27\Delta m^2_{31}(eV^2)L(km)/E(GeV)$ and $A = \sqrt{2}G_F N_e 2E/\Delta m^2_{31}$.

For $\bar{\nu}_\mu \to \bar{\nu}_e$, $\underbrace{P_{\sin\delta} \to -P_{\sin\delta}}_{\text{CP asymmetry}}$, $\underbrace{A \to -A}_{\text{matter asymmetry}}$

Osc. vs L/E

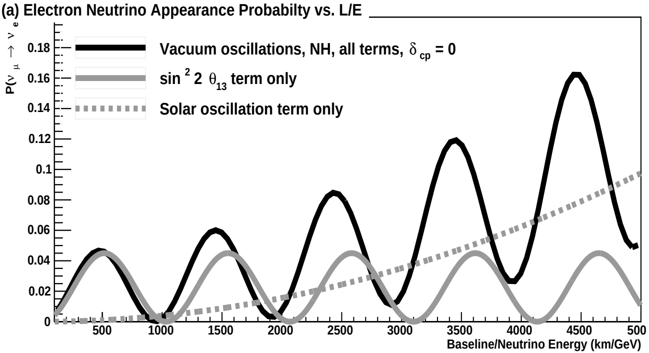
$\sin^2 2\theta_{13} = 0.09$, $\sin^2 \theta_{23} = 0.5$, $\Delta m^2_{31} = \pm 2.4 \times 10^{-3}\mathrm{eV}^2$

ν Exercise: Use ROOT and reproduce the plots shown below

The $\nu_\mu \rightarrow \nu_e$ oscillation probability maxima occur at

$$\frac{L\ (\mathrm{km})}{E_n(\mathrm{GeV})} = \left(\frac{\pi}{2}\right)\frac{(2n-1)}{1.27 \times \Delta m^2_{31}(\mathrm{eV}^2)} \approx (2n-1) \times \frac{515\ \mathrm{km}}{\mathrm{GeV}}$$

Oscillations in vacuum - different terms ($\delta_{CP} = 0$)

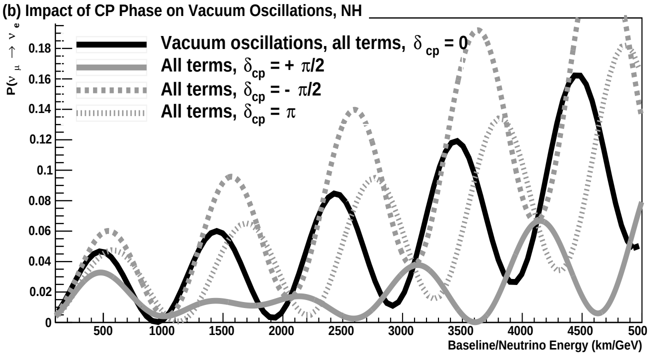


ν Exercise: Use ROOT and reproduce the plots shown below

The $\nu_\mu \to \nu_e$ oscillation probability maxima occur at

$$\frac{L\ (\mathrm{km})}{E_n(\mathrm{GeV})} = \left(\frac{\pi}{2}\right)\frac{(2n-1)}{1.27\times\Delta m^2_{31}(\mathrm{eV}^2)} \approx (2n-1)\times\frac{515\ \mathrm{km}}{\mathrm{GeV}}$$

Impact of δ_{CP} on oscillations in vacuum, $\Delta m^2_{31} > 0$ (NH)



Osc. vs L/E

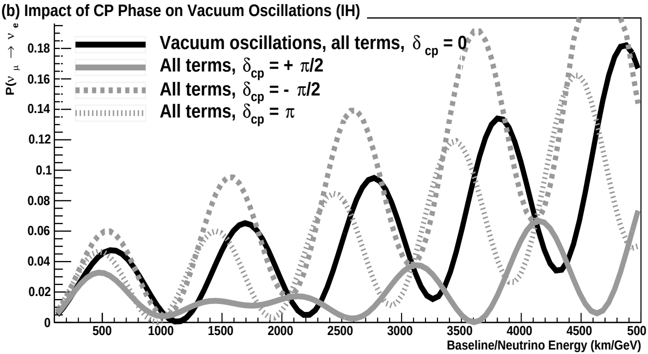
$\sin^2 2\theta_{13} = 0.09, \sin^2\theta_{23} = 0.5, \Delta m^2_{31} = \pm 2.4\times 10^{-3}\mathrm{eV}^2$

ν Exercise: Use ROOT and reproduce the plots shown below

The $\nu_\mu \to \nu_e$ oscillation probability maxima occur at

$$\frac{L\ (\mathrm{km})}{E_n(\mathrm{GeV})} = \left(\frac{\pi}{2}\right)\frac{(2n-1)}{1.27\times\Delta m^2_{31}(\mathrm{eV}^2)} \approx (2n-1)\times\frac{515\ \mathrm{km}}{\mathrm{GeV}}$$

Impact of δ_{CP} on oscillations in vacuum, $\Delta m^2_{31} < 0$ (IH)



Osc. vs L/E

$\sin^2 2\theta_{13} = 0.09, \sin^2\theta_{23} = 0.5, \Delta m^2_{31} = \pm 2.4 \times 10^{-3}\text{eV}^2$

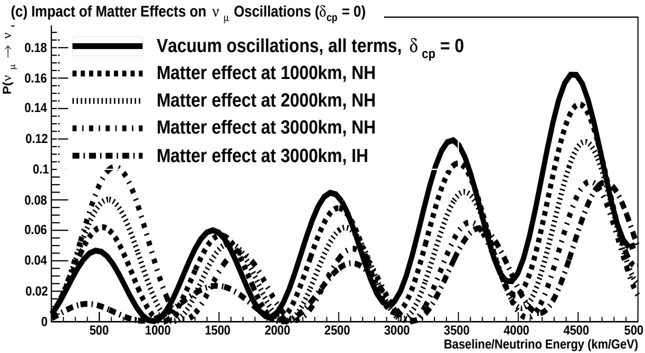
The Little Neutral One
Mary Bishai Brookhaven National Laboratory
History of ν
Discovery of ν
Neutrino Flavor
ν Oscillations
Neutrino Mixing
Neutrinos in the 21st Century
Daya Bay
NOνA
DUNE/LBNF
Summary

ν Exercise: Use ROOT and reproduce the plots shown below

The $\nu_\mu \rightarrow \nu_e$ oscillation probability maxima occur at

$$\frac{L\ (\text{km})}{E_n(\text{GeV})} = \left(\frac{\pi}{2}\right)\frac{(2n-1)}{1.27 \times \Delta m^2_{31}(\text{eV}^2)} \approx (2n-1) \times \frac{515\ \text{km}}{\text{GeV}}$$

Impact of matter effect on ν_μ oscillations ($\delta_{\text{CP}} = 0$)

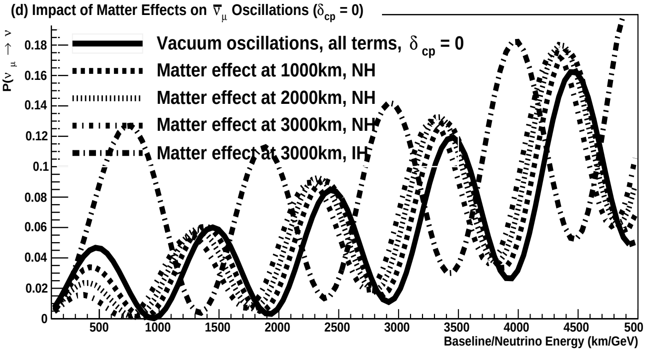


ν Exercise: Use ROOT and reproduce the plots shown below

The $\nu_\mu \rightarrow \nu_e$ oscillation probability maxima occur at

$$\frac{L\ (\mathrm{km})}{E_n(\mathrm{GeV})} = \left(\frac{\pi}{2}\right) \frac{(2n-1)}{1.27 \times \Delta m^2_{31}(\mathrm{eV}^2)} \approx (2n-1) \times \frac{515\ \mathrm{km}}{\mathrm{GeV}}$$

Impact of matter effect on $\bar{\nu}_\mu$ oscillations ($\delta_{\mathrm{CP}} = 0$)



Expected Appearance Signal Event Rates

ν Exercise: The total number of electron neutrino appearance events expected for a given exposure from a muon neutrino source as a function of baseline is given as

$$N_{\nu_e}^{\rm appear}(L) = \int \Phi^{\nu_\mu}(E_\nu, L) \times P^{\nu_\mu \to \nu_e}(E_\nu, L) \times \sigma^{\nu_e}(E_\nu) dE_\nu$$

Assume the neutrino source produces a flux that is constant in energy and using only the dominant term in the probability(no matter effect)

$$\begin{aligned}
\Phi^{\nu_\mu}(E_\nu, L) &\approx \frac{C}{L^2}, \quad C = \text{number of } \nu_\mu/\text{m}^2/\text{GeV}/\text{sec at 1 km} \\
P^{\nu_\mu \to \nu_e}(E_\nu, L) &\approx \underbrace{\sin^2\theta_{23}\sin^2 2\theta_{13}\sin^2(1.27\Delta m_{31}^2 L/E_\nu)}_{P_0} \\
\sigma^{\nu_e}(E_\nu) &= 0.7\times 10^{-42}(\text{m}^2/\text{GeV}/N)\times E_\nu, \quad E_\nu > 1\ \text{GeV}
\end{aligned}$$

Prove that the rate of ν_e appearing integrated over a constant range of L/E is independent of baseline for $L > 500$ km!

Expected Appearance Signal Event Rates

$$N_{\nu_e}^{\mathrm{appear}}(L) \propto \text{constant term} \times \int \frac{\sin^2(ax)}{x^3}\,dx,$$
$$x \equiv L/E_\nu,\ \ a \equiv 1.27\Delta m^2_{31}\ \mathrm{GeV/(eV^2.km)}$$

ν Exercise:

$C \approx 1 \times 10^{17}\ \nu_\mu/\mathrm{m}^2/\mathrm{GeV/yr}$ at 1 km (from 1MW accelerator)
$\sin^2 2\theta_{13} = 0.084$, $\sin^2\theta_{23} = 0.5$, $\Delta m^2_{31} = 2.4 \times 10^{-3}\mathrm{eV}^2$

Calculate the rate of ν_e events observed per kton of detector integrating over the region $x = 100$ km/GeV to 2000 km/GeV. Use ROOT to do the integral!

Expected Appearance Signal Event Rates

$$N_{\nu_e}^{\mathrm{appear}}(L) \propto \text{constant term} \times \int \frac{\sin^2(ax)}{x^3}dx,$$

$$x \equiv L/E_\nu, \ \ a \equiv 1.27\Delta m^2_{31}\ \mathrm{GeV/(eV^2.km)}$$

ν Exercise:

$C \approx 1 \times 10^{17}\ \nu_\mu/\mathrm{m^2/GeV/yr}$ at 1 km (from 1MW accelerator)
$\sin^2 2\theta_{13} = 0.084$, $\sin^2\theta_{23} = 0.5$, $\Delta m^2_{31} = 2.4 \times 10^{-3}\mathrm{eV^2}$

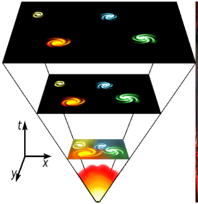
Calculate the rate of ν_e events observed per kton of detector integrating over the region $x = 100$ km/GeV to 2000 km/GeV. Use ROOT to do the integral!

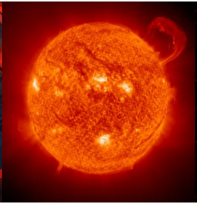
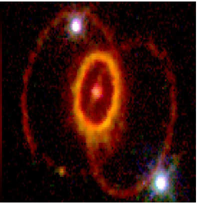
$$N_{\nu_e}^{\mathrm{appear}}(L) \approx (2 \times 10^6 \mathrm{events/kton/yr}) \cdot (\mathrm{km/GeV})^2 \int_{x_0}^{x_1} \frac{\sin^2(ax)}{x^3}dx,$$

$$\boxed{N_{\nu_e}^{\mathrm{appear}}(L) \sim \mathcal{O}(20 - 30)\ \mathrm{events/kton/yr}}$$

Sources of Neutrinos (Summary)

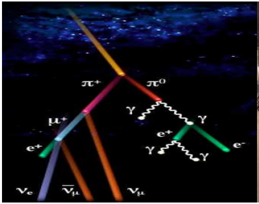
Big Bang | **Reactors** | **Sun** | **SuperNova**

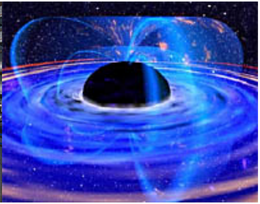


Big Bang	Reactors	Sun	SuperNova
10^{-4} eV	few MeV	0.1-14 MeV	$\sim$ 10 MeV
$56/\text{cm}^3$	$10^{21}/\text{GW}_{th}/\text{s}$	$10^{10}/\text{cm}^2/\text{s}$	$10^9/\text{cm}^2/\text{s}$

Atmosphere | **Accelerators** | **Extragalactic**



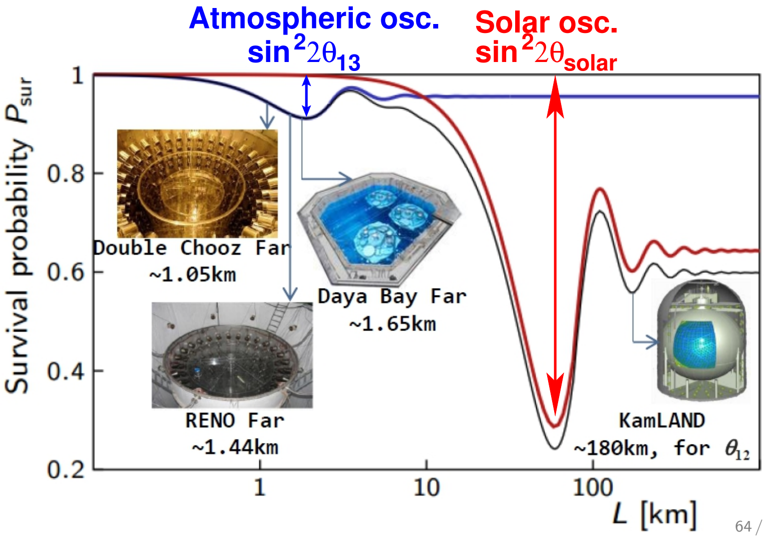



Atmosphere	Accelerators	Extragalactic
$\sim$ 1 GeV	1-20 GeV	TeV-PeV
$\text{few}/\text{cm}^2/\text{s}$	$10^6/\text{cm}^2/\text{s}/\text{MW}$ (at 1km)	varies

NEUTRINO EXPERIMENTS OF THE 21^{st} CENTURY

More Reactor $\bar{\nu}_e$: The 3rd Mixing Amplitude (θ_{13})

$\sin^2\theta_{13}$ = fraction of ν_e in ν_3 state, $\sin^2\theta_{12}$ = fraction of ν_e in ν_2 state



The Daya Bay Reactor Complex

Reactor Specs:

Located 55km north-east of Hong Kong.

Initially: 2 cores at Daya Bay site + 2 cores at Ling Ao site = 11.6 GW$_{th}$

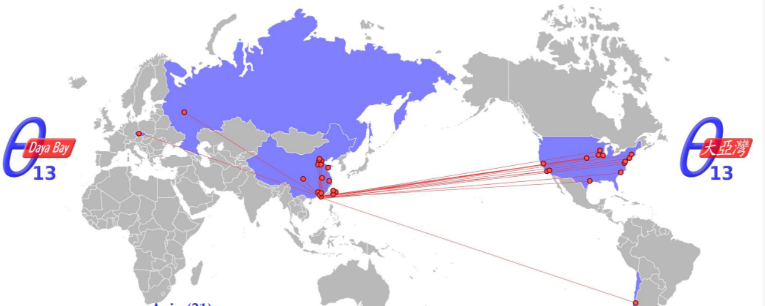
By 2011: 2 more cores at Ling Ao II site = 17.4 GW$_{th}$ $\Rightarrow$ top five worldwide

1 GW$_{th}$ = $2 \times 10^{20} \bar{\nu}_e$/second

Deploy multiple near and far detectors

Reactor power uncertainties $< 0.1\%$

The Daya Bay Collaboration : 231 Collaborators



Asia (21)
Beijing Normal Univ., CNG, CIAE, Dongguan Polytechnic, ECUST, IHEP, Nanjing Univ., Nankai Univ., NCEPU, Shandong Univ., Shanghai Jiao Tong Univ., Shenzhen Univ., Tsinghua Univ., USTC, Xian Jiaotong Univ., Zhongshan Univ., Chinese Univ. of Hong Kong, Univ. of Hong Kong, National Chiao Tung Univ., National Taiwan Univ., National United Univ.

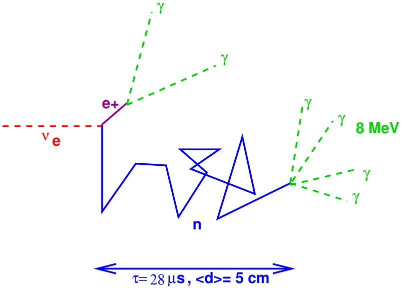
Europe (2)
Charles University, JINR Dubna

North America (17)
Brookhaven Natl Lab, CalTech, Illinois Institute of Technology, Iowa State, Lawrence Berkeley Natl Lab, Princeton, Rensselaer Polytechnic, Siena College, UC Berkeley, UCLA, Univ. of Cincinnati, Univ. of Houston, UIUC, Univ. of Wisconsin, Virginia Tech, William & Mary, Yale

South America (1)
Catholic Univ. of Chile

Detecting Neutrinos from the Daya Bay Reactors

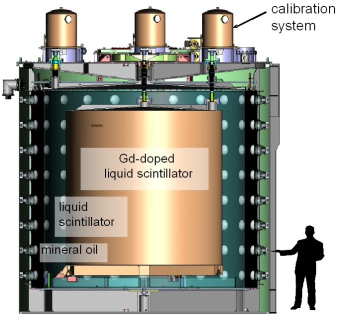
The active target in each detector is liquid scintillator loaded with 0.1% Gd

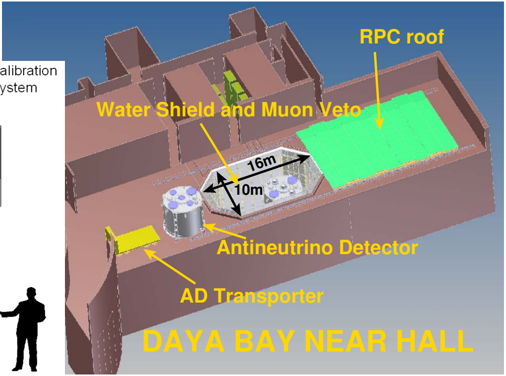


- $\bar{\nu}_e + p \rightarrow n + e^+$
- $e^+ + e^- \rightarrow \gamma\gamma$ (2X 0.511 MeV $+T_{e^+}$, prompt)
- $n + p \rightarrow D + \gamma$ (2.2 MeV, $\tau \sim$ 180μs). **OR**
- $n + Gd \rightarrow Gd* \rightarrow Gd + \gamma$'s (8 MeV, $\tau \sim 28\mu s$).

$\Rightarrow$ delayed co-incidence of e^+ conversion and n-capture ($>$ 6 MeV) with a specfic energy signature

The Daya Bay Experimental Apparatus

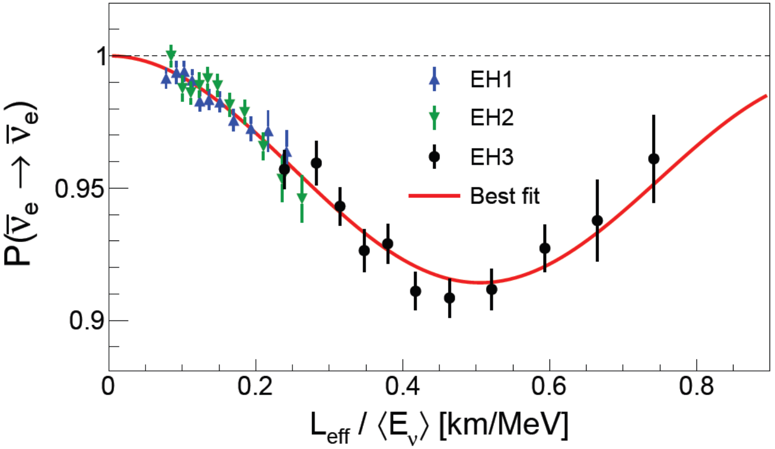




- **Multiple "identical" detectors at each site.**
- **Manual and multiple automated calibration systems per detector.**
- **Thick water shield to reduce cosmogenic and radiation bkgds.**

	DYB	LA	Far
Event rates/20T/day	840	740	90

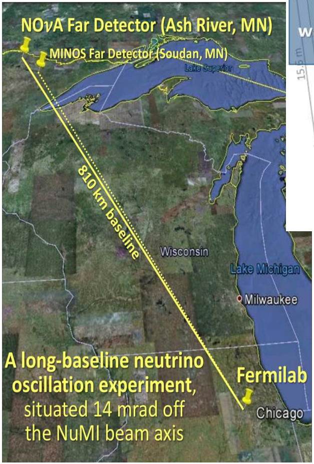
Daya Bay Measurement of Non-zero θ_{13}

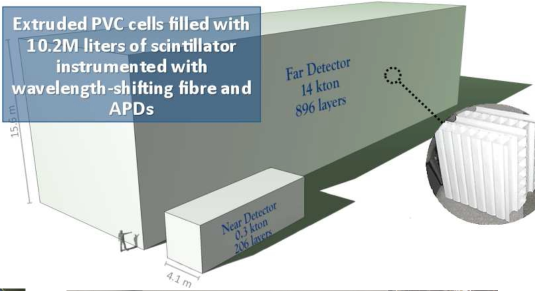


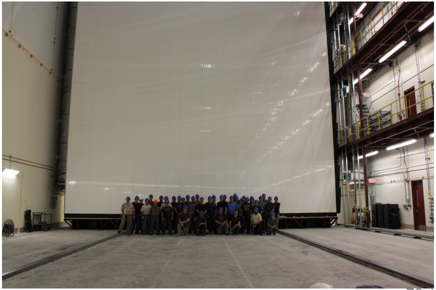
First to discover non-zero θ_{13} (2012) and currently most precise result:

$$\sin^2 2\theta_{13} = 0.084 \pm 0.005$$

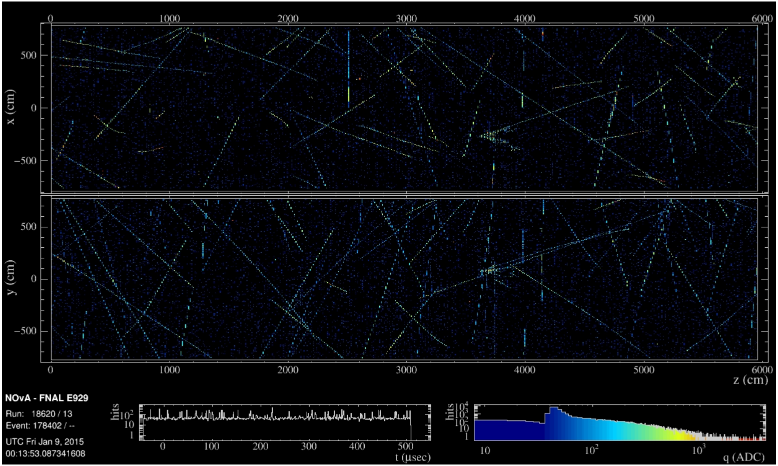
The NOνA Experiment



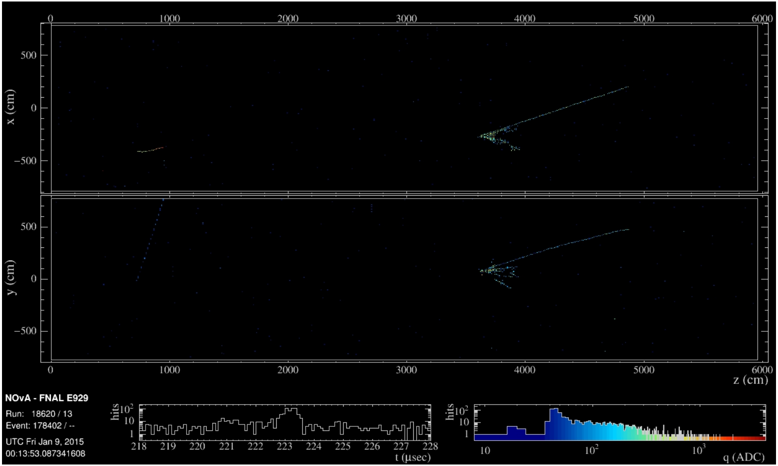




Neutrino Events in NOνA



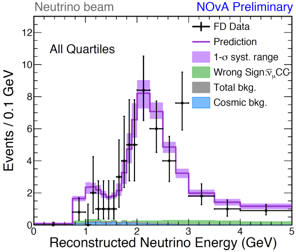
Neutrino Events in NOνA



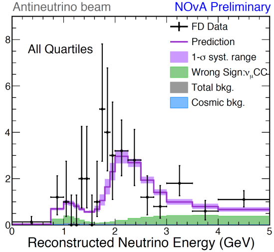
The Little Neutral One
Mary Bishai Brookhaven National Laboratory
History of ν
Discovery of ν
Neutrino Flavor
ν Oscillations
Neutrino Mixing
Neutrinos in the 21st Century
Daya Bay
NOνA
DUNE/LBNF
Summary

Results from NOνA (2018)

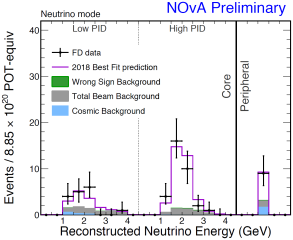
$\nu_\mu \to \nu_\mu$



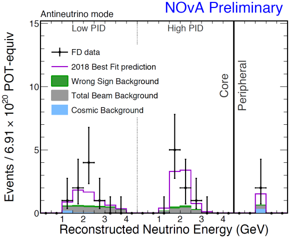
$\bar{\nu}_\mu \to \bar{\nu}_\mu$



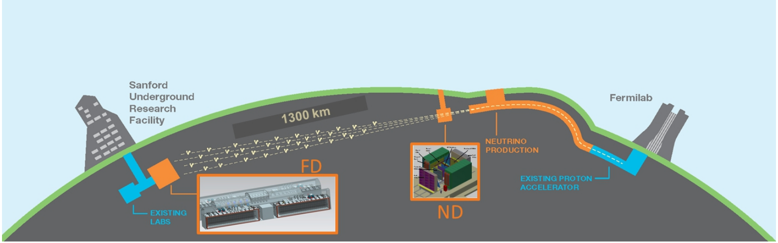
$\nu_\mu \to \nu_e$



$\bar{\nu}_\mu \to \bar{\nu}_e$



The Deep Underground Neutrino Experiment



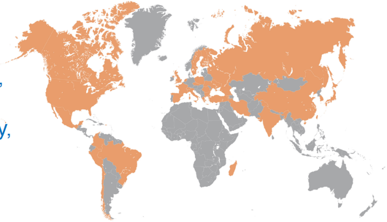
- **A very long baseline experiment:** **1300km from Fermilab in Batavia, IL to the Sanford Underground Research Facility (former Homestake Mine) in Lead, SD.**
- **A highly capable near detector at Fermilab.**
- **A very deep (1 mile underground) far detector: massive 40-kton Liquid Argon Time-Projection-Chamber with state-of-the-art instrumentation.**
- **High intensity *tunable* wide-band neutrino beam from LBNF produced from upgraded MW-class proton accelerator at Fermilab.**

The DUNE Scientific Collaboration

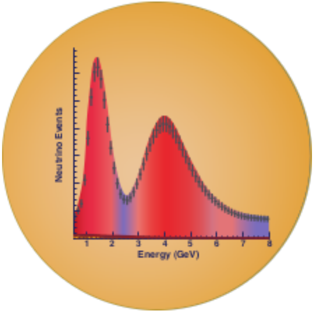
As of Jan 2018:

1061 collaborators from $\underbrace{\text{175 institutions}}^{\text{60 \% non-US}}$ in 31 nations

Armenia, Brazil, Bulgaria, Canada, CERN, Chile, China, Colombia, Czech Republic, Finland, France, Greece, India, Iran, Italy, Japan, Madagascar, Mexico, Netherlands, Paraguay, Peru, Poland, Romania, Russia, South Korea, Spain, Sweden, Switzerland, Turkey, UK, Ukraine, USA

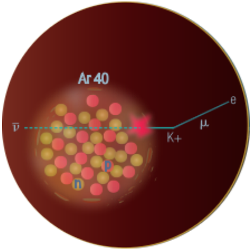


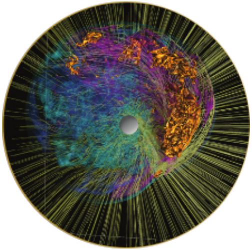
Scientific Objectives of DUNE



1. **precision measurements of the parameters that govern $\nu_\mu \rightarrow \nu_e$ oscillations; this includes precision measurement of the third mixing angle θ_{13}, measurement of the charge-parity (CP) violating phase δ_{CP}, and determination of the neutrino mass ordering (the sign of $\Delta m^2_{31} = m^2_3 - m^2_1$), the so-called mass hierarchy**
2. **precision measurements of the mixing angle θ_{23}, including the determination of the octant in which this angle lies, and the value of the mass difference, —Δm^2_{32}—, in $\nu_\mu \rightarrow \nu_{e,\mu}$ oscillations**

Scientific Objectives of DUNE

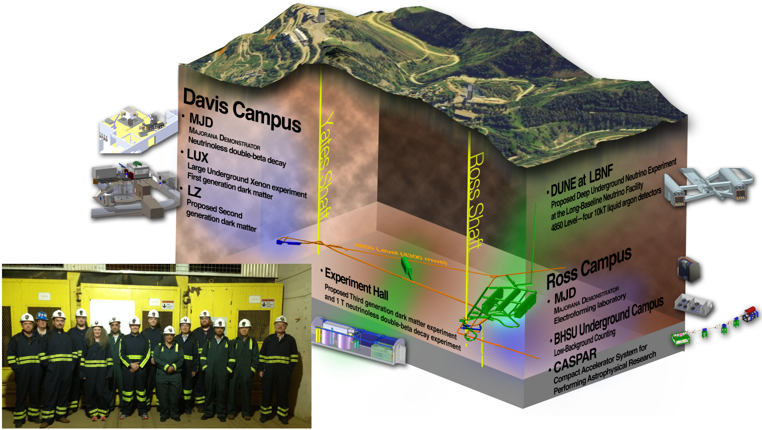




3. search for proton decay, yielding significant improvement in the current limits on the partial lifetime of the proton (τ/BR) in one or more important candidate decay modes, e.g., $p \rightarrow K^{+}\overline{\nu}$

4. detection and measurement of the neutrino flux from a core-collapse supernova within our galaxy, should one occur during the lifetime of DUNE

The Sanford Underground Research Facility

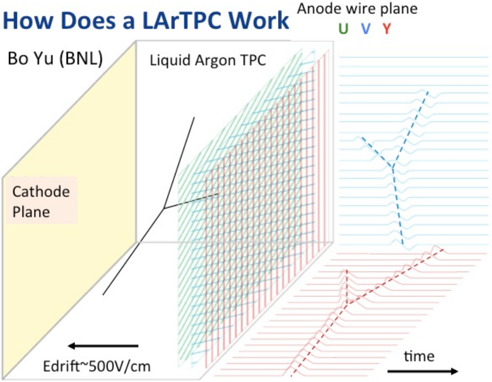


Experimental facility operated by the state of South Dakota. LUX (dark matter) and Majorana ($0\nu - 2\beta$) demonstrator operational expts at 4850-ft level. Chosen as site of G2 dark matter experiment

The DUNE Far Detector

A large cryogenic liquid Argon detector located a mile underground in the former Homestake Mine with a mass of at least 40 kilo-tons is used to image neutrino interactions with unprecedented precision:

Single Phase LArTPC



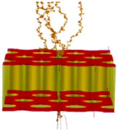
The DUNE prototype wireplane

The DUNE Far Detector

Dual Phase LArTPC

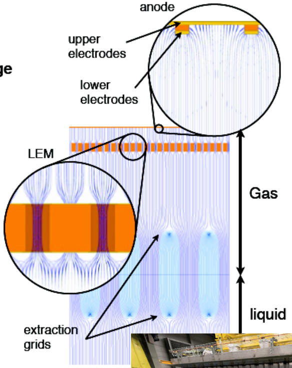
4.) **Charge collection on a 2D anode readout** (symmetric unipolar signals with two orthogonal views)

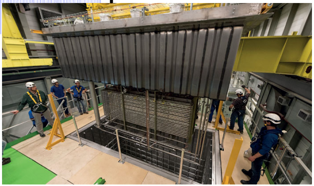
3.) **Charge multiplication in the holes of the Large Electron Multiplier (LEM)**



2.) **Drift electrons are efficiently emitted into the gas phase**

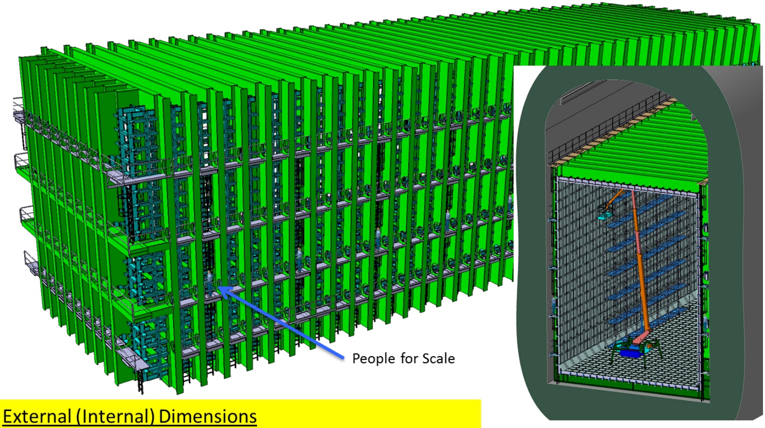
1.) **Ionization electrons drift towards the liquid argon surface**





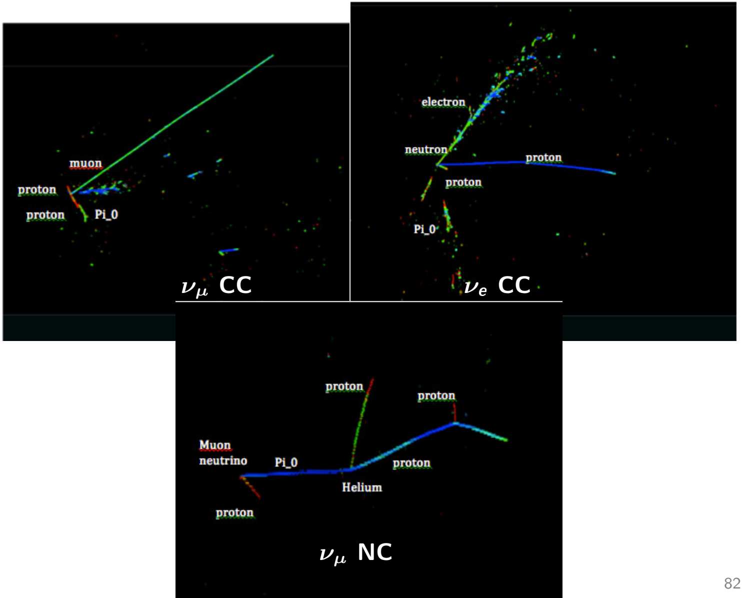
The DUNE Far Detector

The 40-kton (fiducial) detector is constructed of four modules with a total mass of 17.4 kton each.



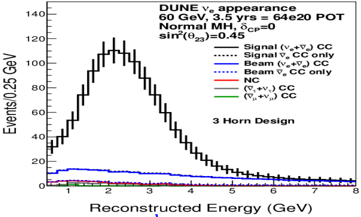
External (Internal) Dimensions
19.1m (16.9m) W x 18.0m (15.8m) H x 66.0m (63.8m) L

Reconstructed Neutrino Interactions in a LArTPC

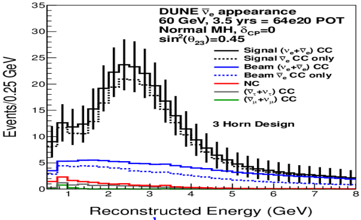


DUNE Event Spectra

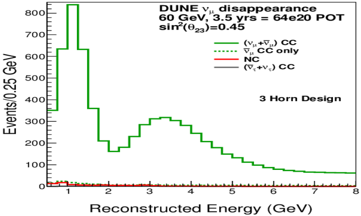
Exposure: 150 kT.MW.yr (equal $\nu/\bar{\nu}$) 1MW.yr $= 1 \times 10^{21}$ p.o.t at 120 GeV. ($\sin^2 2\theta_{13} = 0.085$, $\sin^2 \theta_{23} = 0.45$, $\delta m^2_{31} = 2.46 \times 10^{-3}$ eV^2)



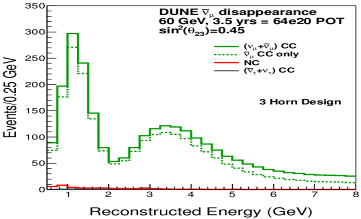
930 ν_e, 5 $\bar{\nu}_e$, 204 ν_e^{beam}, 17 NC, 19 ν_τ, 3 ν_μ



154 $\bar{\nu}_e$, 32 ν_e, 98 ν_e^{beam}, 7 NC, 8 ν_τ, 1 ν_μ



8329 ν_μ, 192 $\bar{\nu}_\mu$, 72 NC, 29 ν_τ

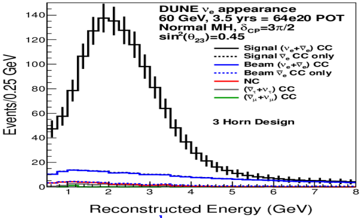


2420 $\bar{\nu}_\mu$, 791 ν_μ, 33 NC, 13 ν_τ

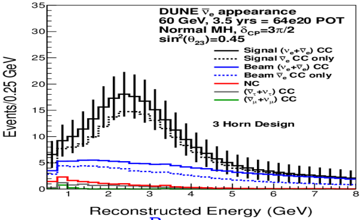
Simultaneous fit to all four samples. Richness of spectral information in both ν_μ and $\bar{\nu}_\mu$ $\Rightarrow$ explicit demonstration of CPV

DUNE Event Spectra

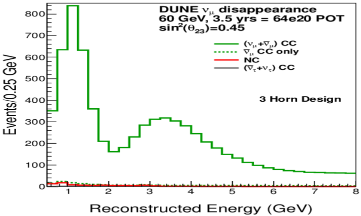
Exposure: 150 kT.MW.yr (equal $\nu/\bar{\nu}$) 1MW.yr = 1×10^{21} p.o.t at 120 GeV. ($\sin^2 2\theta_{13} = 0.085,\ \sin^2\theta_{23} = 0.45,\ \delta m^2_{31} = 2.46 \times 10^{-3}\ \mathrm{eV}^2$)



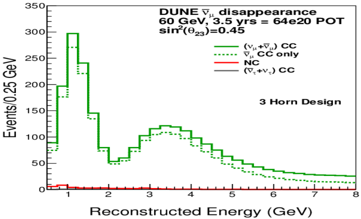
1171 ν_e,3 $\bar{\nu}_e$, 204 ν_e^{beam}, 17 NC, 19 ν_τ, 3 ν_μ



94 $\bar{\nu}_e$,39 ν_e, 98 ν_e^{Beam}, 7 NC, 8 ν_τ, 1 ν_μ



8329 ν_μ, 192 $\bar{\nu}_\mu$, 72 NC, 29 ν_τ



2420 $\bar{\nu}_\mu$, 791 ν_μ, 33 NC, 13 ν_τ

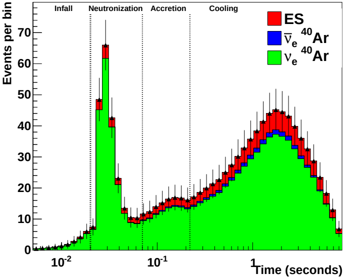
Simultaneous fit to all four samples. Richness of spectral information in both ν_μ and $\bar{\nu}_\mu$ $\Rightarrow$ explicit demonstration of CPV

Possible Supernova Signature in DUNE

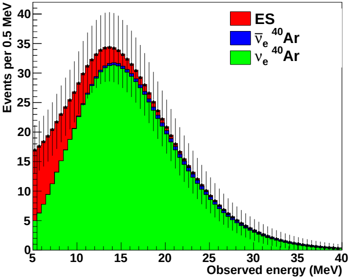
Liquid argon is particularly sensitive to the ν_e component of a supernova neutrino burst:

$$\nu_e +^{40} \mathrm{Ar} \rightarrow e^- +^{40} \mathrm{K}^*, \tag{1}$$

Expected time-dependent signal in 40 kton of liquid argon for a Supernova at 10 kpc:



Time distribution



Energy spectrum (time integrated)

LBNF/DUNE Schedule

- **2017: Cavern excavation begins**
- **2018: DUNE prototypes (single & dual phase) operational in test beam at CERN**
- **2019: Technical design review (beam and far detectors) by US-DOE and international funding energies. Conceptual design for near detector ready.**
- **2021: First 10kton FD module (single phase) installation**
- **2023: Second FD module (single or dual phase) installation**
- **2024: Data taking (non-beam) starts with 20 kton operational**
- **2026: First beam operations at 1.2 MW**

SUMMARY

Summary

- **Neutrinos have been at the forefront of fundamental disco veries ein particle physics for decades.**
- **Discoveries of neutrino properties like the very small mass, large almost maximal mixing, are the** *ONLY direct evidence for physics beyond the Standard Model of particle physics, and new hidden symmetries.*
- **Results from the current generation of accelerator based neutrino experiments hint (inconclusively) at large matter/anti-matter asymmetries.**
- **The future T2HK and LBNF/DUNE project are ambitious multi-national neutrino experiments designed to probe matter/anti-matter asymmetries, neutrino oscillations and cosmological neutrinos with unprecedented precision.**

THANK YOU

Click for Neutrino rap!!